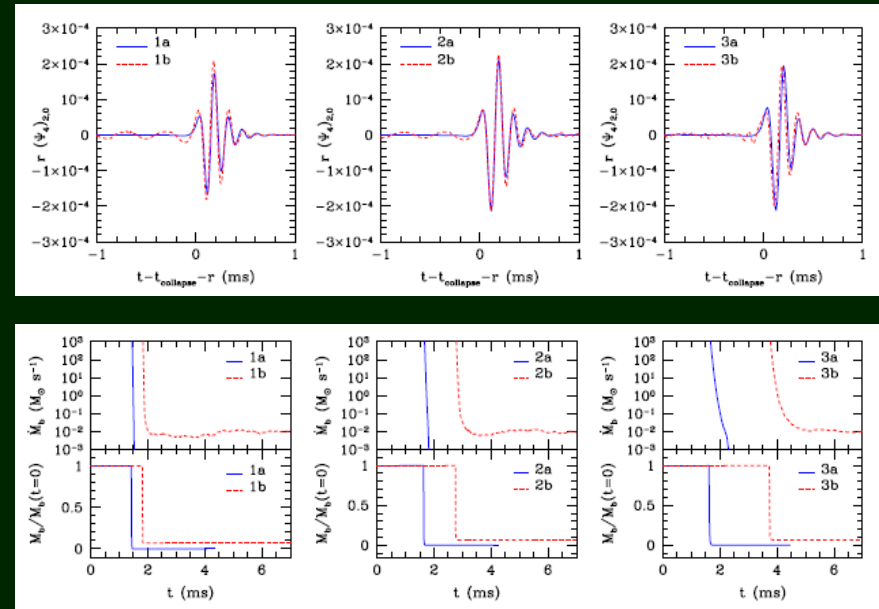
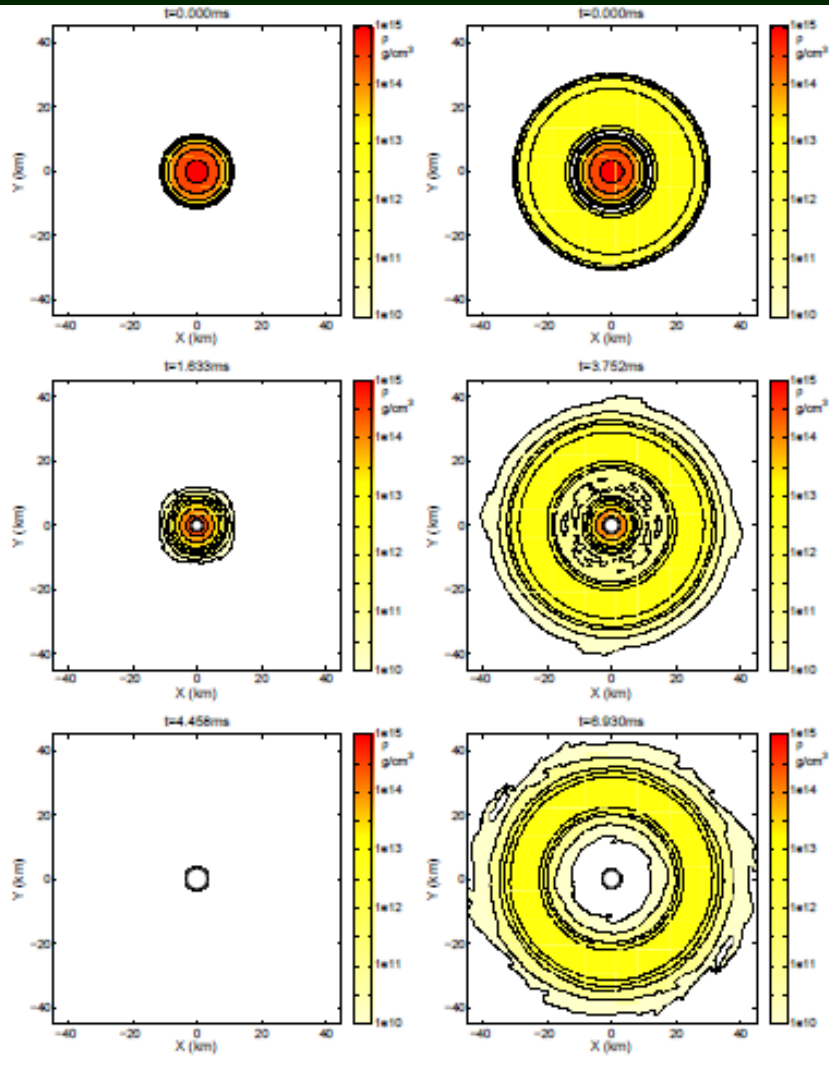


# Black holes: Introduction

# NS to BH



The authors studied collapse from NS to BH. Calculations were done for two cases: with and without massive (7%) disc. If a disc is present then such objects can appear as sGRB. GW signal is weak, and so they are a subject for the third generation of detectors.

# Main general surveys

- [astro-ph/0610657](#) [Neven Bilic](#) BH phenomenology
- [astro-ph/0604304](#) [Thomas W. Baumgarte](#) BHs: from speculations to observations
- [hep-ph/0511217](#) [Scott A. Hughes](#) Trust but verify: the case for astrophysical BHs
- [arXiv: 0907.3602](#) [Josep M. Paredes](#) Black holes in the Galaxy
- [arXiv: 1003.0291](#) [S.-N. Zhang](#) Astrophysical Black Holes in the Physical Universe
- [arXiv: 1312.6698](#) [Narayan, McClintock](#) Observational Evidence for Black Holes
- [arXiv: 1711.10256, 1810.07032, 1810.07041](#) [C. Bambi](#)  
Astrophysical black holes: several reviews
- [arXiv: 1808.01507](#) [Eric Curiel](#) The Many Definitions of a Black Hole
- [arXiv: 1809.09130](#) [Michela Mapelli](#) Astrophysics of stellar black holes

# BHs as astronomical sources

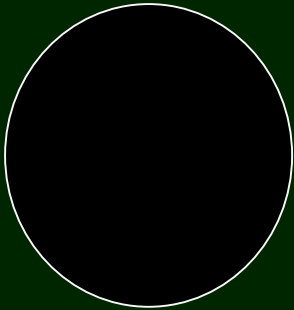
- **Primordial BHs.**  
Not discovered, yet. Only upper limits (mostly from gamma-ray observations).
- **Stellar mass BHs.**  
There are more than twenty good candidates in close binary systems.  
Accretion, jets. Observed at all wavelengths.  
Isolated stellar mass BHs are not discovered up to now.  
But there are interesting candidates among microlensing events.
- **Intermediate mass BHs.**  
Their existence is uncertain, but there are good candidates among ULX.  
Observed in radio, x-rays, and optics.
- **Supermassive BHs.**  
There are many (dozens) good candidates with mass estimates.  
In the center of our Galaxy with extremely high certainty there is supermassive BH.  
Accretion, jets, tidal disruptions of normal stars.  
Observed at all wavelengths.



# Что такое черная дыра?

## Для физика

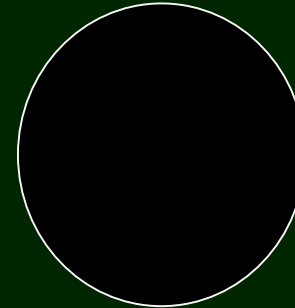
Обладает определенными внутренними свойствами



Объект, обладающий горизонтом.

