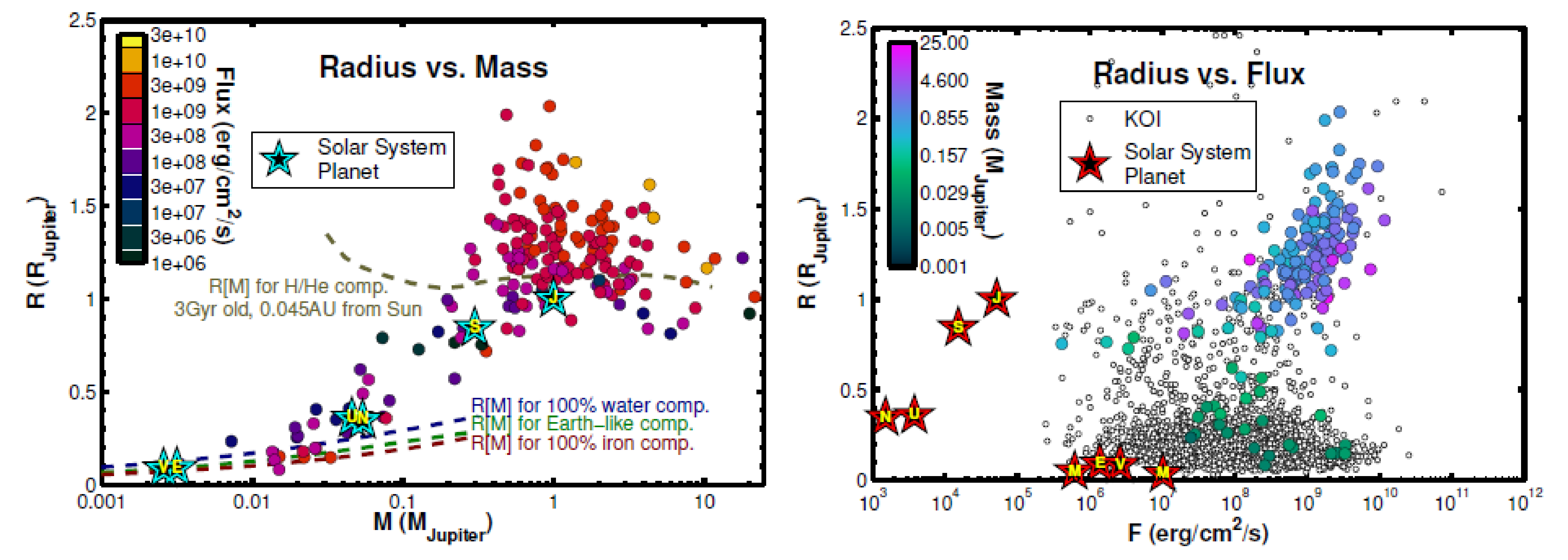
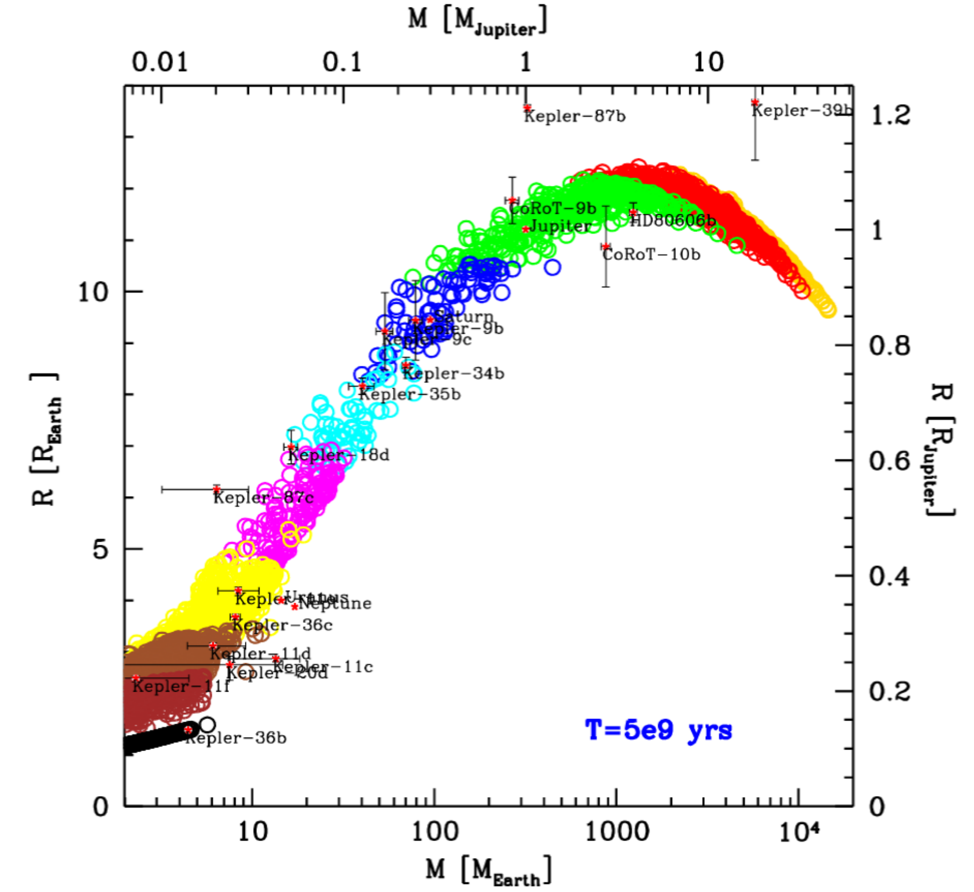
Internal structure and atmospheres of planets

SERGEI POPOV

Sizes and masses



Radius vs. mass



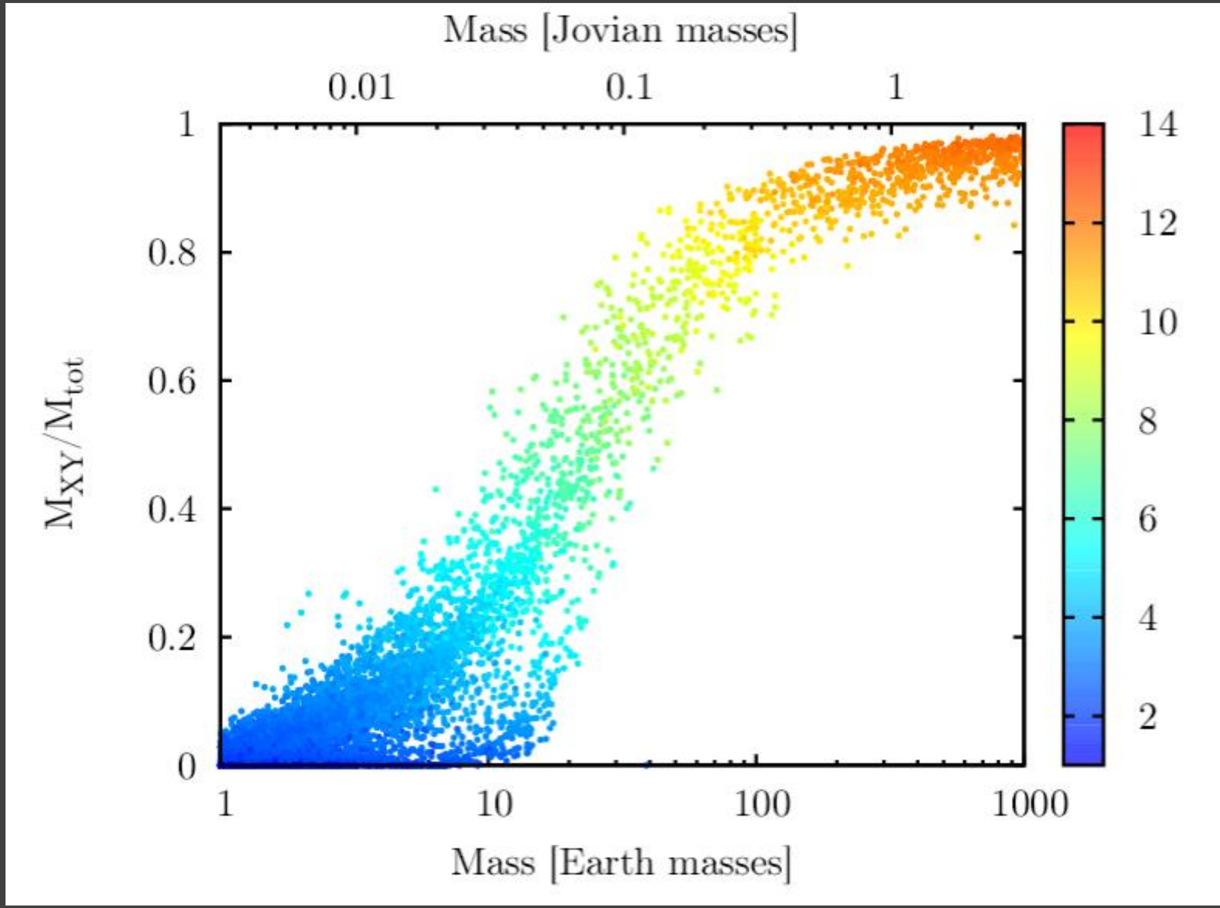
Results of modeling.

Old (relaxed) planets.
Planets ages are usually determined due to
stellar ages (1803.03125, 1804.02214).

Colors correspond to different fractions of light elements.

1604.07558

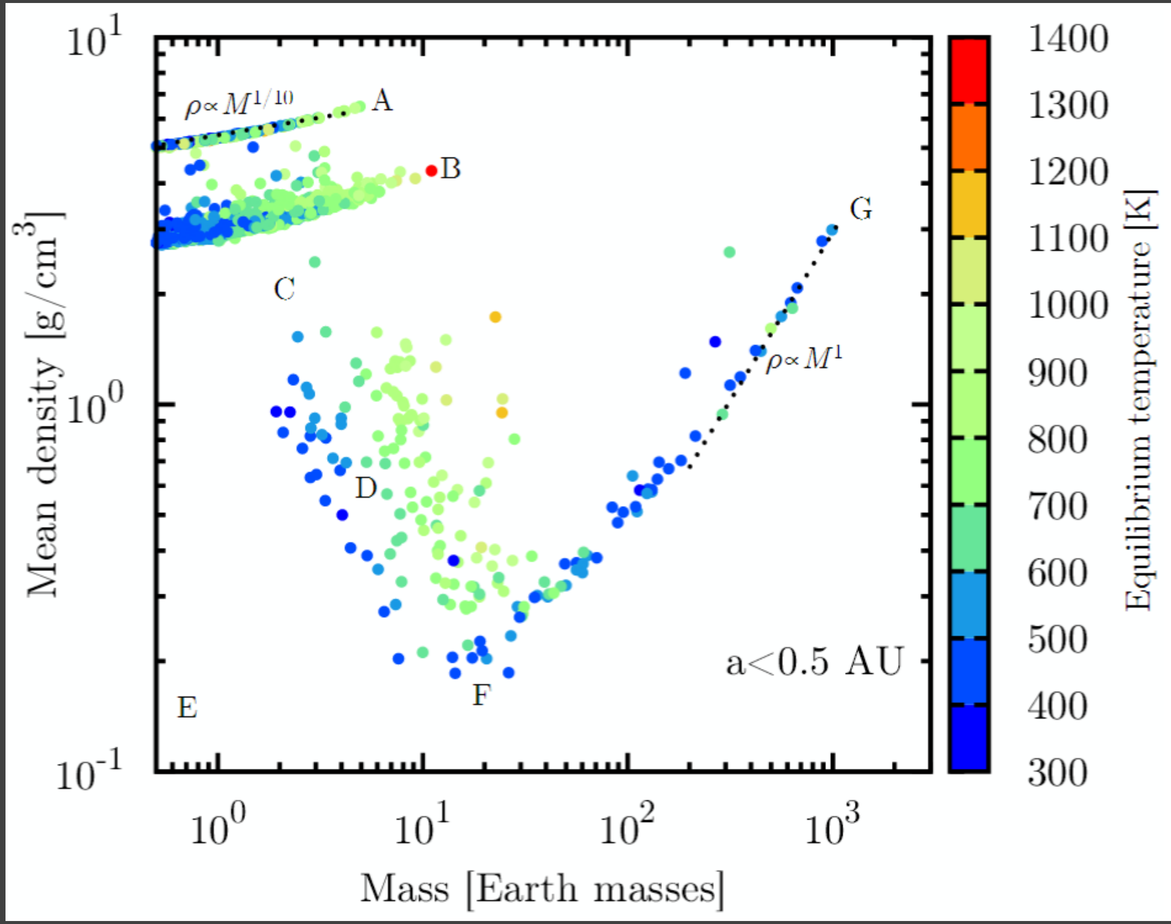
Light elements contribution



Results of modeling.

Different slopes above and below ~100 Earth masses are due to different regimes of gas accretion.

Density and mass

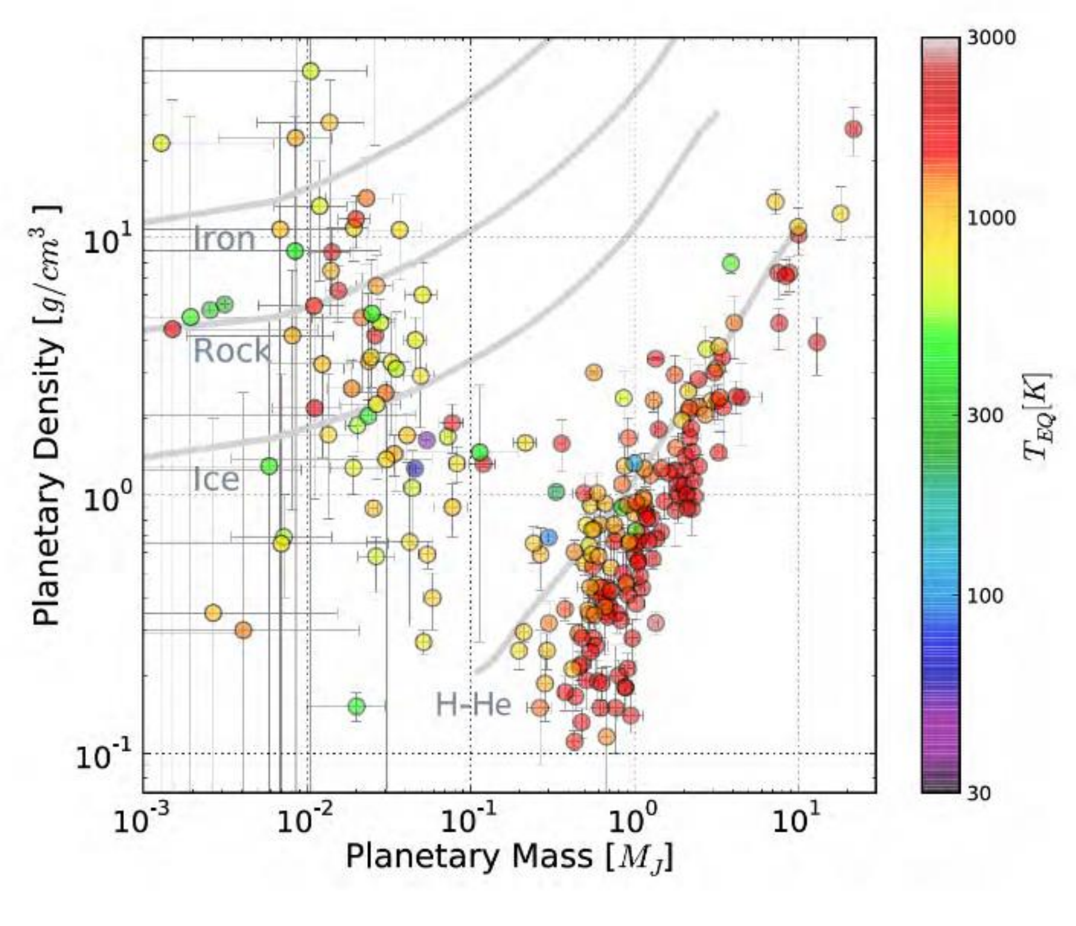


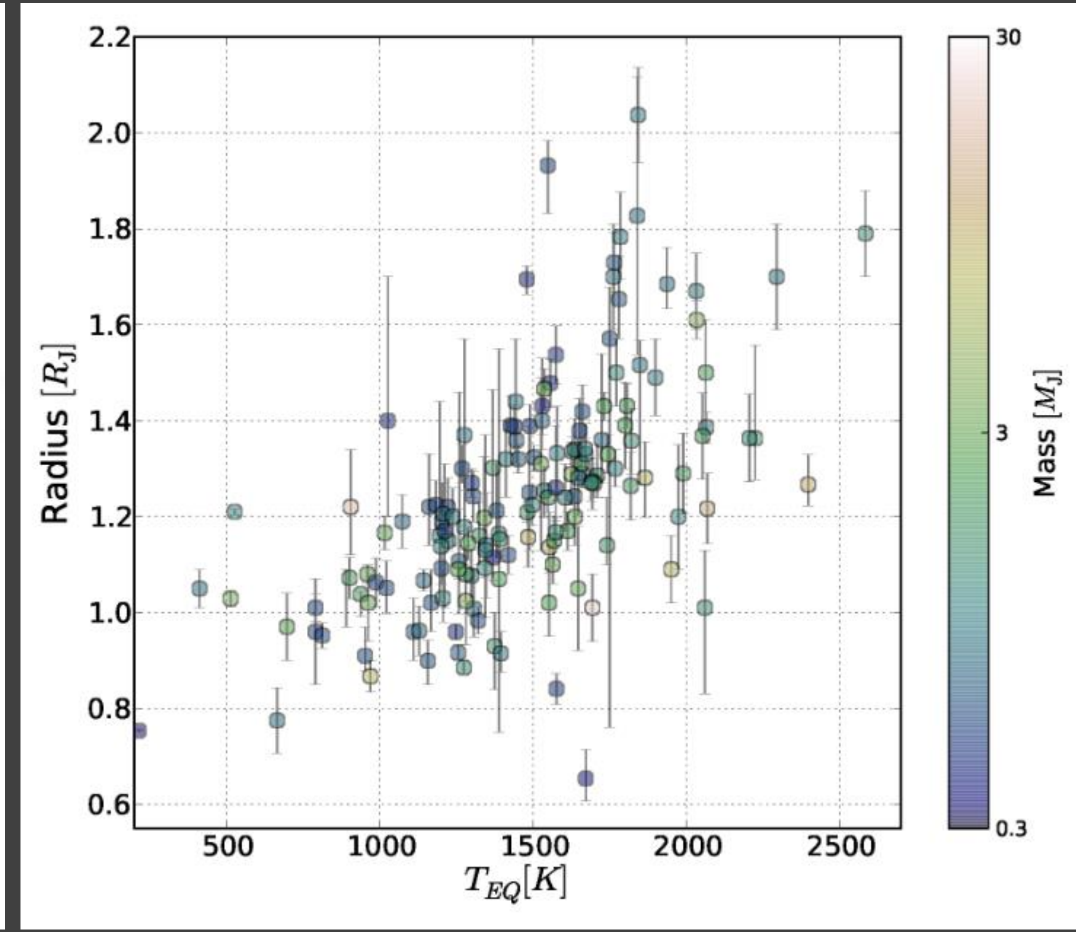
Results of modeling.

Old (5 Gyrs) planets.

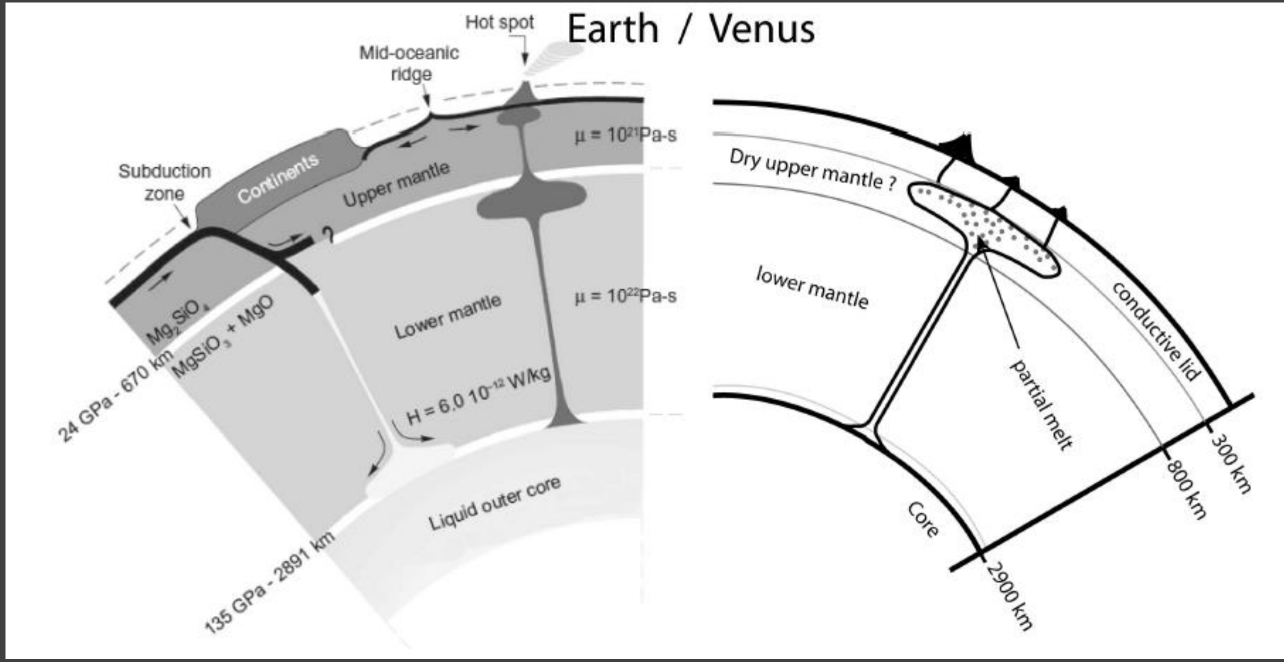
A – solid iron-stone
B – solid ice
C – evaporating
D – low-mass planets with large cores, but with significant fraction of H and He
E – forbidden zone (evaporating)
F – transition to giants
G - giants

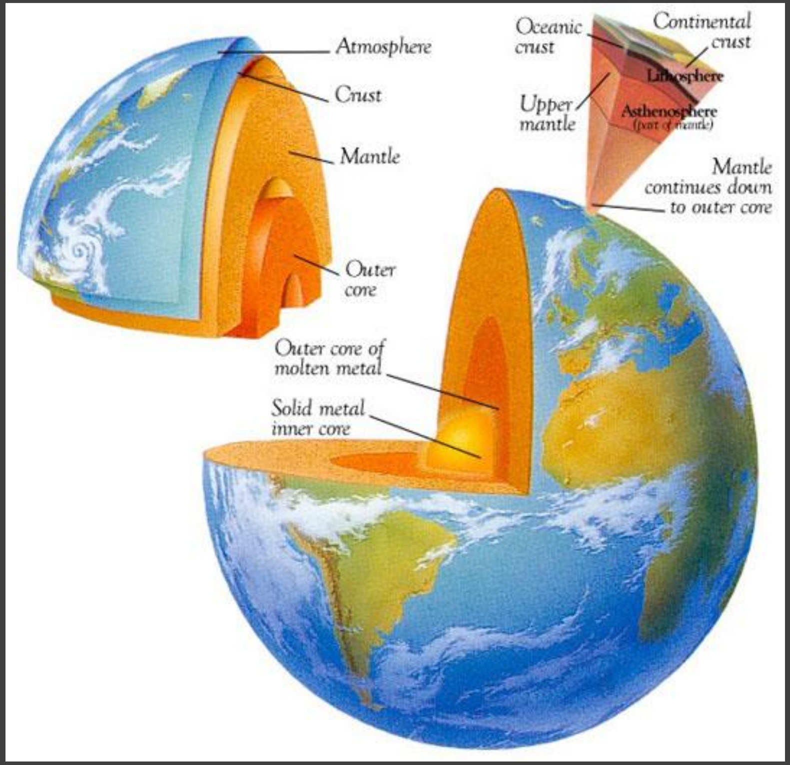
Mass-density. Observations. Heating.





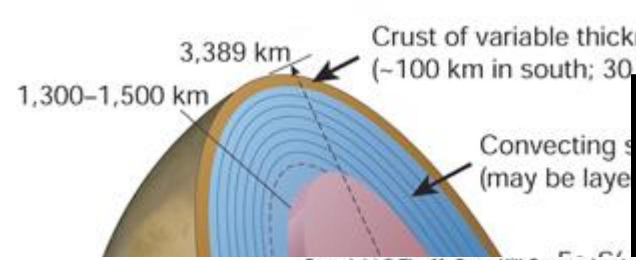
Planet structure

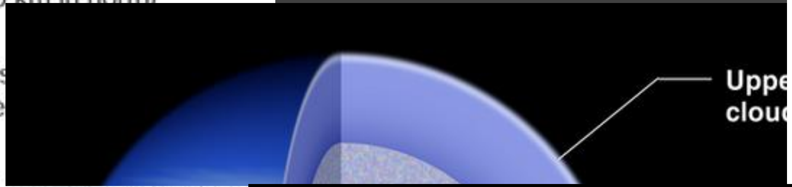


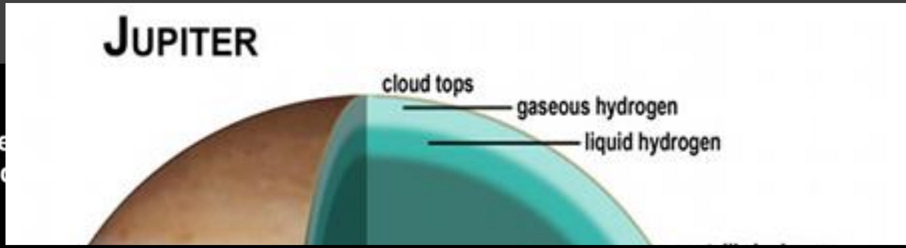


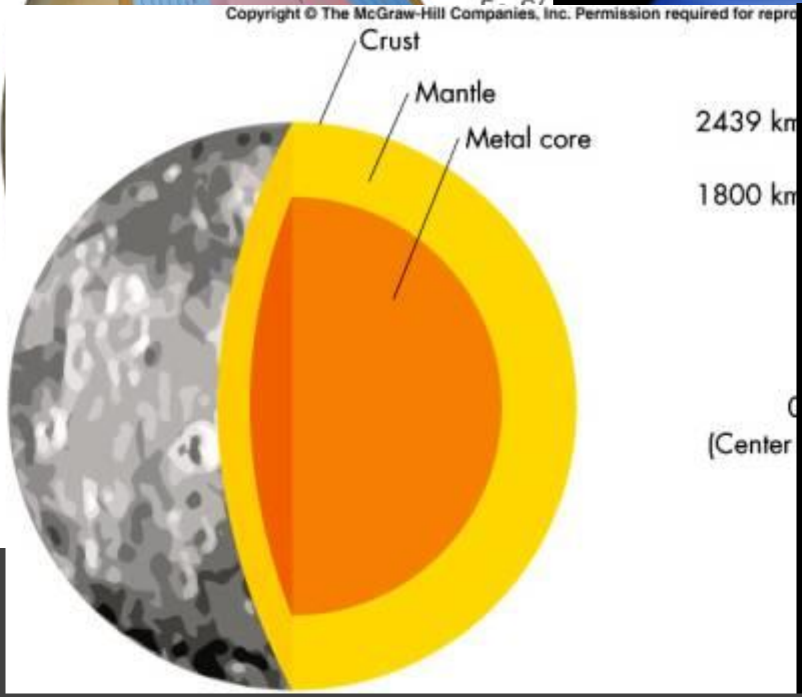
Even about the Earth we do not know many details of the internal structure. Data about other planets is very incomplete and indirect.

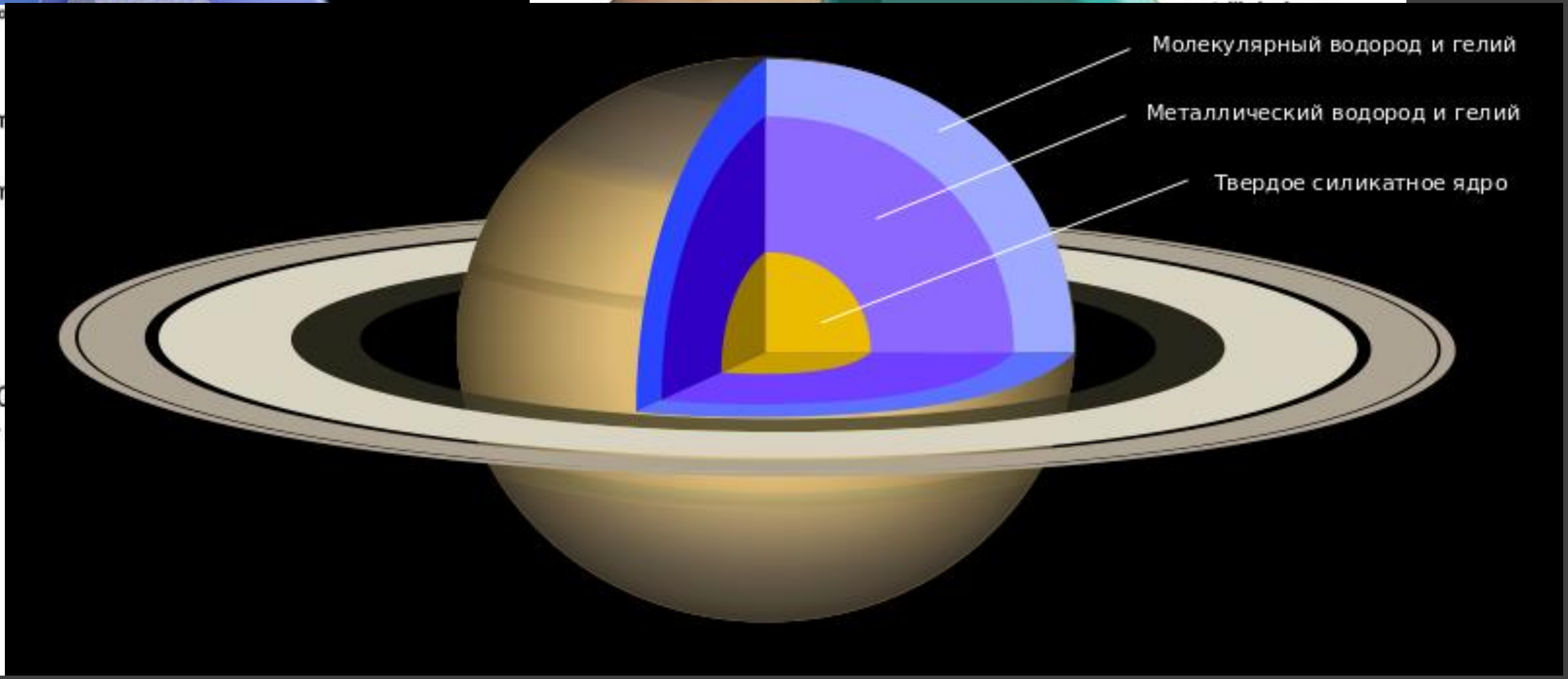
Structure of the Solar system planets



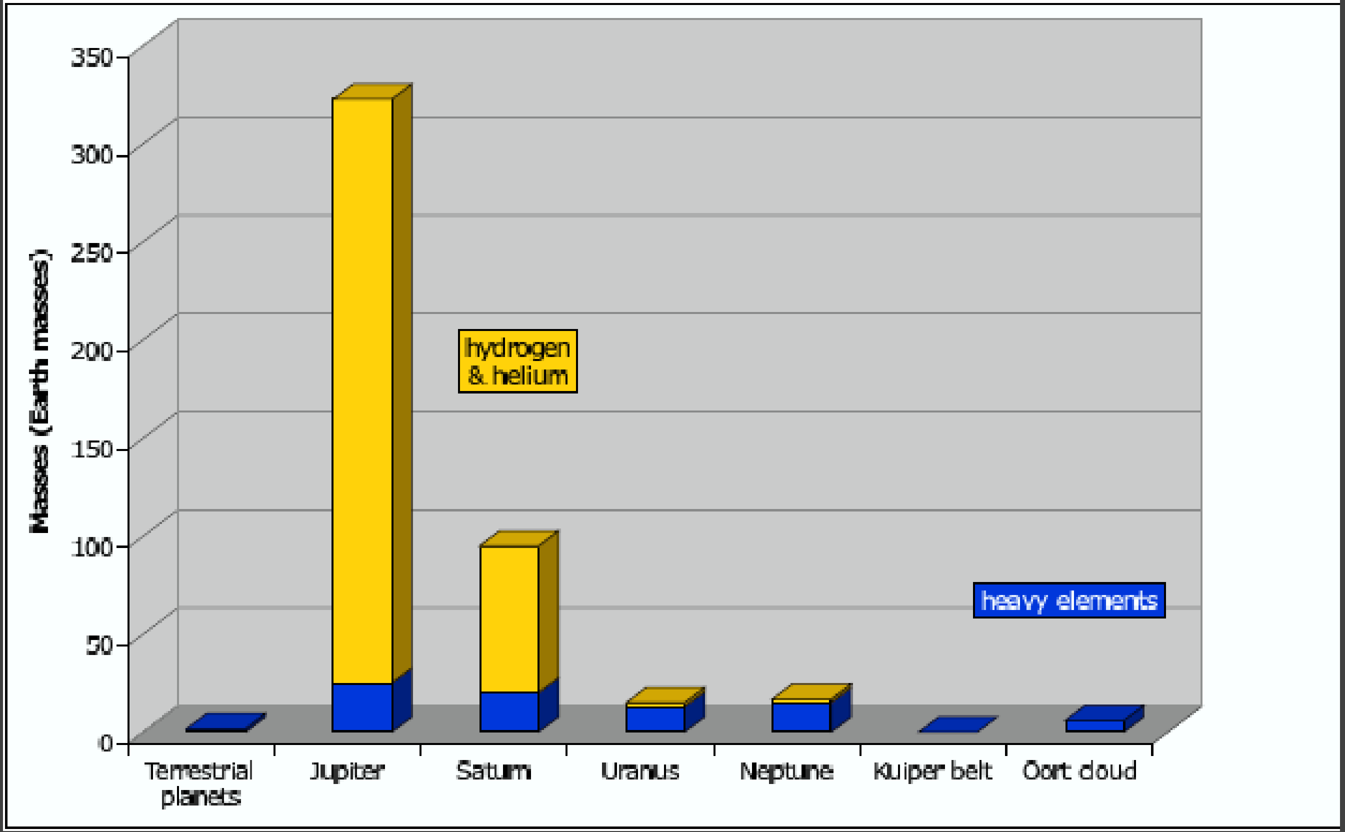








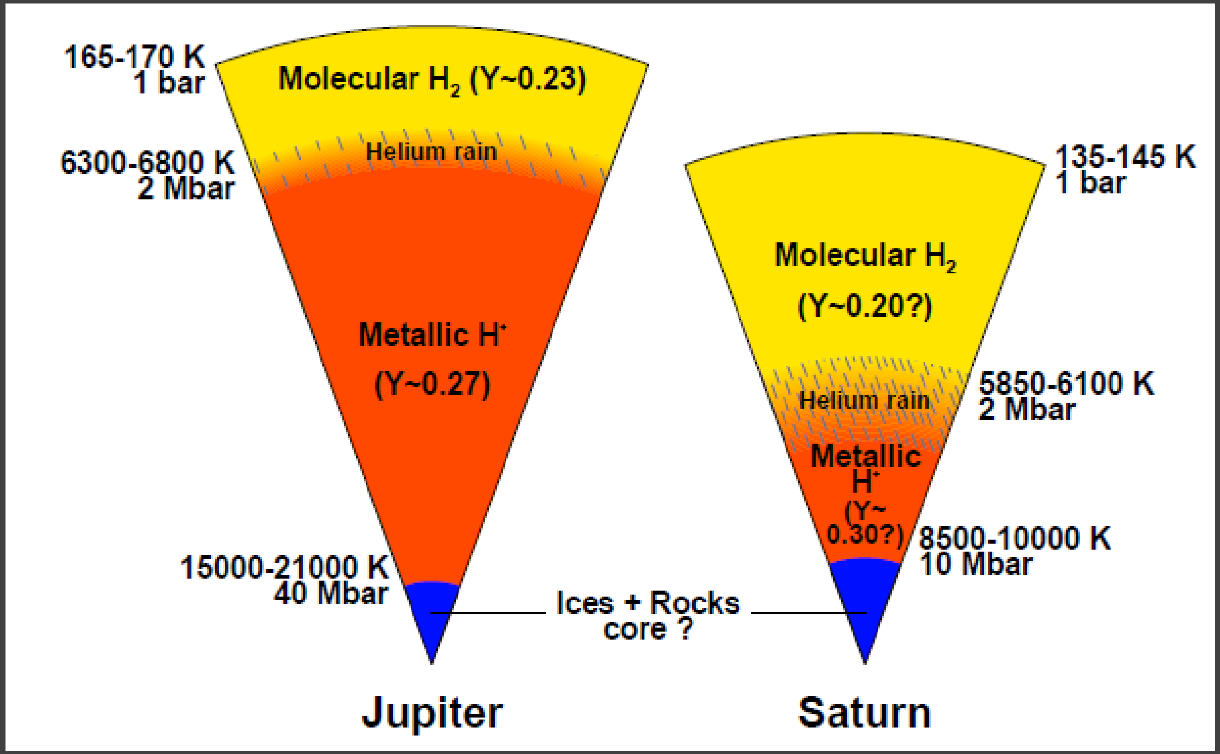
What Solar system planets are made of?



Except Jupiter and Saturn planets are mostly made of elements heavier than helium.

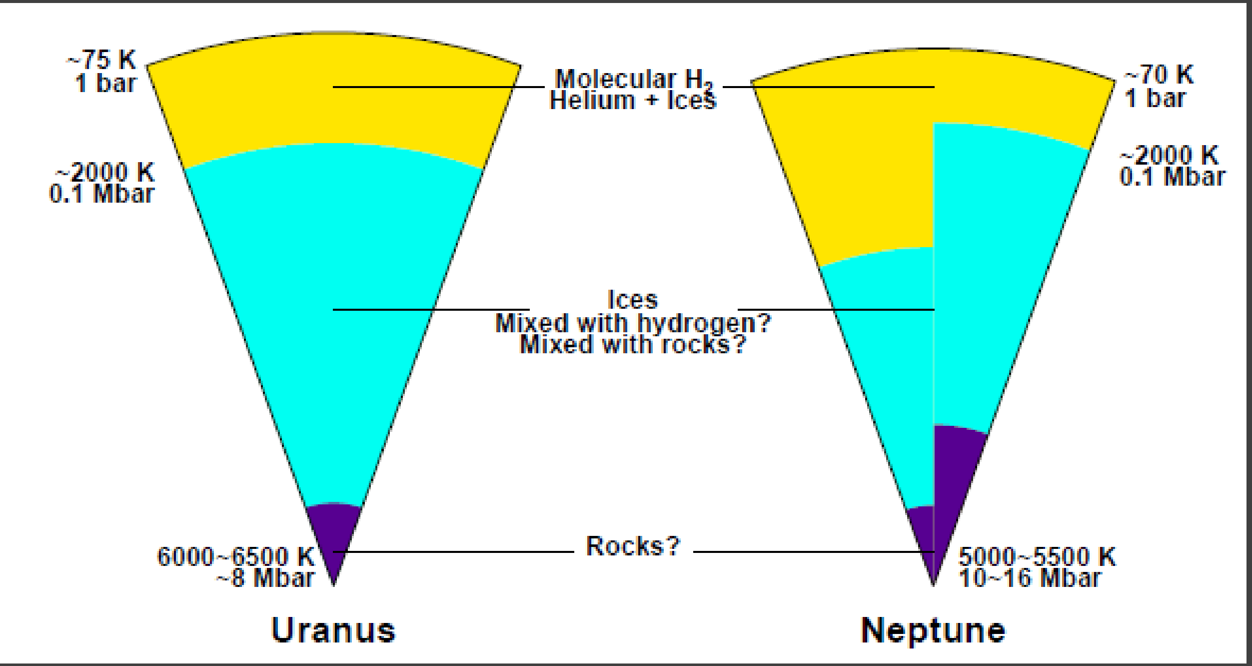
Even icy-giants – Neptune and Uranus, - are mainly made not of H+He.

