Kick velocity

Why do neutron stars move so rapidly?
Stars vs. Neutron Stars

Progenitor have velocities about \( \sim 10-30 \text{ km/s} \)
PSRs have velocities \( \sim 100-500 \text{ km/s} \)
Pulsar velocity distribution

Normal stars have velocities \( \sim 10-30 \text{ km/s} \).

Already in 70s it became clear that PSRs have high spatial velocities (\( >10 \text{ km/s} \)).

A breakthrough happened in 1994 when Lyne and Lorimer in a seminal paper in Nature showed that velocities are even higher than it was thought before – hundreds km/s.

Note, that the observed distribution is much different from the initial one. To derive the later it is necessary to calculate a model.
SN explosions should not be symmetric!

\[ E_{\text{total}} \approx 3 \times 10^{53} \text{ erg} \]

Most of energy is carried away by neutrinos.

~Few % asymmetry in energy release can produce a strong kick up to 1000 km/s.

**Main kick mechanisms**

- Asymmetric mass ejection (Shklovsky 1970)
- Asymmetric neutrino emission (Chugai 1984)

Asymmetric mass ejection includes three mechanisms:

- gravitational pull due to asymmetric matter
- asymmetric neutrino emission due to matter distribution
- asymmetric matter jets (Khokhlov et al. 1999)

*Leonid Ozernoy in 1965 discussed asymmetry of SN explosions in the context of GW radiation.*
SN and kick explosion mechanisms

<table>
<thead>
<tr>
<th>Mechanism</th>
<th>Time scale</th>
<th>$V_{\text{max}}$, km s$^{-1}$</th>
<th>Alignment (spin and $V$)</th>
<th>Main recent refs.</th>
</tr>
</thead>
<tbody>
<tr>
<td>Hydrodynamical</td>
<td>0.1 s</td>
<td>$\sim (100 - 200)$</td>
<td>random</td>
<td>Lai et al. (2001)</td>
</tr>
<tr>
<td>$\nu$-driven</td>
<td>$\sim$ few s</td>
<td>$\sim 50 , B_{15}$</td>
<td>parallel</td>
<td>Lai et al. (2001)</td>
</tr>
<tr>
<td>Electromagnetic rocket</td>
<td>long</td>
<td>$1400 , R_{10}^2 , P_{\text{ms}}^{-2}$</td>
<td>parallel</td>
<td>Lai et al. (2001), Huang et al. (2003)</td>
</tr>
<tr>
<td>Binary disruption (without add. kick)</td>
<td>$&lt;&lt; P_{\text{orb}}$</td>
<td>$\sim 1000$</td>
<td>perpendicular</td>
<td>Iben &amp; Tutukov (1996)</td>
</tr>
<tr>
<td>NS instability</td>
<td>few ms</td>
<td>$\sim 1000$</td>
<td>perpendicular</td>
<td>Colpi &amp; Wasserman (2002), Imshennik &amp; Ryazhskaya (2004)</td>
</tr>
<tr>
<td>Magnetorotational</td>
<td>0.2 s – minutes</td>
<td>$\sim 300$</td>
<td>quasirandom</td>
<td>Moiseenko et al. (2003), Ardeljan et al. (2004)</td>
</tr>
</tbody>
</table>

For neutrino emission: $V_{\text{kick}} = \varepsilon \, E_{\text{tot}} / M_c \sim 1000$ km/s ($\varepsilon/0.1$) ($E_{\text{tot}}/10^{53}$ erg). Also it depends on the magnetic field.
Kick modeling

Recently, new results on the origin of NS and BH kicks have been obtained:

• Neutrino-triggered asymmetric magnetorotational mechanism arXiv:1110.1041
• Hydrodynamic Origin of Neutron Star Kicks arXiv: 1112.3342
• Three-dimensional neutrino-driven supernovae arXiv:1210.8148
• BH kicks arXiv:1203.3077

A review on SNae properties and explosion mechanisms: arXiv:1210.4921
To kick or not to kick?

Up to mid-90s it was not clear if kicks are absolutely necessary.