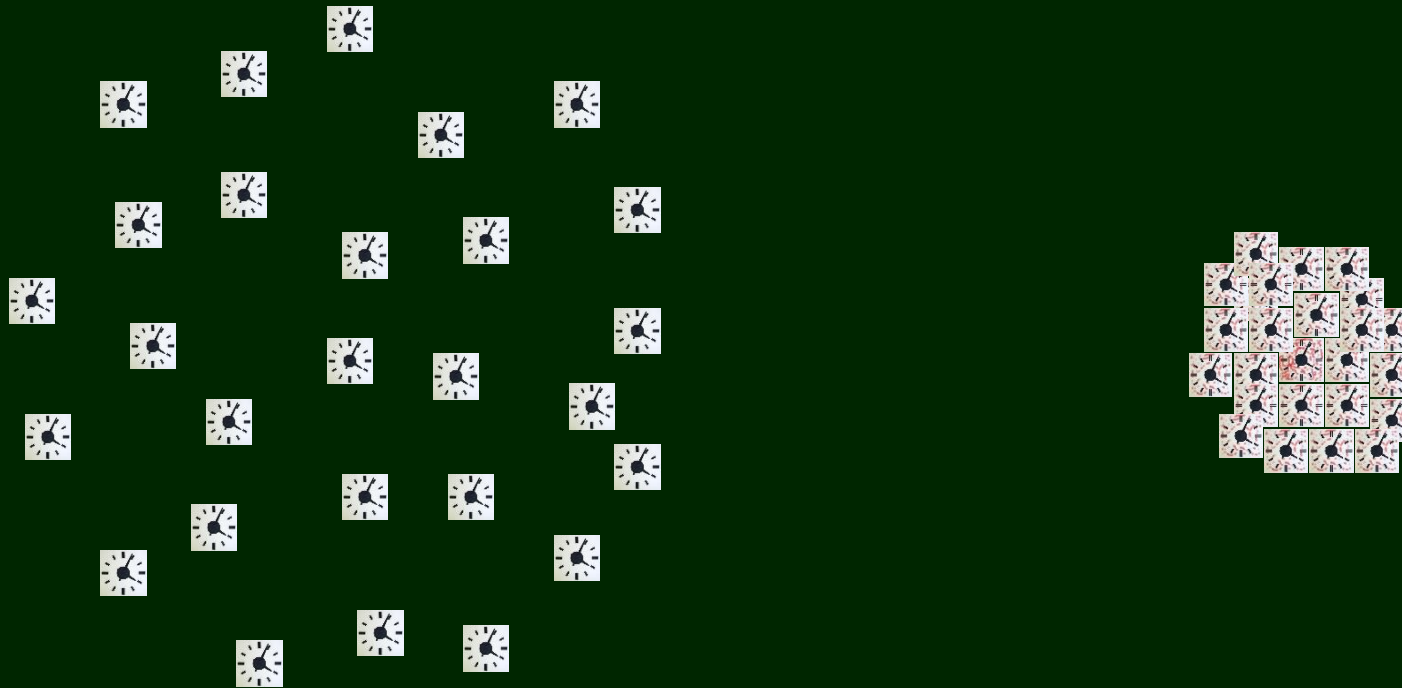
## Для астронома

Обладает определенными внешними проявлениями



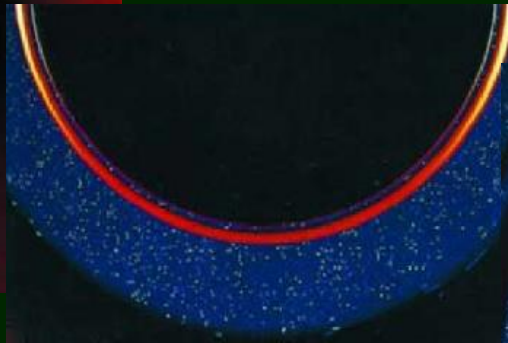
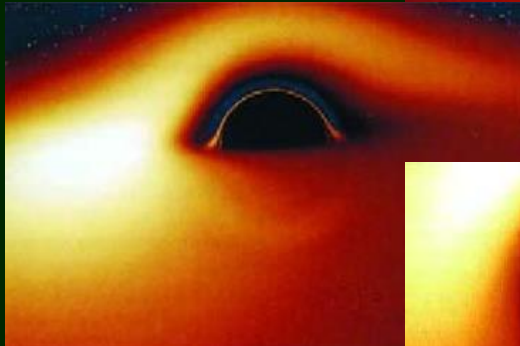
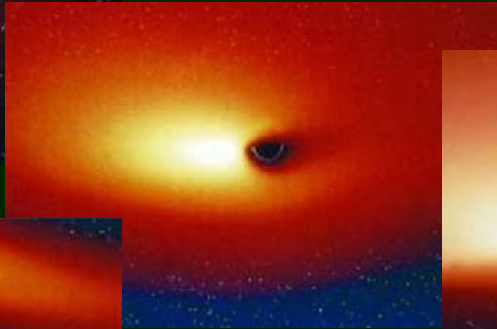
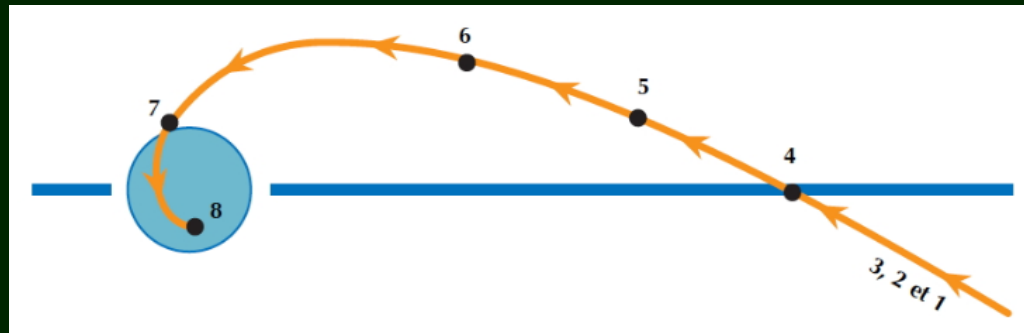
Компактное (размер горизонта) массивное тело, не проявляющее признаков наличия поверхности, и чьи недра недоступны для наблюдений.

# Коллапс облака

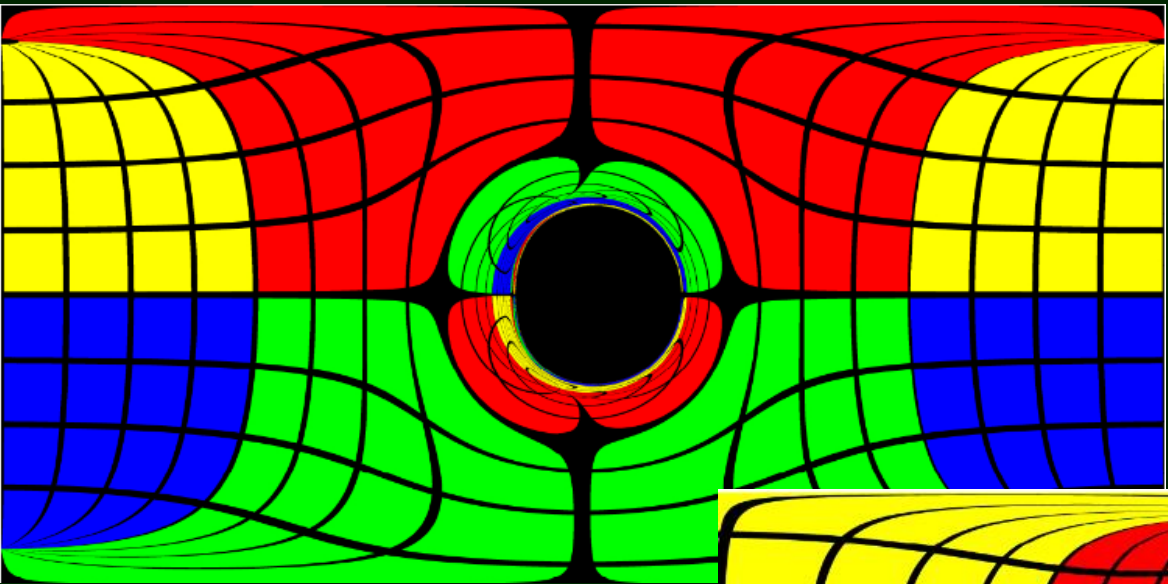


Мы всегда видим часы в центре,  
но они все краснее и краснее...