Structure of giant planets

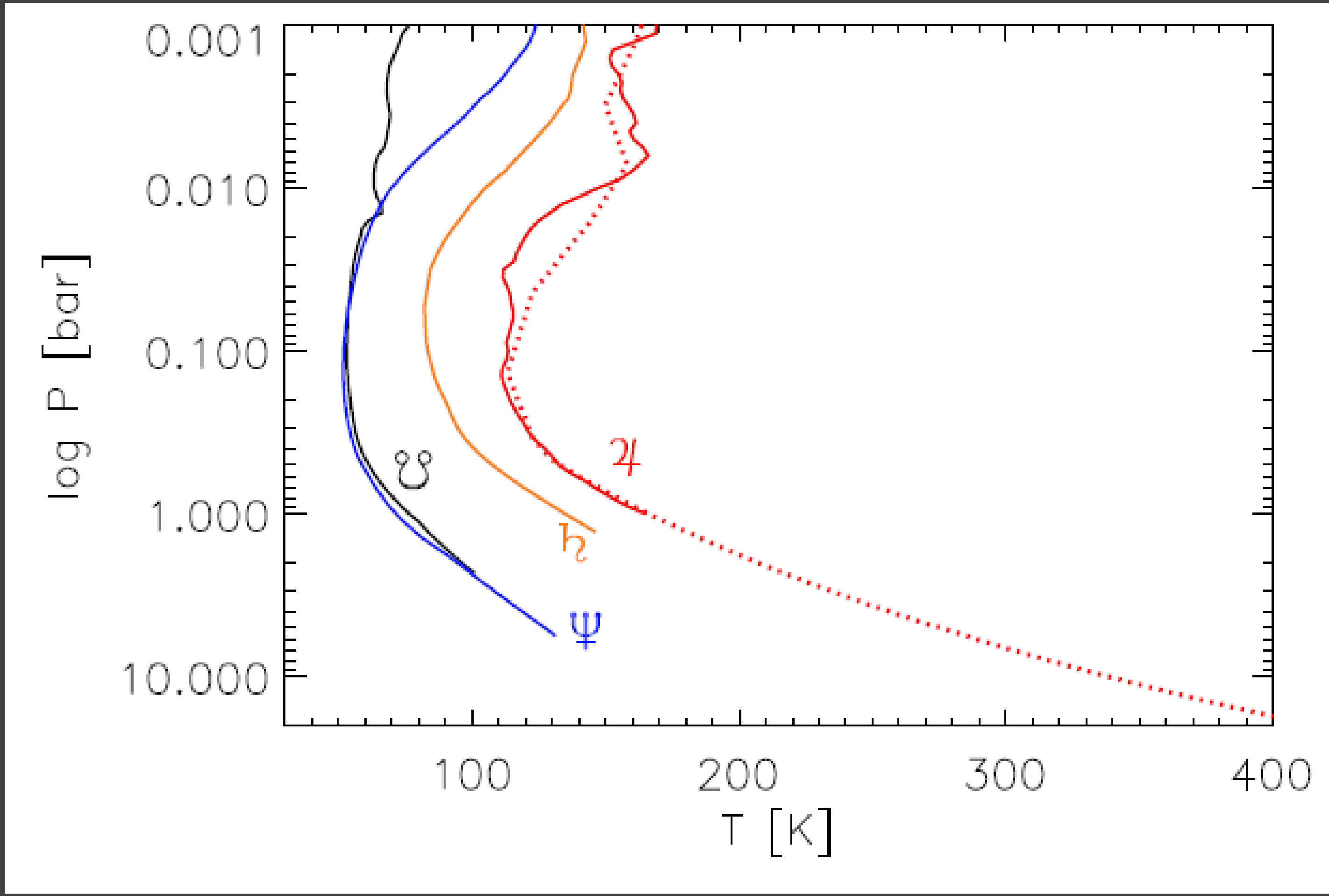


See 1812.07436 for a recent detailed review.

Except Uranus giant planets might not have solid cores.
However, there cores are made of heavy elements.
And so often they are called made of rocks.

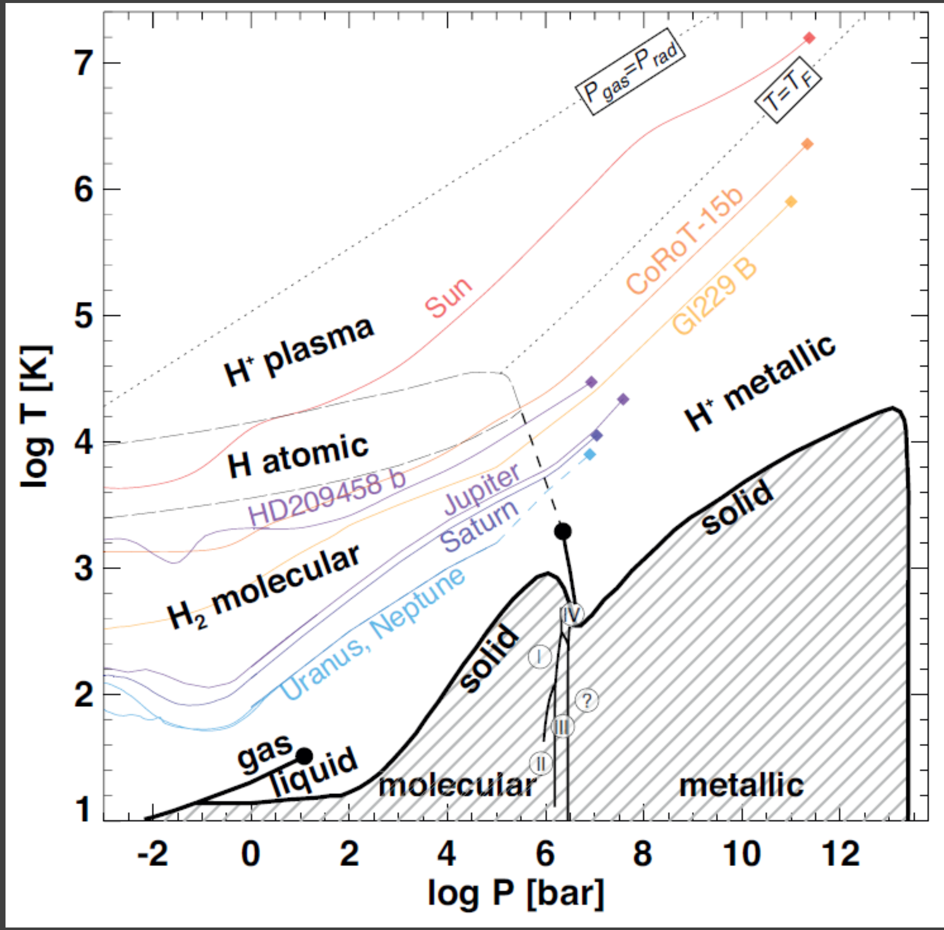


Temperature and pressure in atmospheres of giants



For Jupiter direct data are available due to Galileo probe measurements.

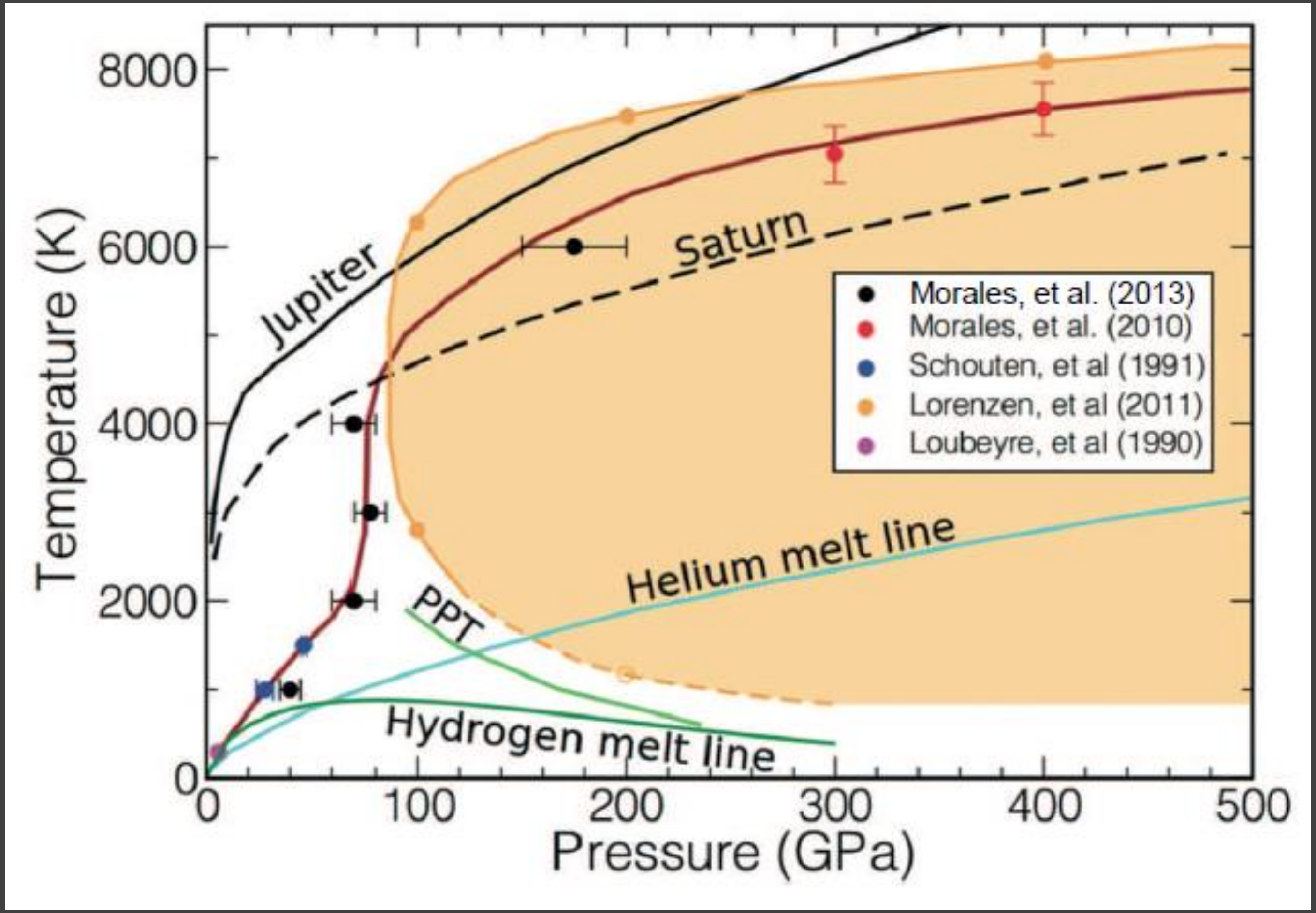
Hydrogene equation of state



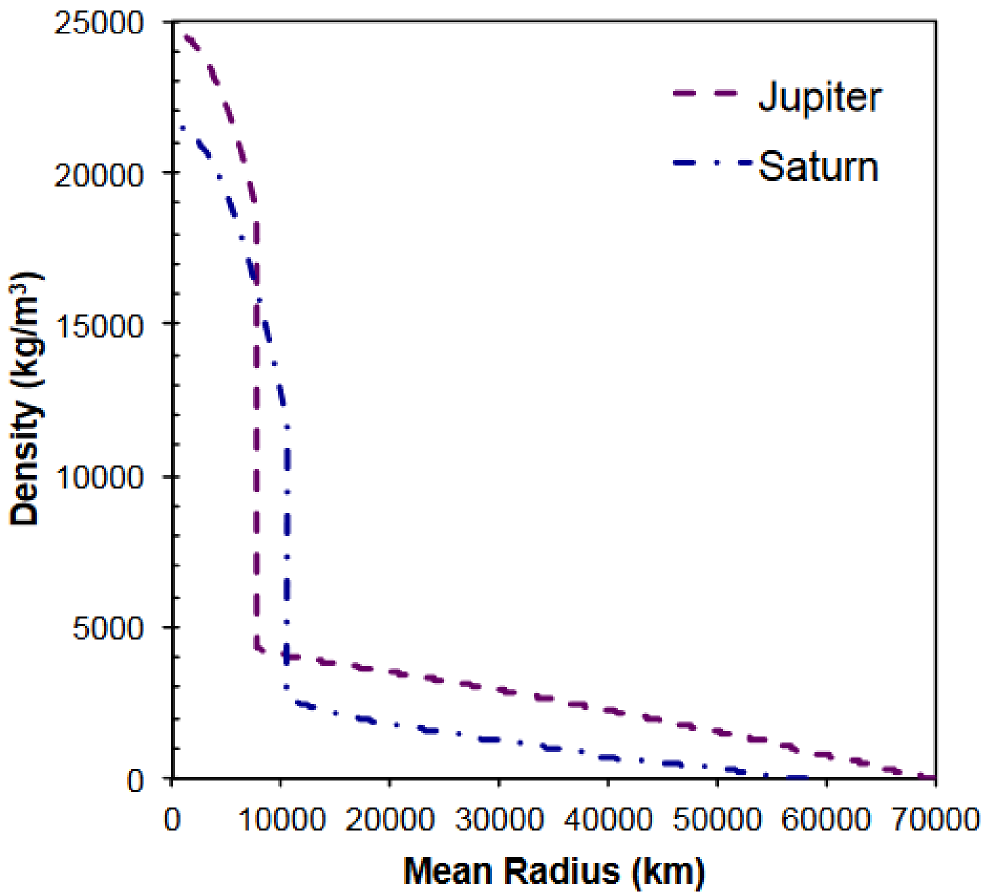
Still, there are important uncertainties even for the hydrogen equation of state.

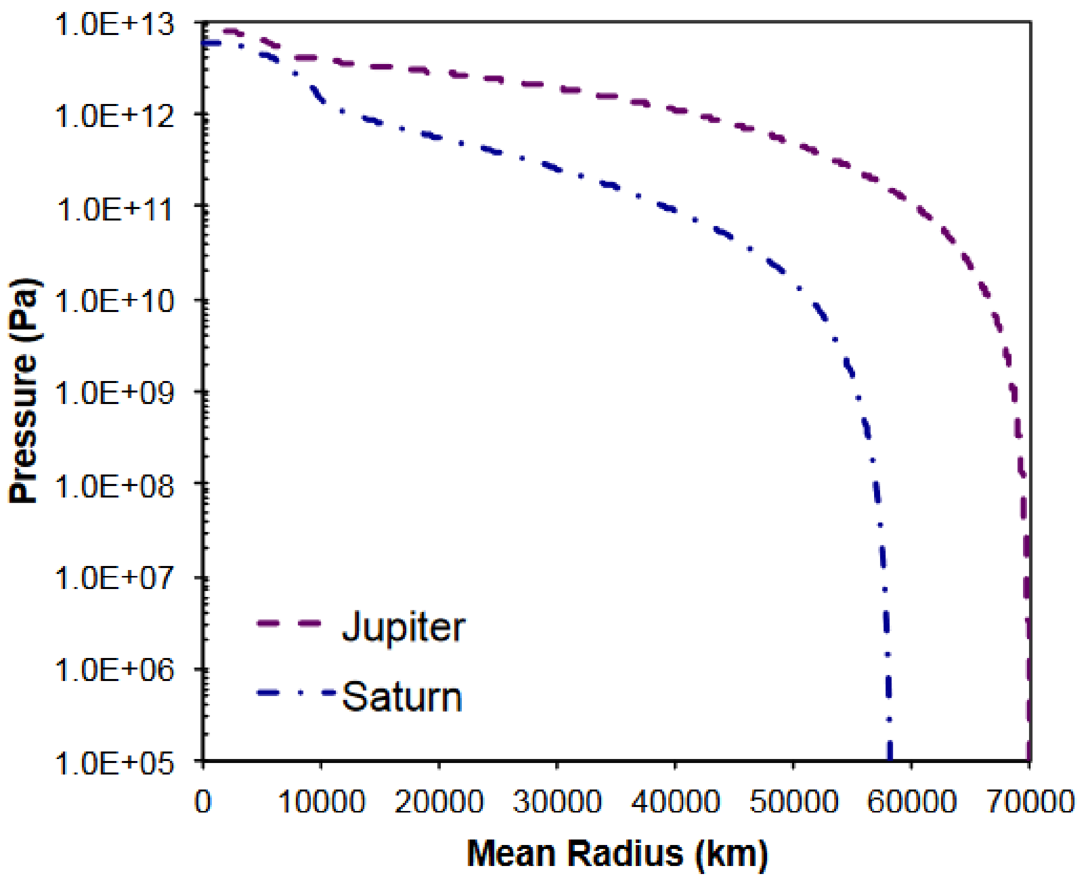
Some regimes have been never measured in laboratories.

Hydrogen plus helium mixture

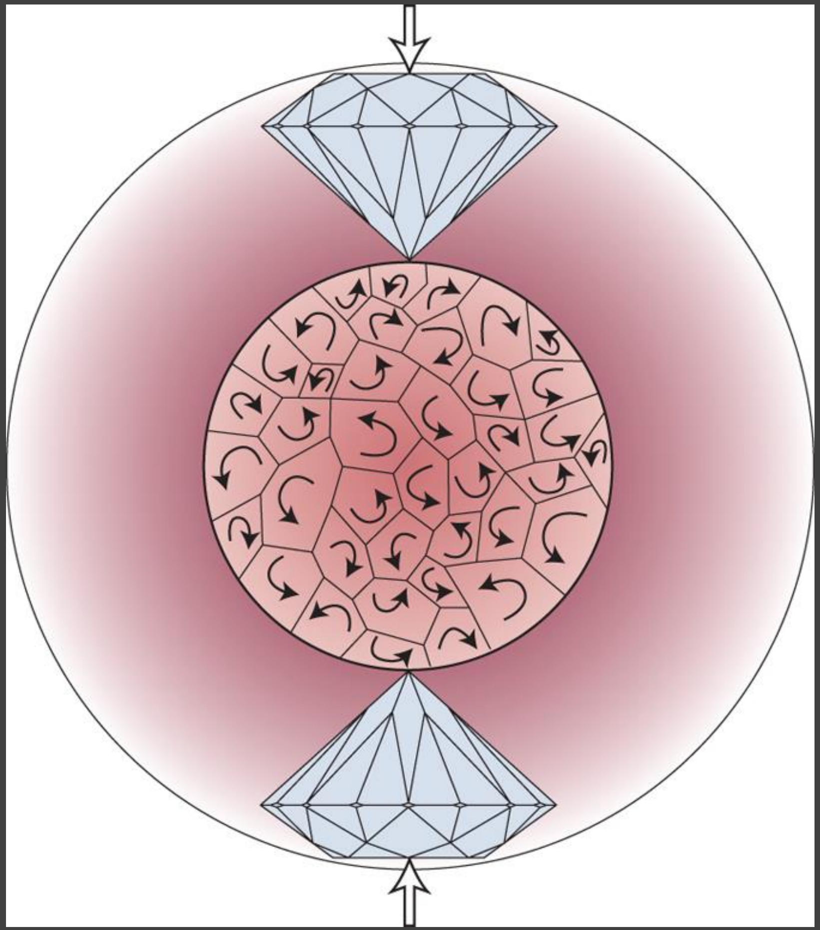


Density and pressure

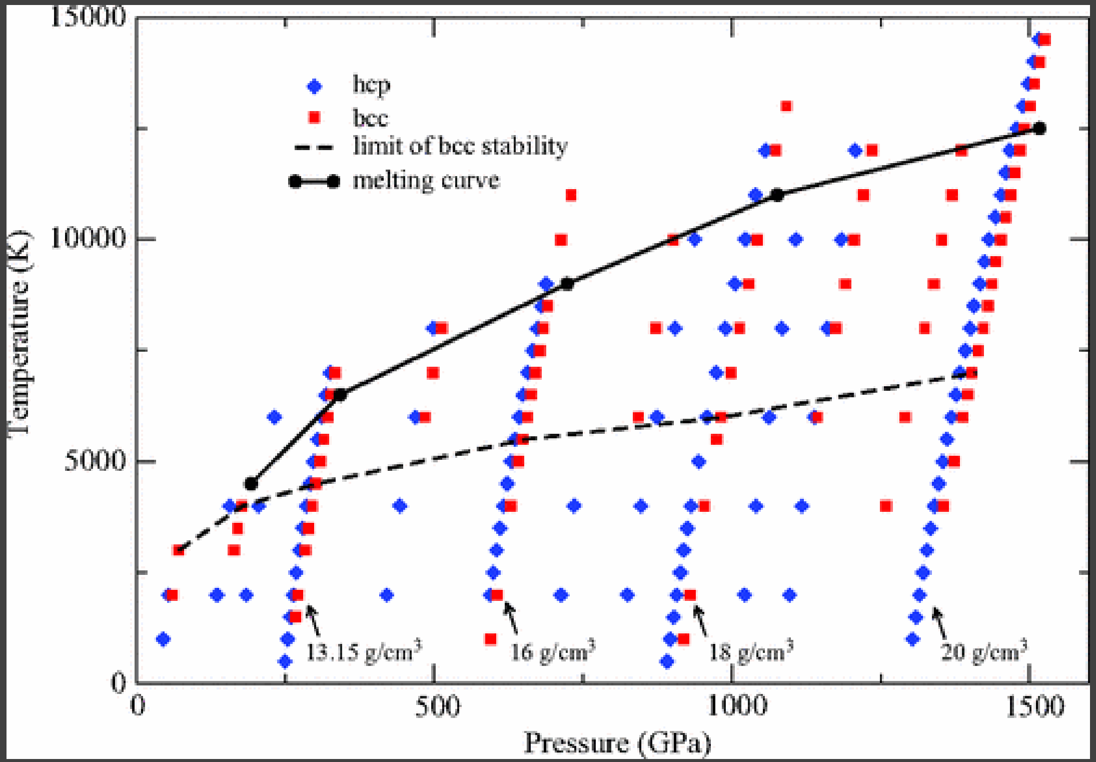




Diamond anvil cells



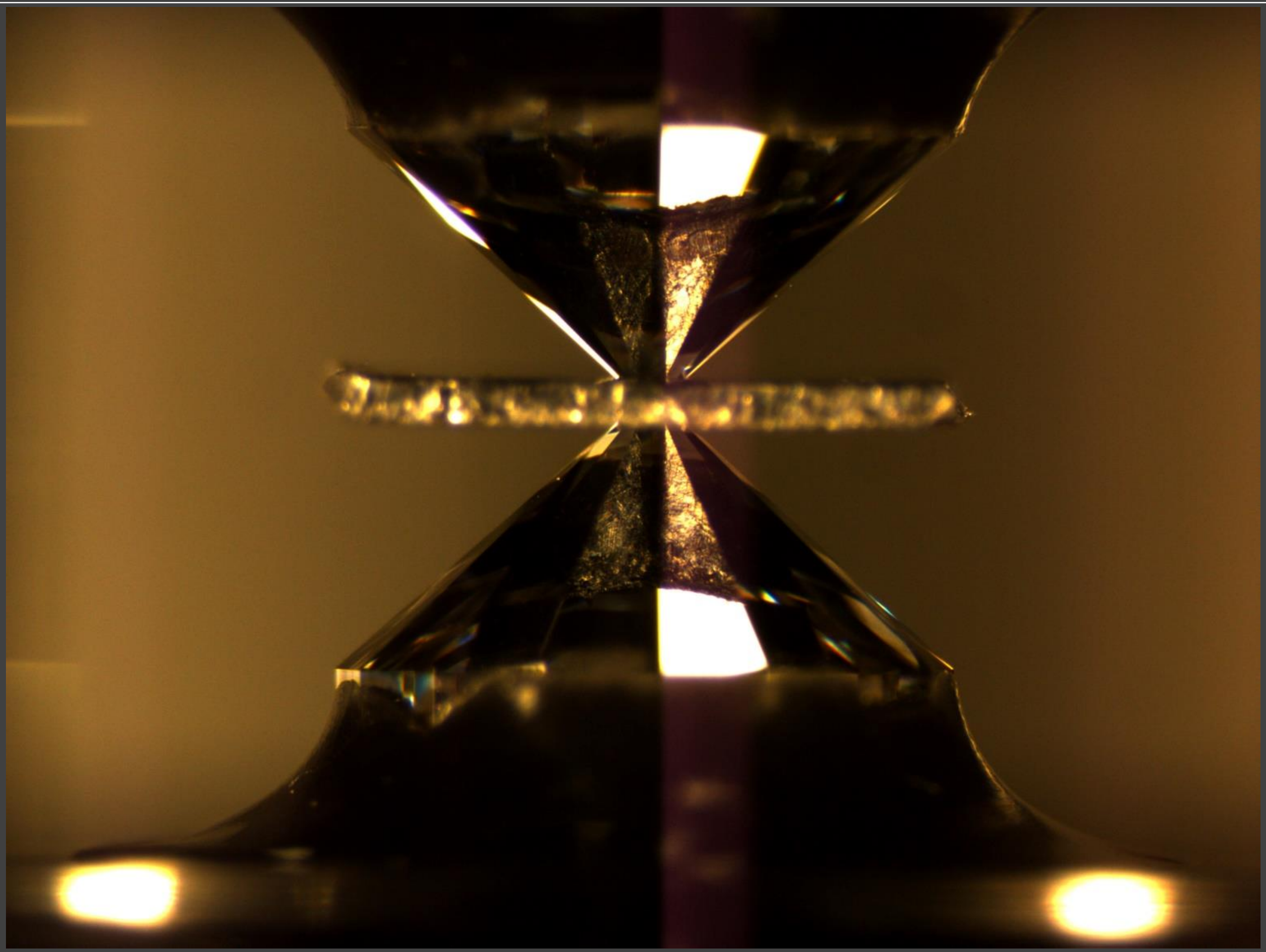
Diamond cells are used to reach high pressures in laboratory experiments. However, it is not enough, and in many cases we have to base only on numerical models.



Merkel (2013)

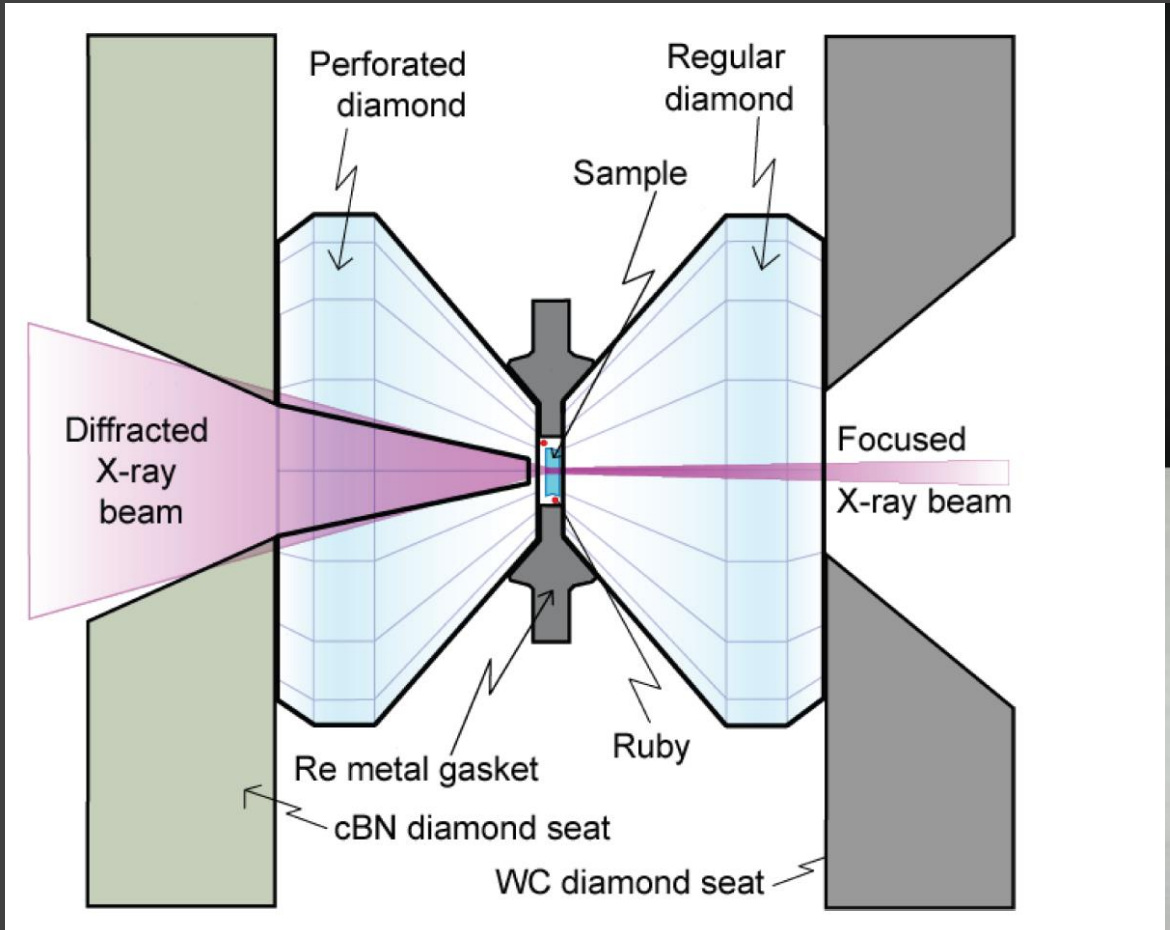
Bouchet et al. (2013)

Diamond cell



http://serc.carleton.edu/NAGTWorkshops/mineralogy/mineral_physics/diamond_anvil.html

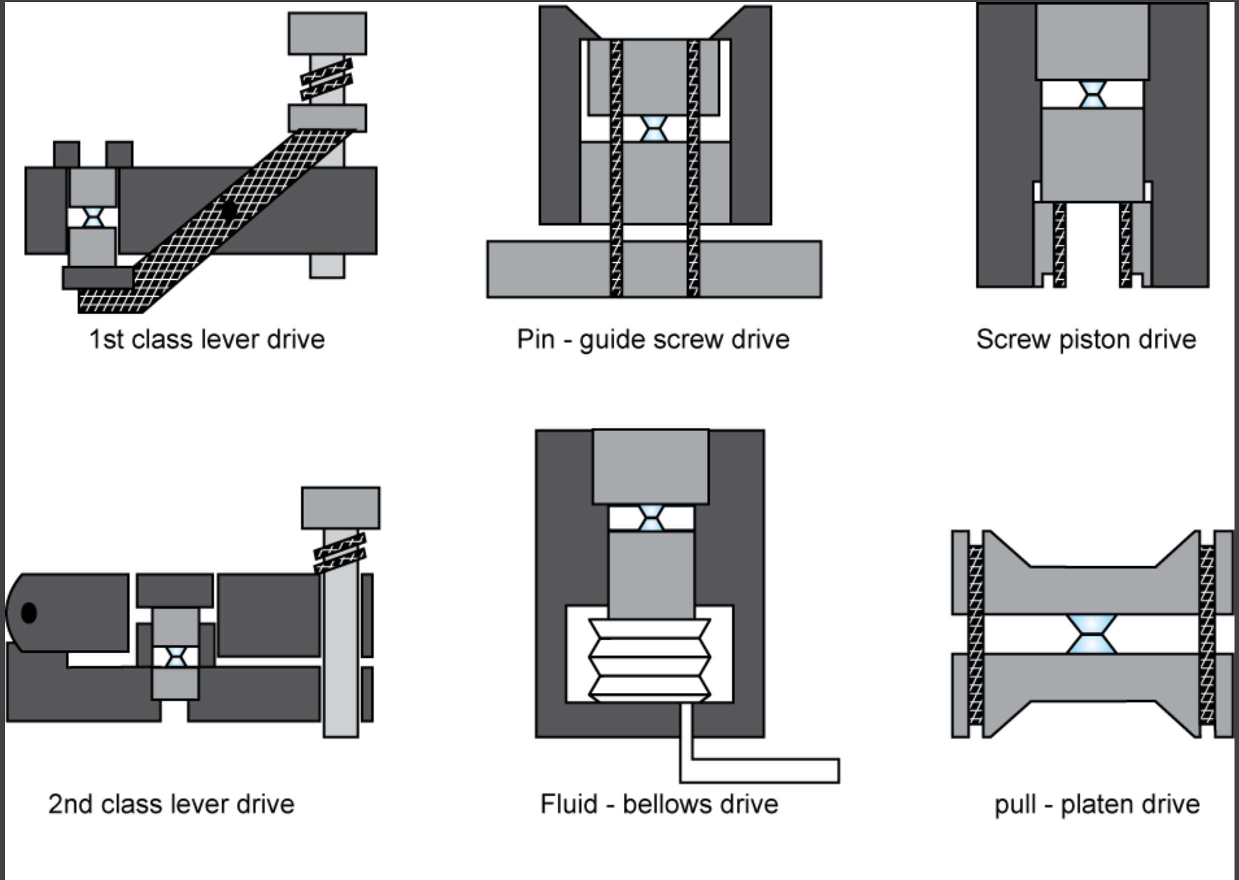
Scheme of the experiment



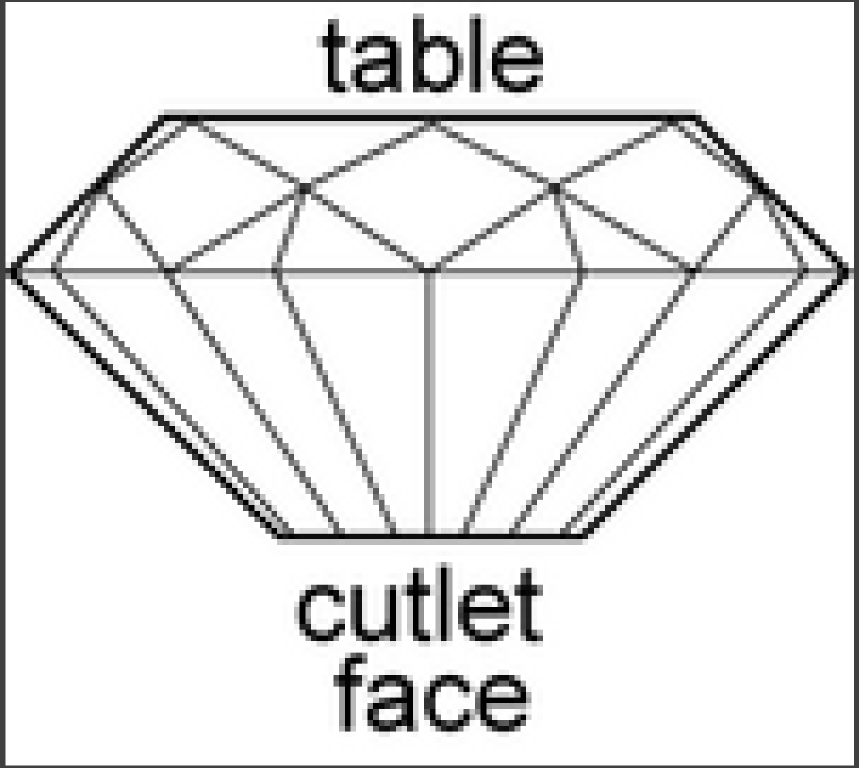
http://serc.carleton.edu/NAGTWorkshops/mineralogy/mineral_physics/diamond_anvil.html

How to press?