- Tademaru (rocket) mechanism
- Binary disruption (Blaaw mechanism)
- Core fragmentation (Berezinski et al., Imshennik)

However, some discoveries directly point to necessity of natal kicks.
Direct evidence

1. High-velocity NSs and binaries
2. Spin inclination in binaries and geodetic precession

Orbit inclination relative to a normal star equator can be measured due to:
- orbital precession due to spin-orbit interaction (Kaspi et al. 1996)
- circumstellar disc inclination (Prokhorov, Postnov 1997)

The most spectacular 3D velocity measurements for NSs are related to nebulae around these objects.

The transversal velocity can be measured by proper motion observations of radio pulsars and other neutron stars.

For binaries large velocities are measured (Cir X-1: Johnston et al. 1999).
Double pulsar PSR J0737-3039

Pulsar A’s spin is tilted from the orbital angular momentum by no more than 14 degrees at 95% confidence; pulsar B’s -- by $130 \pm 1$ degrees at 99.7% confidence.

This spin-spin misalignment requires that the origin of most of B’s present-day spin is connected to the supernova that formed pulsar B. The spin could be thought of as originating from the off-center nature of the kick.

See also 1302.2914 about probably near-zero for the pulsar A.
Many kick velocity distributions are proposed

Three popular models:

- Arzoumanian, Chernoff, Cordes (2002)
- Hobbs et al. (2005)
- Faucher-Giguer and Kaspi (2006)

Note the difference: We observe present day velocities with selection and evolutionary effects, but we are interested in the velocity at birth!

(Faucher-Giguer, Kaspi 2006)
Bimodal distribution

Arzoumanian et al. 2002

Scheck et al. 2006
Hydrodynamical models

2D simulations

Acceleration of a NS is mainly due to gravitational pull of the anisotropic ejecta

Scheck et al. See 3D calculations in 1010.0167
Spin-kick alignment resulting from a neutrino-driven explosion launched from a phase of strong spiral-SASI activity.

While the explosion starts by equatorial expansion, the final NS kick is determined by the slower mass ejection in the polar directions.

The NS is accelerated by the gravitational attraction of the mass in these more slowly expanding, dense regions. In the cartoon the NS is pulled more strongly towards the northern direction and therefore opposite to the (southern) hemisphere where the explosion is more powerful.
Spatial velocity and spin axis are nearly coincident. *Nearly* is important: there is some misalignment.
Crab and Vela are not the only cases, but are the best studied ones. Spin-velocity correlation (in direction) is reported for many radio pulsars. For some of them pulsar wind nebula observations are used, for some only direction of proper motion and polarization properties can be used.
Some set of PSRs with known spin-velocity orientation

<table>
<thead>
<tr>
<th>2D Pulsars</th>
<th></th>
<th></th>
<th></th>
<th></th>
</tr>
</thead>
<tbody>
<tr>
<td>B0628−28</td>
<td>318</td>
<td>+61/−64</td>
<td>5 ± 4</td>
<td></td>
</tr>
<tr>
<td>B0740−28</td>
<td>259</td>
<td>+190/−149</td>
<td>7 ± 5</td>
<td></td>
</tr>
<tr>
<td>B0823+26</td>
<td>189</td>
<td>+55/−34</td>
<td>21 ± 7</td>
<td></td>
</tr>
<tr>
<td>B0835−41</td>
<td>170</td>
<td>±30</td>
<td>13 ± 11</td>
<td></td>
</tr>
<tr>
<td>B0919+06</td>
<td>506</td>
<td>±80</td>
<td>32 ± 17</td>
<td></td>
</tr>
<tr>
<td>B1133+16</td>
<td>639</td>
<td>+38/−35</td>
<td>22 ± 2</td>
<td></td>
</tr>
<tr>
<td>B1325−43</td>
<td>597</td>
<td>±254</td>
<td>31 ± 22</td>
<td></td>
</tr>
<tr>
<td>B1426−66</td>
<td>150</td>
<td>+40/−24</td>
<td>5 ± 9</td>
<td></td>
</tr>
<tr>
<td>B1449−64</td>
<td>219</td>
<td>+55/−18</td>
<td>1 ± 3</td>
<td></td>
</tr>
<tr>
<td>B1508+55</td>
<td>1082</td>
<td>+103/−90</td>
<td>23 ± 7</td>
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<tr>
<td>B1642−03</td>
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<td>+34/−32</td>
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<tr>
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<td>347</td>
<td>+48/−57</td>
<td>7 ± 8</td>
<td></td>
</tr>
<tr>
<td>B1842+14</td>
<td>512</td>
<td>+51/−50</td>
<td>5 ± 15</td>
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<tr>
<td>B1929+10</td>
<td>173</td>
<td>+4/−5</td>
<td>16 ± 2</td>
<td></td>
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<tr>
<td>B2045−16</td>
<td>304</td>
<td>+39/−38</td>
<td>3 ± 6</td>
<td></td>
</tr>
<tr>
<td>IC 443</td>
<td>250</td>
<td>±50</td>
<td>45 ± 10</td>
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</table>