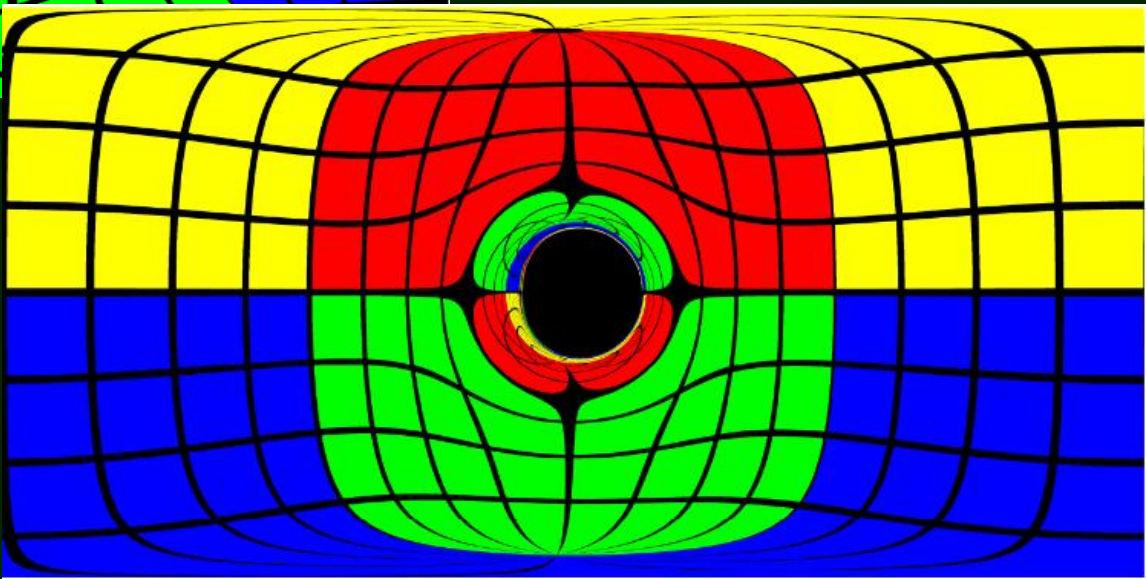
# Falling into a BH



# BH virtual reality



Observers at  $10 R_g$ .  
Left: at rest.  
Right (bottom): free-falling.



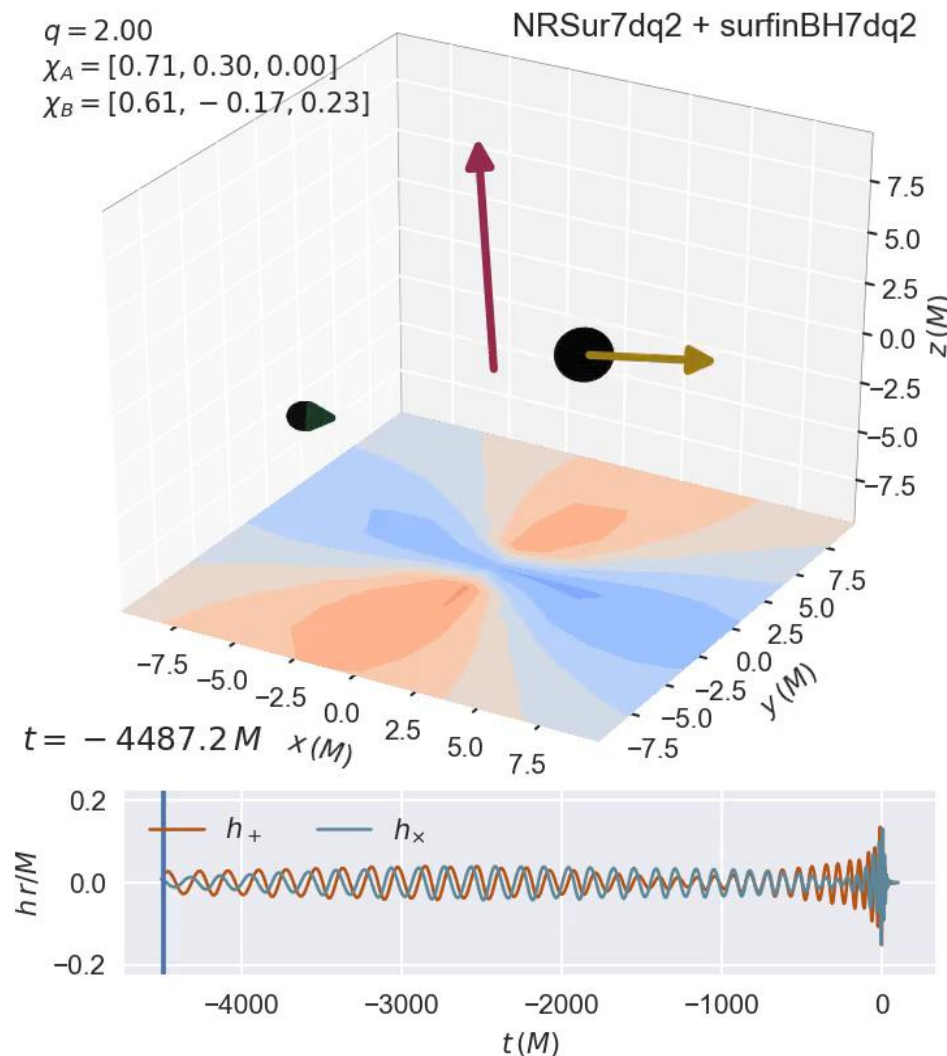
# Video for Sgr A\*



<https://www.youtube.com/watch?v=SXN4hvp977s>

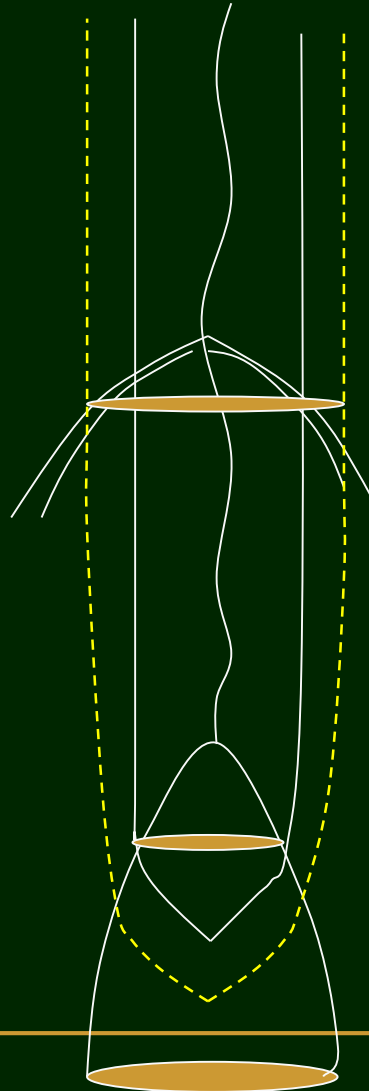
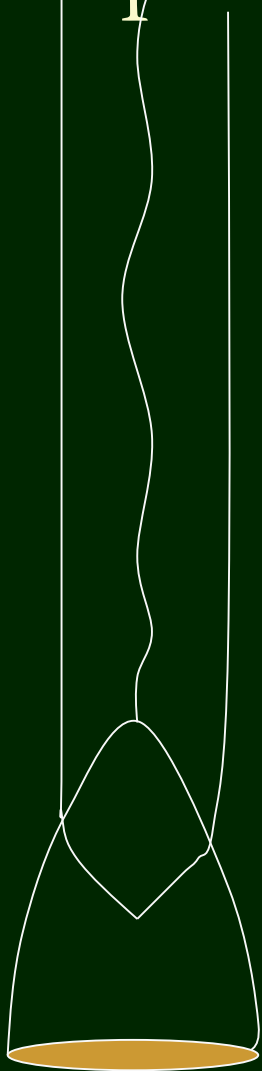
<https://blackholecam.org/>