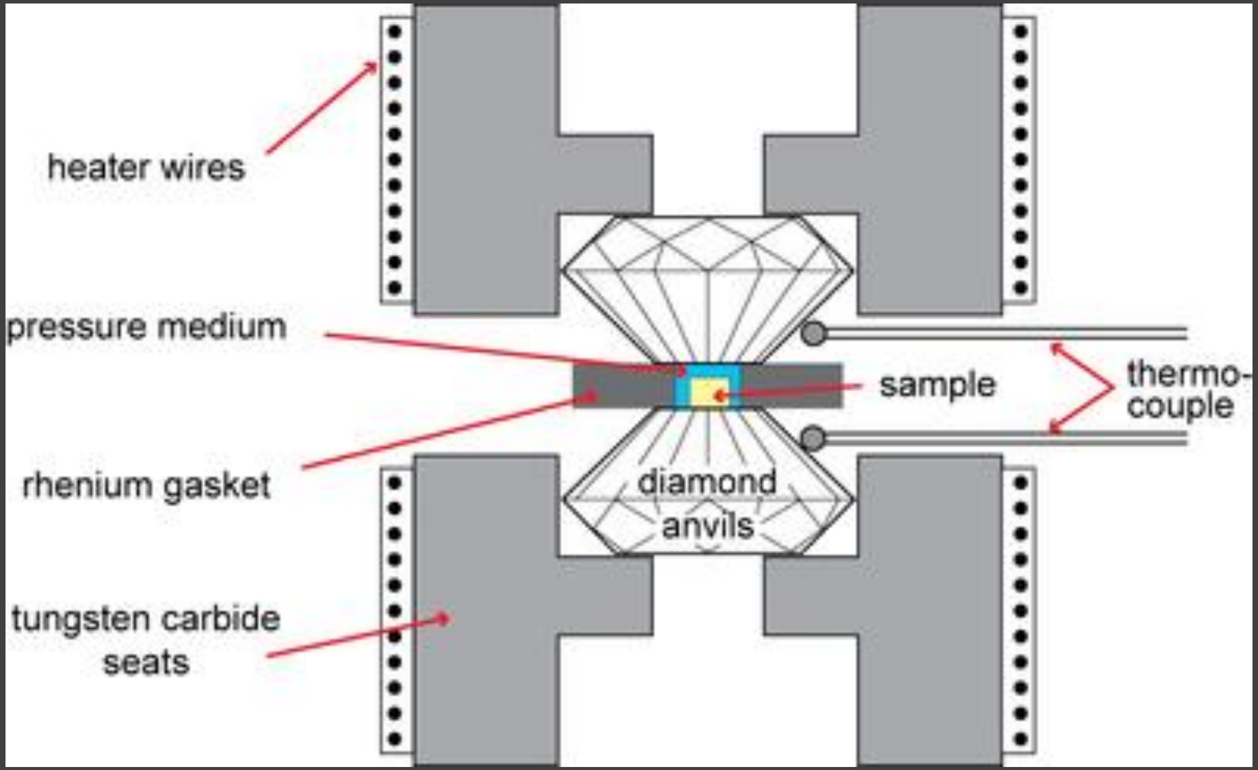
http://serc.carleton.edu/NAGTWorkshops/mineralogy/mineral_physics/diamond_anvil.html



six basis ways of providing force in the DAC (Bassett, 1979)

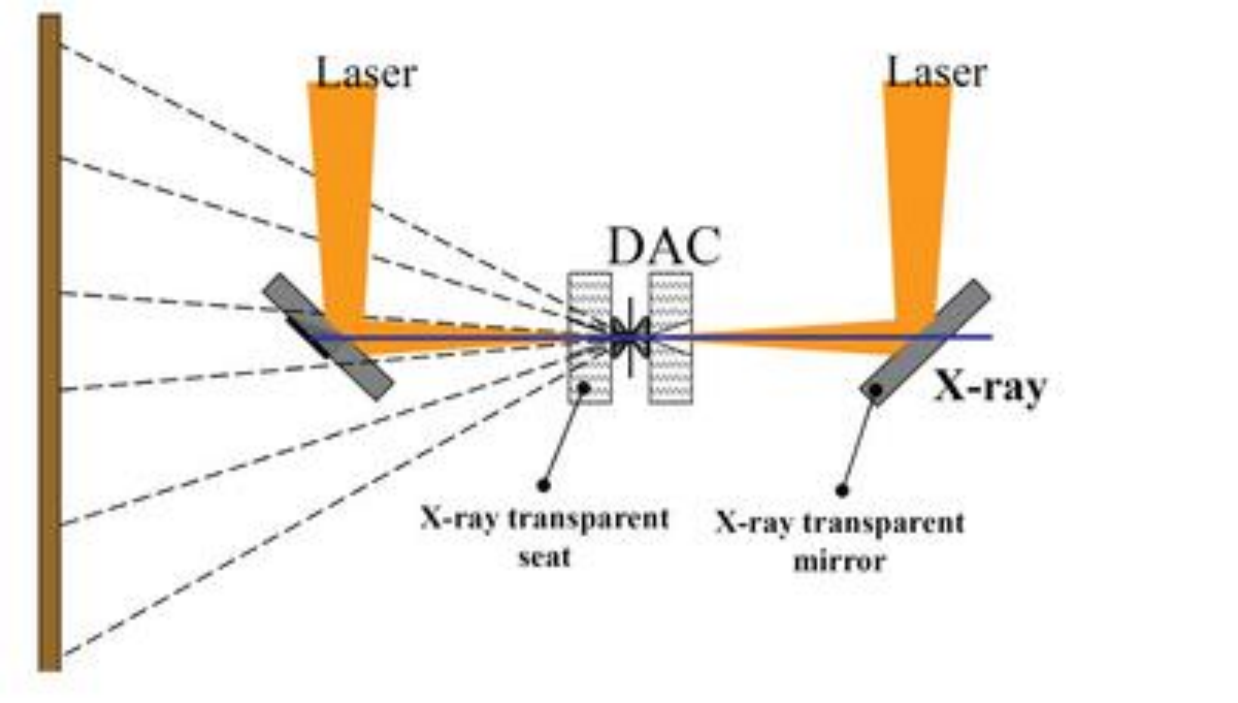


How to heat the matter



Up to 1300K

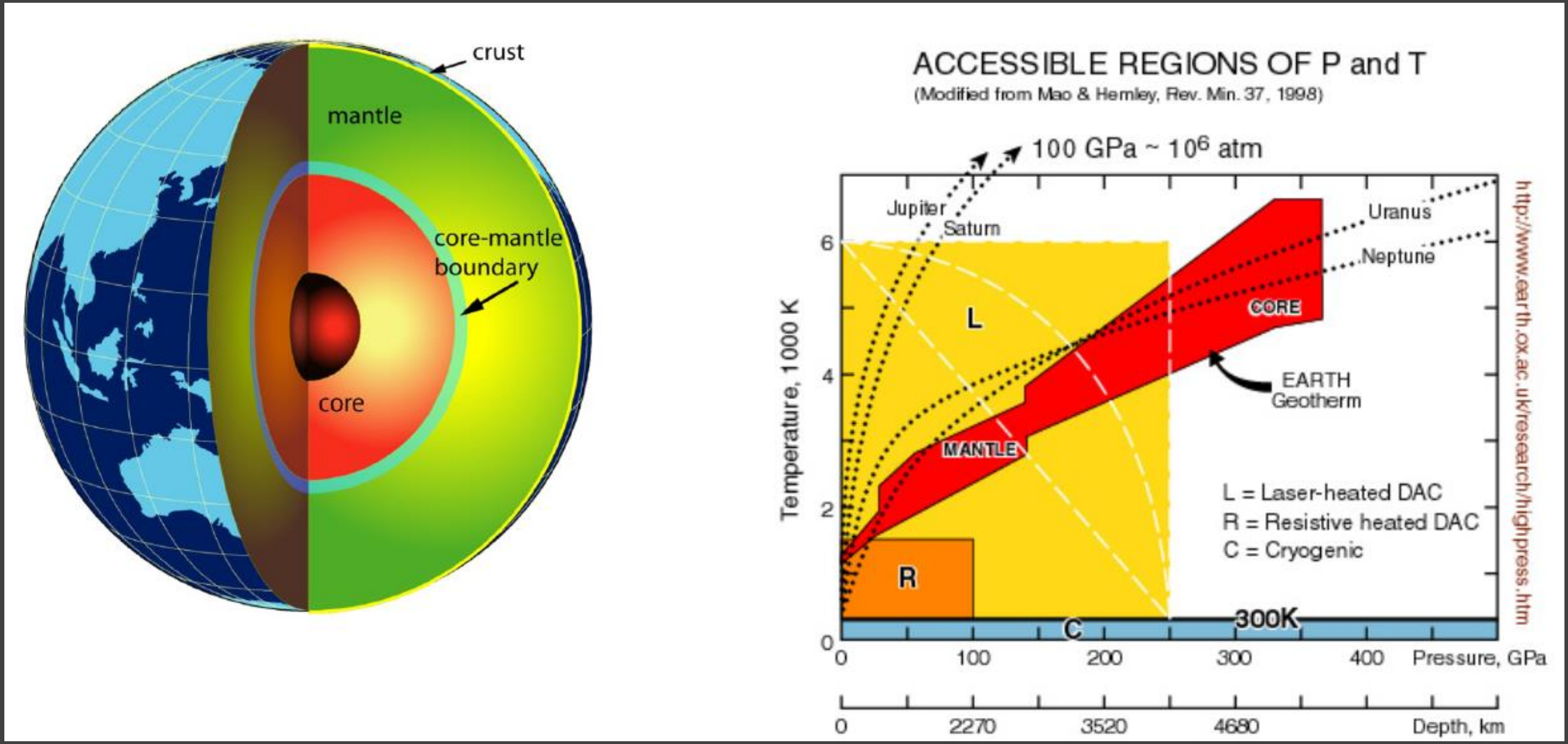
Electric current (for lower temperatures)
or laser (for higher temperature).



Above 1300K

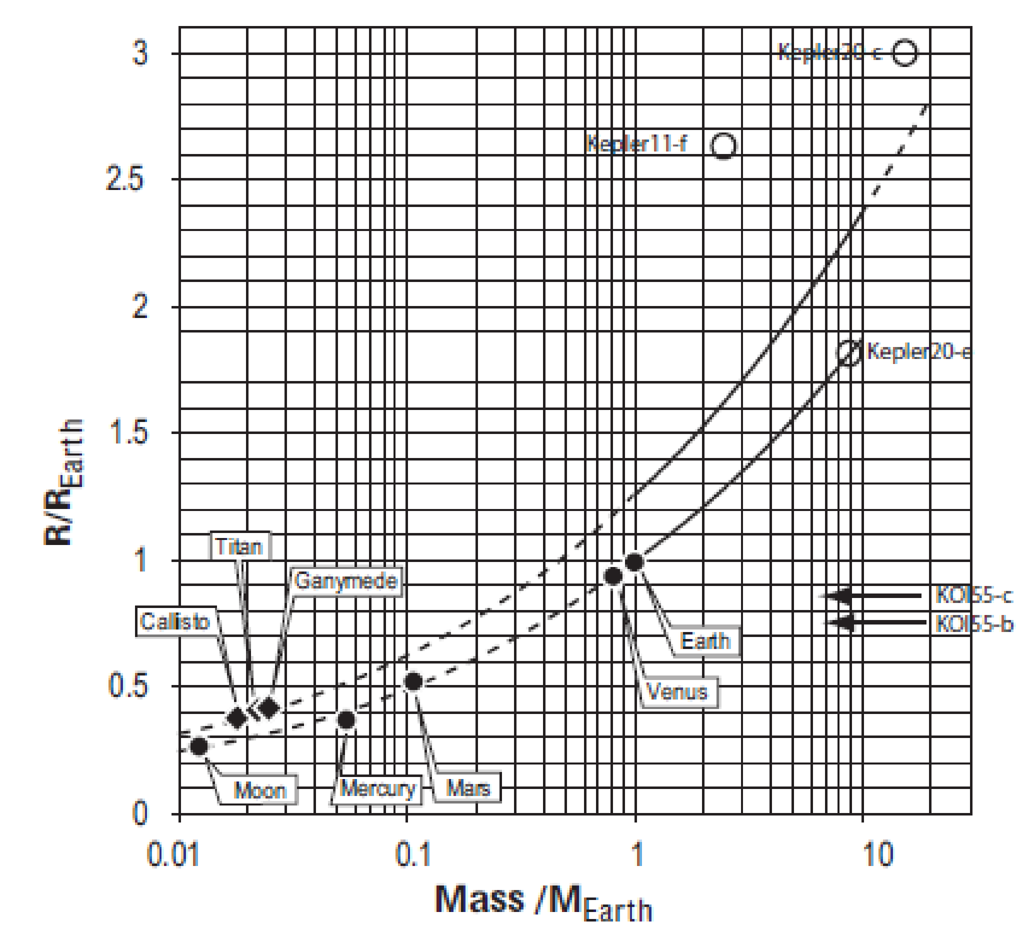
http://serc.carleton.edu/NAGTWorkshops/mineralogy/mineral_physics/diamond_anvil.html

Comparison with conditions in the Earth



http://mini.physics.sunysb.edu/~pstephens/downloadable/ehm_dac.pdf

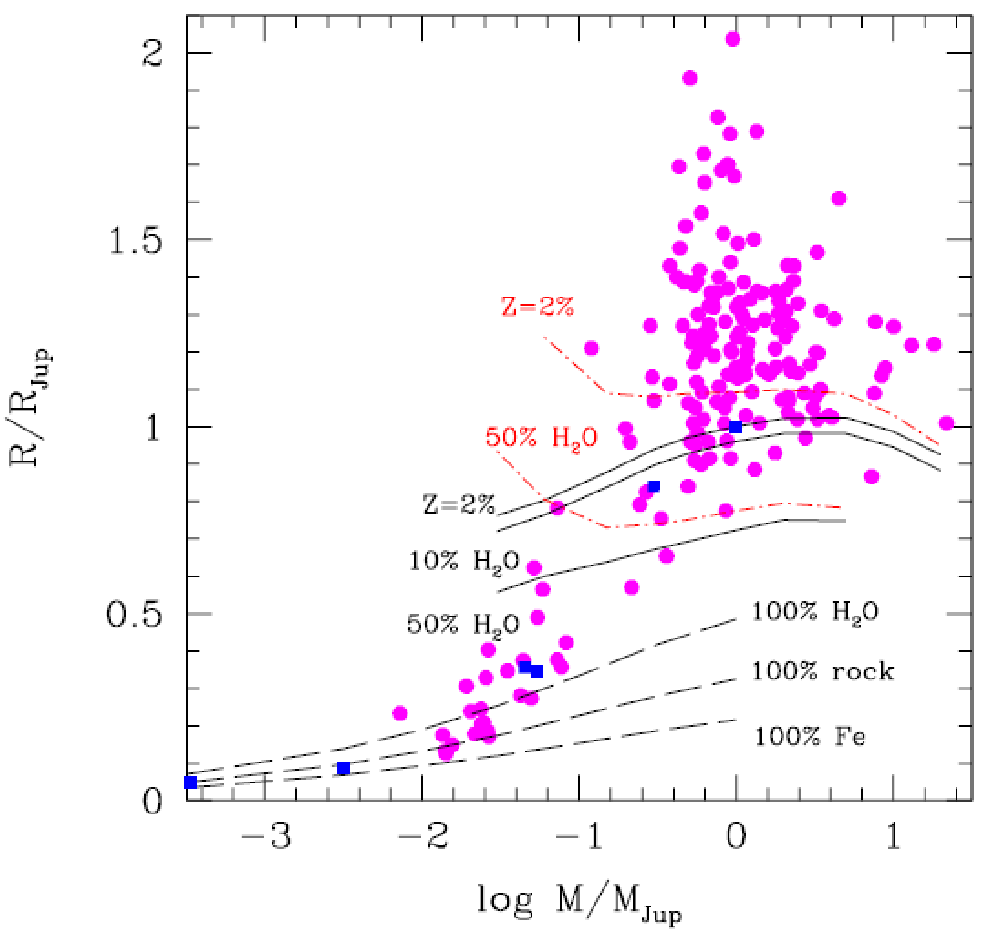
Mass-radius models for planets



Relatively simple model
based on just 8 key elements.

Good results for Solar system planets.

Mass-radius diagram for exoplanets



Planet radius, of course, depends on its composition.

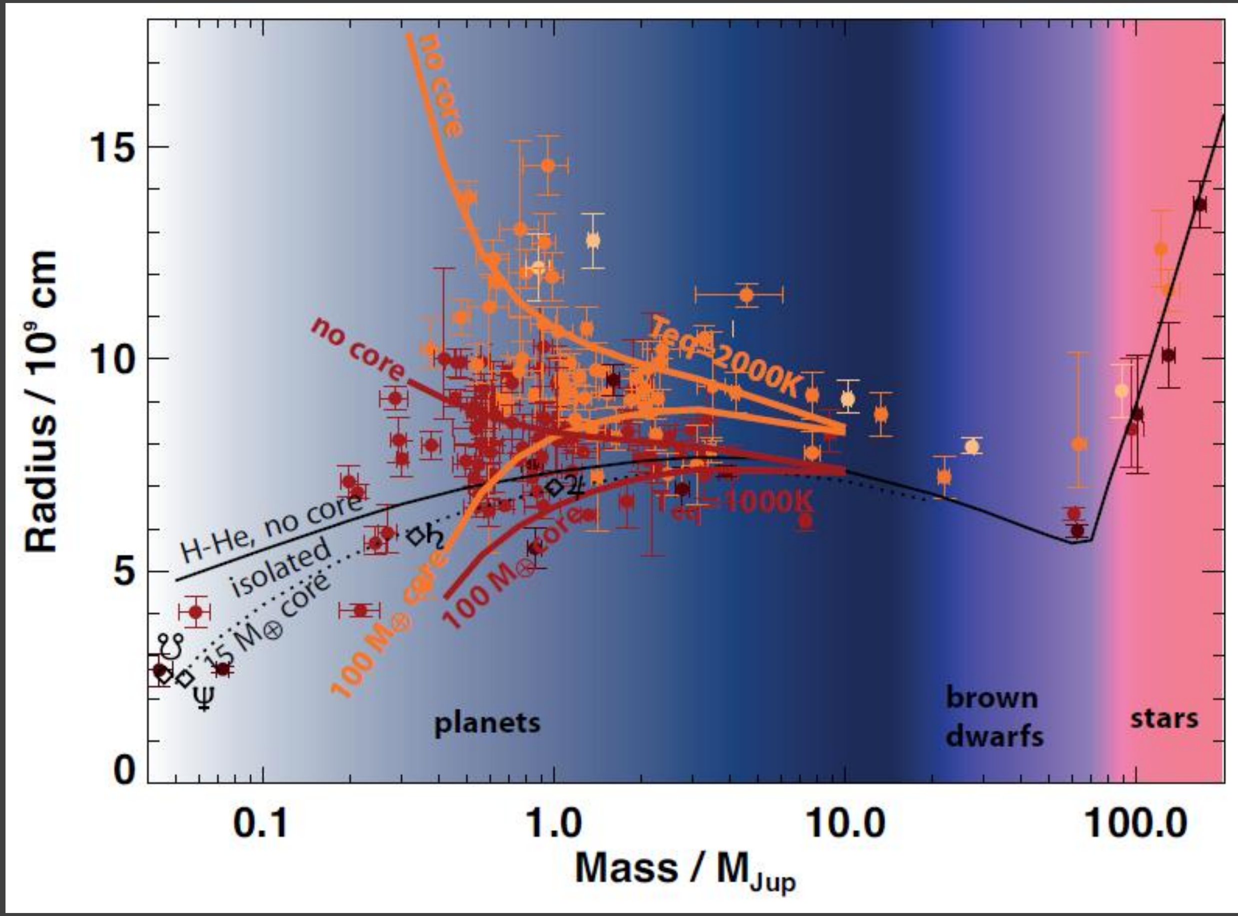
Light planets typically do not have extended gas envelopes.

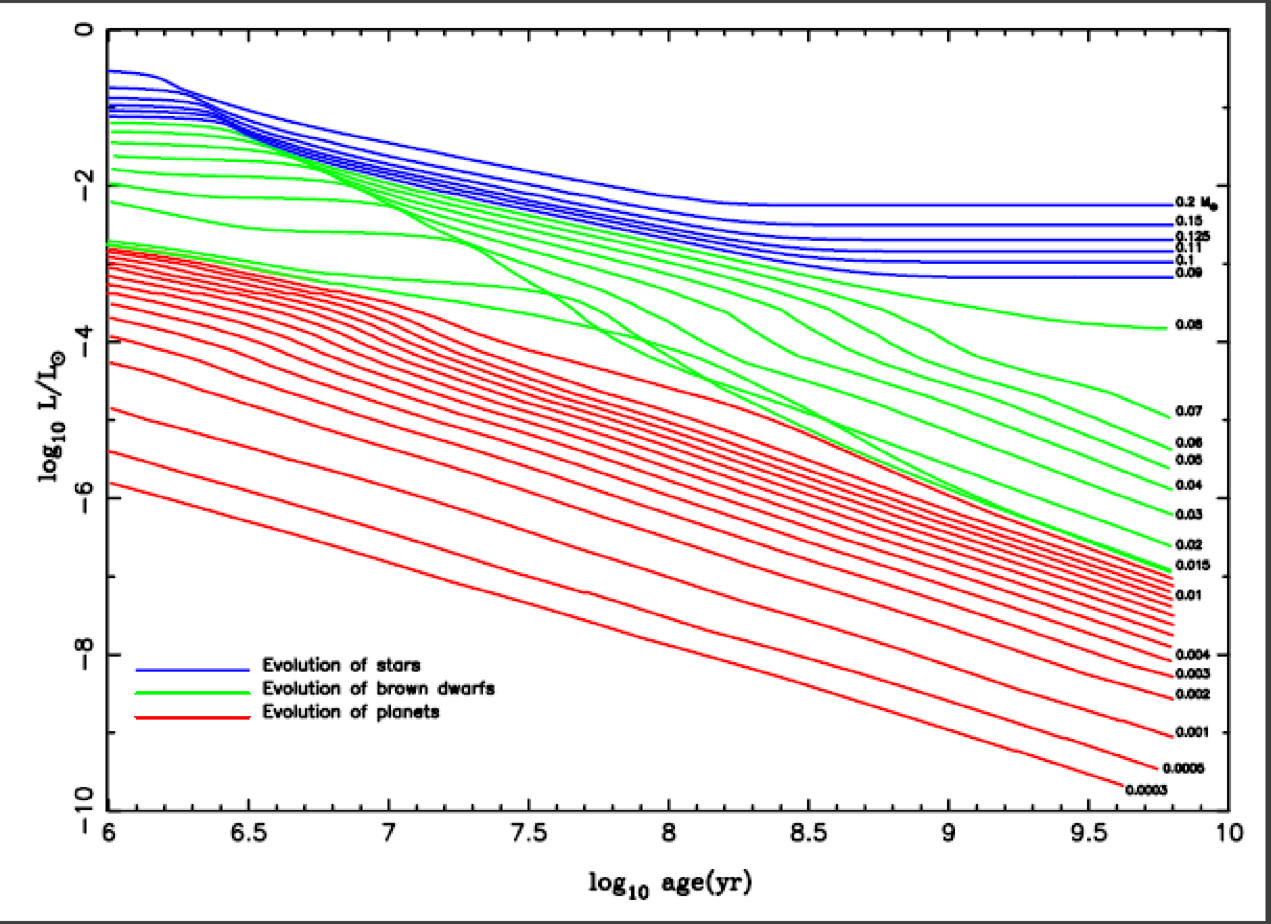
Oppositely, giant planets might hath very thick gas envelopes.

Very often data on mass and radius
can be explained by different combinations of ingredients.

Solid and long-dashed lines (in black) are
for non-irradiated models.
Dash-dotted (red) curves correspond to
irradiated models at 0.045 AU from a Sun.