<table>
<thead>
<tr>
<th>3D Pulsars</th>
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<th></th>
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<th></th>
</tr>
</thead>
<tbody>
<tr>
<td>J0205+6449</td>
<td>838</td>
<td>±251</td>
<td>21 ± 10</td>
<td></td>
</tr>
<tr>
<td>B0531+21</td>
<td>140</td>
<td>±8</td>
<td>26 ± 3</td>
<td></td>
</tr>
<tr>
<td>J0537−6910</td>
<td>634</td>
<td>±50</td>
<td>3 ± 5</td>
<td></td>
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<tr>
<td>J0538+2817</td>
<td>407</td>
<td>+116/−74</td>
<td>12 ± 4</td>
<td></td>
</tr>
<tr>
<td>B0540−69</td>
<td>1300</td>
<td>±612</td>
<td>34 ± 33</td>
<td></td>
</tr>
<tr>
<td>B0833−45</td>
<td>61</td>
<td>±2</td>
<td>10 ± 2</td>
<td></td>
</tr>
<tr>
<td>B1706−44</td>
<td>645</td>
<td>±194</td>
<td>35 ± 10</td>
<td></td>
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<tr>
<td>J1833−1034</td>
<td>125</td>
<td>±30</td>
<td>16 ± 15</td>
<td></td>
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<tr>
<td>B1951+32</td>
<td>273</td>
<td>±11</td>
<td>18 ± 5</td>
<td></td>
</tr>
</tbody>
</table>

[Ng & Romani 2007]
Recent data on radio pulsars

The tendency is clear, but it is only a tendency.

Rankin (2007)

Johnston et al. (2007)

New data and discussion in 1502.05270
Alignment

47 pulsars with well-determined parameters.
Why?

Kick can be confined in some angle around the spin axis. Typical cones must be $\leq 10^\circ$ (see, for example, Kuranov et al. 2009).

Kick mechanism can be operative for a long time (many spin periods), so that its influence is average. Typical duration must be 1-10 sec.
Kicks in binary evolution
Influence of kicks on binaries

Kicks can both – destroy and save – binaries!

If a more massive star is about to explode, and the lost mass is larger than the sum of the remnant mass and the secondary mass, then the system should be destroyed. But favorable kick can save it.

It is very important for NS+NS, NS+BH, BH+BH binaries. Because typically in the second SN the more massive star is exploding. So, kicks save systems to produce GW.
Parameters of binaries after kicks

Kicks significantly influence binary parameters (for example, eccentricity distribution). This is specially important for systems which survived the second explosion (NS+NS).

There are examples, when a NS rotates “in a wrong direction”, i.e. its orbital motion is in the direction opposite to the spin of the second companion.