# Binary BH visualization



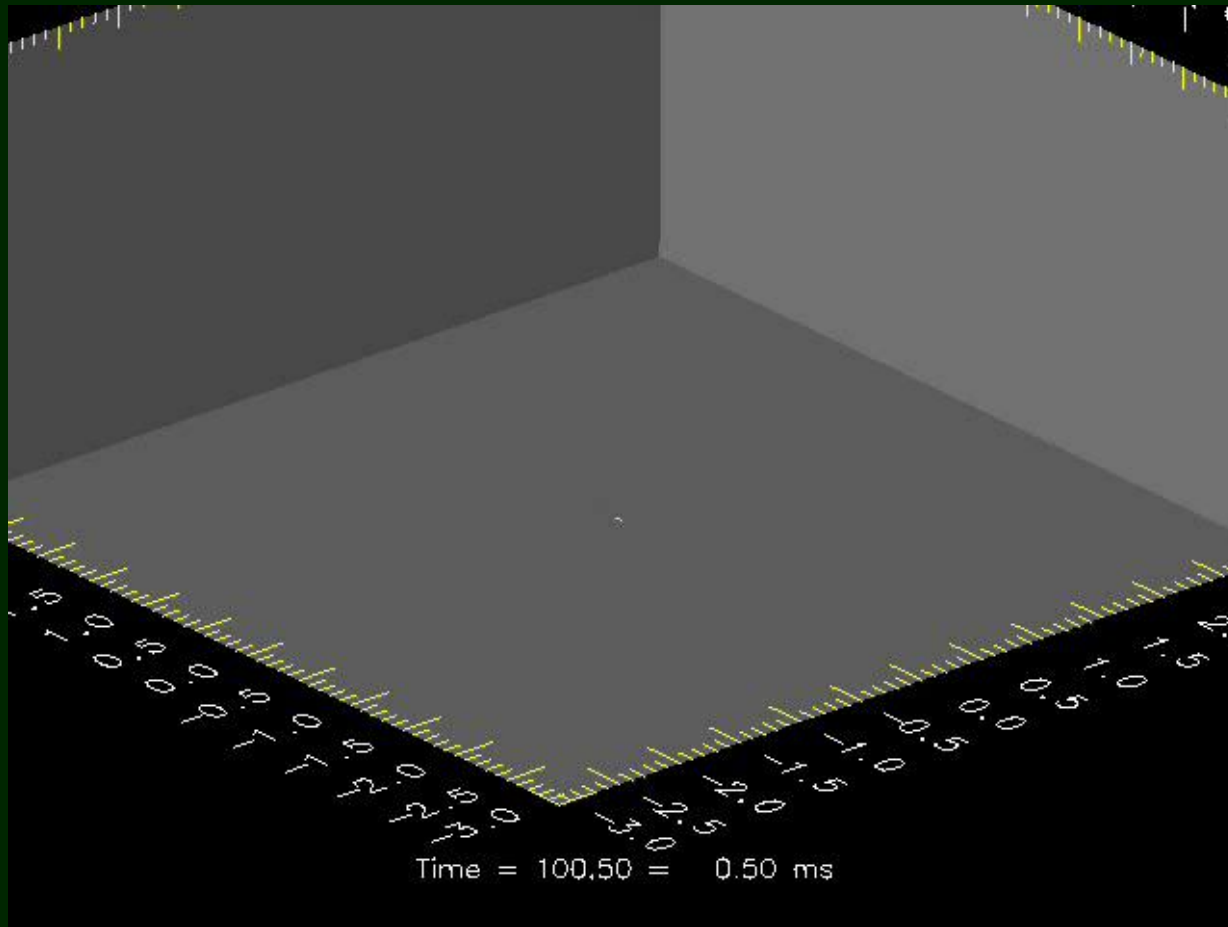
The black holes are shown as oblate spheres, with arrows indicating their spins. The orbital angular momentum is indicated by the pink arrow at the origin. The colors in the bottom-plane shows the value of the plus polarization of the GW as seen by an observer at that location; red means positive and blue means negative, notice the quadrupolar pattern of the radiation. In the subplot at the bottom, we show the plus and cross polarizations as seen from the camera viewing angle.

# Горизонты

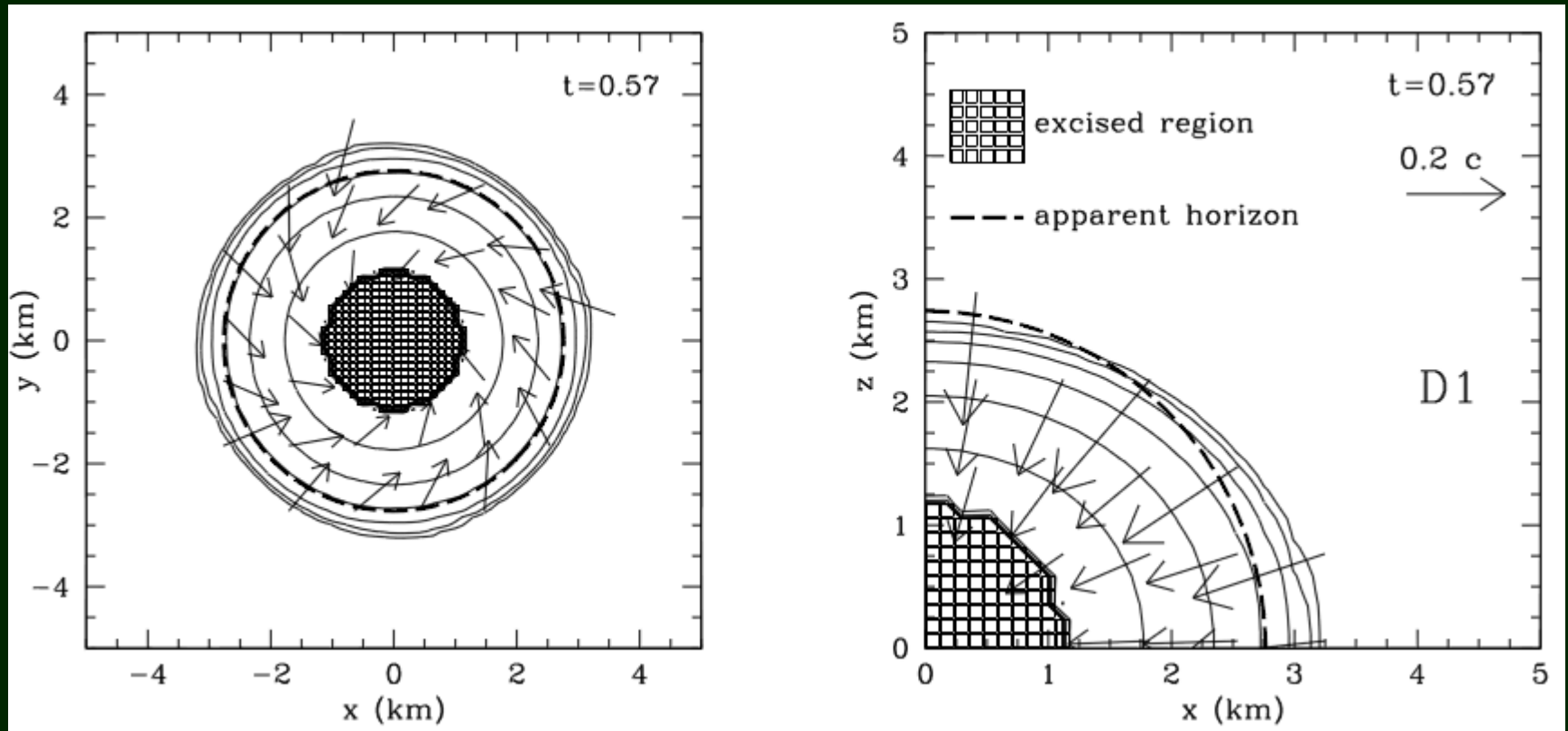


Теорема об отсутствии волос:  
«Черные дыры не имеют волос».  
Черная дыра полностью описывается  
массой и вращением.