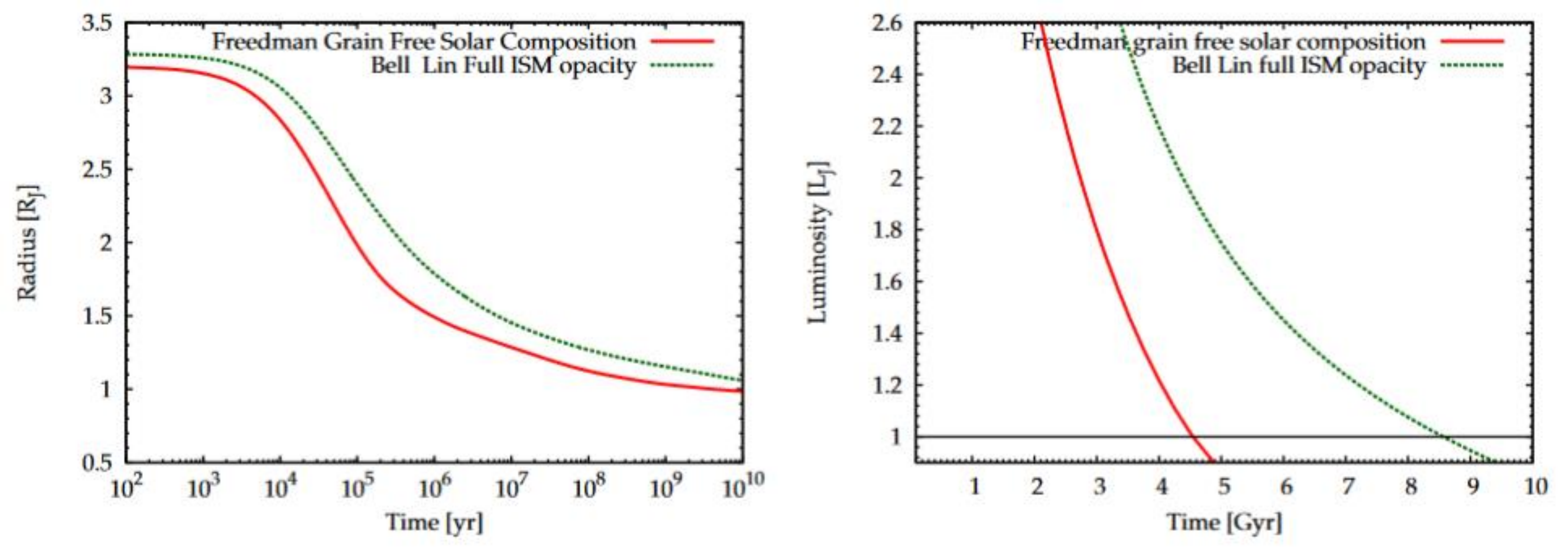
Theory vs. observations



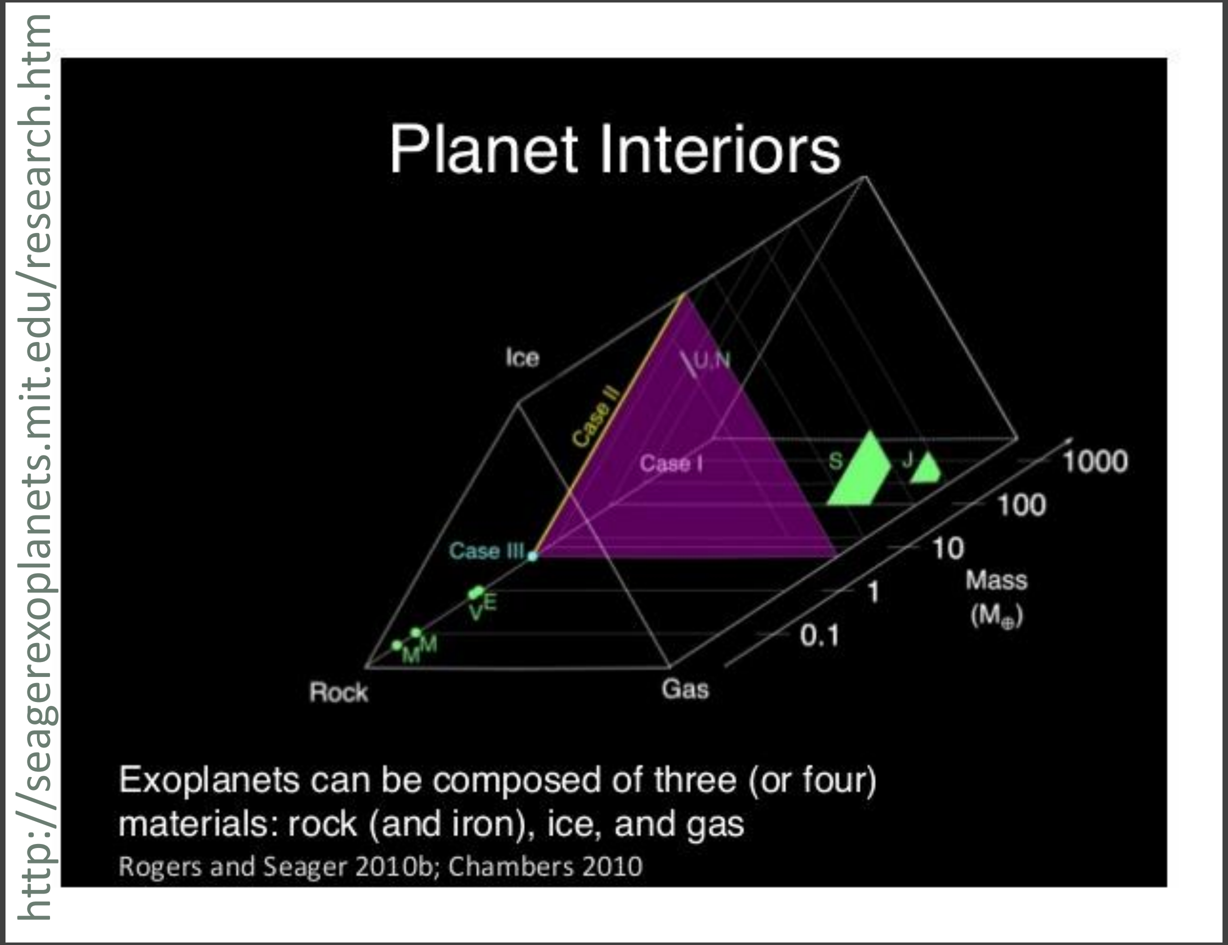


1405.3752

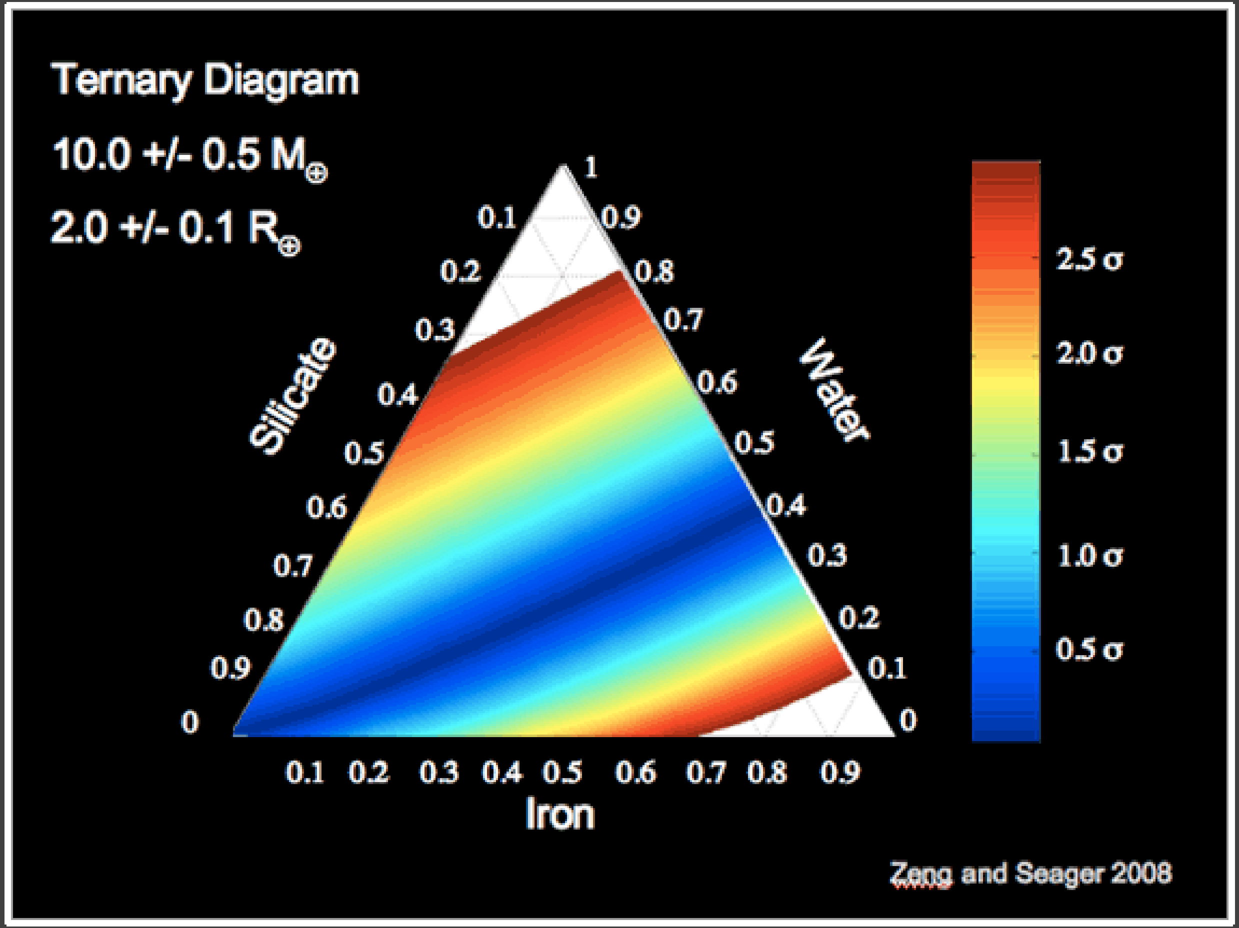
Evolution of giant planets



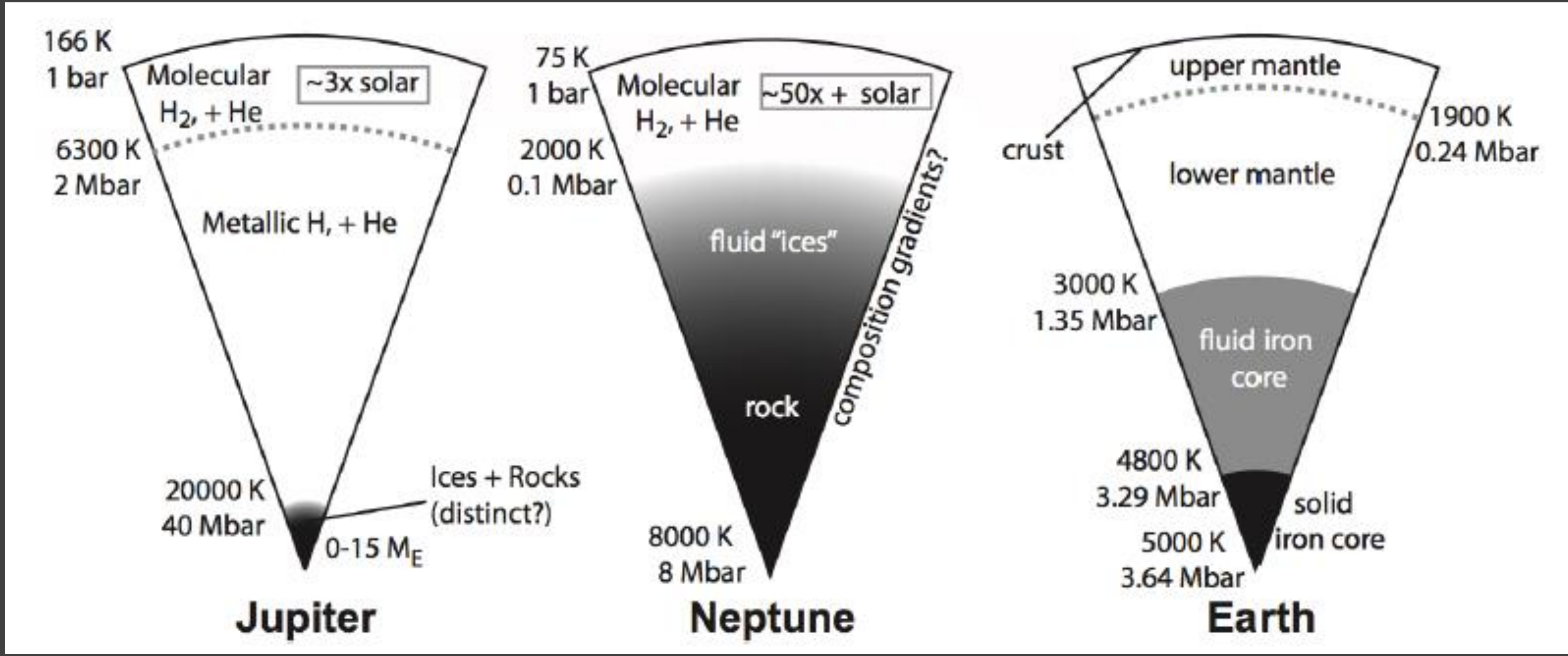
Three main ingredients: gas, ice, rock



http://seagerexoplanets.mit.edu/research.htm

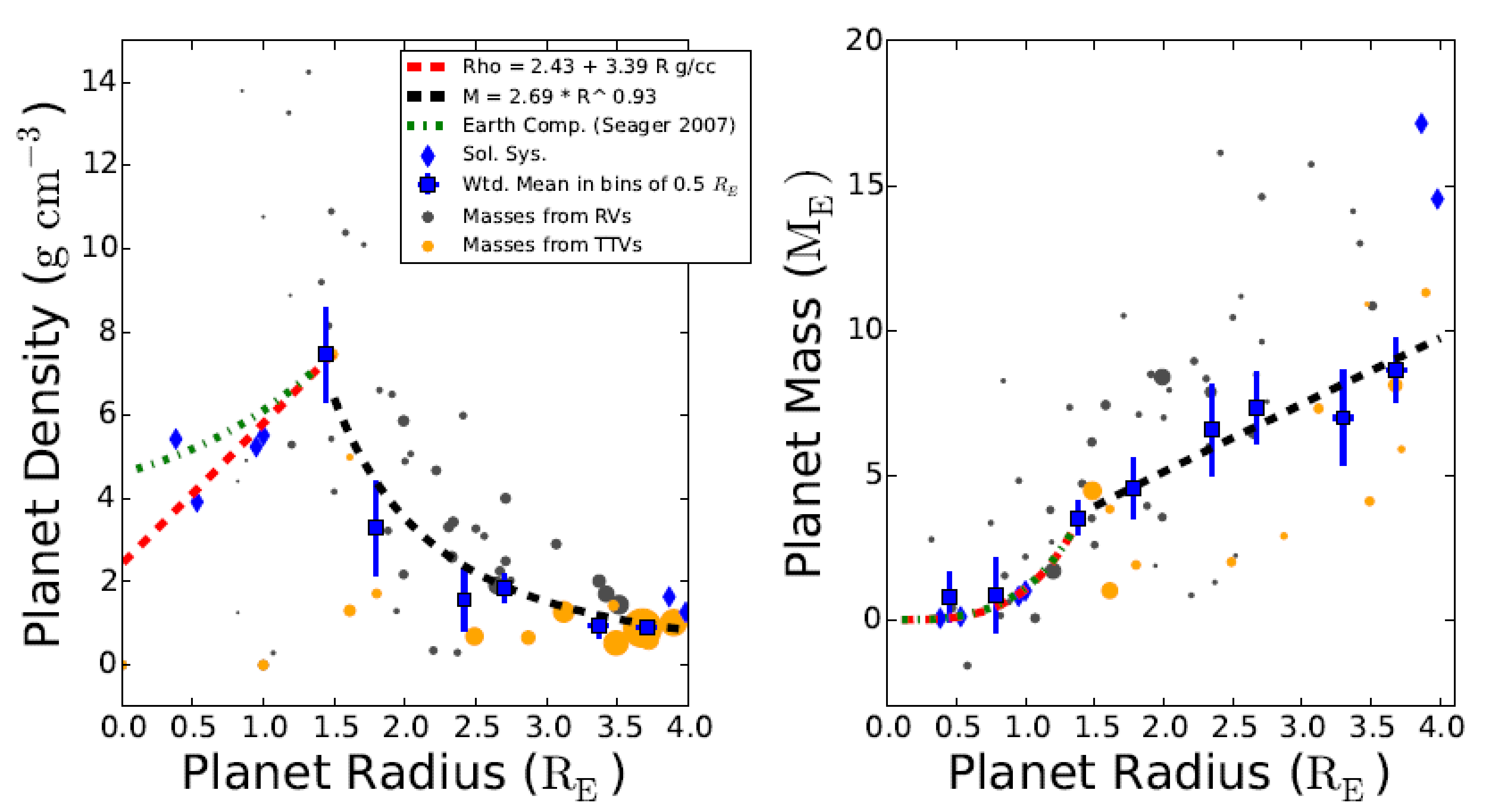


Three main types of planets

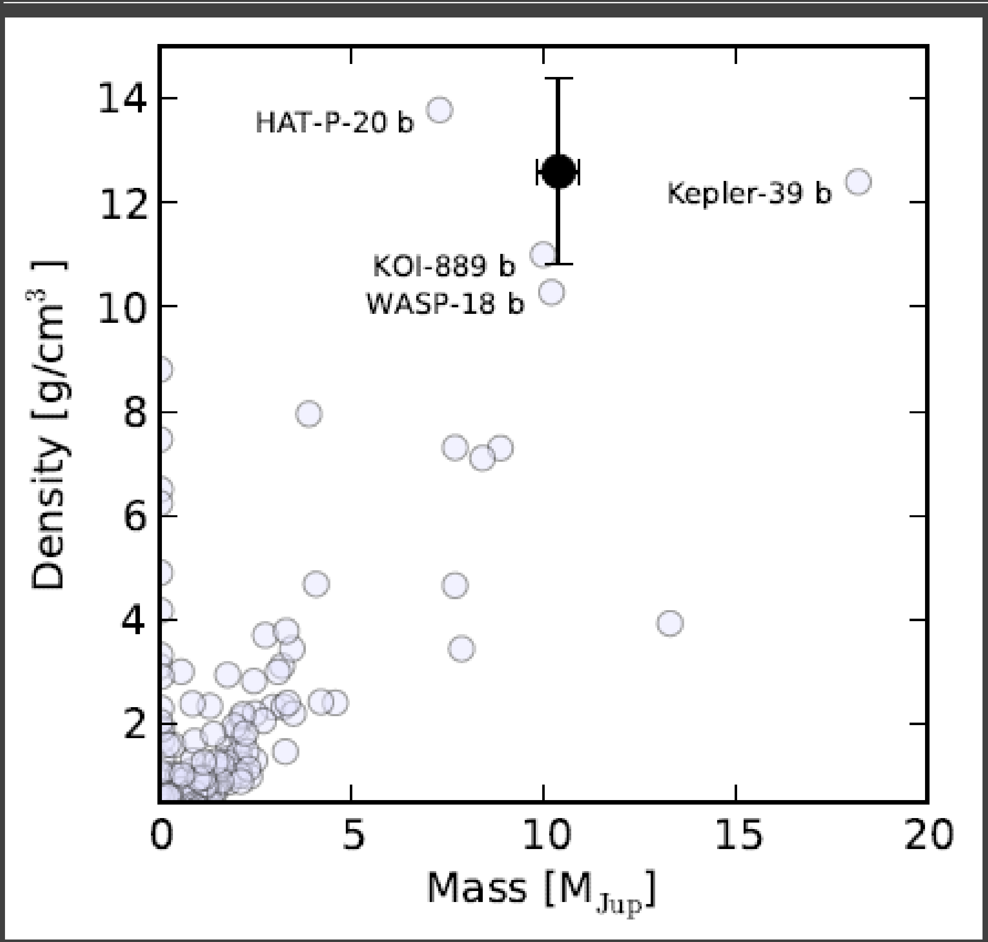


Gas giants
H/He | Icy giants
H/He+ice+core | Solid planets
Si, Mg, Fe, O, C

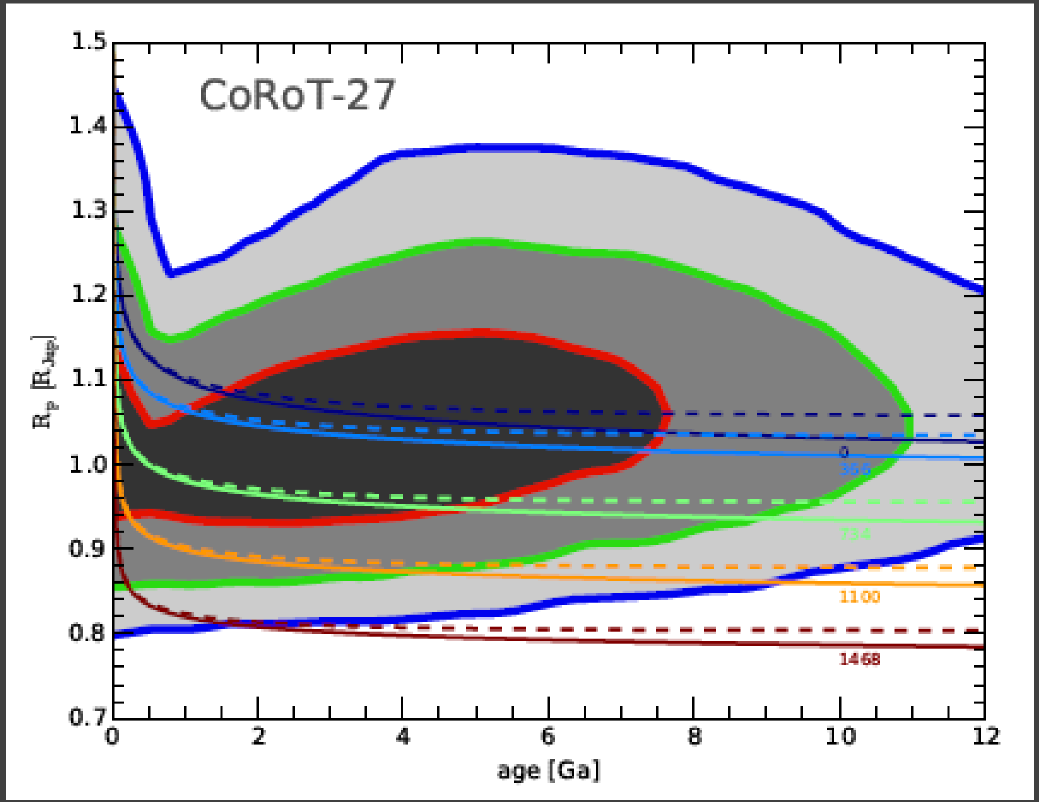
Thick atmospheres for $M>4M_{Earth}$



Corot-27b. Dense planet



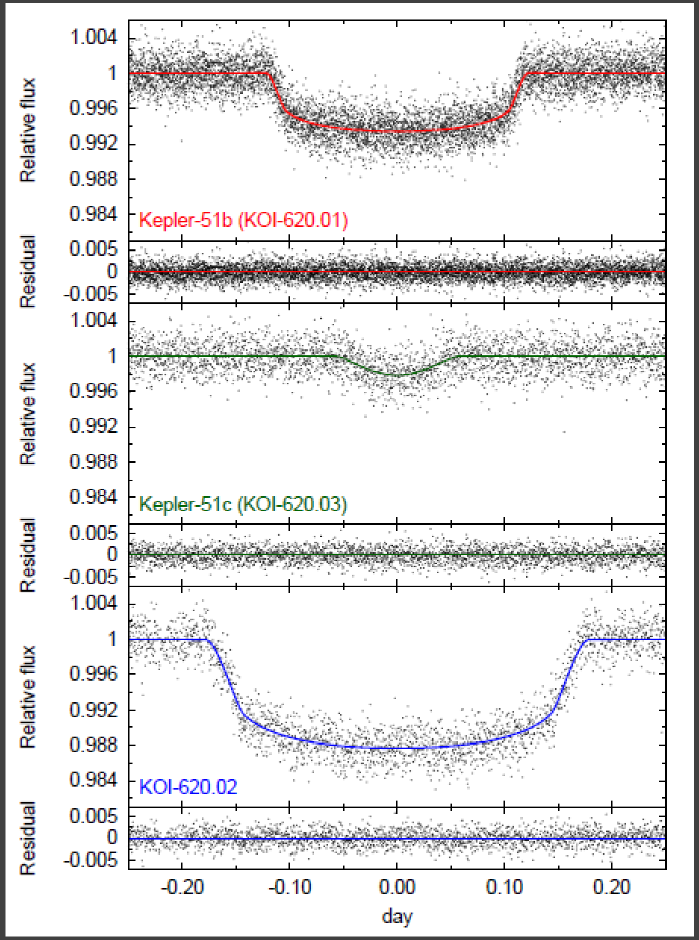
Orbital period 3.6 days.
Solar-like star



Now the densest is KELT-1b: 23.7+/- 4 gcm^{-3}. 1808.04533

1401.1122

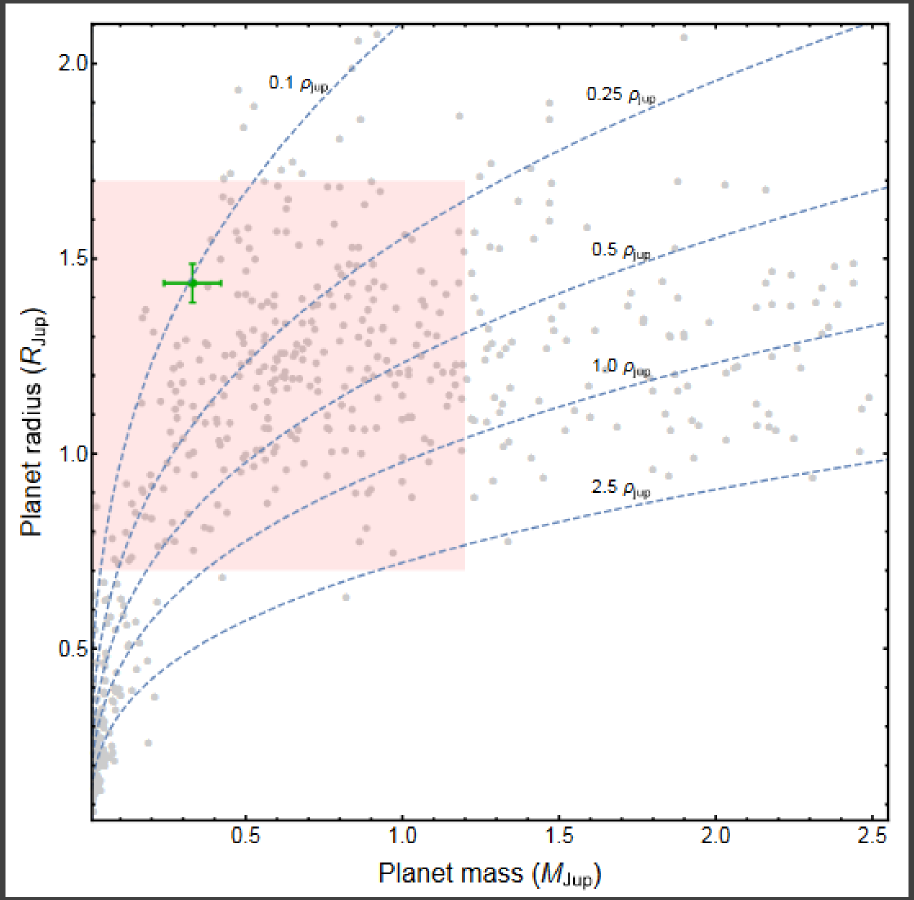
Kepler-51. Crumbly planets.



Solar type star.
Three planets with masses 2-8 M_{earth}
and low densities: <0.05 g/cm^3

Orbital periods 45-130 days.

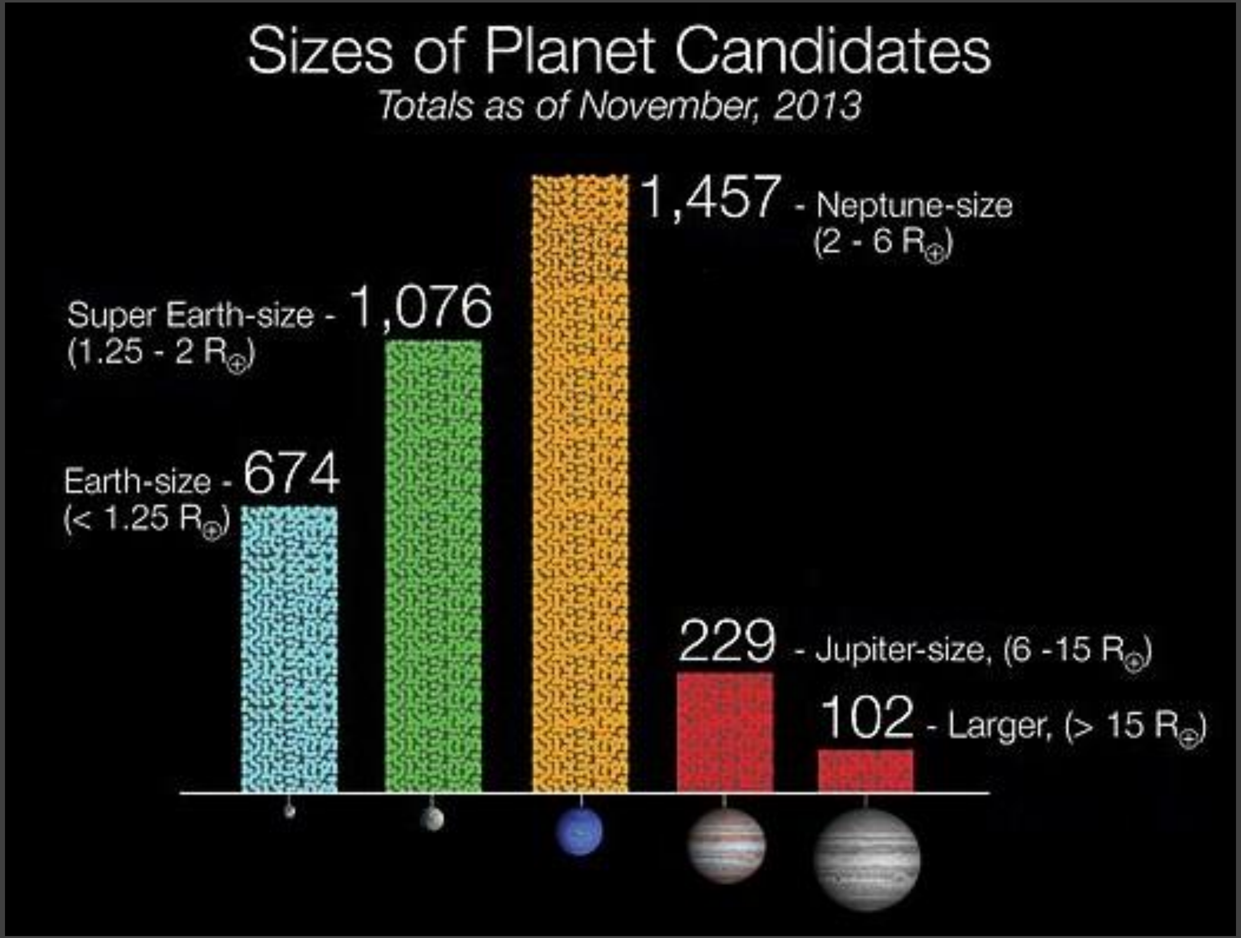
Inflated hot jupiter

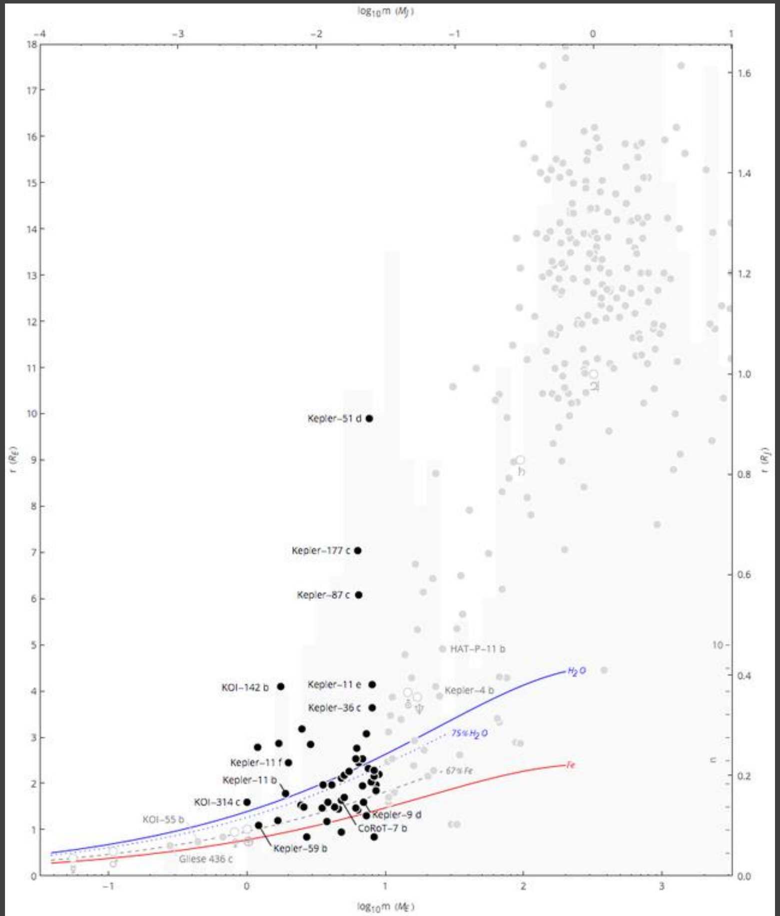


Mass and radius measured together.
Grazing transit.
Density 0.1-0.17 g/cm3

Superearths. Diversity of properties.

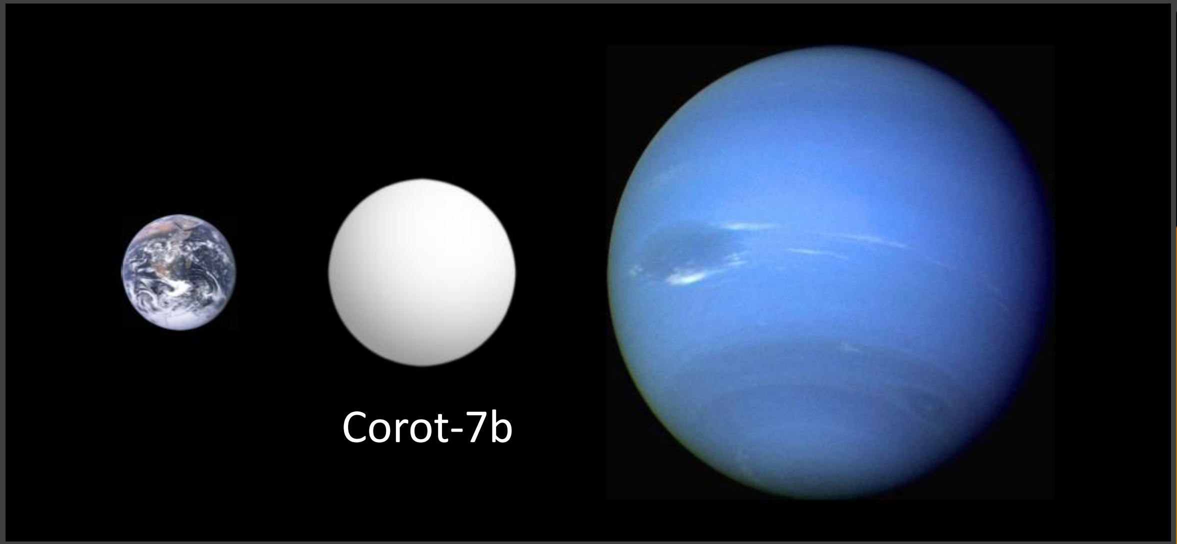
Wikipedia

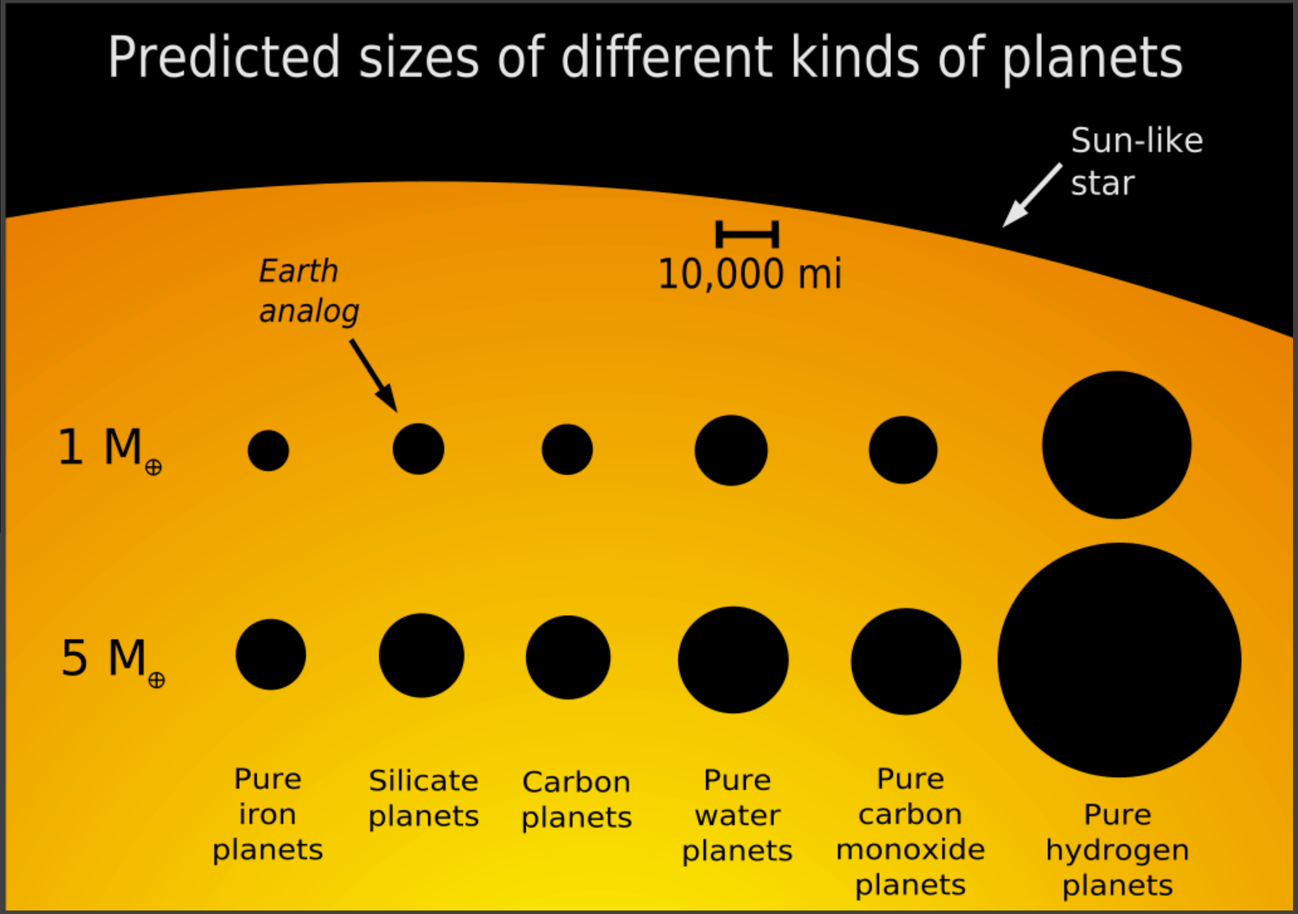




Sizes of superearths

Wikipedia

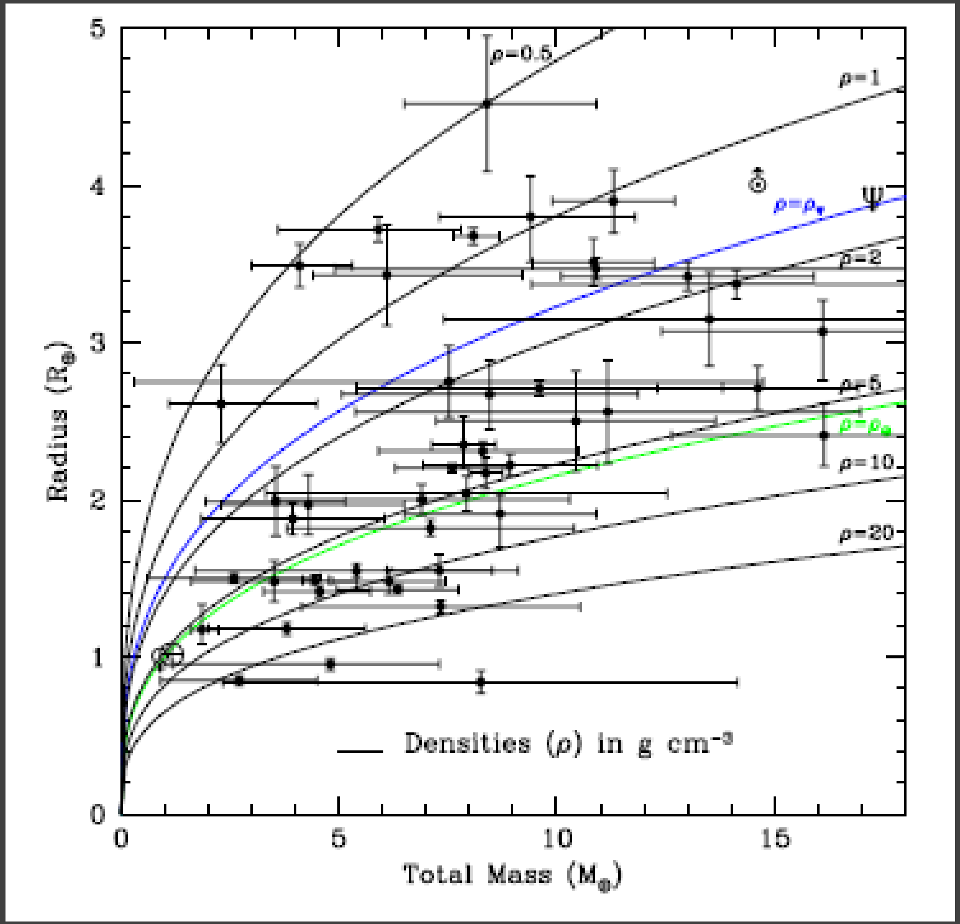




Typical radii 1-4 of the Earth
I.e., between the Earth and Neptune).

Sometimes low density planets in the range are called mini-Neptunes.

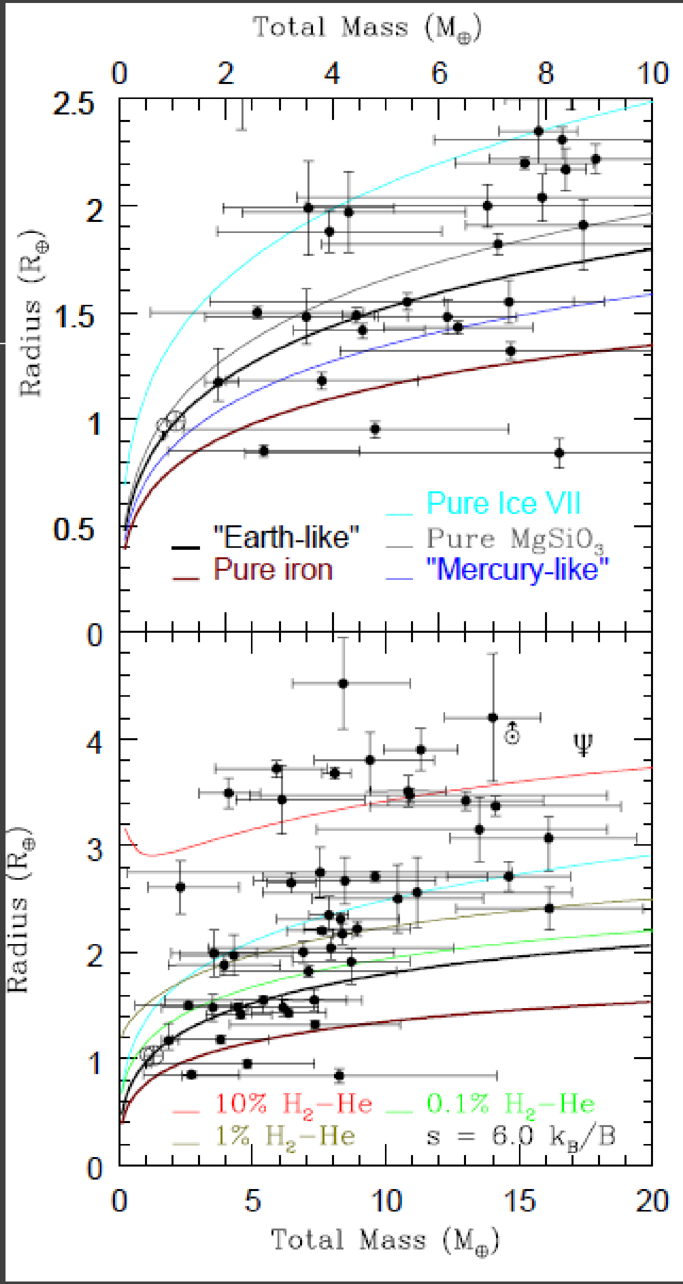
Superearths: mass-radius



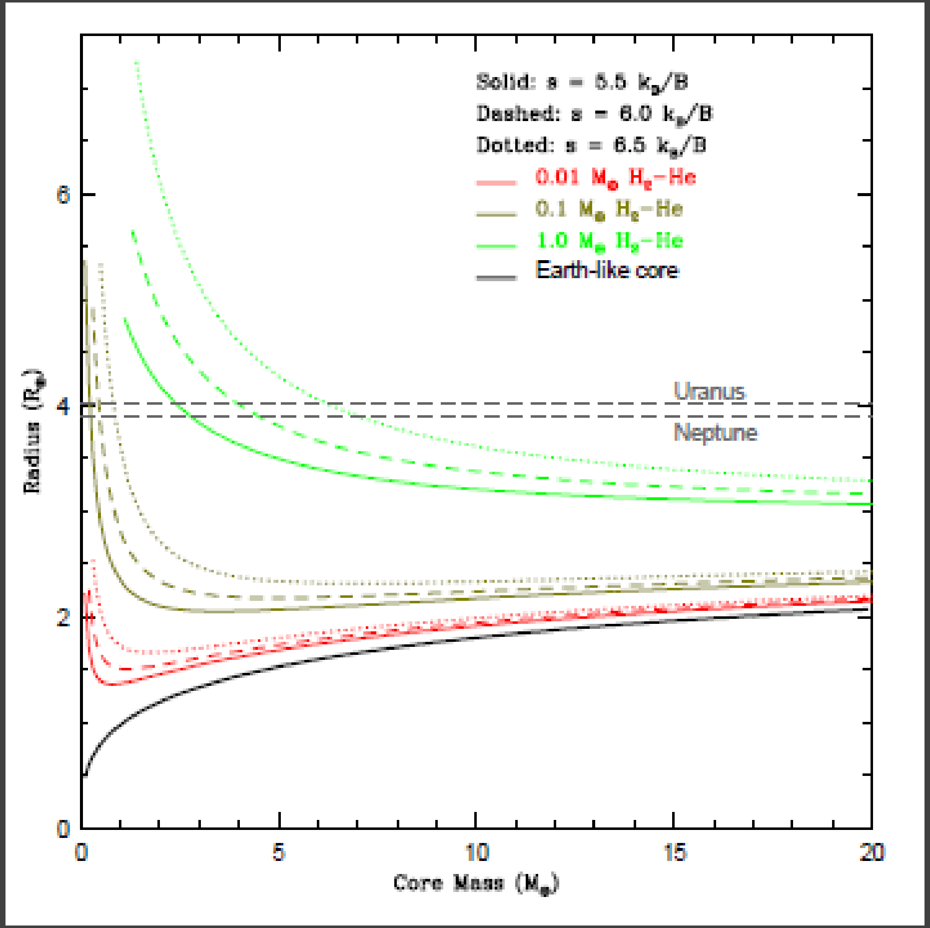
Superearths are very numerous planets.
Only those with well-determined mass and radius are shown.

Inner cores can consist either of rocks (and iron) or of ices.

Some of superearths obviously have thick gas envelopes.
This is a challenge to formation models.



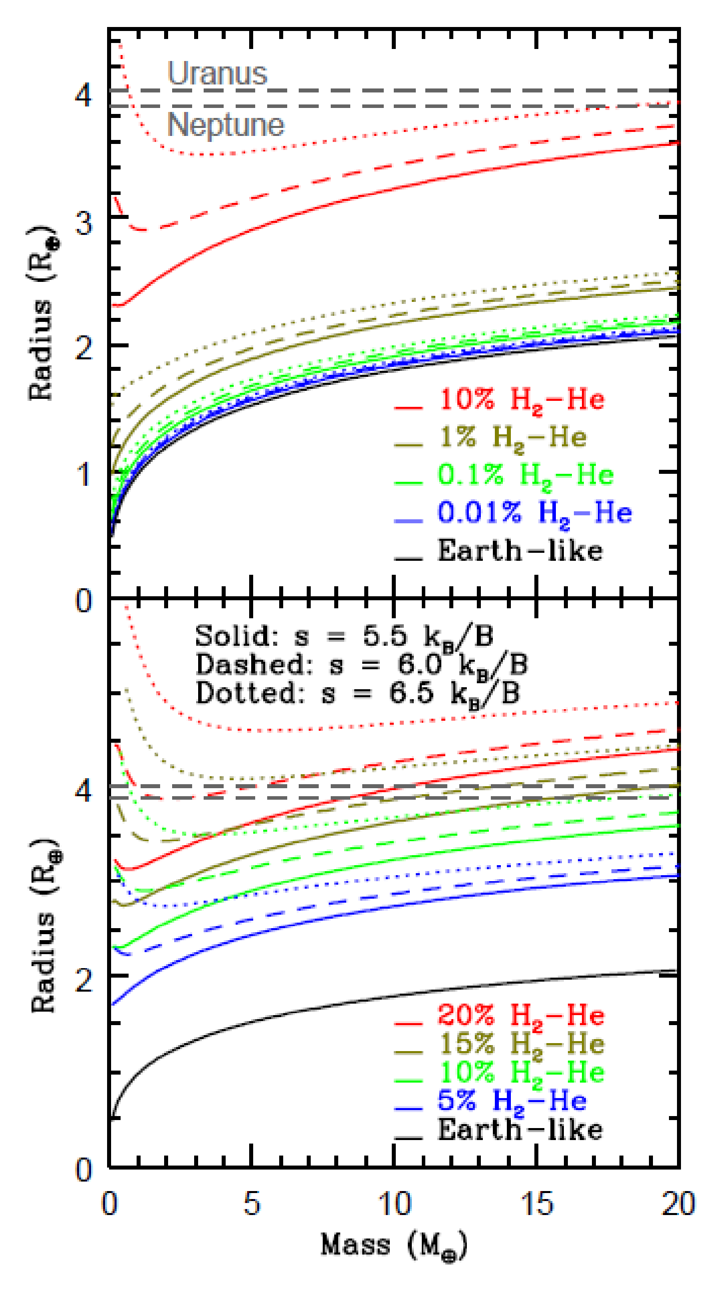
Superearths models



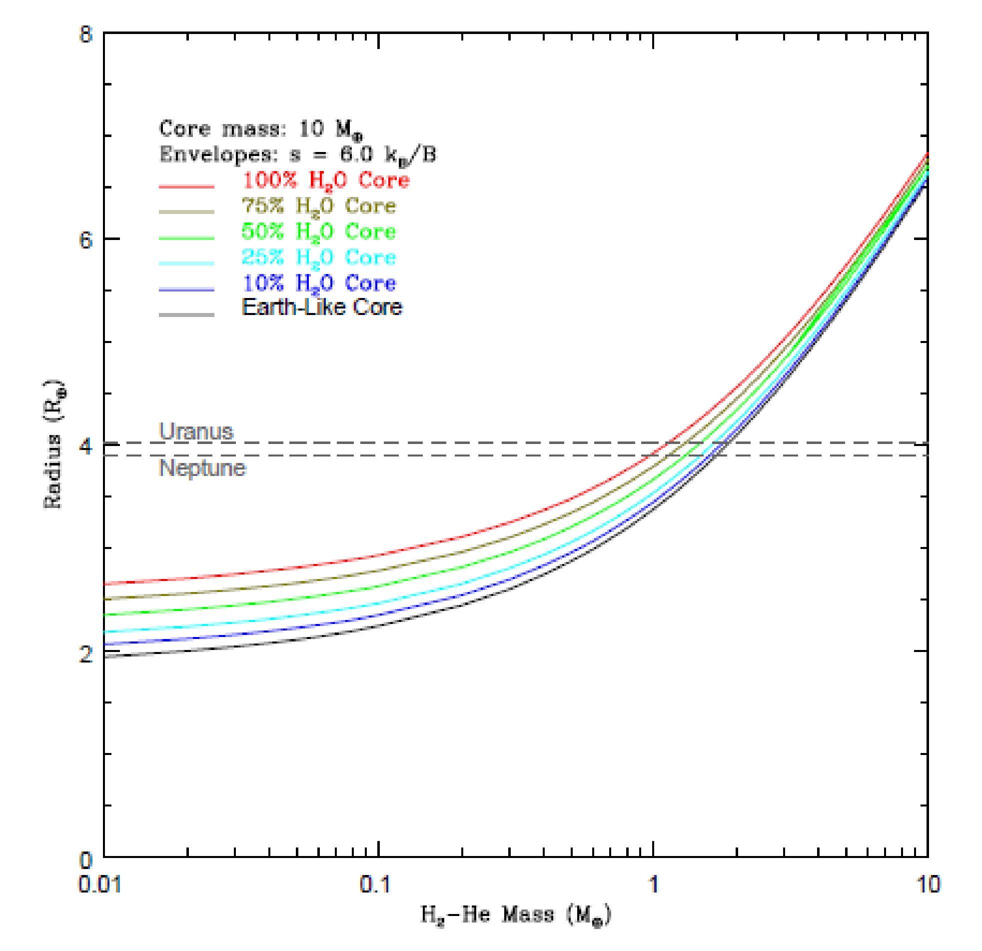
For less massive planets parameters are mainly determined by the core.
For more massive – by the outer envelope.

Heating can be also important.

Results are shown for planets with solid earth-type cores.



Just add water



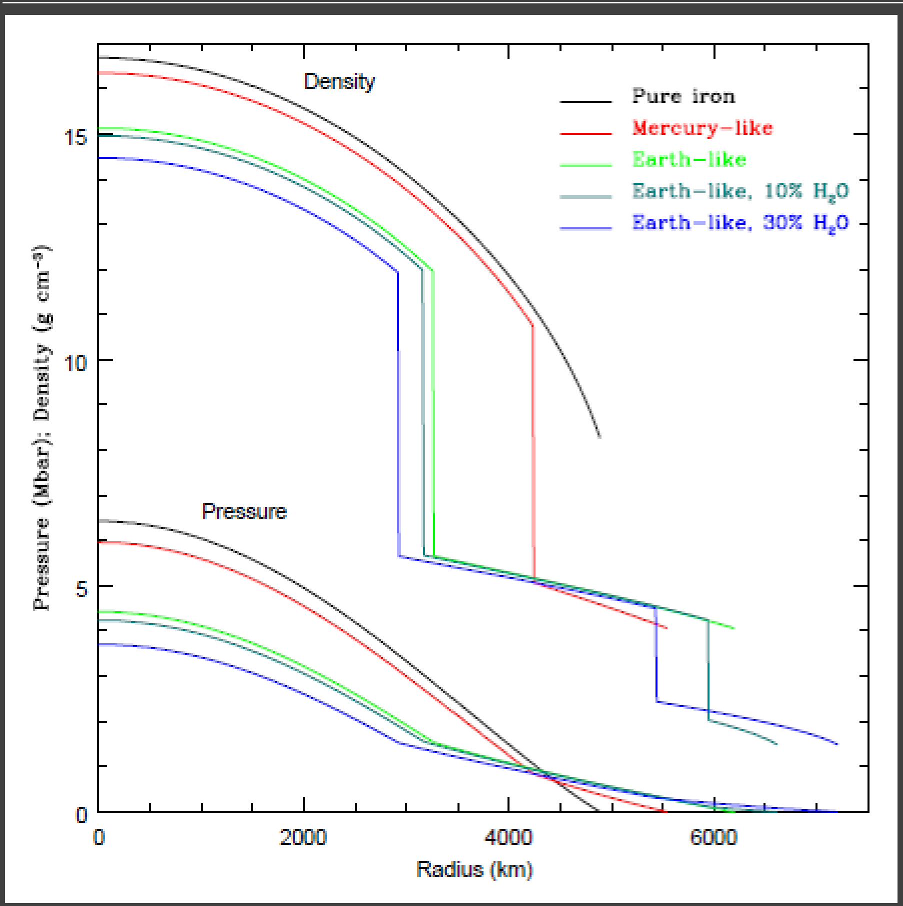
Let us fix the planet mass and change the fraction of ice.

Here water is added as an ice layer above a solid (rocky) core.