For detailed description see Postnov, Yungelson (astro-ph/0701059) pp. 18-22.

\[
\frac{a_f}{a_i} = \left[ 2 - \chi \left( \frac{w_x^2 + w_z^2 + (V_i + w_y)^2}{V_i^2} \right) \right]^{-1}
\]

\[
1 - e^2 = \chi \frac{a_i}{a_f} \left( \frac{w_z^2 + (V_i + w_y)^2}{V_i^2} \right)
\]

\[
\cos \theta = \frac{\vec{J}_f \cdot \vec{J}_i}{|\vec{J}_f| \cdot |\vec{J}_i|},
\]

\[
\cos \theta = \frac{V_i + w_y}{\sqrt{w_z^2 + (V_i + w_y)^2}}.
\]

\[
V_f \geq \sqrt{\frac{2}{\chi}}.
\]

\[\chi \equiv \frac{(M_1 + M_2)}{(M_c + M_2)} \geq 1.\]
**e⁻-supernovae with low kicks**

In 80s it was proposed by Nomoto, Miyaji et al. that in some cases a SN explosion can happen due to electron capture by $^{24}\text{Mg}$ and $^{20}\text{Ne}$ (no iron core is formed).

It was noticed (Pfahl et al. 2002, Podsiadlowski et al. 2004; van den Heuvel 2004, 2007) that among Be/X-ray binaries there is a group of systems with small eccentricities. But they suffered one SN explosion and there was no Roche-lobe overflow. This means that kicks in these systems were low. The same is true for some of NS+NS binaries.

The proposed mechanism is related to e⁻-capture SN.

Such explosions can appear not only in binary systems, but in binaries they can be more frequent. Among isolated stars about 4% (up to ~20%!) of SN can be of this type (Poelarends et al. 2008). [It is not clear if they appear among normal PSRs.]

Why kick is low? Uncertain. Low core mass, rapid explosion, low mass ejection…
e⁻-capture SN in binaries

Poelarends et al. 2008

The final size of the helium core depends strongly on the evolution of the hydrogen-rich envelope and on the dredge-up process.
Evolution of $\text{e}^-$-capture SN progenitors

Critical core mass 1.367 solar masses.

For initial stellar masses >11 solar neon is ignited, and later on a Fe-core is formed.
Pulsars with low velocities

Some NSs demonstrate low spatial velocities. Obviously, this is due to low kicks.

3C58.
Low kick velocity.
Projected velocity 30-40 km/s
Kicks as fingerprints

Think about young highly magnetized NSs of different types:

- SGR
- AXP
- RRATs
- Magnificent Seven

Are they relatives?

It is a difficult question, but velocity measurements can give you a hint. Even if fields are decayed, rotation is slowed down, thermal energy is emitted ... if they are relatives – velocity distributions must be identical. Unfortunately, now we do not know the answer.
### Magnetar velocity measurements

<table>
<thead>
<tr>
<th>Source</th>
<th>Velocity Range</th>
<th>Reference</th>
</tr>
</thead>
<tbody>
<tr>
<td>SGR 1806-20</td>
<td>350 +/- 100 km/s</td>
<td>arXiv:1210.8151</td>
</tr>
<tr>
<td>SGR 1900+14</td>
<td>130 +/- 30 km/s</td>
<td>arXiv:1210.8151</td>
</tr>
<tr>
<td>PSR J1550-5418</td>
<td>280 +/- 130 km/s</td>
<td>arXiv:1201.4684</td>
</tr>
<tr>
<td>XTE J1810-197</td>
<td>200 km/s</td>
<td>Helfand et al. (2007)</td>
</tr>
</tbody>
</table>
Record velocities

1. PSR J1357-6429 1600-2000 km/s arXiv:1206.5149 - shown to be wrong

2. IGR J11014-6103 2400-2900 km/s arXiv: 1204.2836
CCO velocities
RX J0822-4300 in the Supernova Remnant Puppis A
672 +/- 115 km/s
arXiv: 1204.3510
Do you play billiard?

Non-central kicks can spin-up a NS. In some cases one can speculate that a new rotation axis is determined mainly by non-central kick. But then velocity - spin period correlation is expected.
Evolution of an isolated NS depends on the intensity of its interaction with the ISM. This intensity depends on the relative velocity of a NS and the ISM.

Will a NS start to accrete from the ISM, or will it stay as Ejector, or Propeller, or will it enter another regime strongly depends on the relative velocity of a NS and the ISM.
Clearly, some NSs are rapid enough to leave the Galaxy.