# Horizons appearance



# ΚΟΛΛΑΠΣ

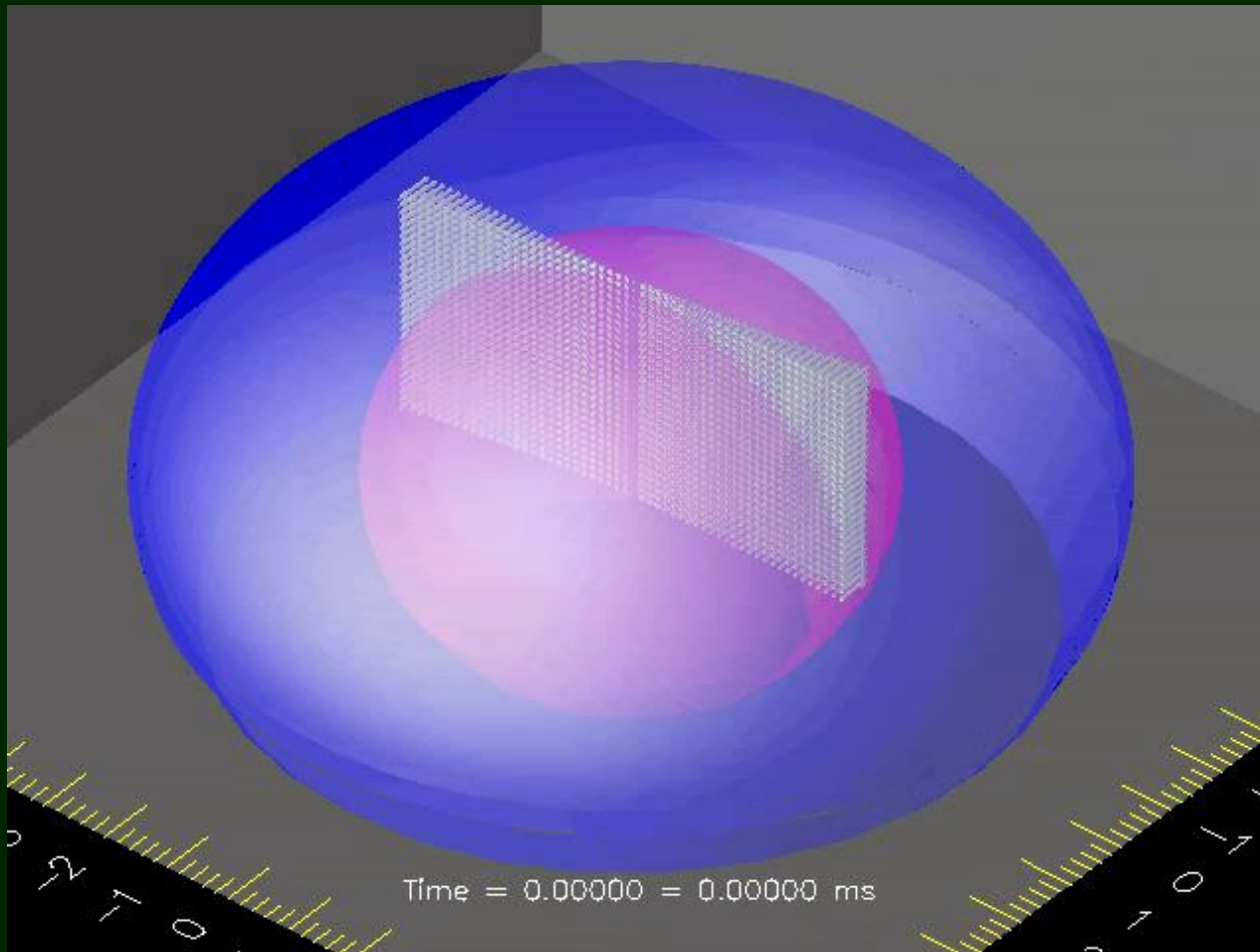


Apparent horizon position is calculated at every time step.

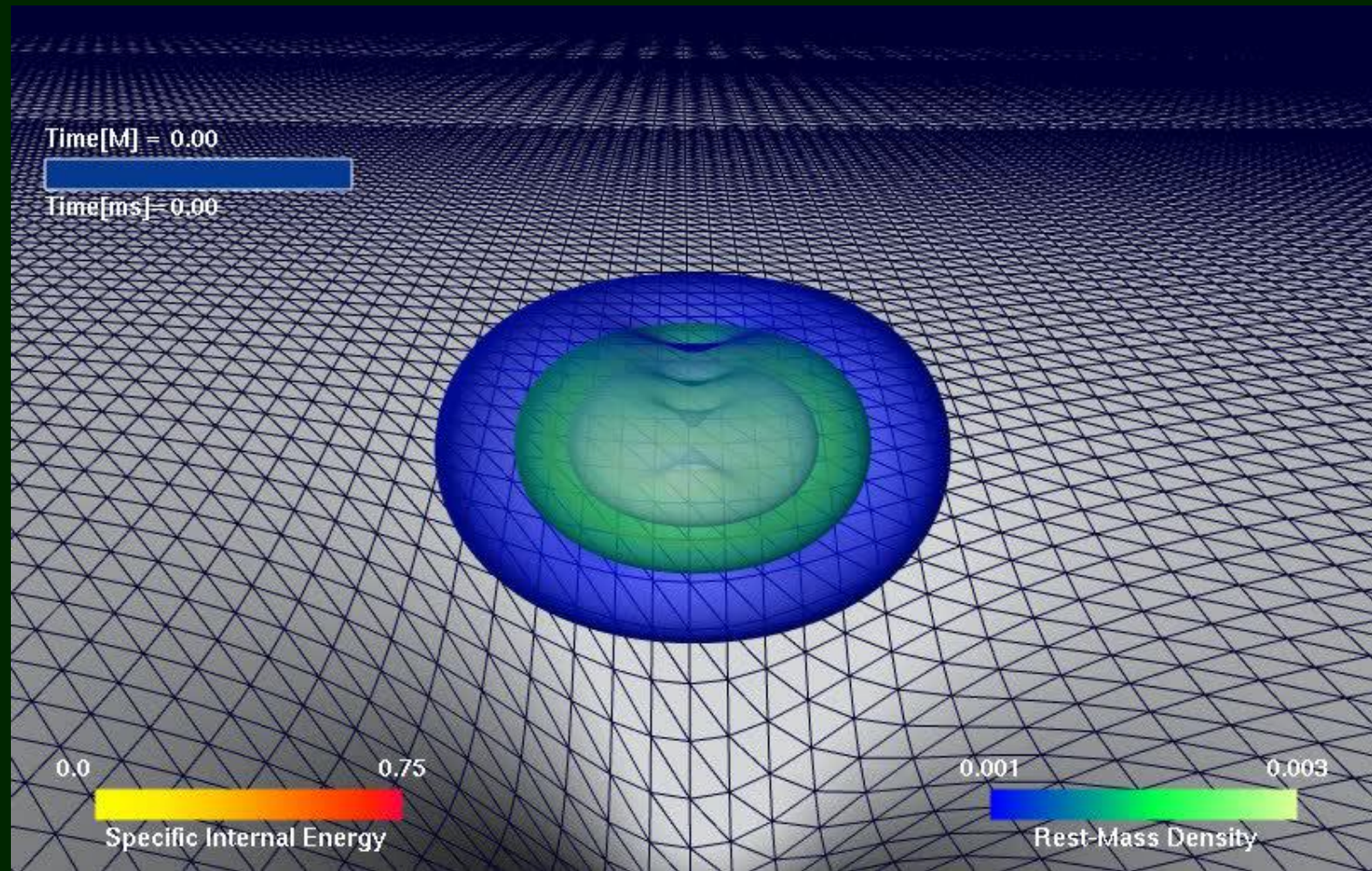
The event horizon (which is growing from zero to its final position and is always outside the apparent horizon) is calculated a posteriori, i.e. after calculations are finished.