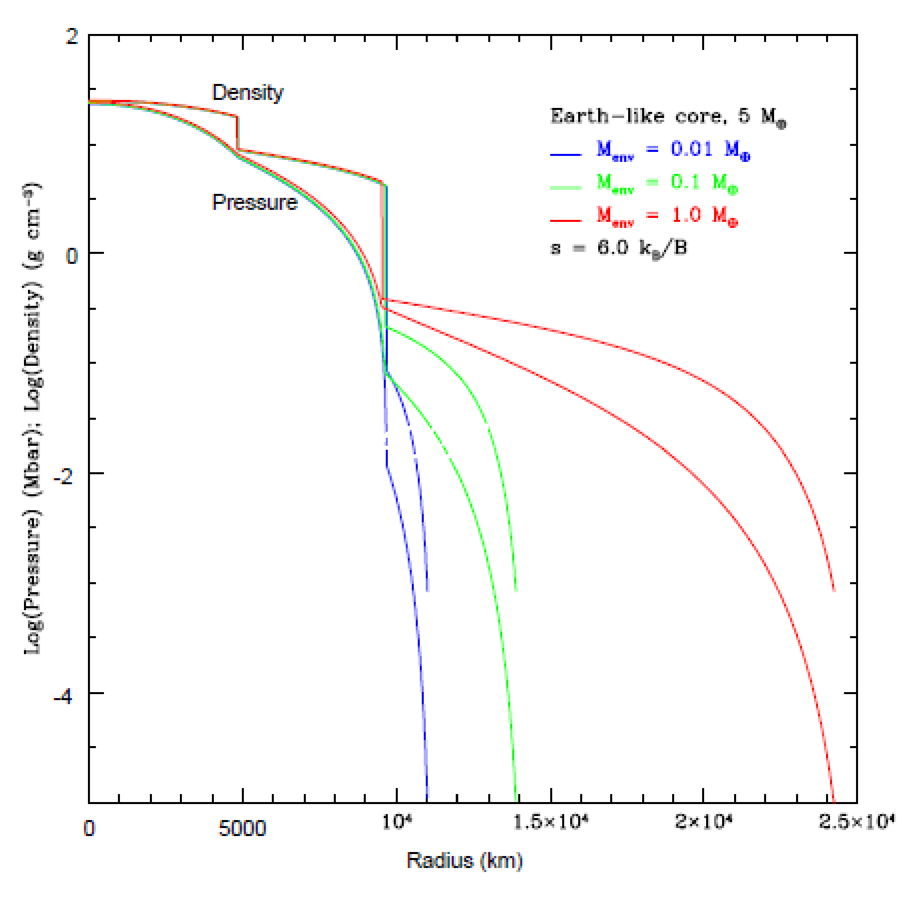
Only for lower masses it is possible to distinguish (by radius measurements) between pure-ice cores and pure-rock cores.

1402.4818

Internal structure

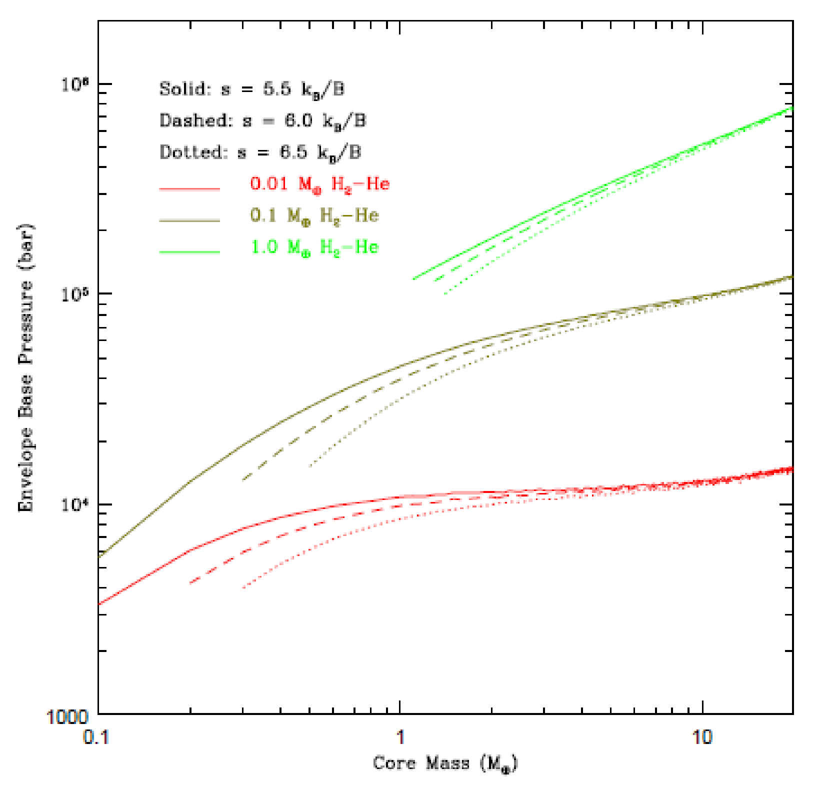


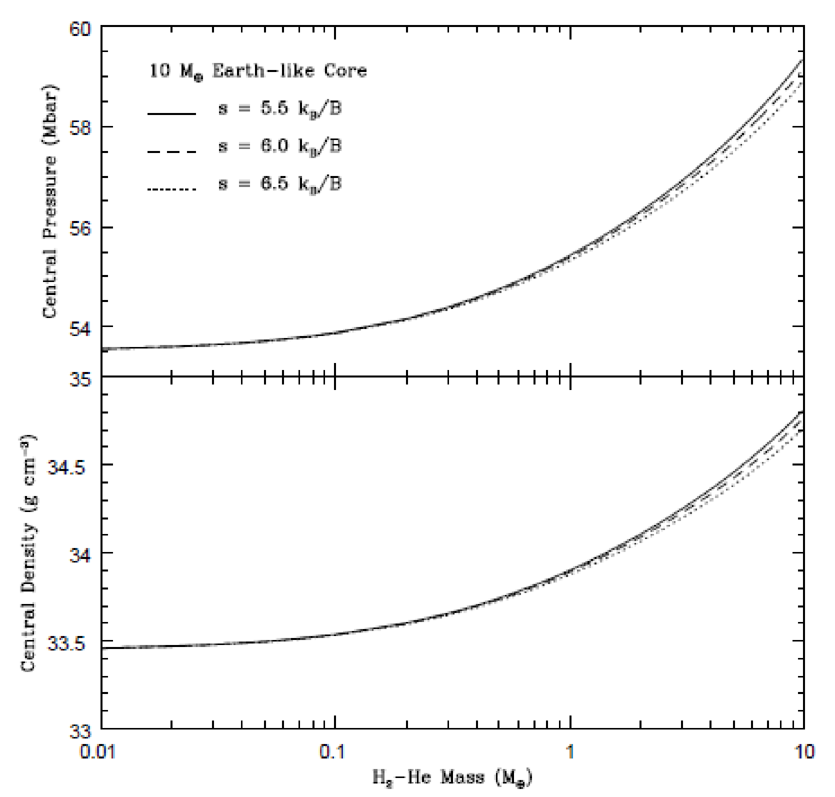
Without an envelope

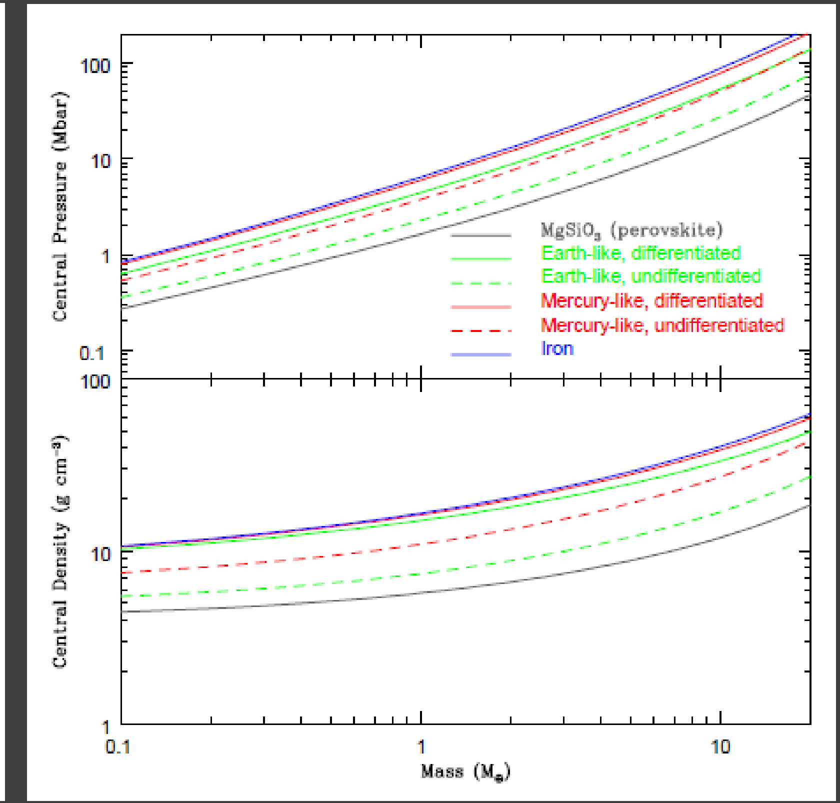


With an envelope

Under pressure

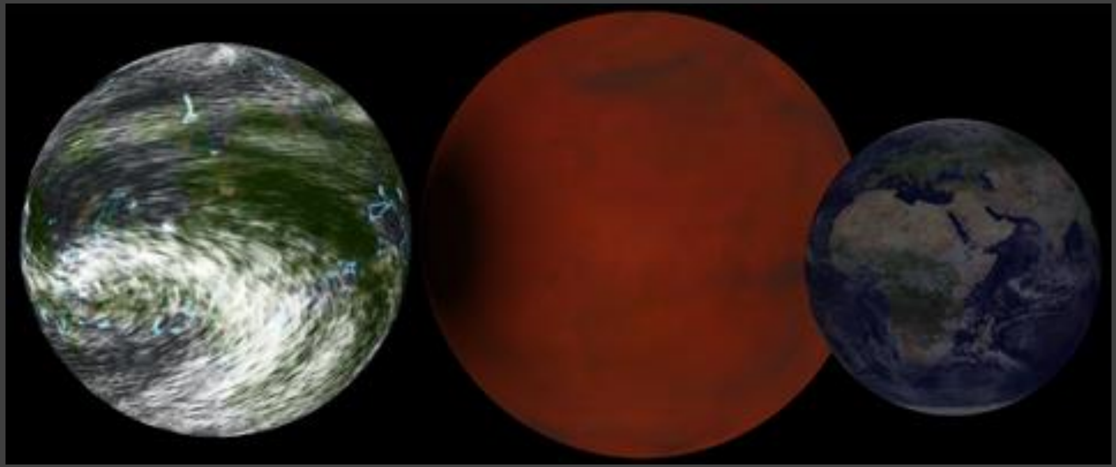




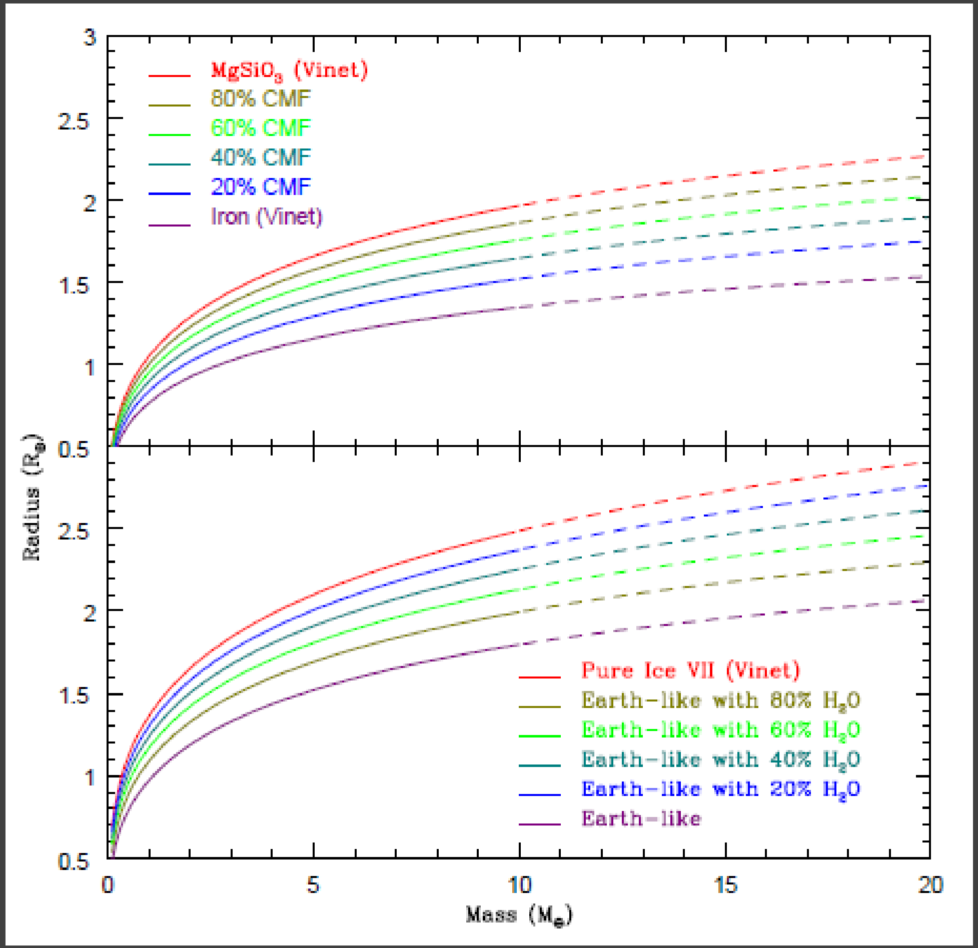


Interiors might have high pressure and density

Soil and water

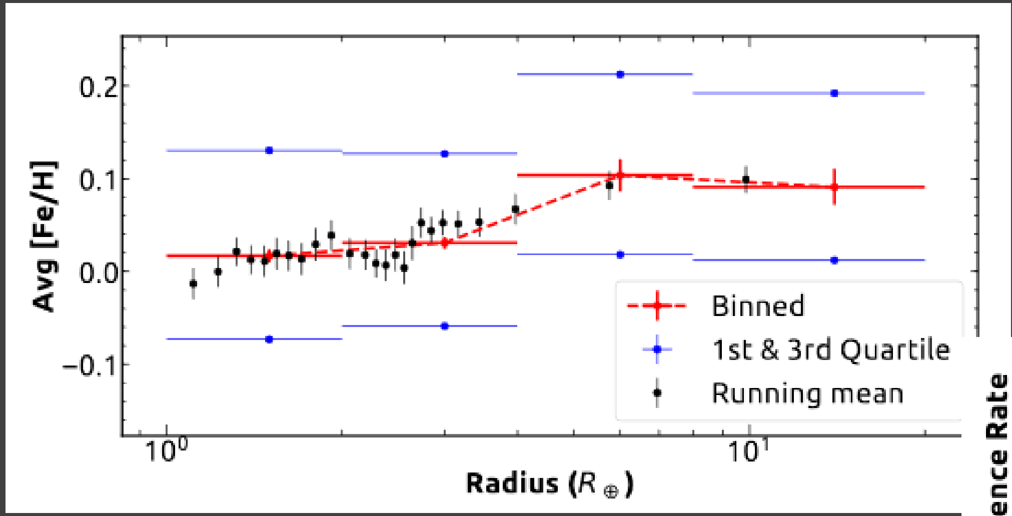


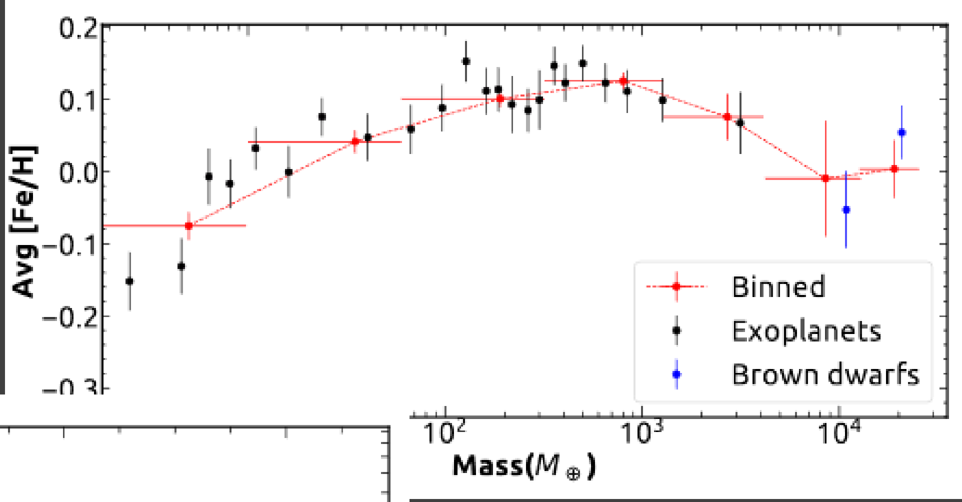
Radius vs. mass for different water content

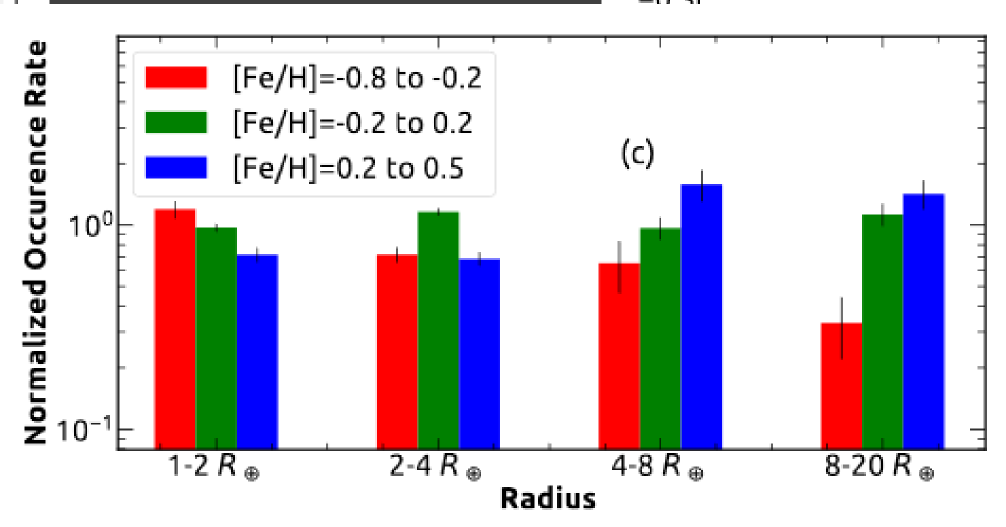


Stellar metallicity and planets

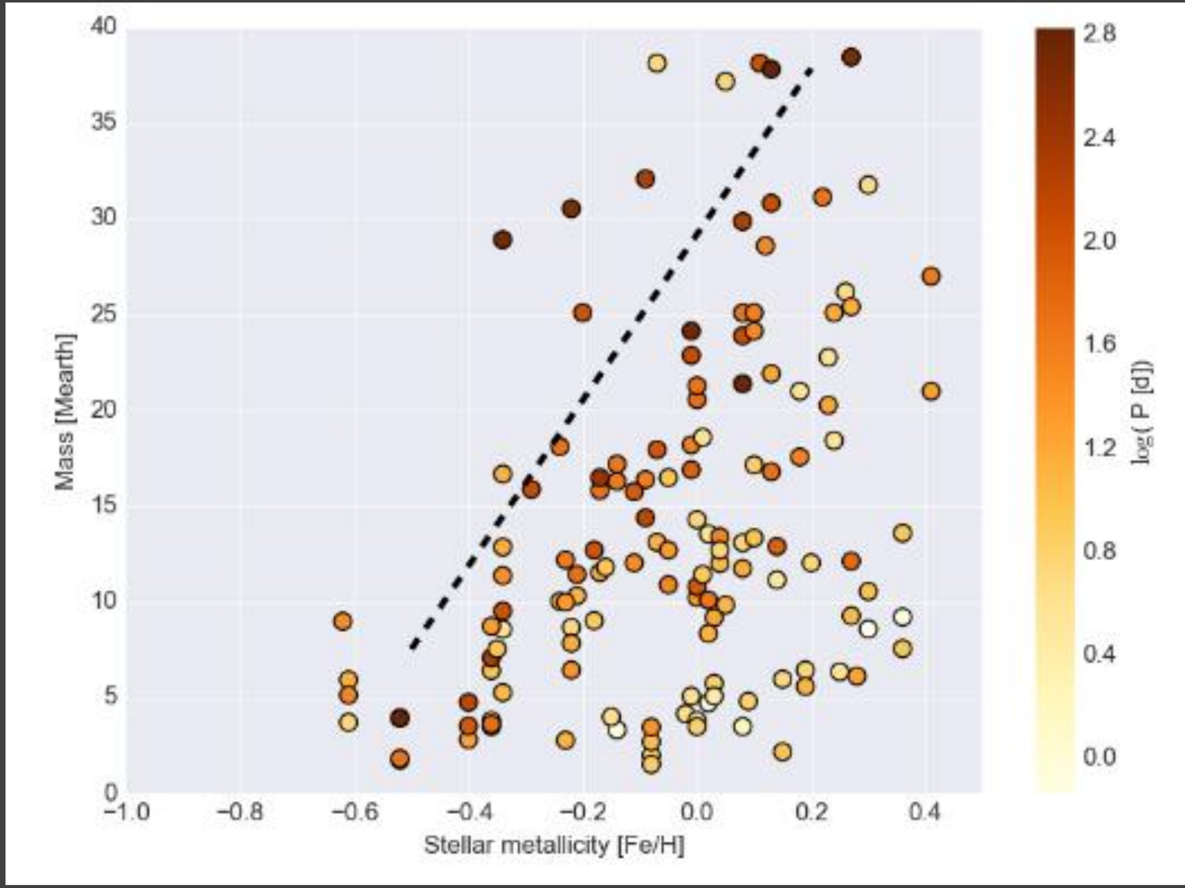
Parameters of planets strongly correlates with stellar metallicity

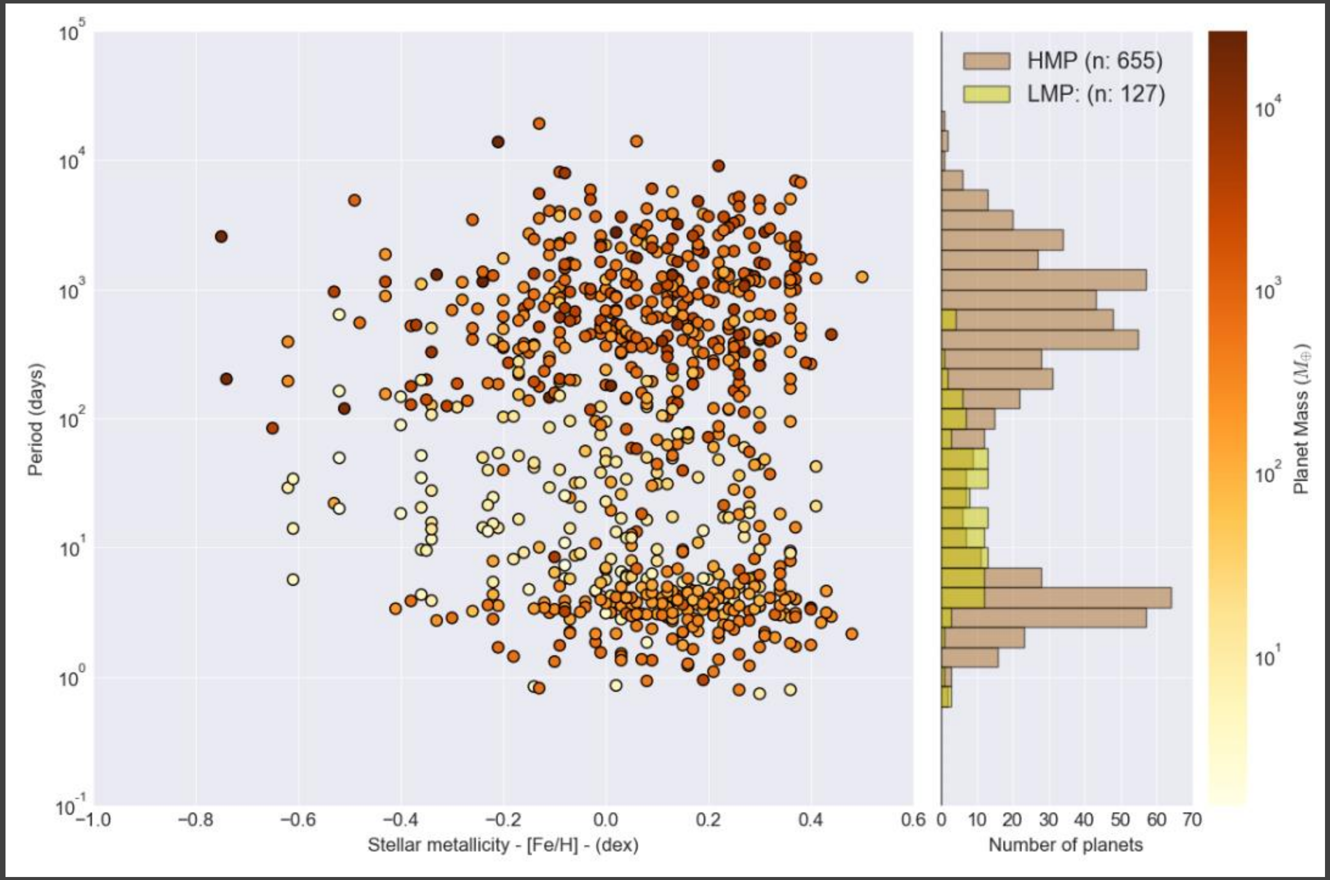






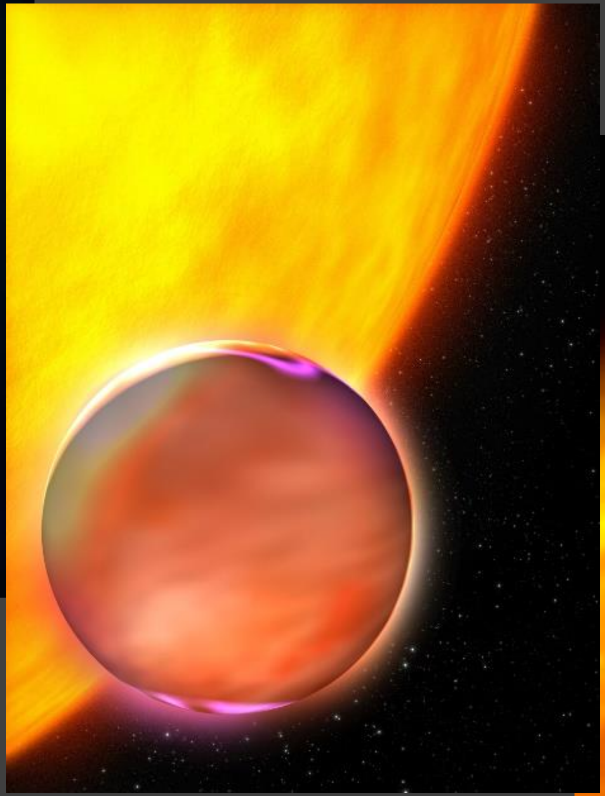
No massive planets around low-metallicity stars

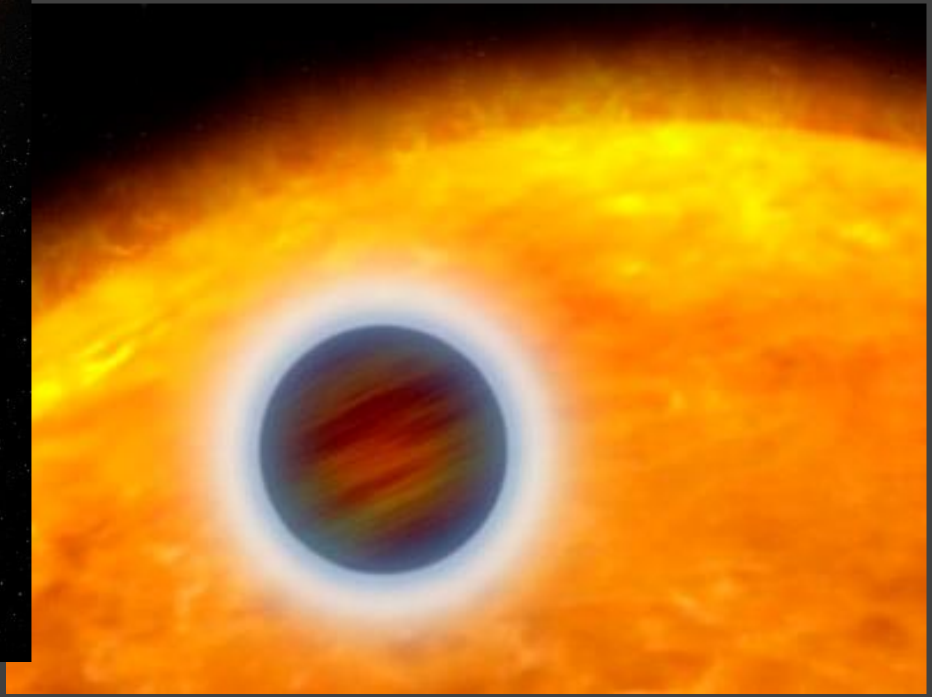




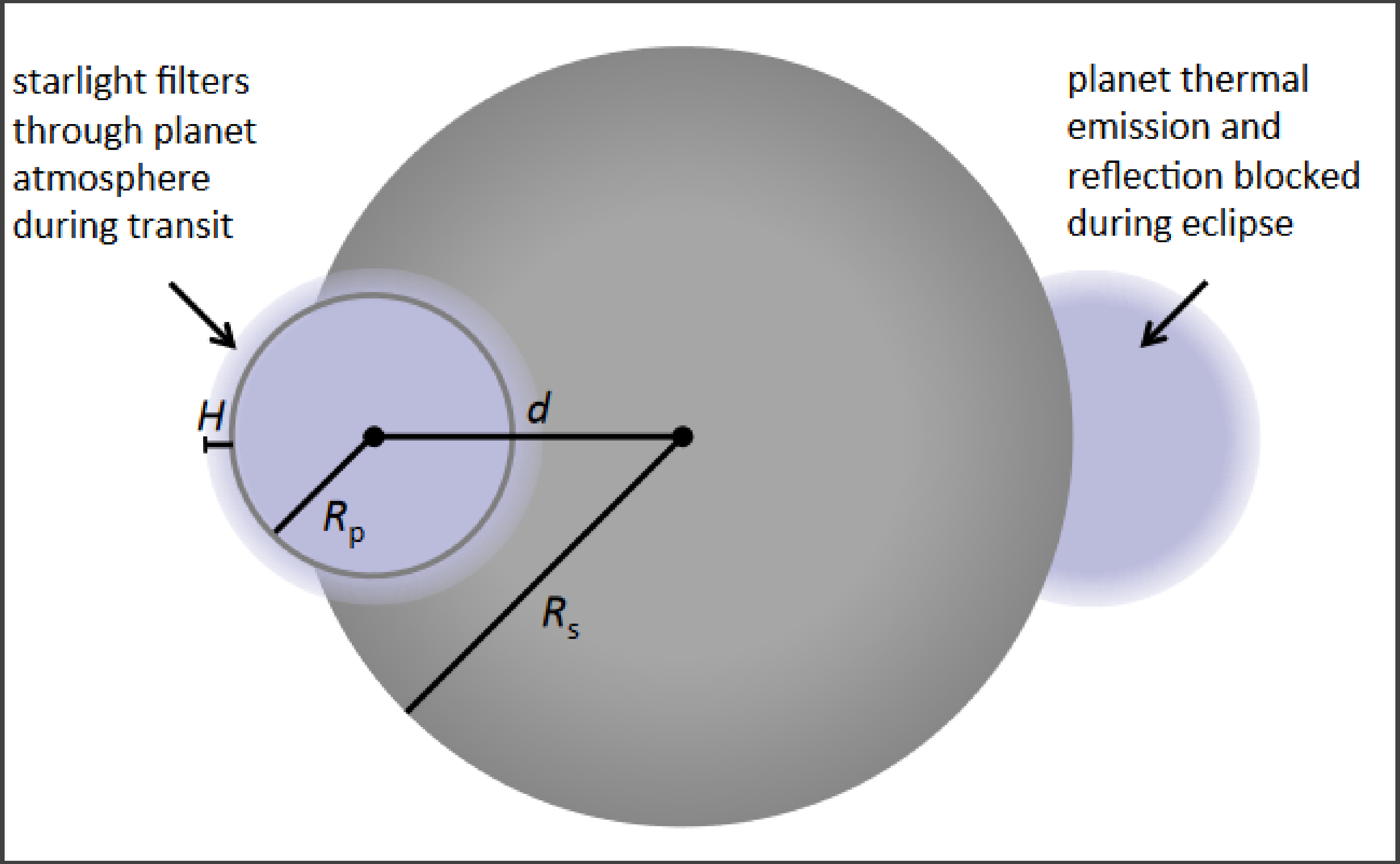
1903.04937

Atmospheres

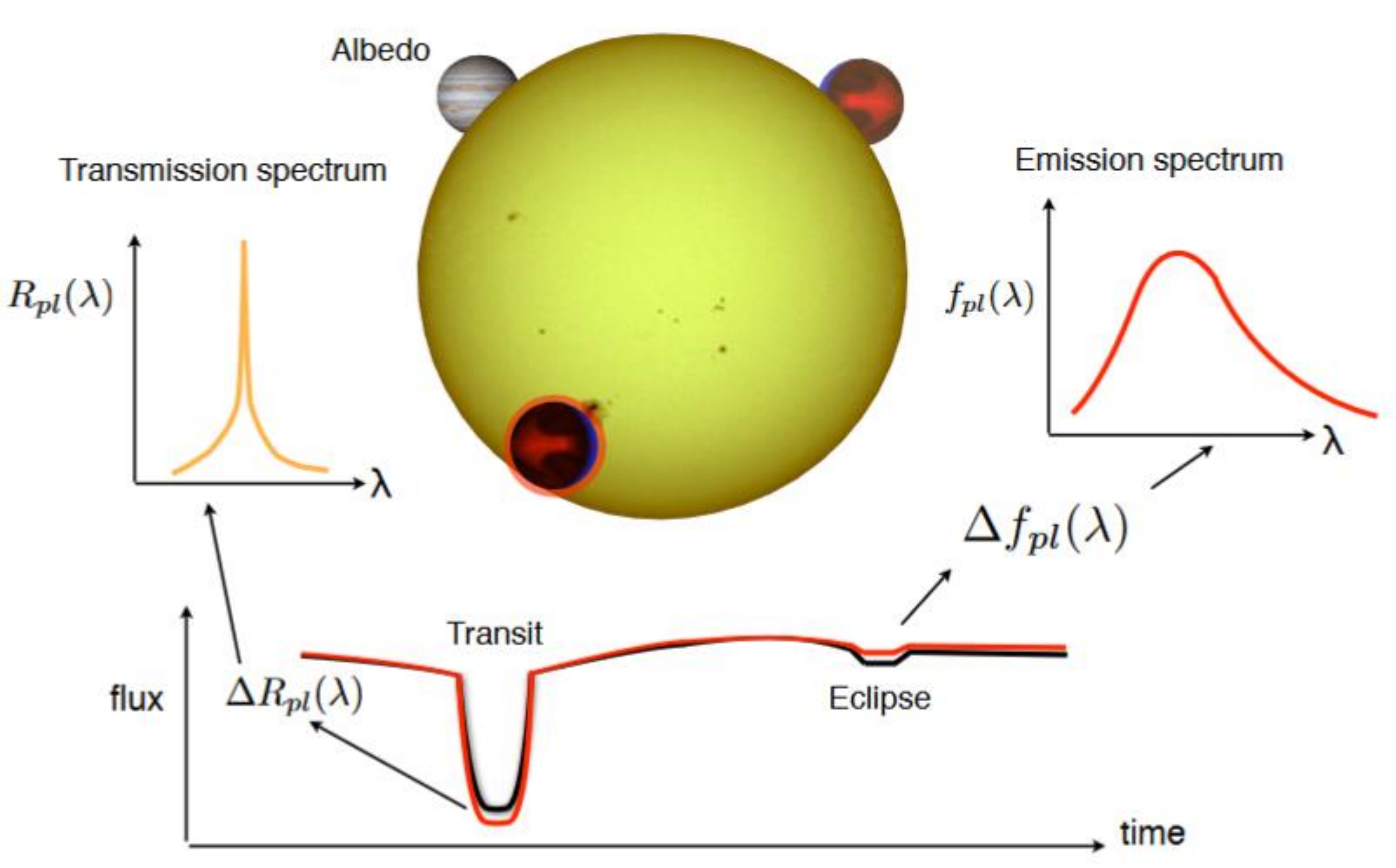




Transits and atmosphere studies



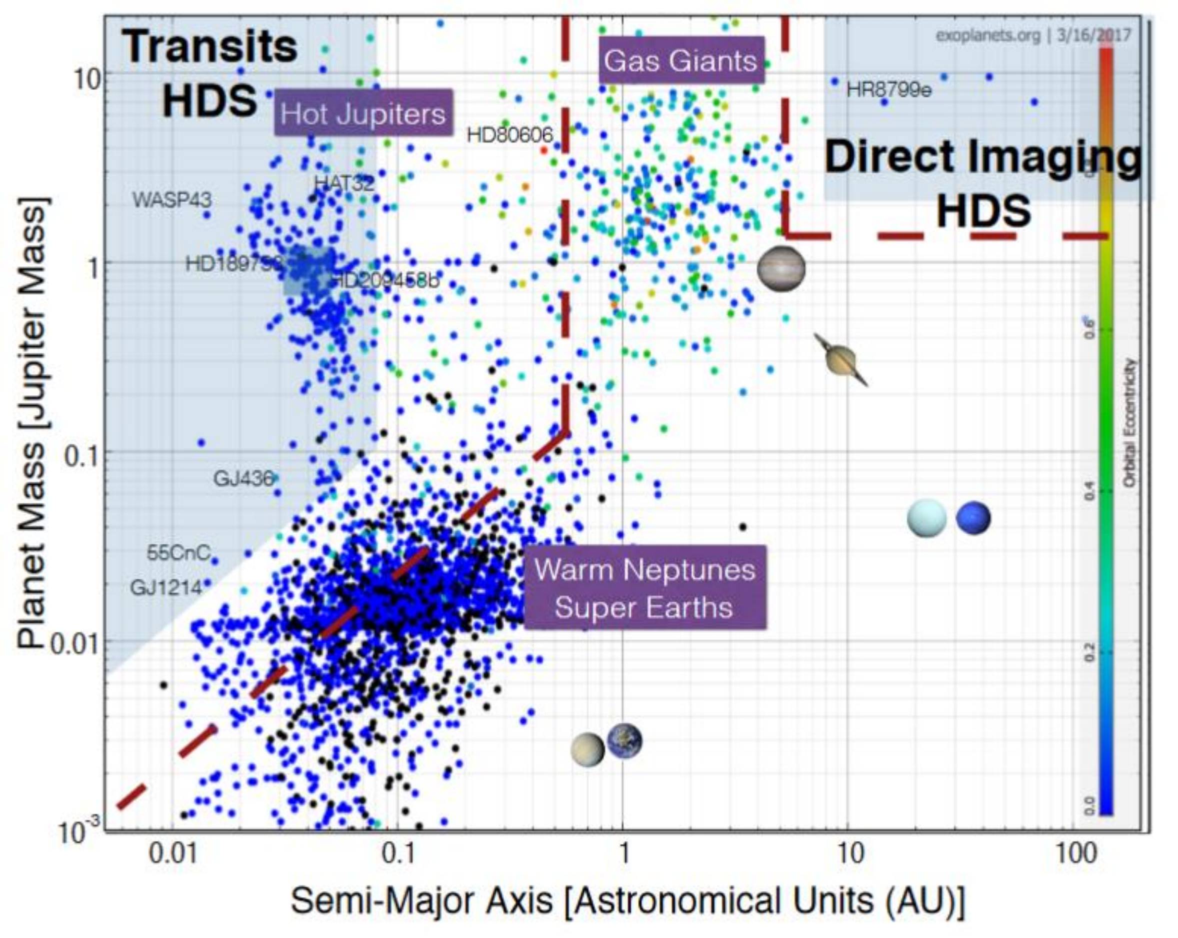
Planet studies during transits



- Integrated properties of the surface (albedo)
- Transmission spectrum
- Emission spectrum
- Mapping

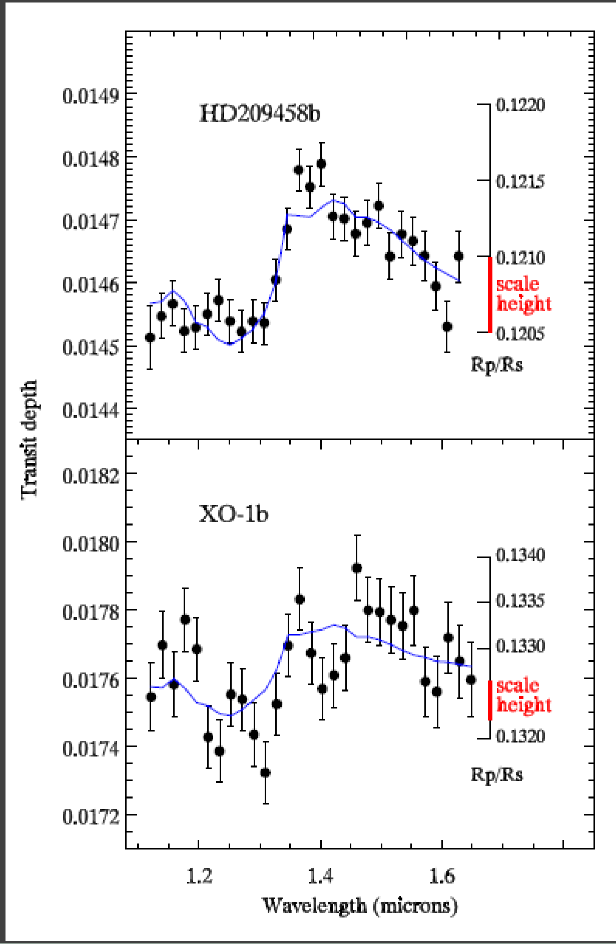
See a review in 1810.04175

Sensitivity of the method



It is easier to detect the signal from planets around M-dwarfs due to a smaller stellar radius.