Z-distribution of PSRs is much wider than the progenitors’ one.
Mass distribution in the Galaxy

Klypin et al. (2002)
“Paczynski” model

Disc+Bulge+Halo
Actually, it is Miyamoto, Nagai (1975) model.
It is simple and popular in NS motion calculations.

At the very center one has to add the central BH potential

\[
\phi_i(R, z) = \frac{GM_i}{\left\{R^2 + [a_i + (z^2 + b_i^2)^{1/2}]^2\right\}^{1/2}}, \quad R^2 = x^2 + y^2
\]

\[
\phi_h = -\frac{GM_c}{r_c} \left[\frac{1}{2} \ln \left(1 + \frac{r^2}{r_c^2}\right) + \frac{r_c}{r} \tan^{-1}\left(\frac{r}{r_c}\right)\right],
\]

\[
M_c \equiv 4\pi \rho_c r_c^3
\]

\[
a_1 = 0, \quad b_1 = 0.277 \text{kpc}, \quad M_1 = 1.12 \times 10^{10} M_\odot, \quad a_2 = 3.7 \text{kpc}, \quad b_2 = 0.20 \text{kpc}, \quad M_2 = 8.07 \times 10^{10} M_\odot, \quad r_c = 6.0 \text{kpc}, \quad M_c = 5.0 \times 10^{10} M_\odot
\]
Examples of old NS distribution
In both models $N = 5 \times 10^8$.
Kick: ACC02.
Potential: Paczynski 1990

NS formation rate is assumed to be proportional to the square of the ISM density at the birthplace.

Formation rate is proportional to $[\exp(-z/75 \text{ pc}) \exp(-R/4 \text{ kpc})]$. 

astro-ph/0305599
X-ray sources in other galaxies

X-ray sources are shifted from the stellar light distribution. This might be due to kicks, especially in the case of NS binaries.

The effect cannot be explained by sources in globular clusters.
Black hole kicks

Do BHs obtain kicks?

- they are more massive
- horizon is formed
- SN mechanism can be different

If before the horizon formation a “protoNS-like” object is formed, then there should be a kick, but smaller (in km/s) due to larger mass.

We do not know isolated BHs, but we know binaries. It is possible to measure velocity.

Knowing just a velocity it is difficult to distinguish kick from dynamical interaction or initially large velocity (for example, a system can be from a globular cluster).

On the mechanism of BH kick see 1306.0007
GROJ1655-40

Kick 45-115 km/s

Willems et al. (2005)
Cyg X-1

For this system the distance is known very precisely. This allows to trace the trajectory back and derive the value of post-SN peculiar velocity.

It is equal to 22-32 km/s. Probably, the BH obtained a moderate kick <77 km/s.
BH binaries in the Galaxy

The situation is not clear when we look at the whole population:
- Distribution for BHs is similar to the one for NS (for kick)
- Modeled distribution for zero kick can explain, roughly, the spatial distribution (against large kick)

Also line-of-site velocities are not high

Black hole kick velocities

Some BHs receive large kicks at birth. Difficult to explain by scaling from NSs.
Large kick is not necessary.
~100 km/s is enough.
IC 10 X-1

Low kick $<130 \text{ km/s}$.

$M_{BH\ Pro}$ – progenitor mass before BH formation

1304.3756
Velocity of BH and NS X-ray binaries

Some BHs might obtain significant kick.

NS binaries kick distribution is compatible with the one derived from PSRs.
Conclusions

- NSs and (most probably) BHs obtain natal kicks
- For NSs kick velocity can be as large as >1000 km/s
- The direction of the kick and rotation are correlated
- Kicks depend on the SN mechanism
- Kicks influence parameters of binaries
- Kicks influence evolution of isolated NSs
Important papers

- Lai astro-ph/0212140 – different kick mechanisms
- ATNF catalogue – database including PSR transversal velocities
- Postnov, Yungelson astro-ph/0701059 – kicks in binaries (pp.18-23)
- Ofek et al. NS spatial distribution. arXiv: 0910.3684