*L. Baiotti, Rezzolla et al.*

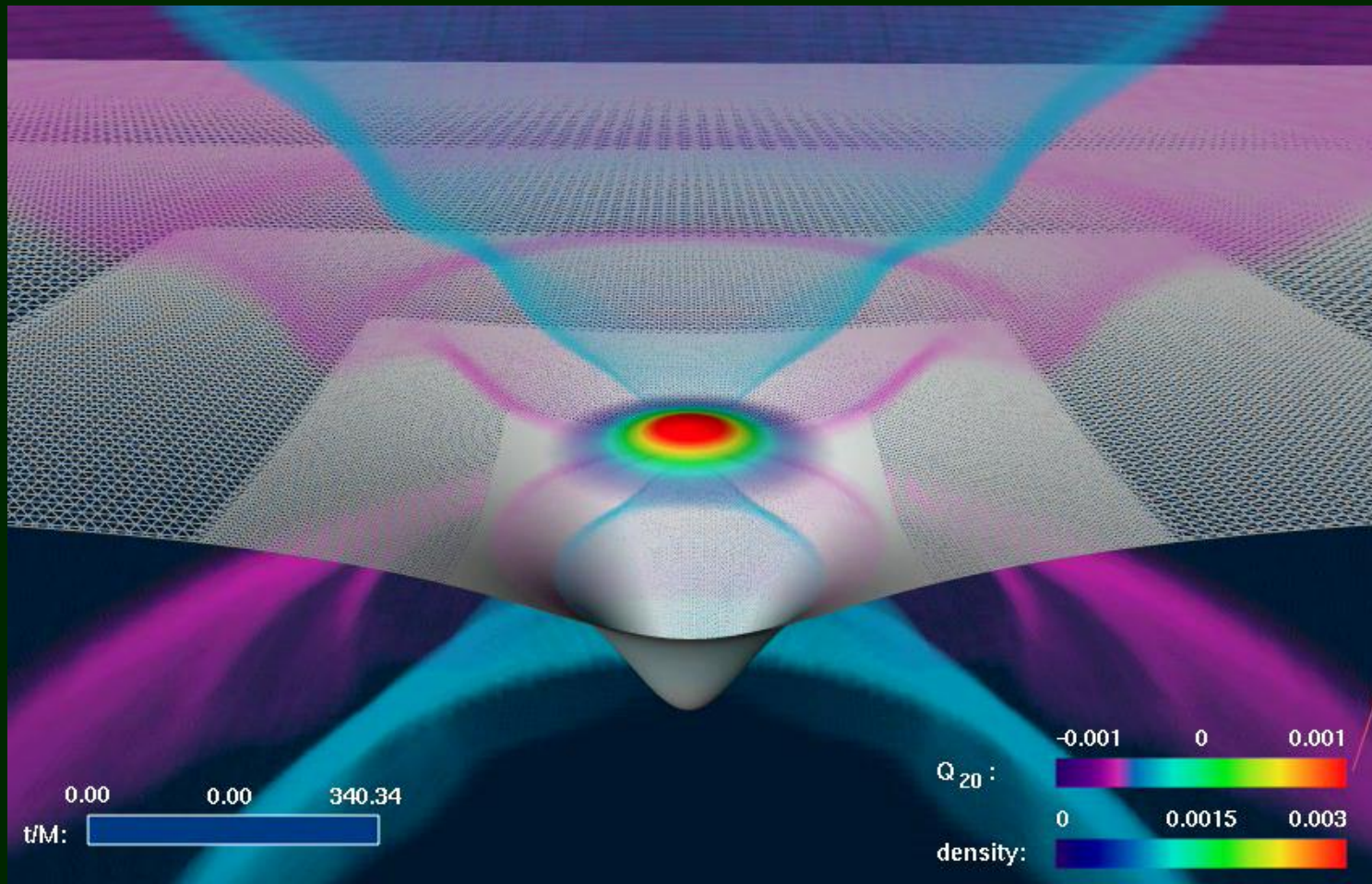
gr-qc/0403029



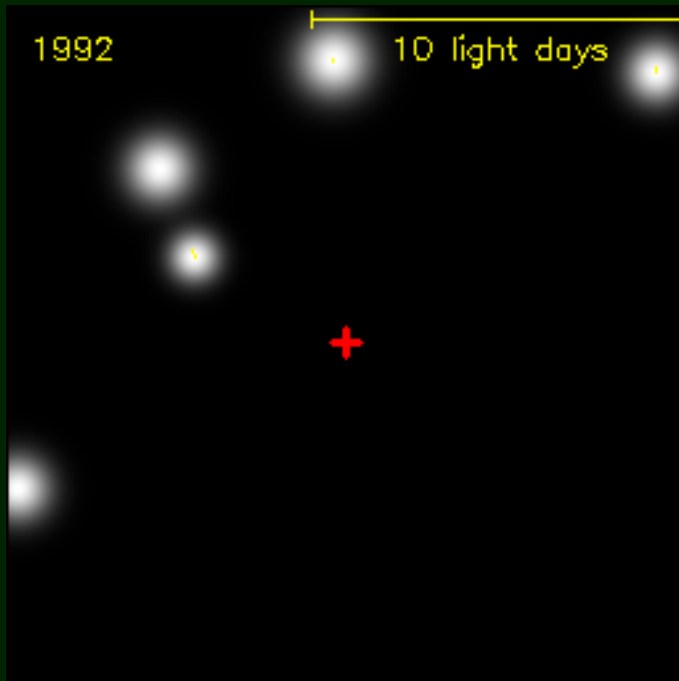
# Collapse of a rotating star



# Collapse and GW emission

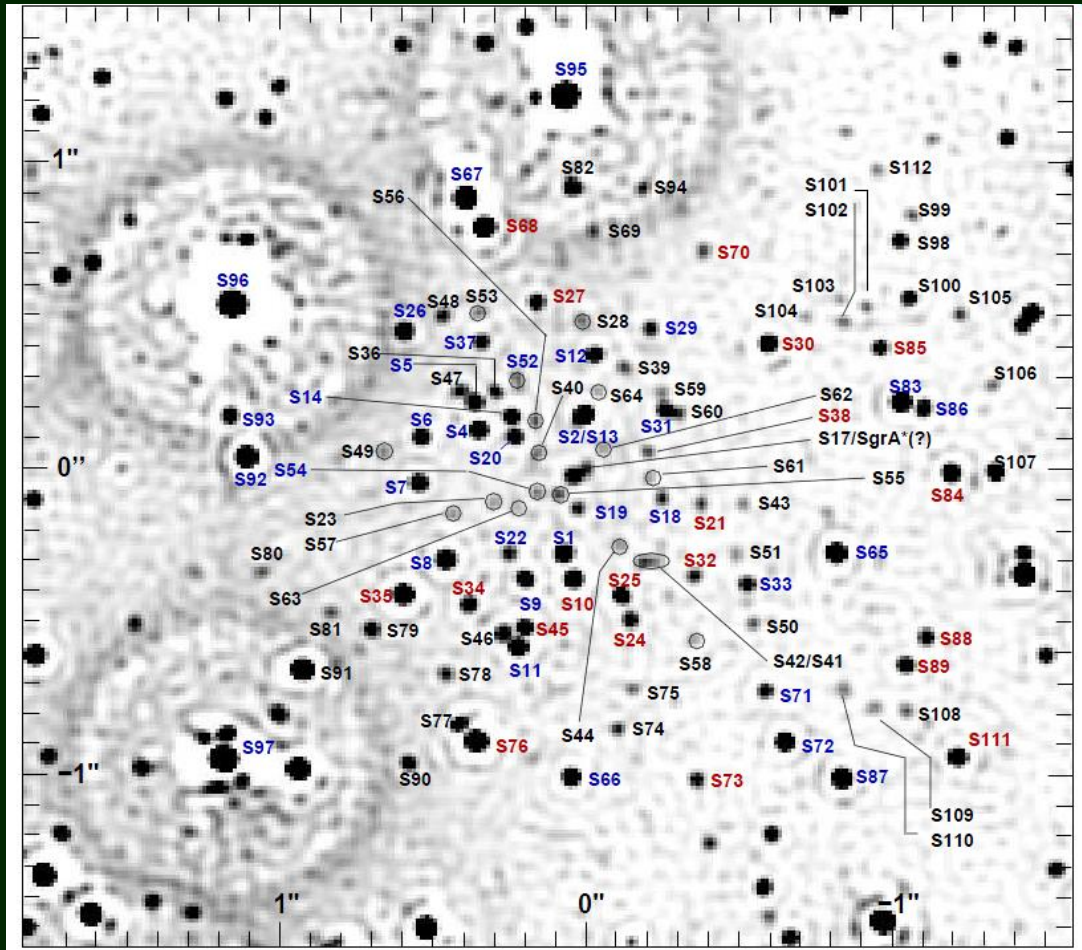


# The most certain BH – Sgr A\*



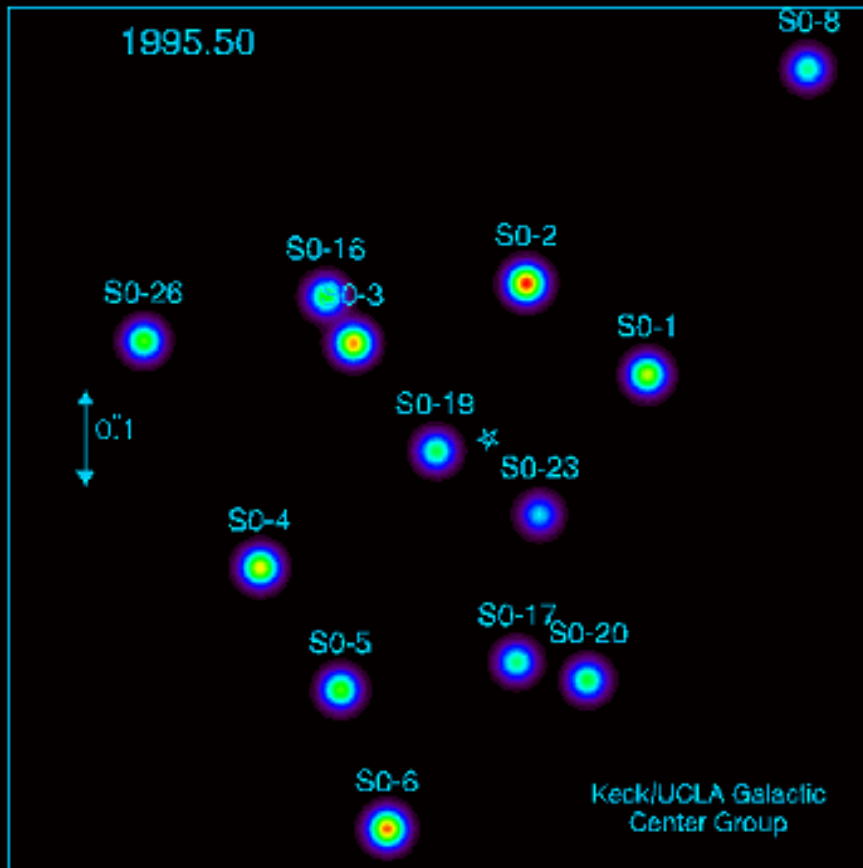
Stellar orbits from 1992 till 2007

(see the reference  
in gr-qc/0506078)