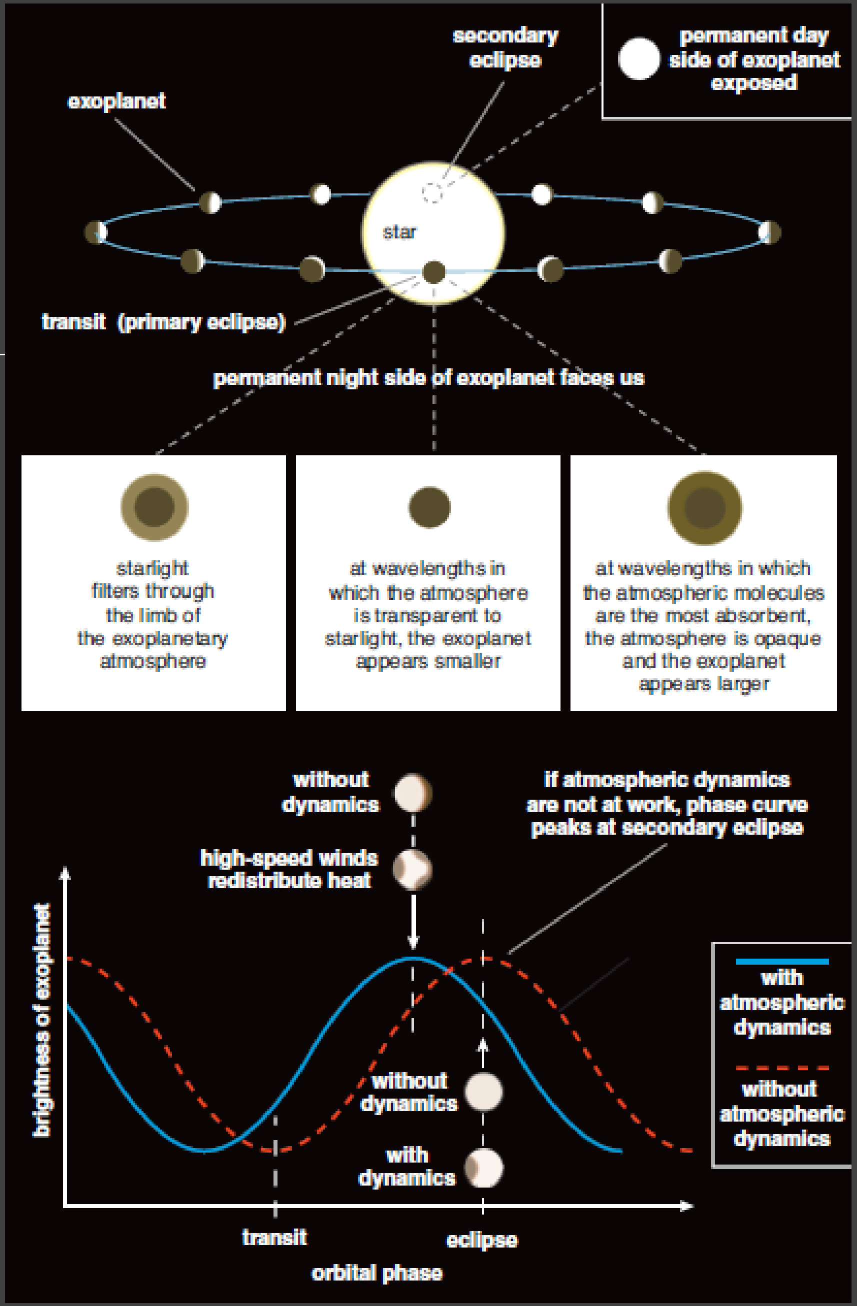
Transits and atmospheres



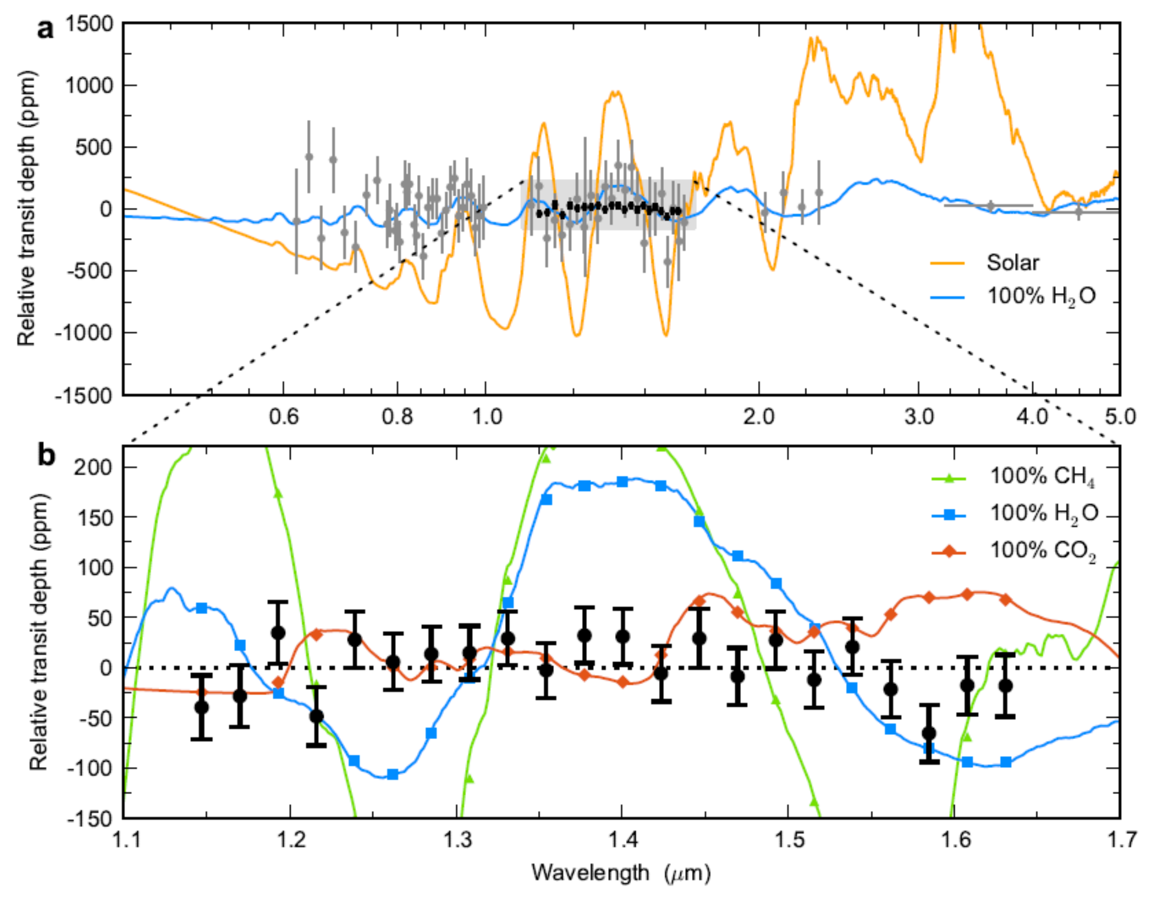
Transit observations in different wavelengths allow to determine properties of the planet atmosphere.

Size can be different in different wavelengths.

In addition, light curve can look different due to atmospheric dynamics. Heat redistribution due to strong winds modifies the flux from the planet.



Featureless spectrum of GJ 1214b



Obscured by clouds.

Hubble space telescope spectrum shows no details.

This is interpreted as the result of the presence of a thick cloud layer in the outer atmosphere of the planet.

Phase dependence

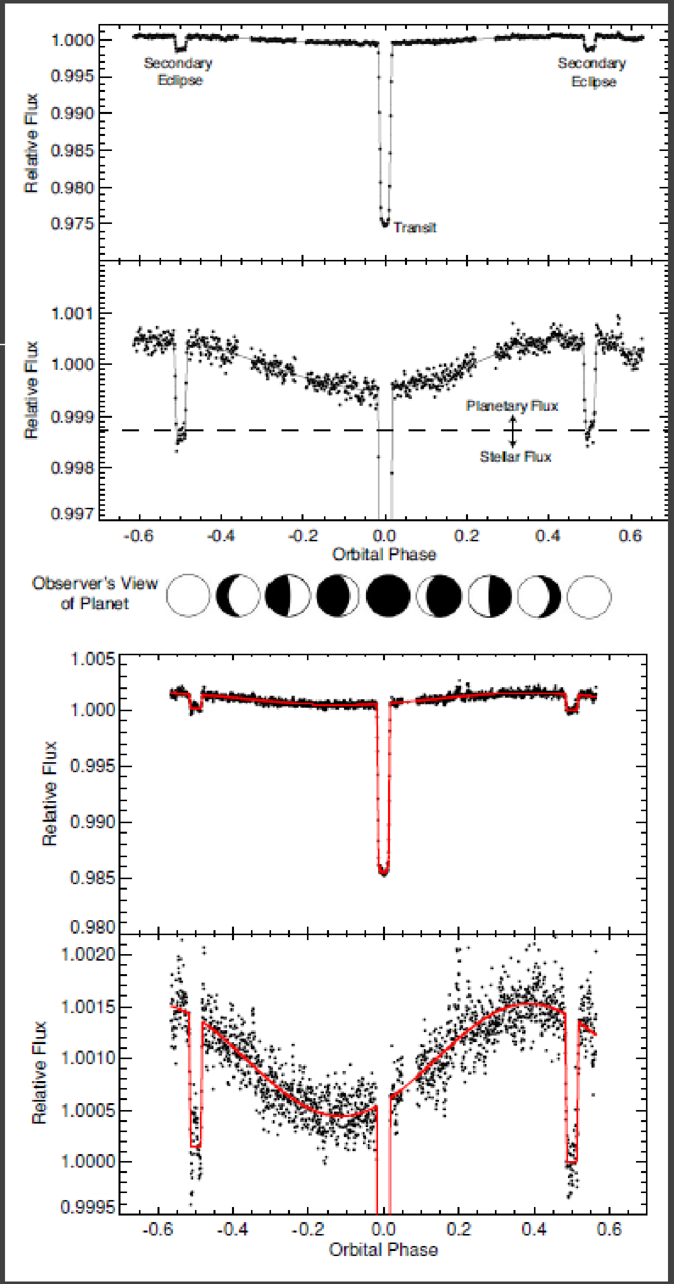
Depending on the phase we observe different parts of a disc.

Results of observations correspond to:

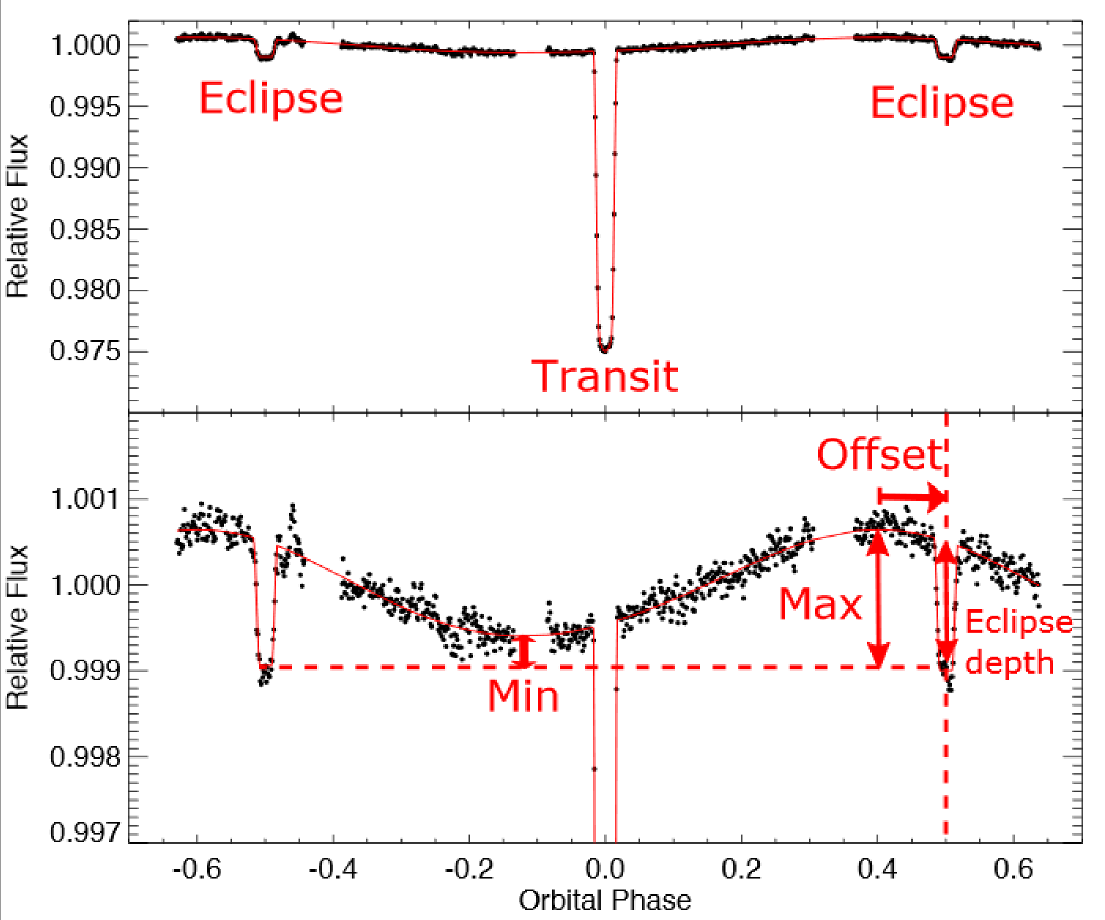
HD 189733b – upper panel;
HD 209458b – lower panel.

Both planets are hot jupiters.

Note, that in the case of HD 209458b planetary disc
is strongly non-symmetric in terms of the emitted flux.

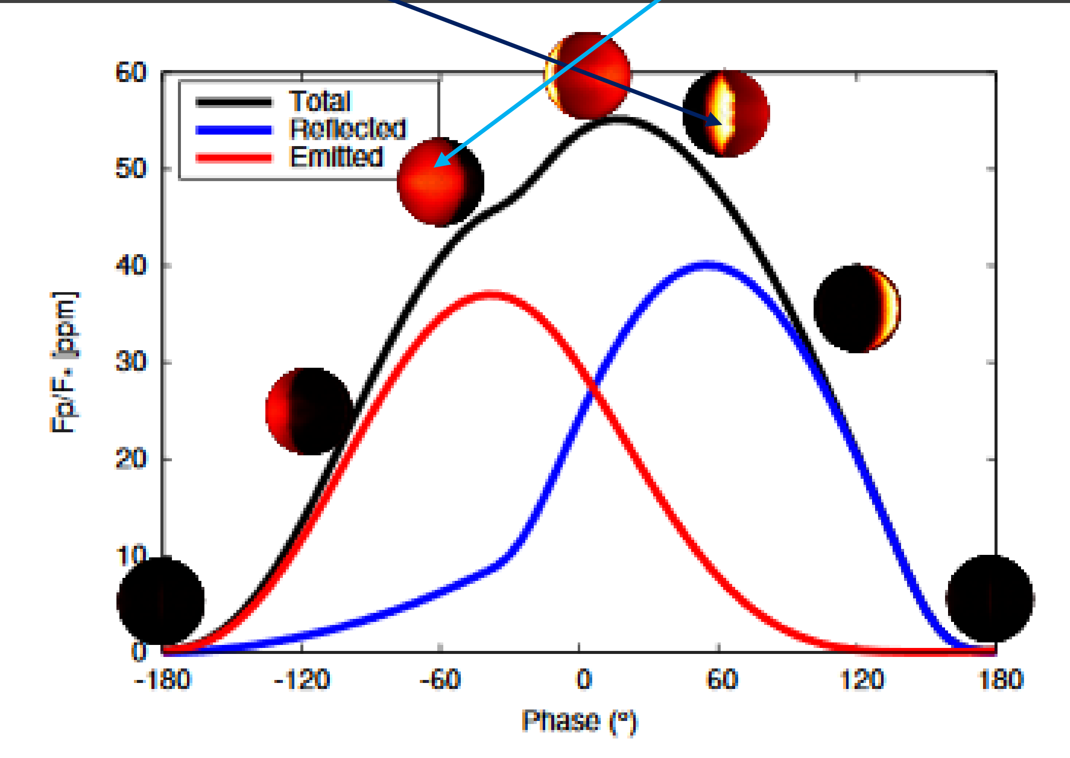


Phase light curves

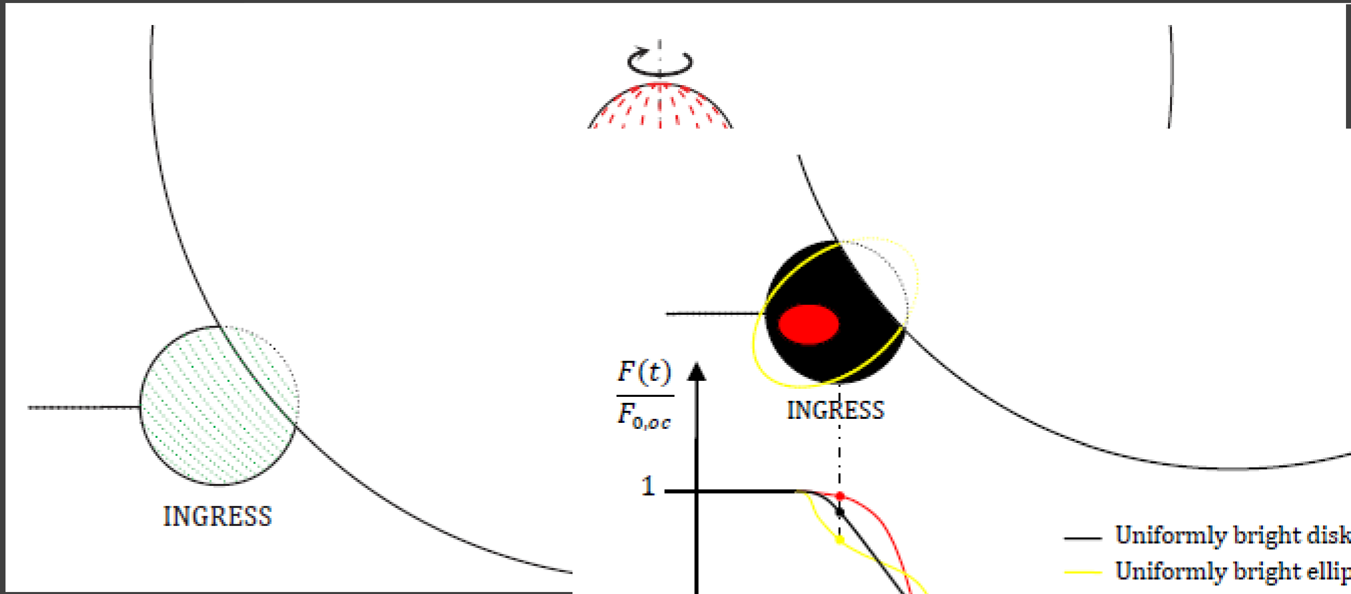


Bright cloud reflect light

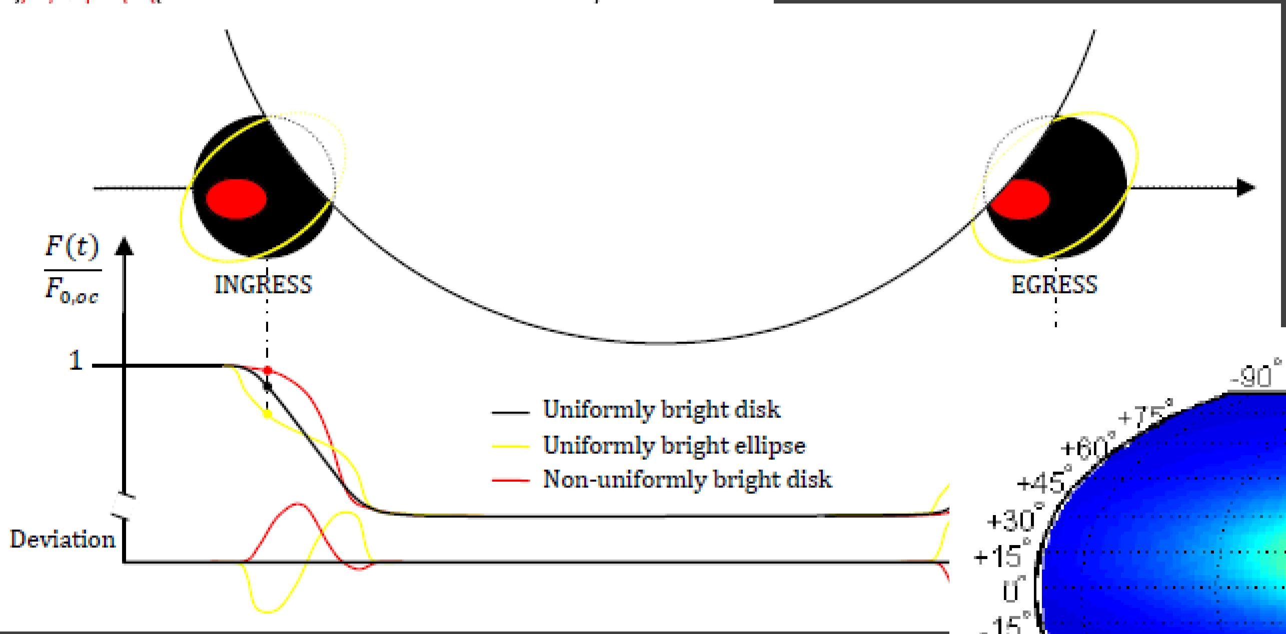
Global winds produce non-symmetric temperature distribution



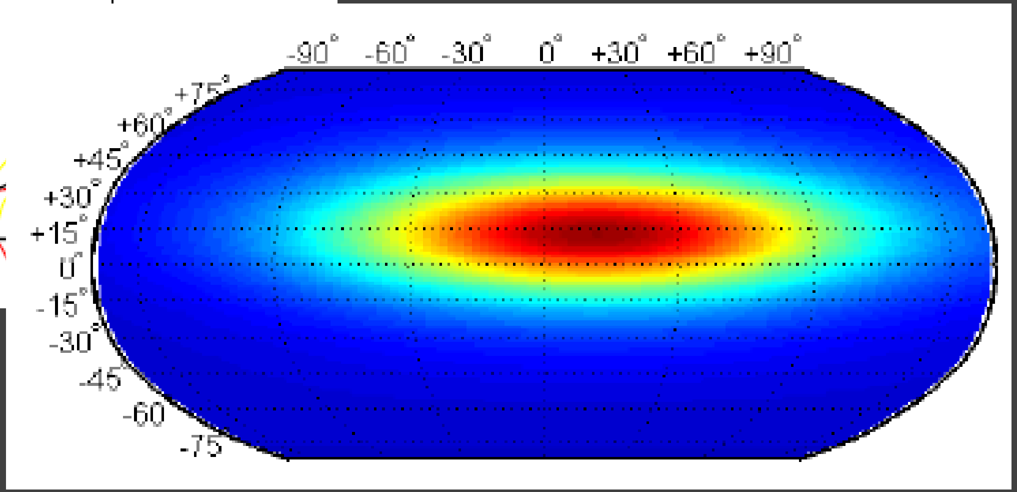
Scanning planetary discs



HD189733b

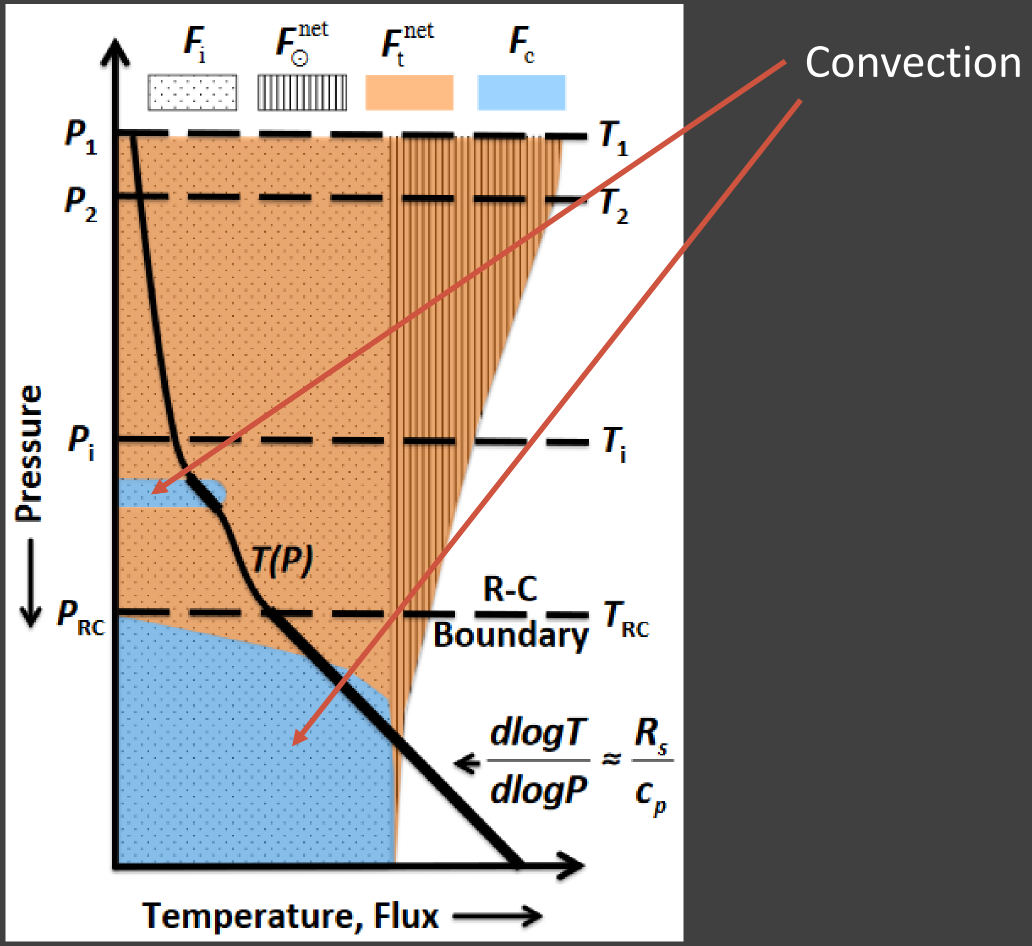


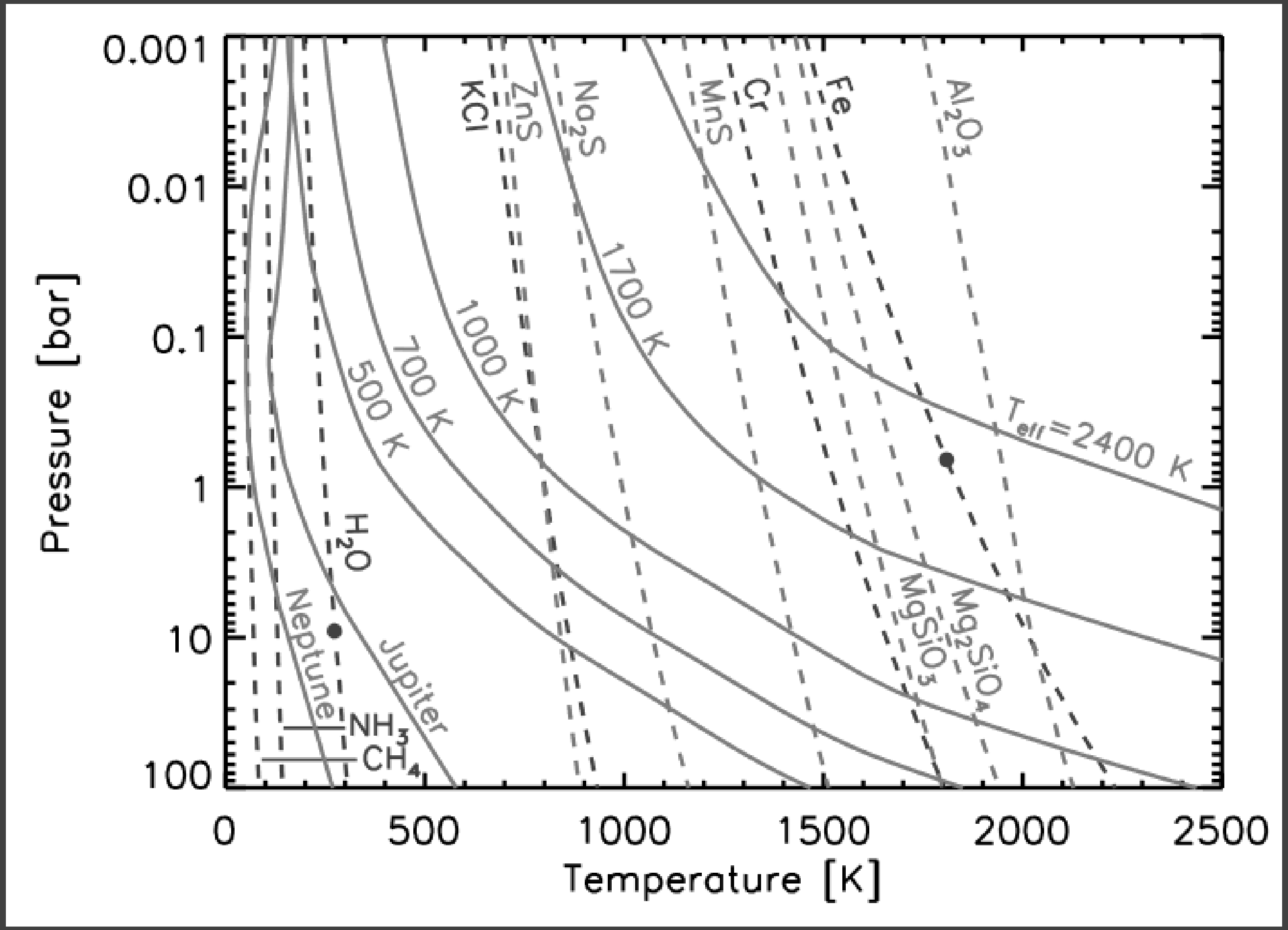
Spitzer space telescope



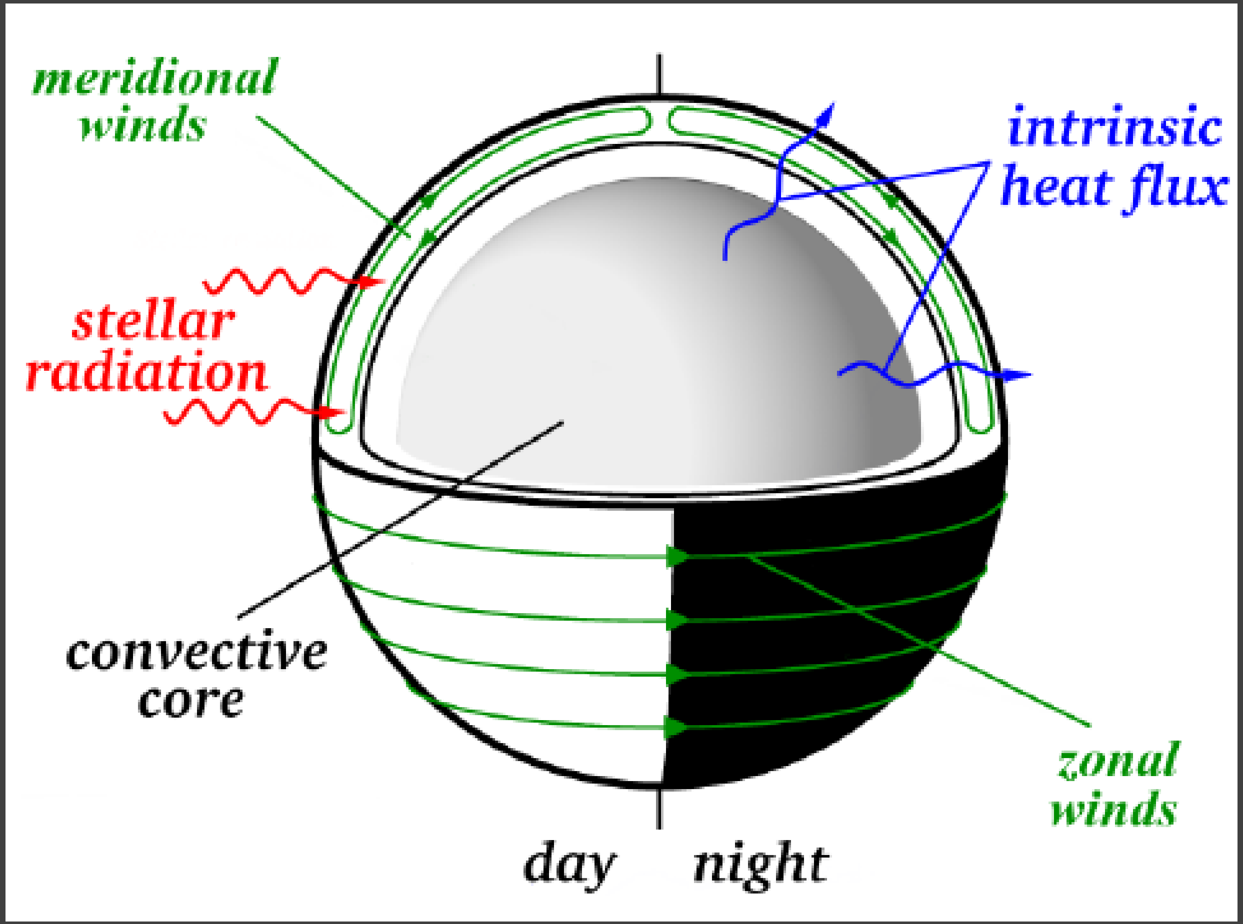
1202.3829

Modeling of planets atmospheres





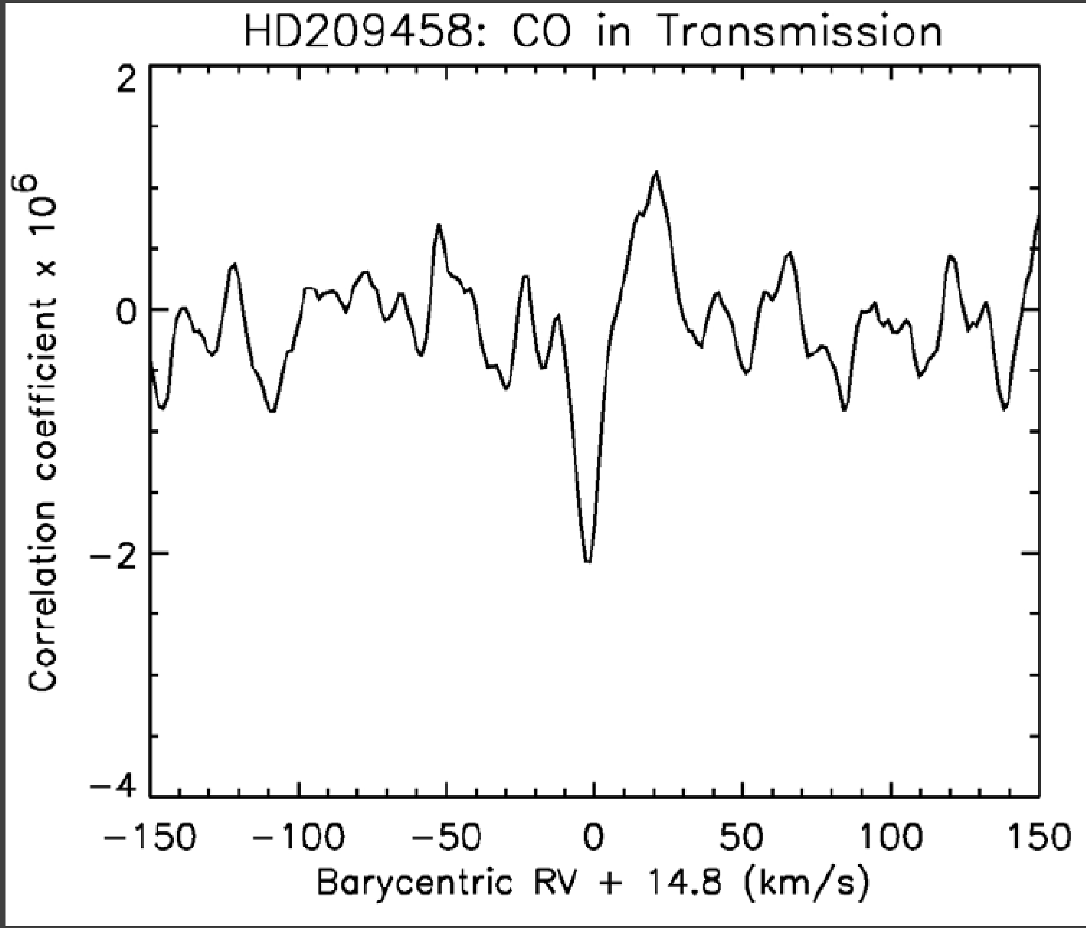
Dynamics of outer layers of hot jupiters



Planet has internal and external heat sources.