arXiv: 0810.4674

... and it becomes more and more certain

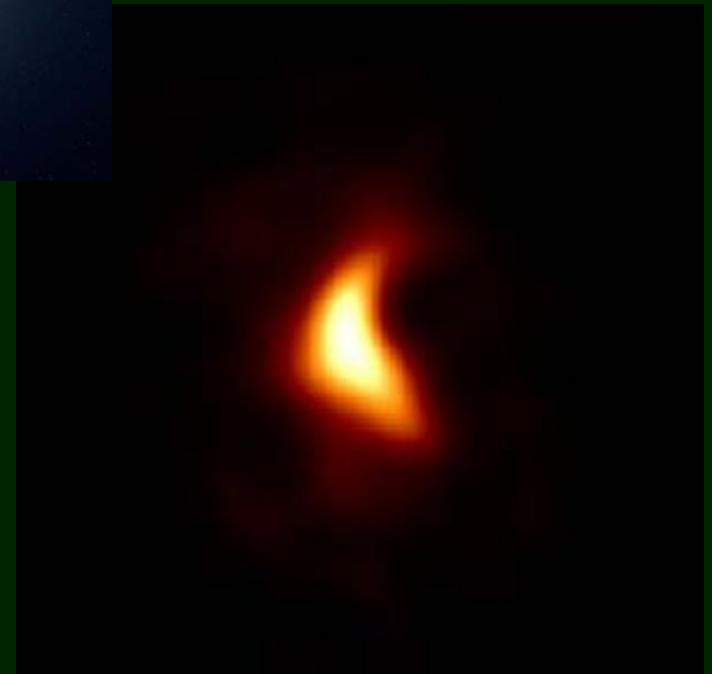
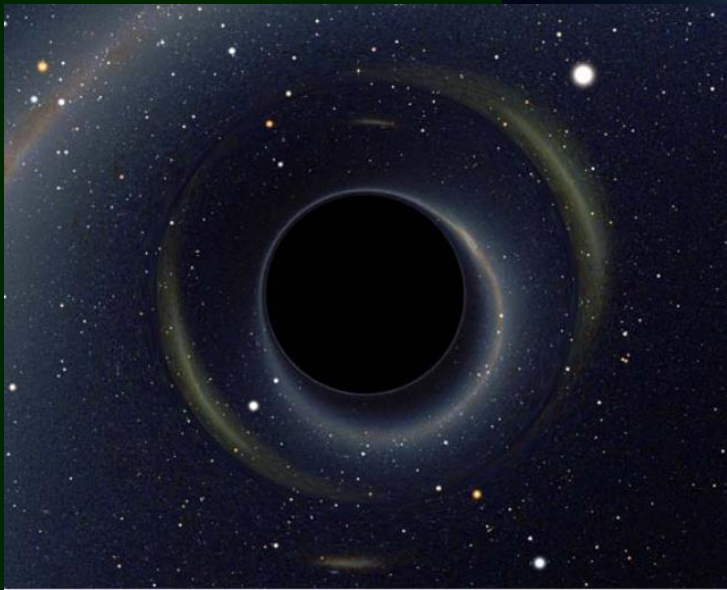
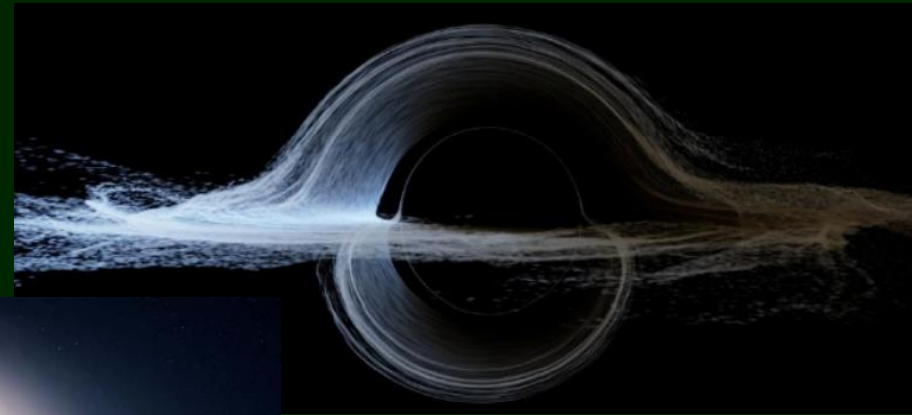
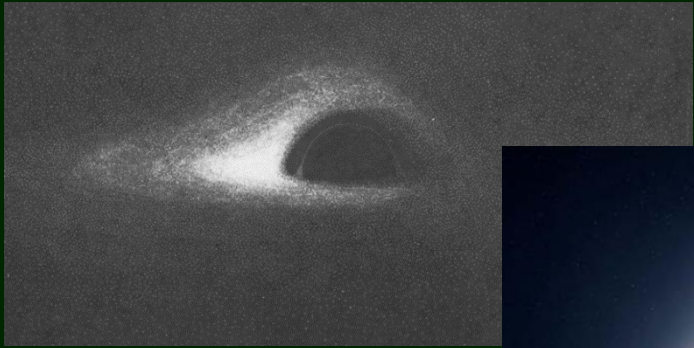


Observations are going on.  
So, the number of stars with  
well measured orbits grows.

$$M_{\text{BH}} \sim 4\text{--}5 \times 10^6 M_{\text{solar}}$$

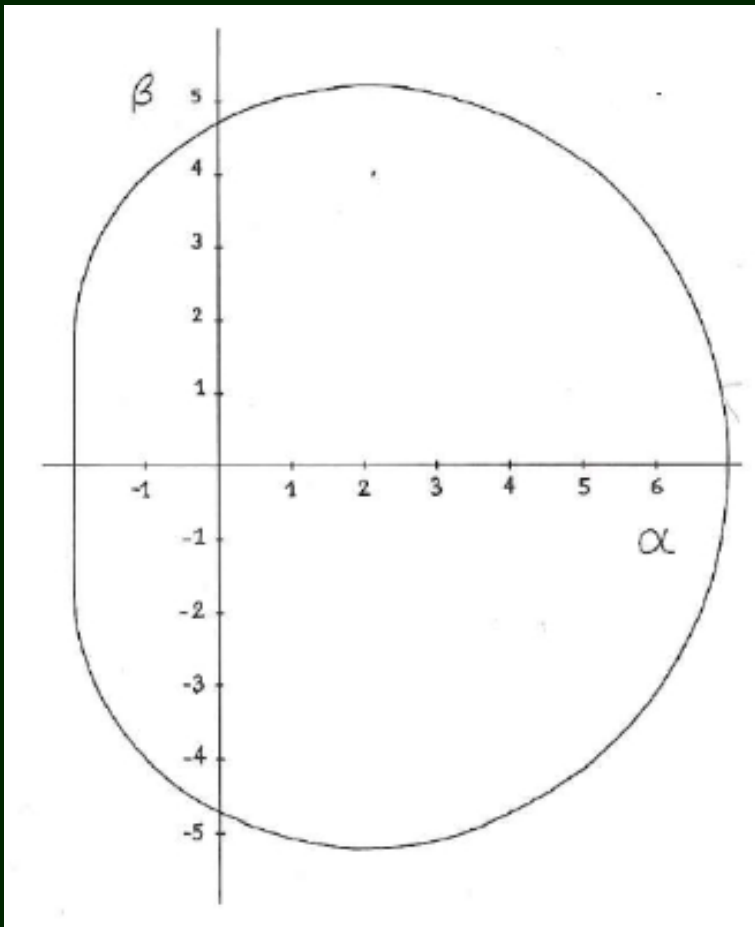
See the reference in [gr-qc/0506078](https://arxiv.org/abs/gr-qc/0506078)  
New data in [arXiv: 0810.4674](https://arxiv.org/abs/0810.4674)  
Recent review - [1311.1841](https://arxiv.org/abs/1311.1841)