This results in violent winds and convection in the outer gas envelope.

Wind on HD 209458

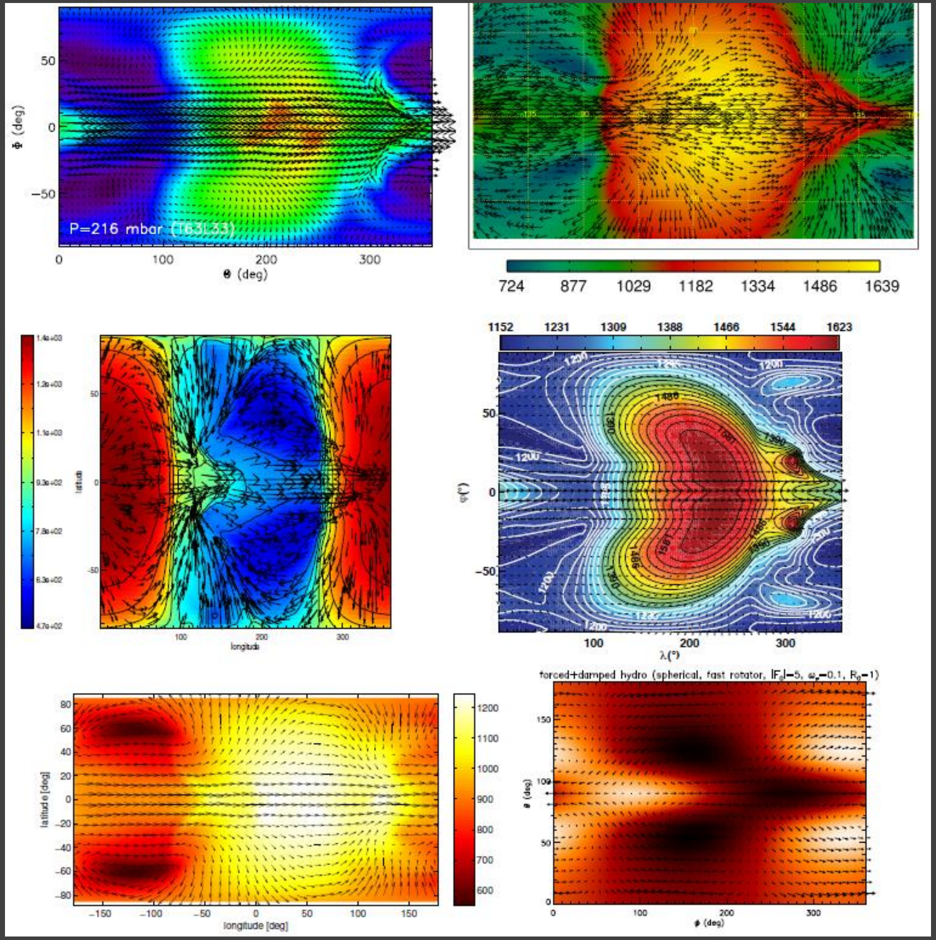


Wind velocity can be directly (!) measured.

The planet is a VERY hot Jupiter.

Wind velocity is ~ 2 km/s
(line is blueshifted by 2 km/s)

Modeling winds on hot jupiters

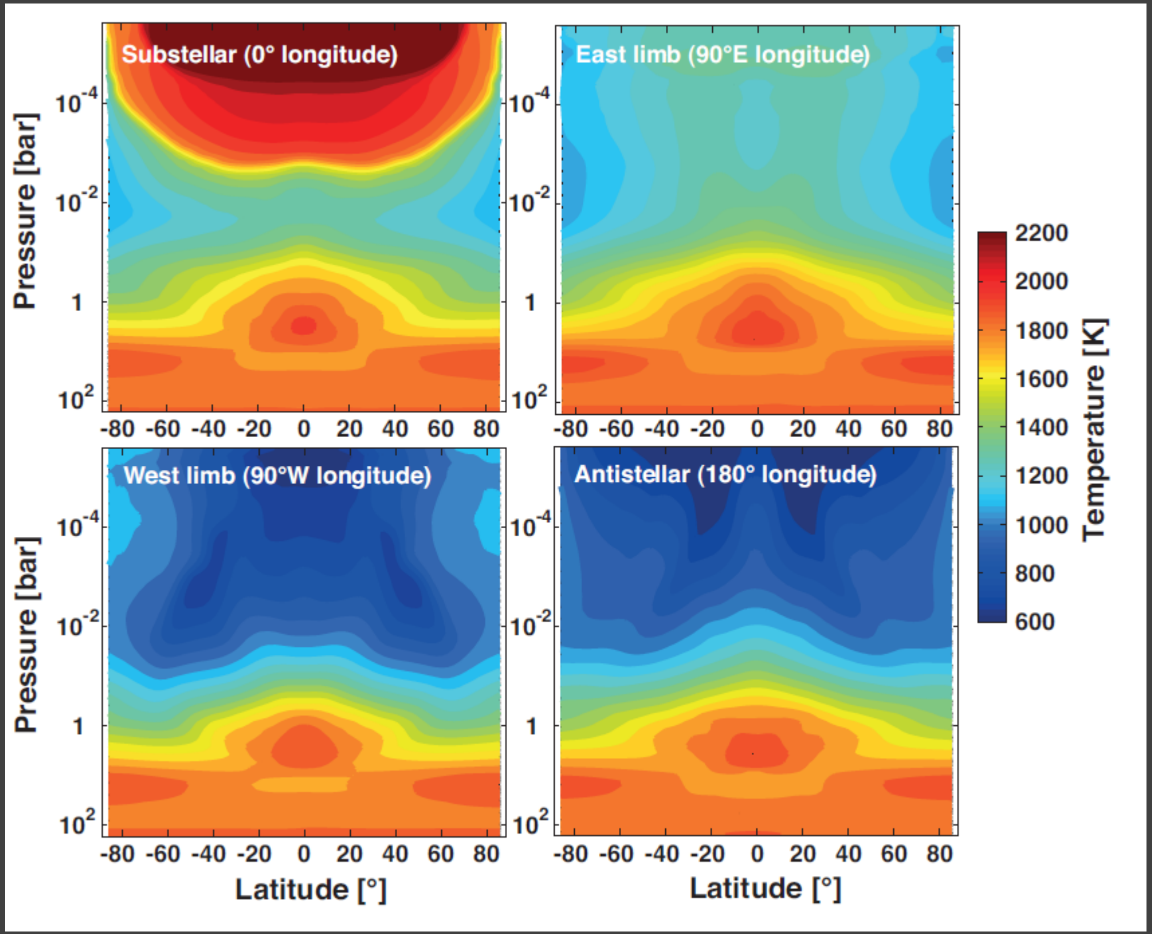


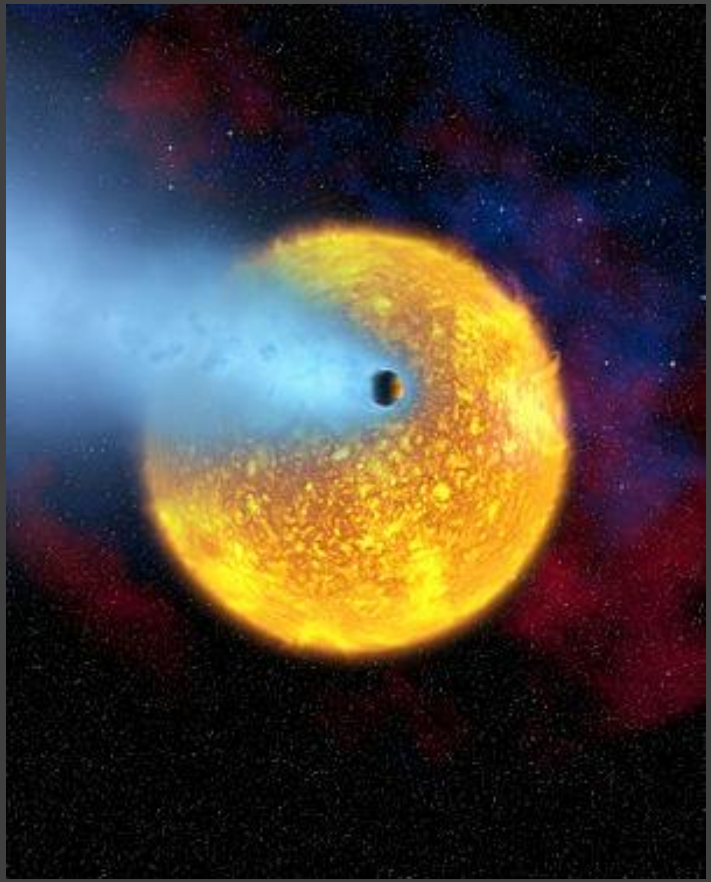
General property:

Strong equatorial wind from the West to the East.

1407.4150

Modeling of HD209458 b

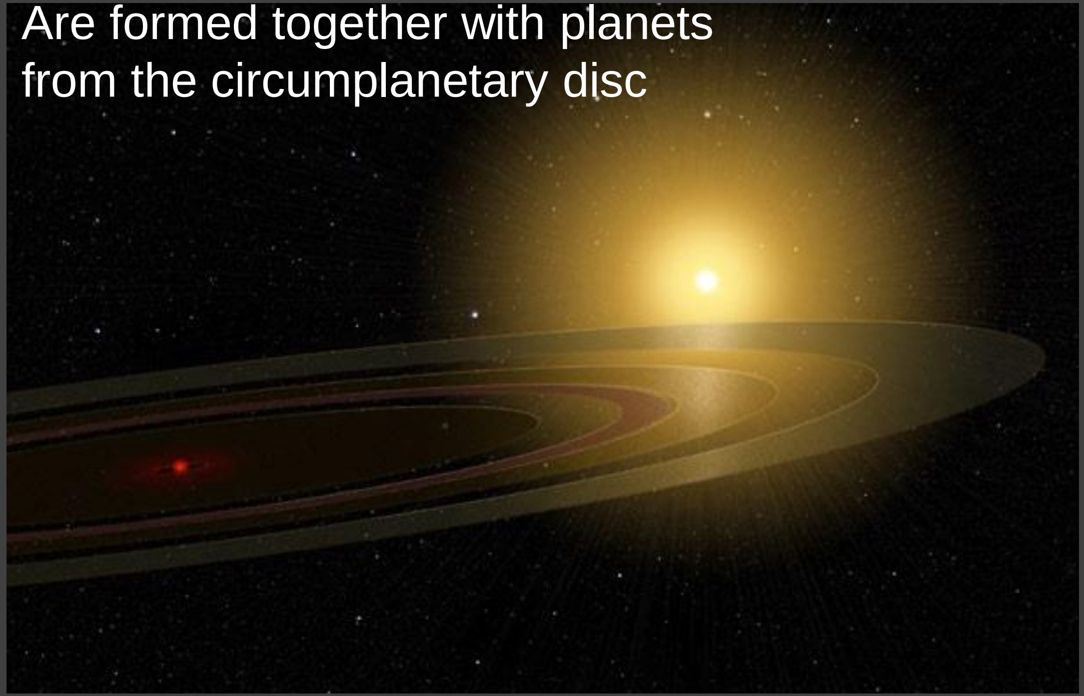




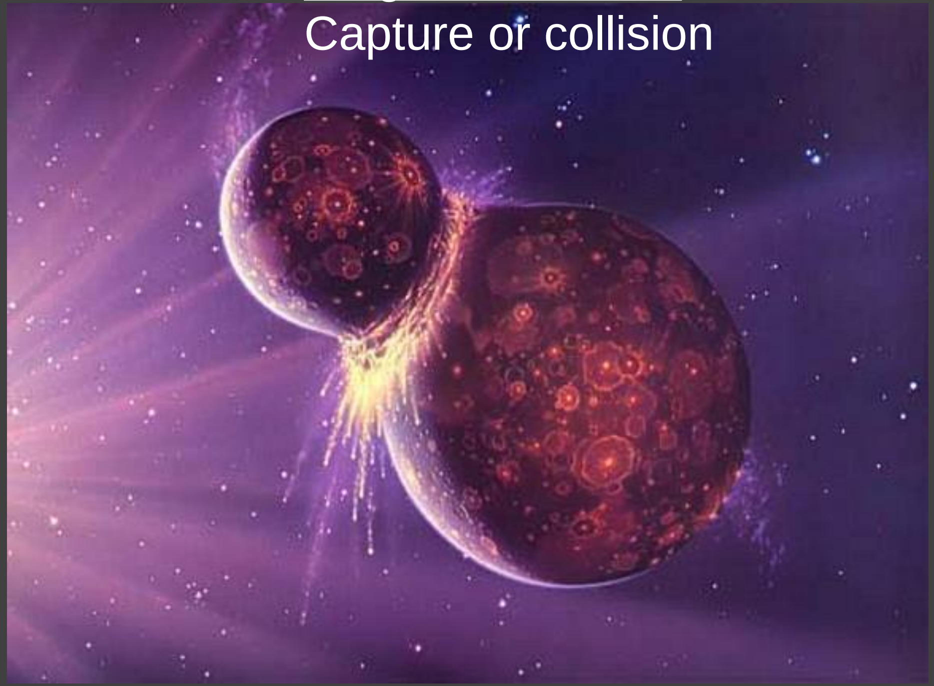
Osiris

Exomoons: how to form

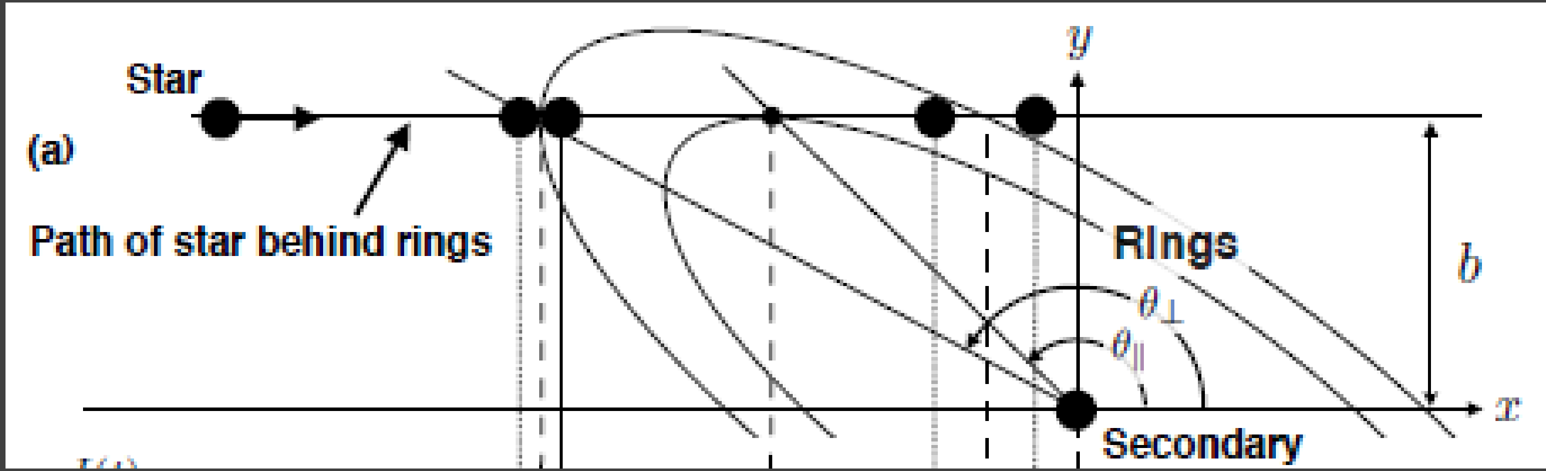
<u>Regular satellites</u>
Are formed together with planets from the circumplanetary disc



<u>Irregular satellites</u>
Capture or collision

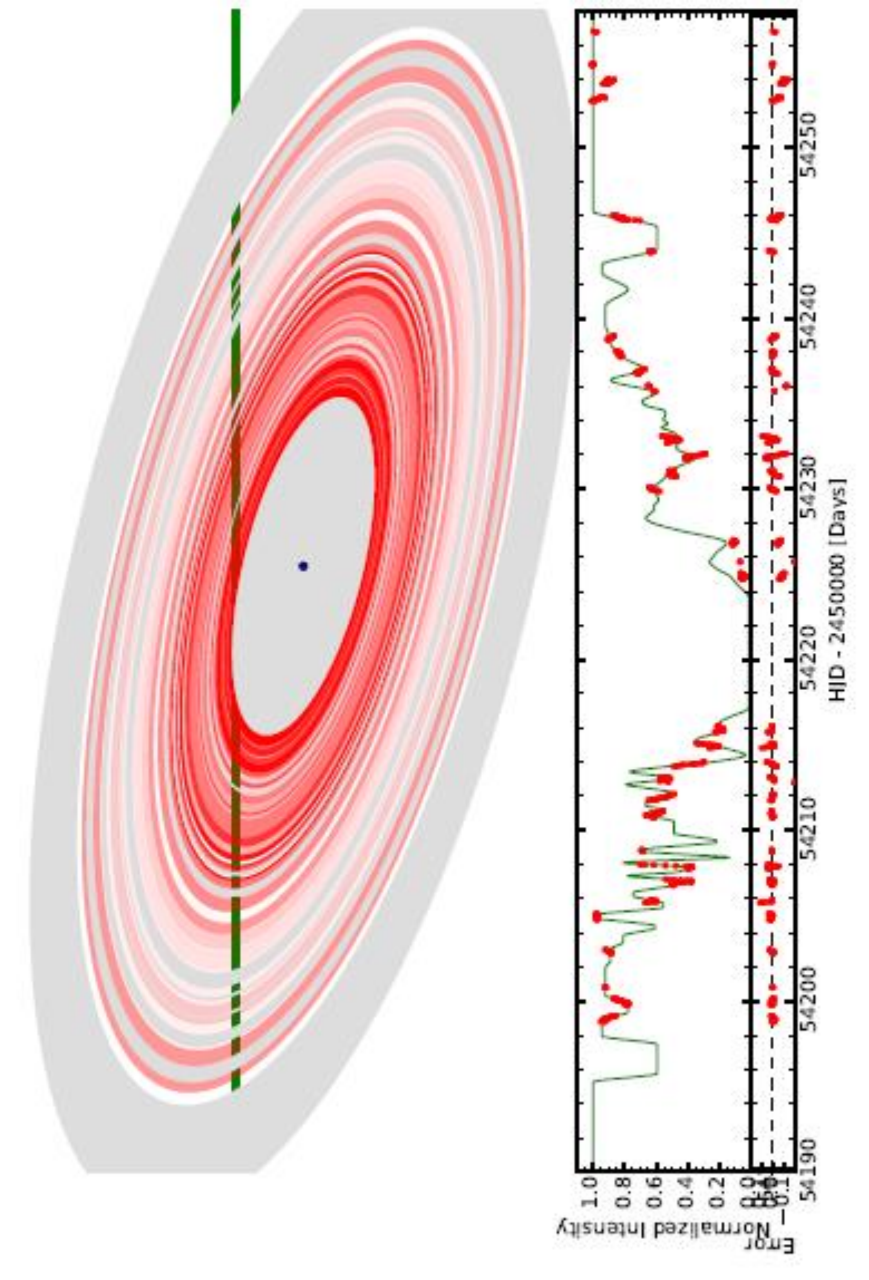


Giant ring system



System of 37 rings extending up to 0.6 AU around a stellar companion.

The star is young (16 Myrs), and so, probably, the system of rings is just forming. Satellites might regulate the shape of the ring system.



Which planets might have detectable satellites?

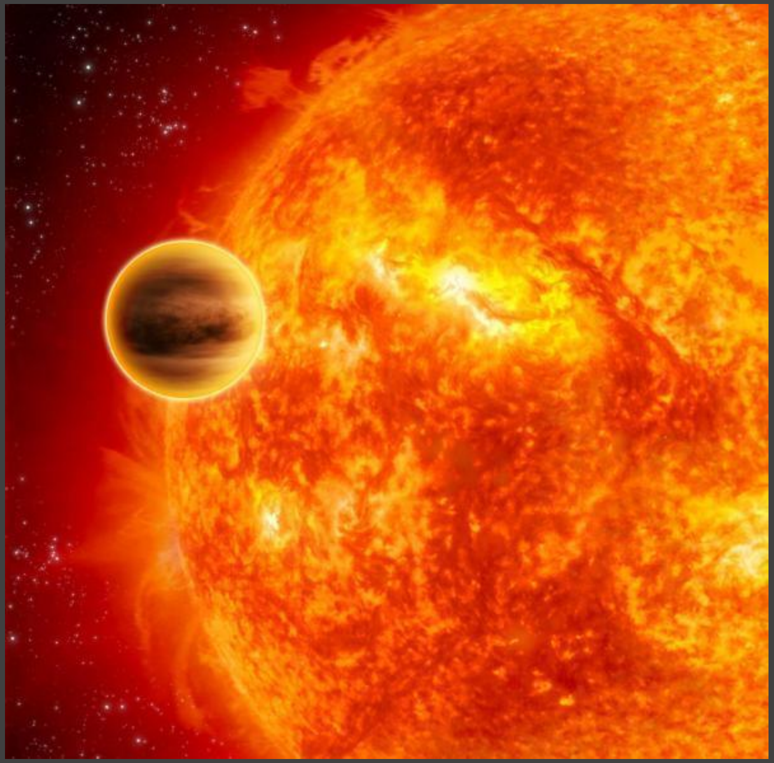
http://www.freemars.org/jeff/planets/Luna/Luna.htm

To be large respect to the host-planet
the satellite might be irregular.

Systems with many planets
are more favorable.

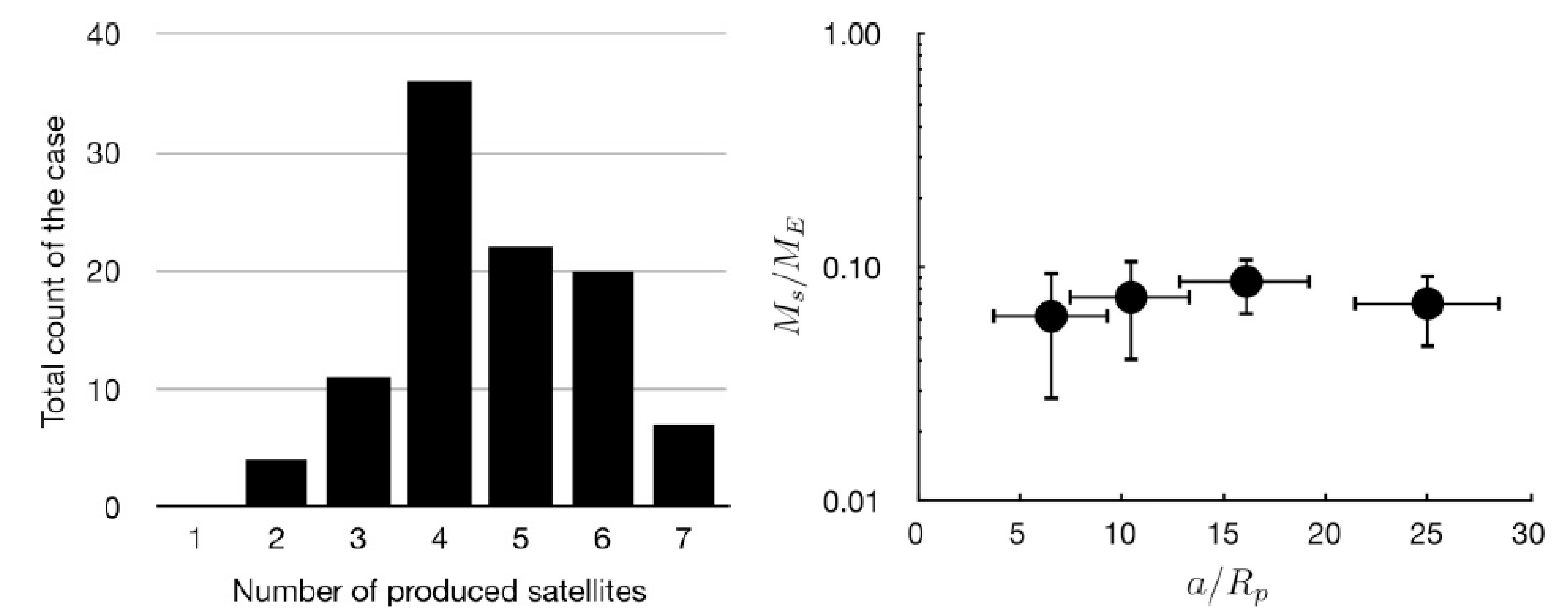
Larger planets have larger moons.

Hot jupiters (and neptunes) can
loose planets during migration.



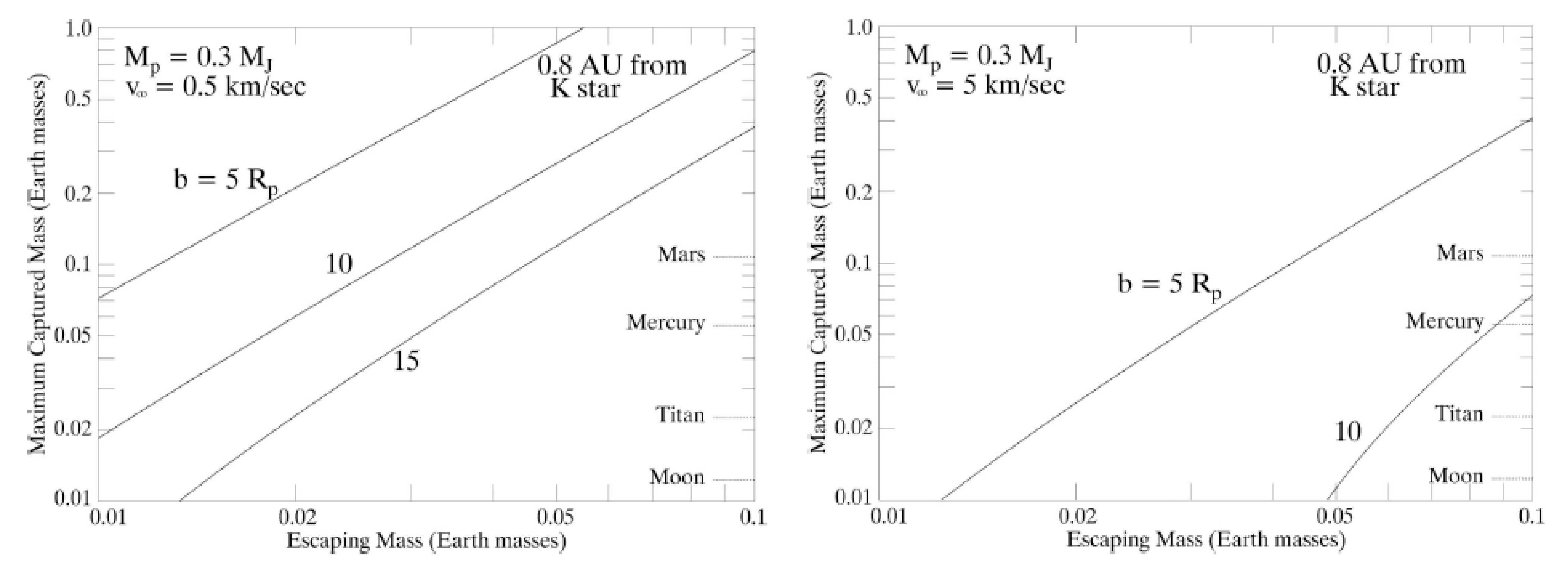
http://minerva.union.edu/parkashv/hotjupiters.html

Modeling satellite formation



A massive planet: 10 $M_{jupiter}$

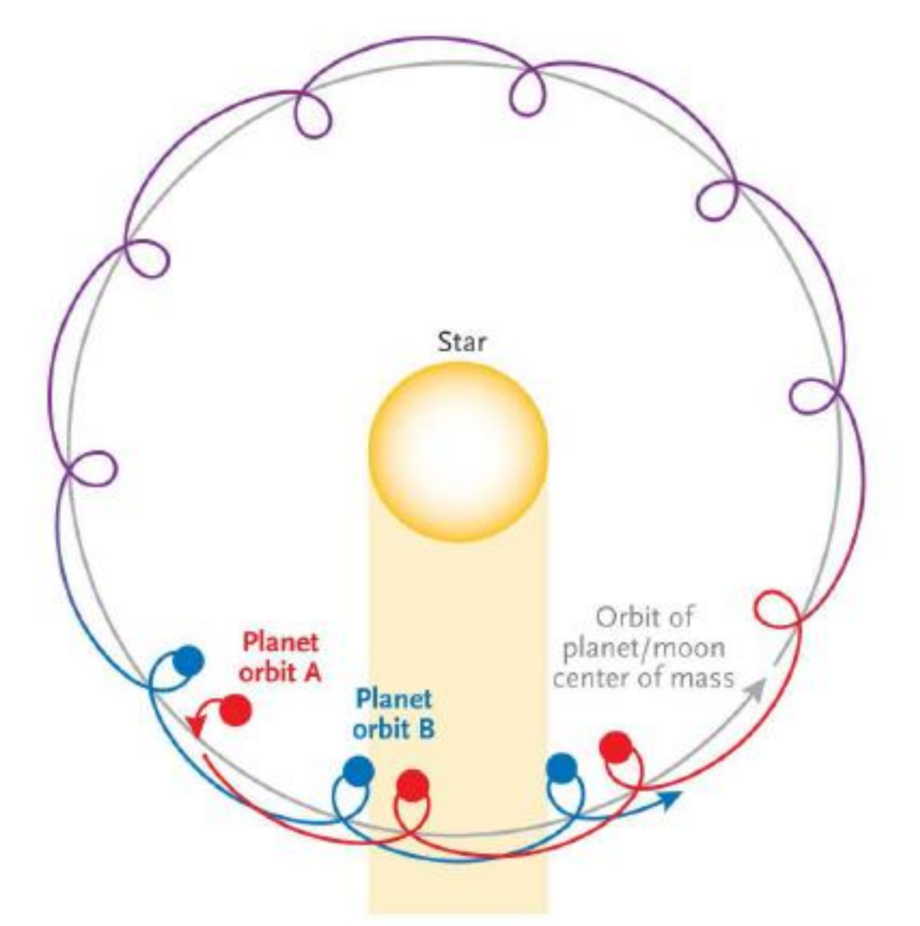
Satellite capture in three-body interaction



Results of modeling of a satellite capture. The body initially had a companion which was lost during three-body interaction. This scenario requires a massive planet. Such interactions can happen in the habitable zone.

1408.6164

How to find an exomoon



Potentially, all methods for exoplanets discovery can work.

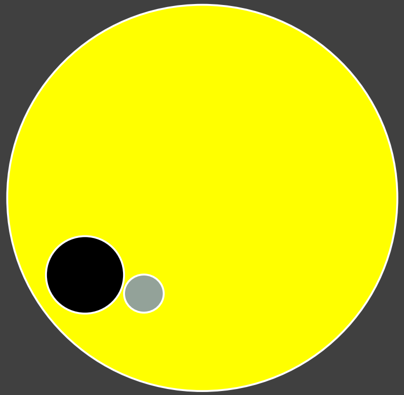
However, presently methods related to transits seems to be more favorable:

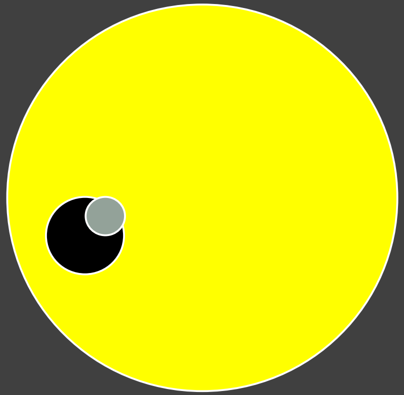
1. TTV

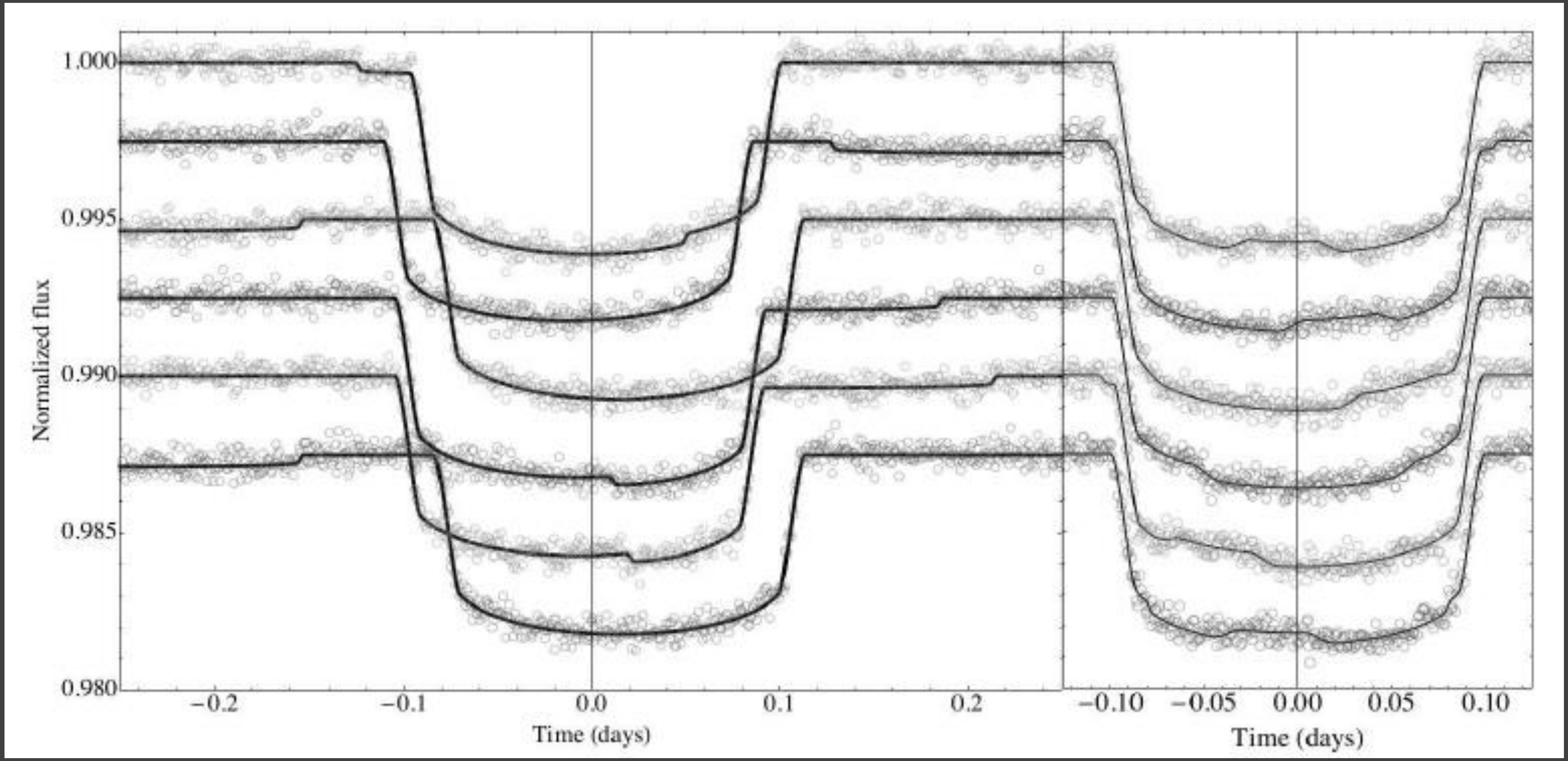
2. TDV

3. Orbital plane changes.

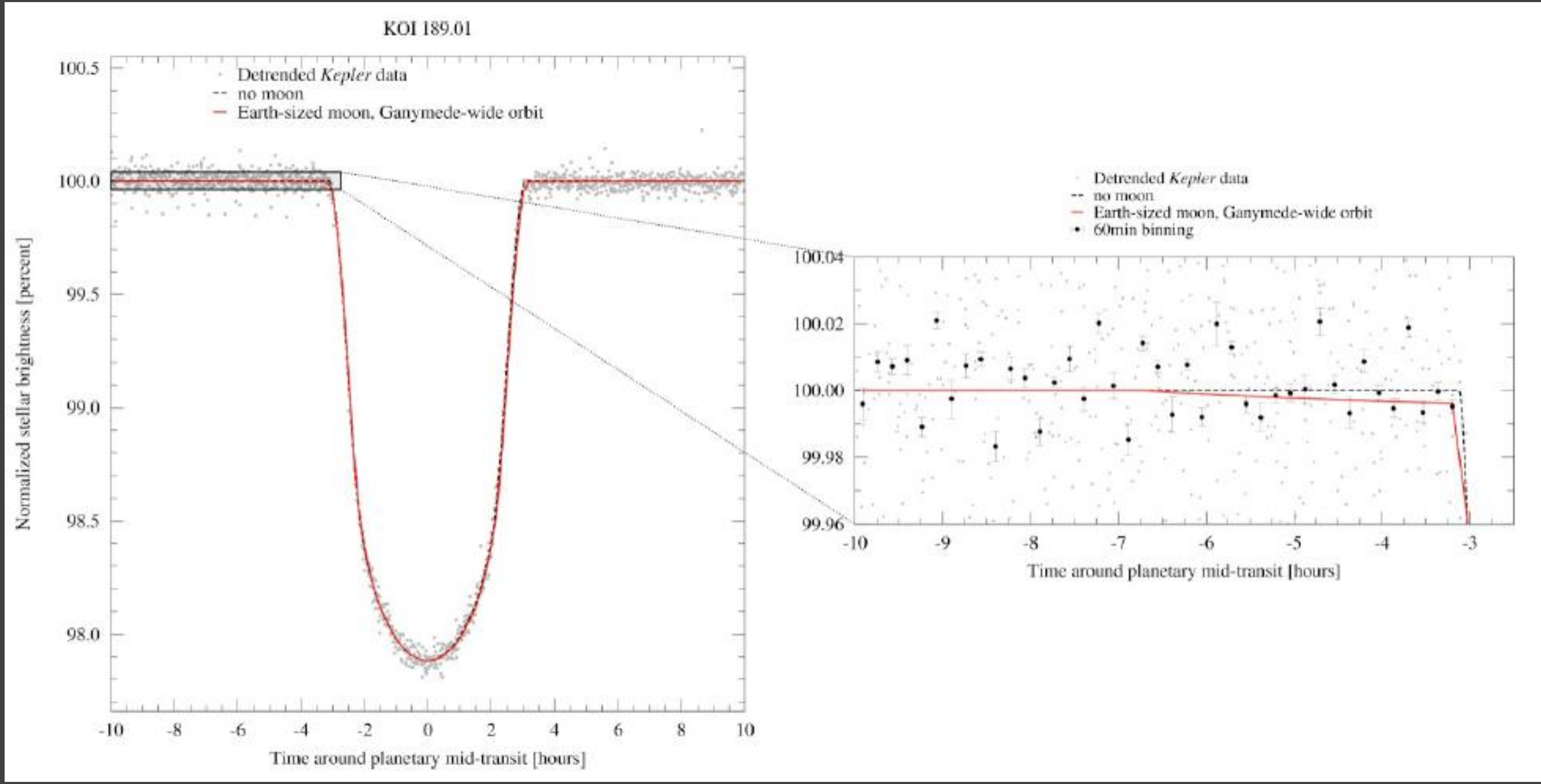
Joint transits



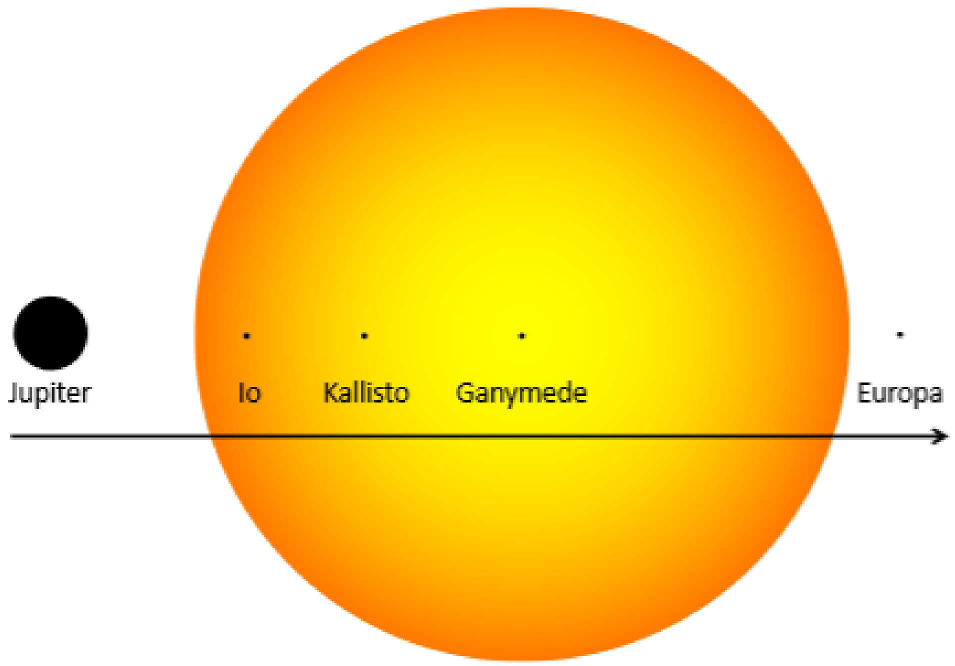


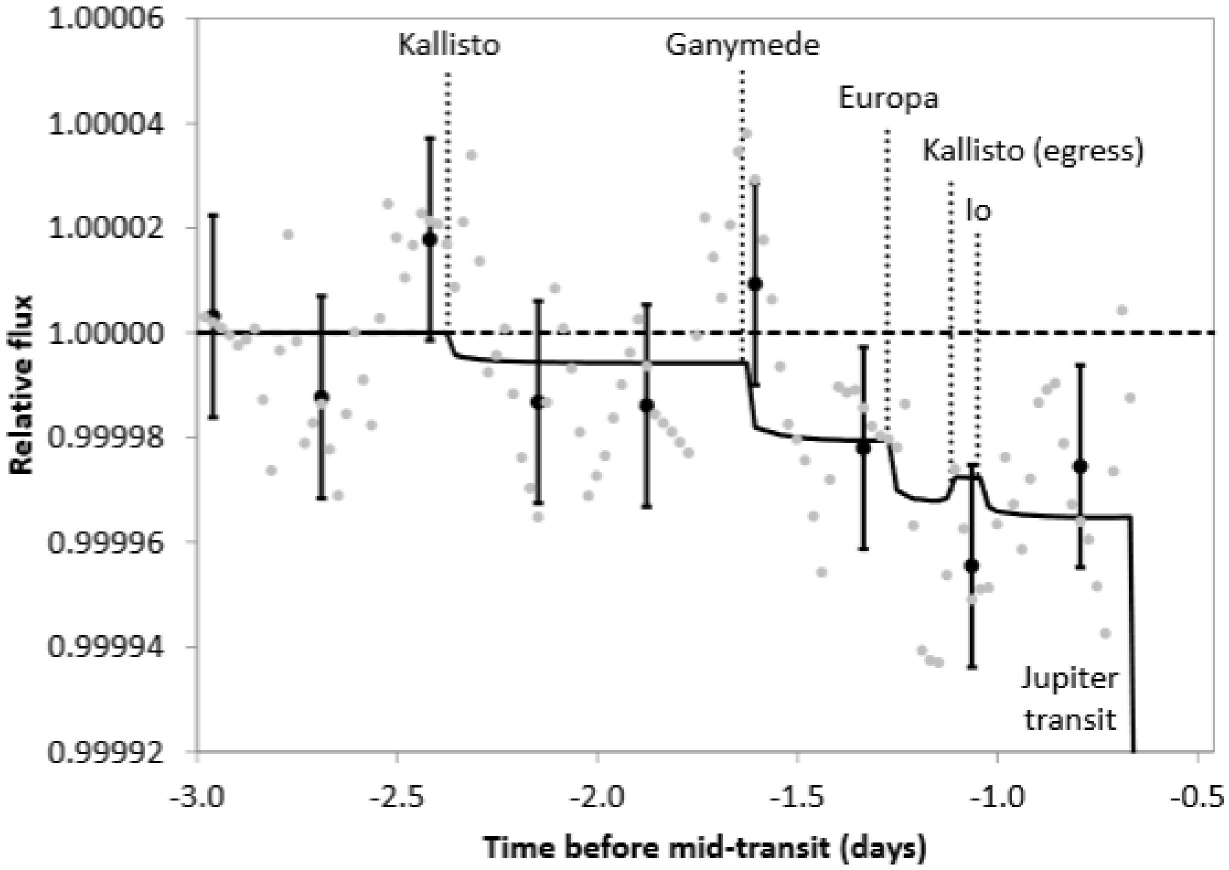


How strong is the effect?

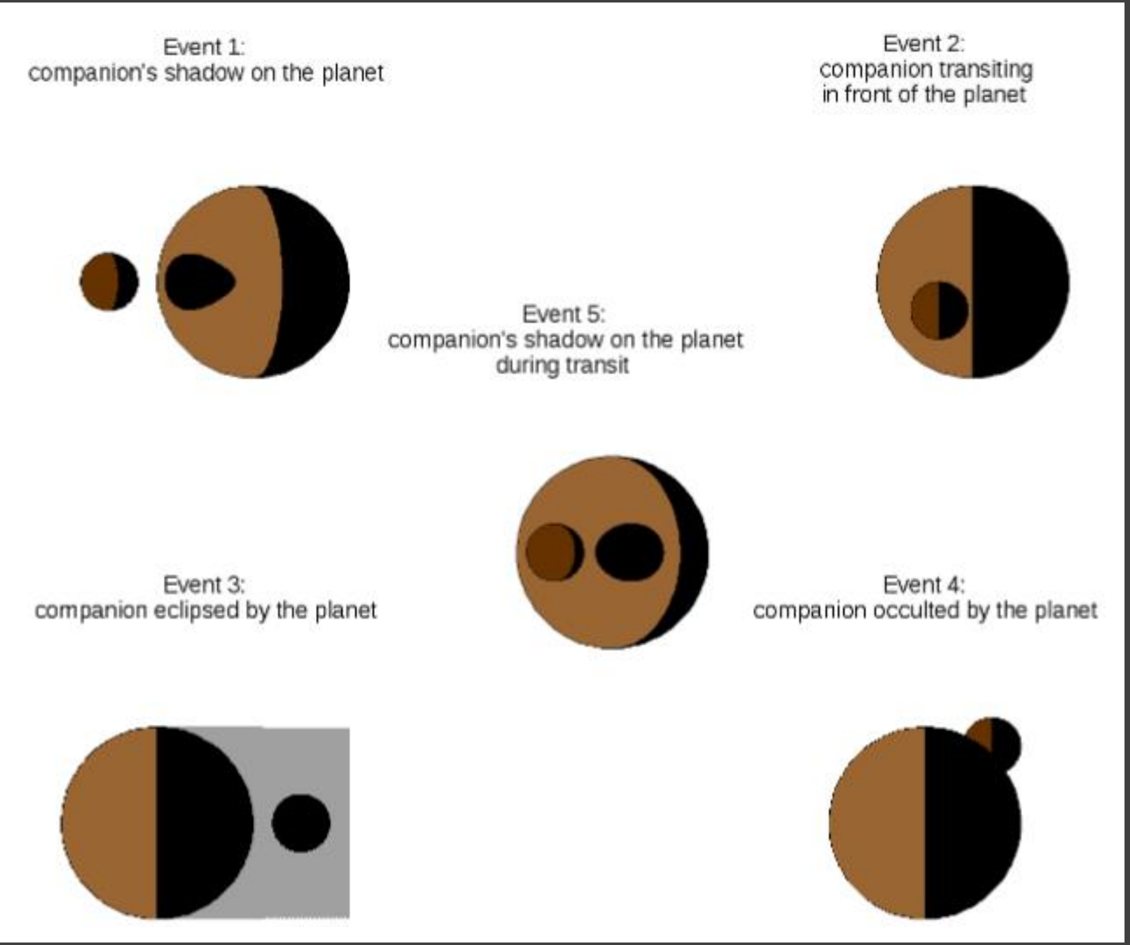


An example: Jupiter with satellites over the Sun





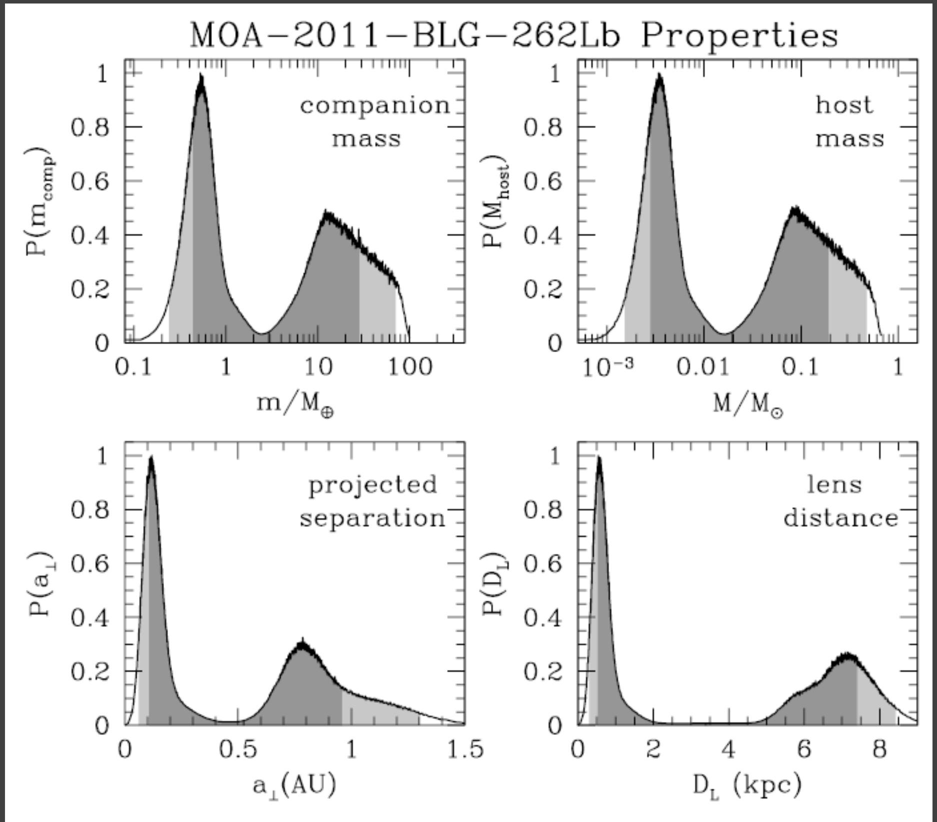
Other ways to see a moon

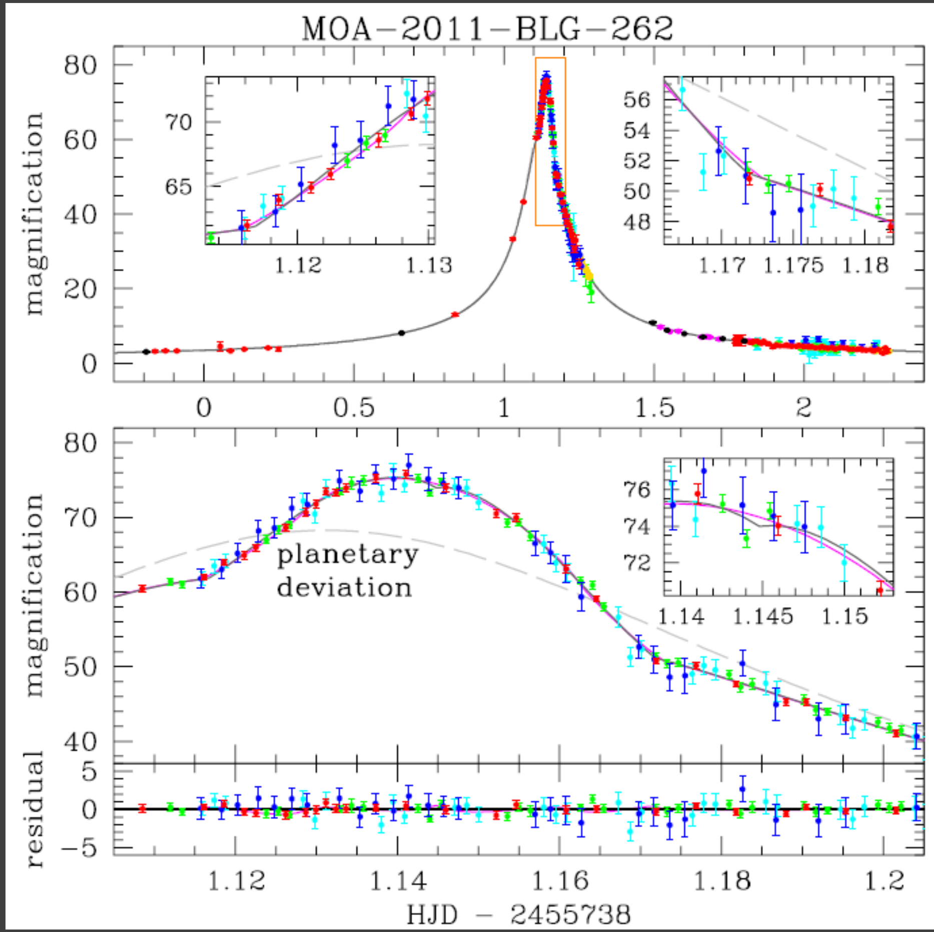


A planet with a moon …but without a star?

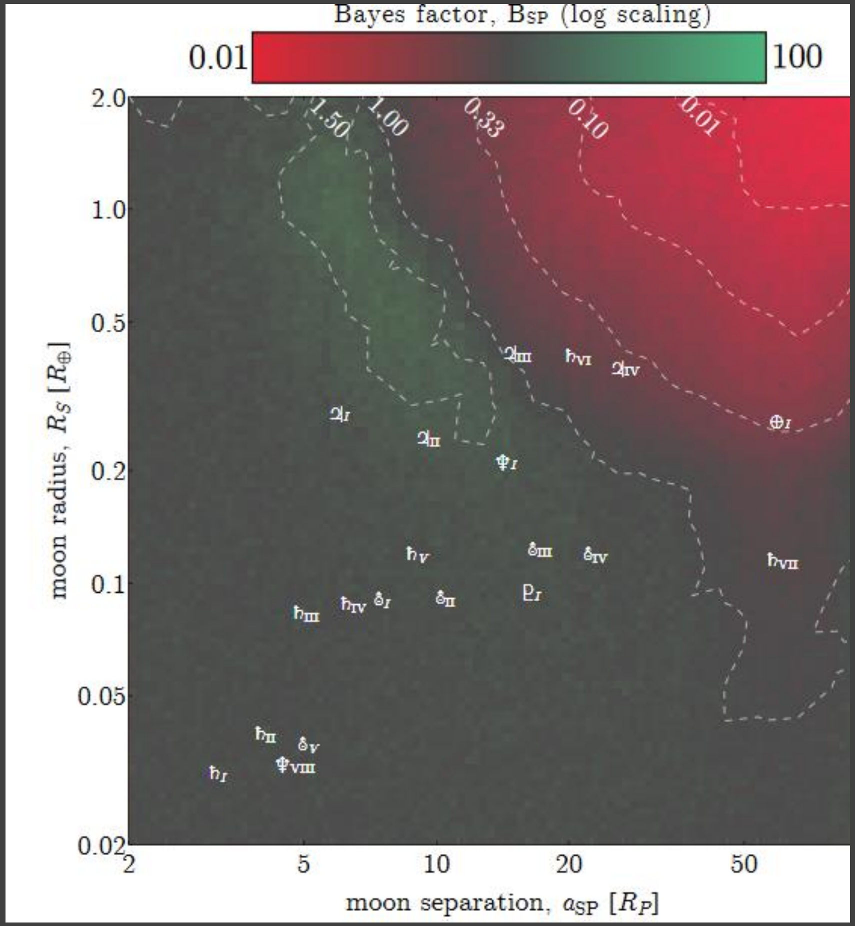
Microlensing.
Two solutions are possible:
1. $0.12M_{sun}+18M_{Earth}$
2. $4M_{Jup}+0.5M_{Earth}$

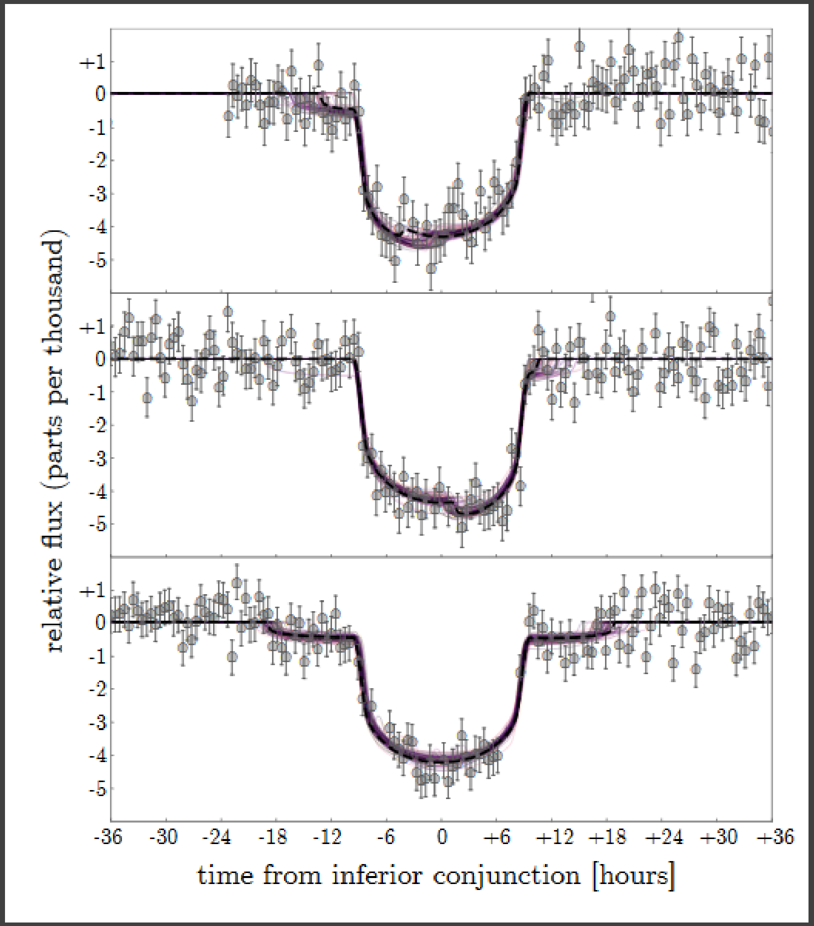
Uncertainty is related to unknown distance





New measurements and a candidate





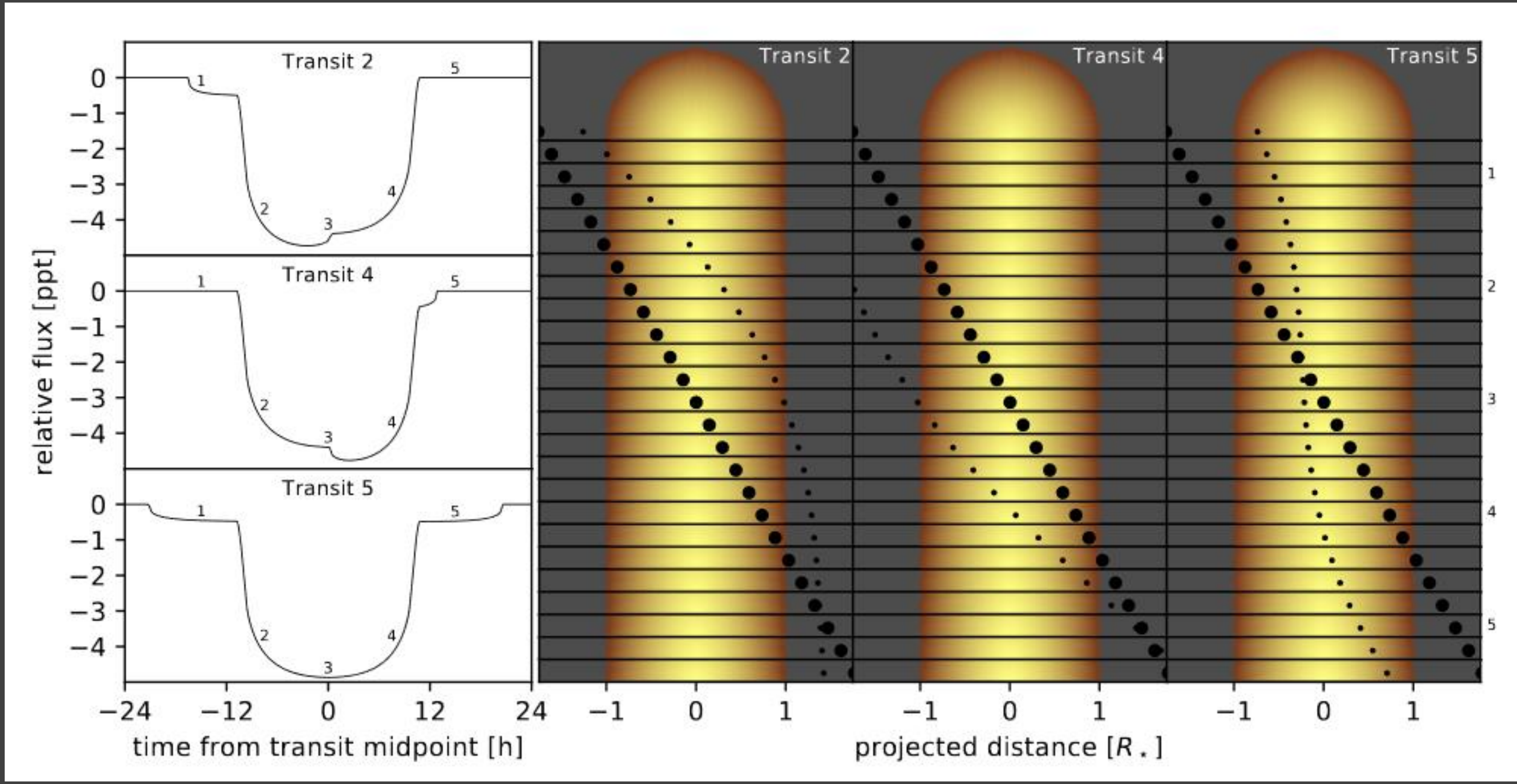
Kepler-1625BI

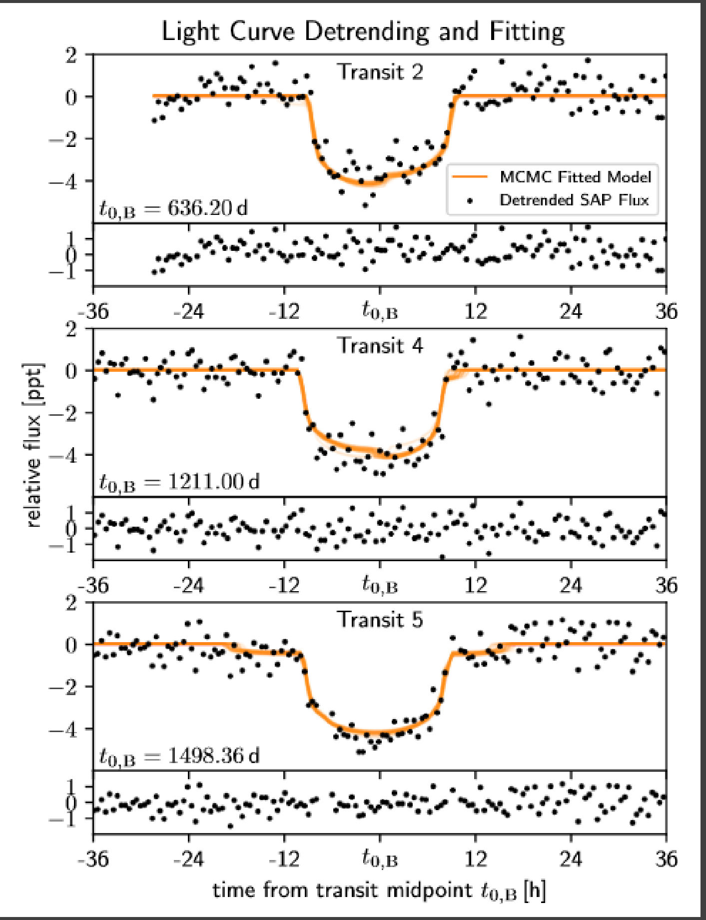
Semimajor axis: 20 planet radii.

Jupiter-like planet.

Planet orbit: 0.8 AU

Confirmation of the candidate





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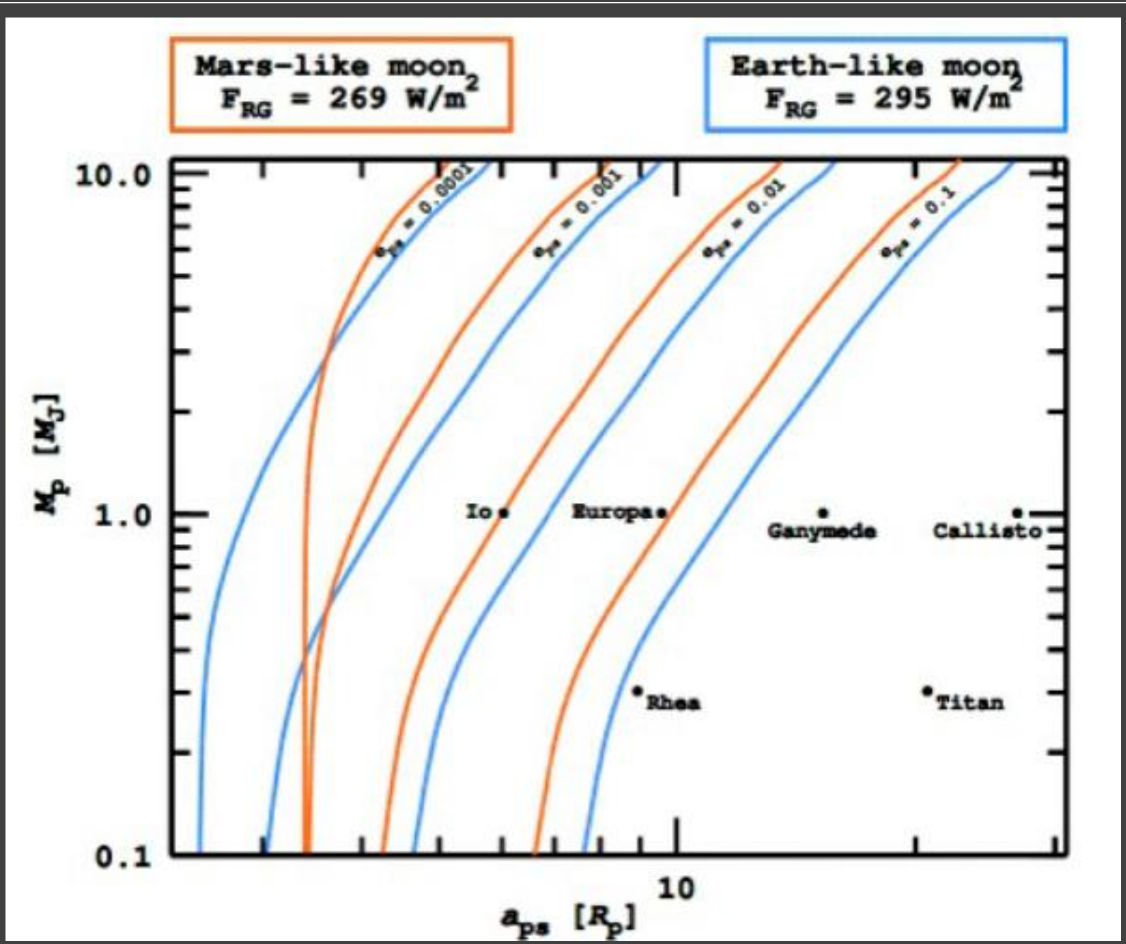
Tidal heating

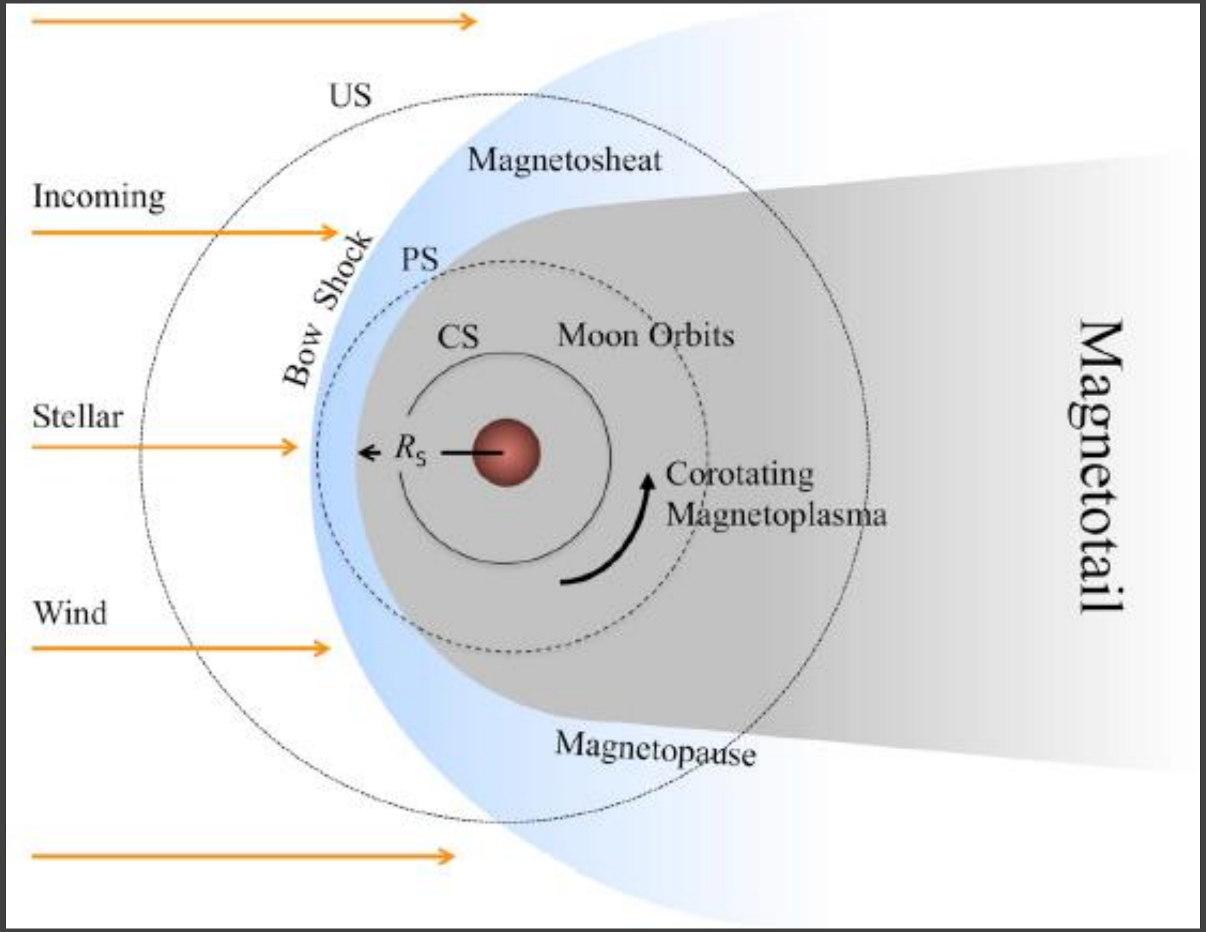


Satellites can be heated by tides.

Effect can be so strong,
that a satellite with an atmosphere
can experience the greenhouse effect.

1408.6164

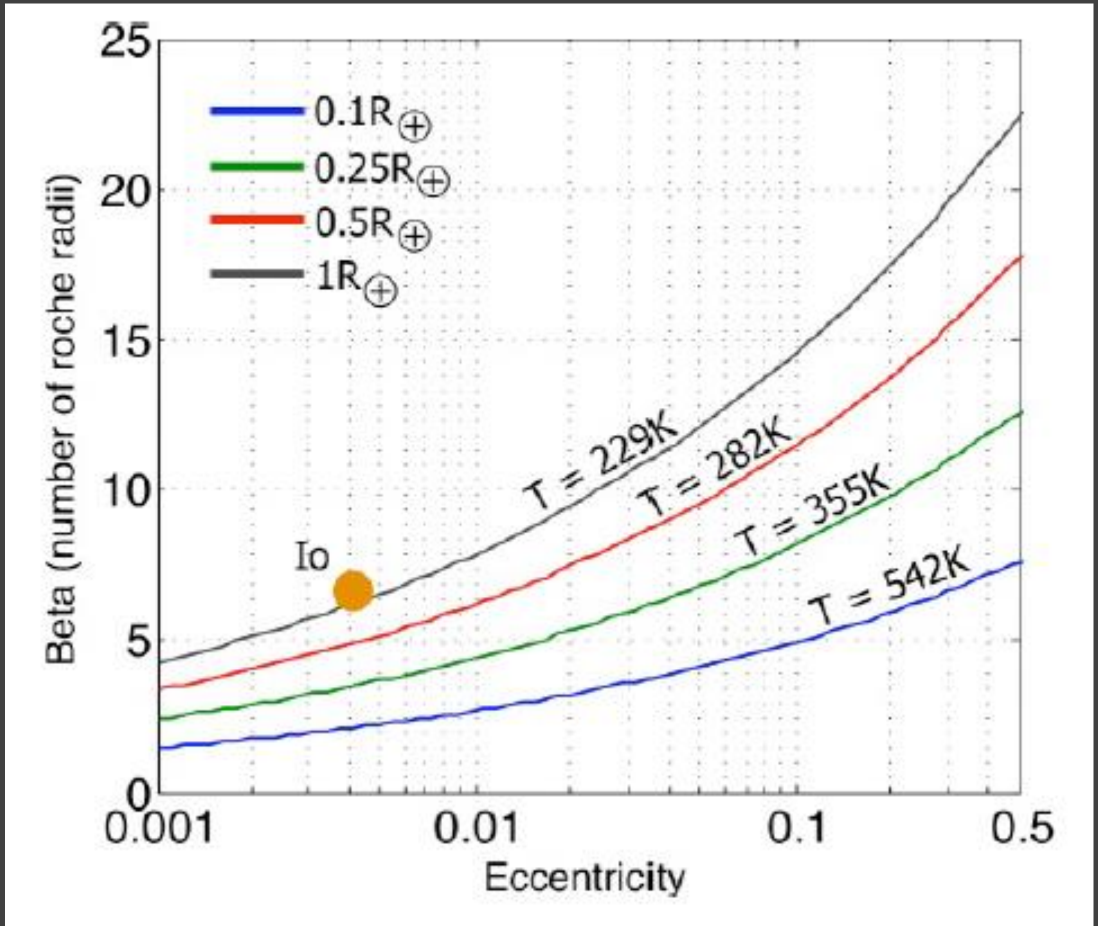
Planetary magnetospheres

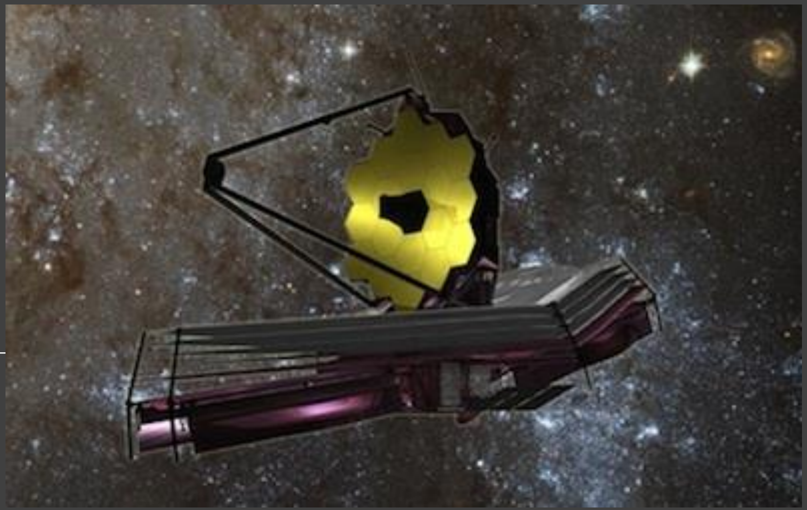


It is argued that magnetic shield
can be important for life.
A satellite can ``use'' the planetary field.

However, if the satellite is too close
to the planet – then tides can heat it up.
If it is too far – it can be out of the magnetosphere.

Can JWST see exomoons?





A satellite might be large (as the Earth) and warm (also as the Earth, at least).

Potentially, such satellites can appear around massive planets far from the star, where it is easier to see them.
A satellite can be heated by tides.

Literature

- arxiv:1604.06092 Exoplanetary Atmospheres - Chemistry, Formation Conditions, and Habitability
- arxiv:1507.03966 Observations of Exoplanet Atmospheres
- arxiv:1401.4738 Planetary internal structures
- arxiv:1312.3323 The Structure of Exoplanets
- arxiv:1501.05685 Exoplanetary Geophysics -- An Emerging Discipline
- arxiv:1701.00493 Illusion and Reality in the Atmospheres of Exoplanets
- arxiv:1411.1740 Seismology of Giant Planets
- arxiv:1709.05941 Exoplanet Atmosphere Measurements from Transmission Spectroscopy
- arxiv:1810.04175 How to characterize the atmosphere of a transiting exoplanet
- arXiv:1904.03190 Exoplanetary atmospheres