# BH visualization



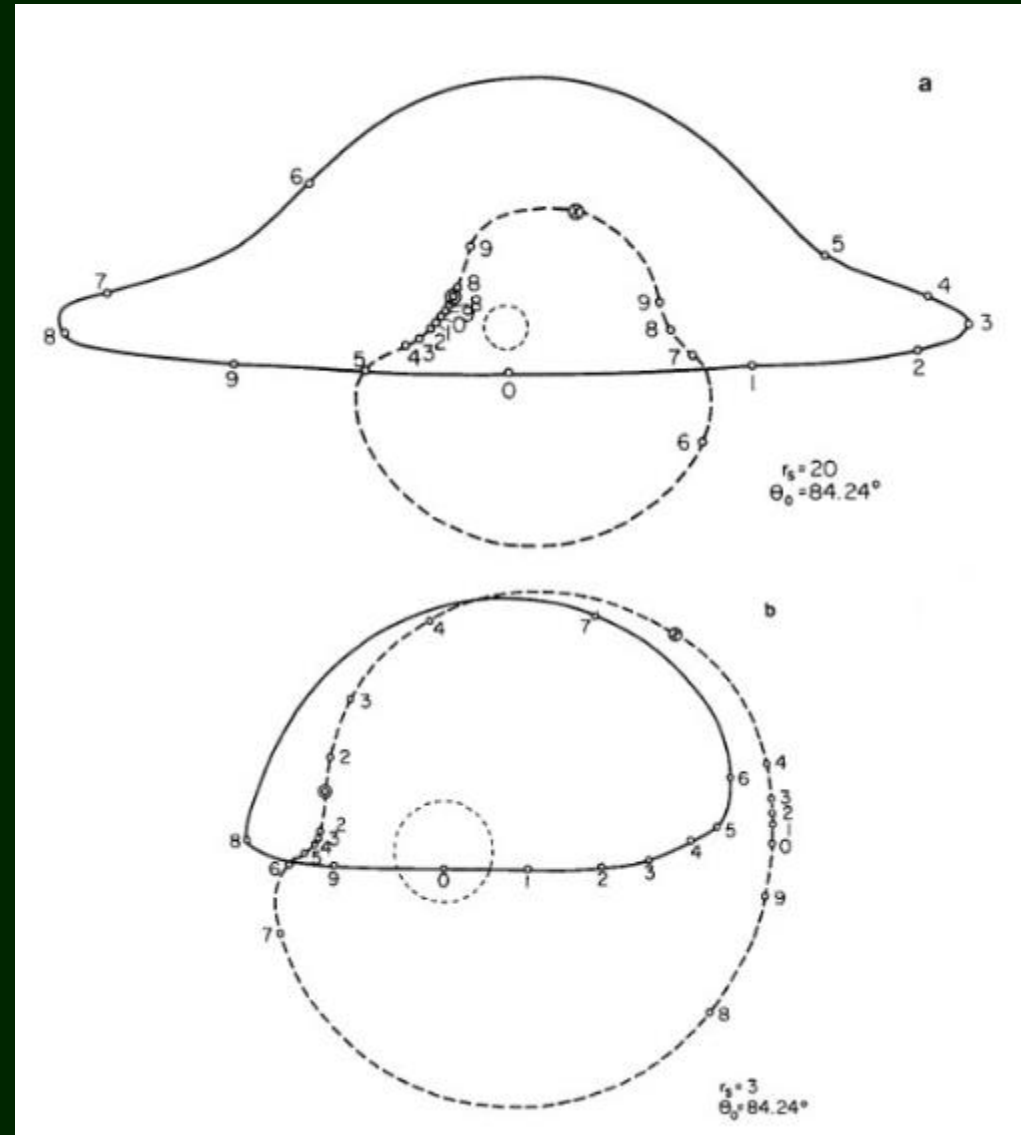
1804.03909, see also 1902.11196

# First calculations



Both for an extremally rotating BH

Cunningham, Bardeen 1973



See a review in 1902.11196

# Supernovae

Schematic representation of the evolutionary stages from stellar core collapse through the onset of the supernova explosion to the neutrino-driven wind during the neutrino-cooling phase of the proto-neutron star.

The horizontal axis gives mass information.

$M_{\text{hc}}$  is the mass of the subsonically collapsing, homologous inner core.

The vertical axis shows corresponding radii.

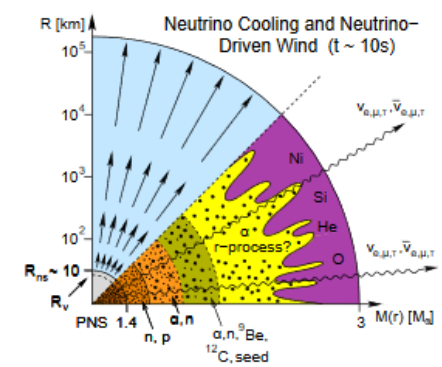
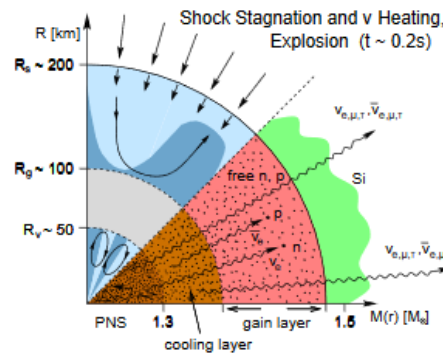
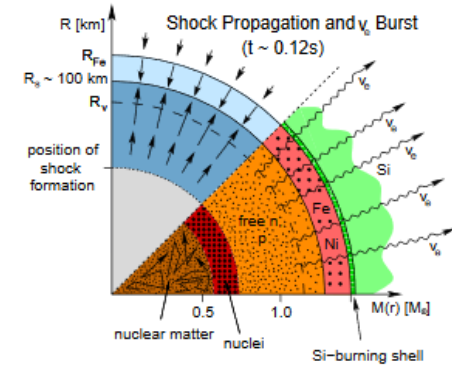
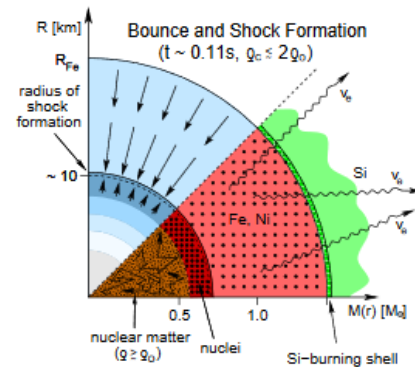
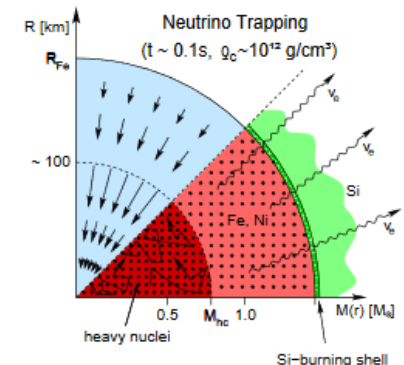
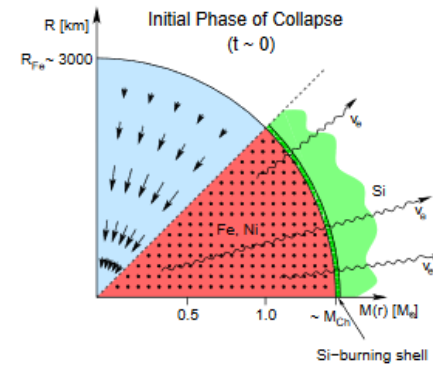
$R_{\text{Fe}}$  - iron core radius;

$R_s$  - shock radius;

$R_g$  - gain radius;

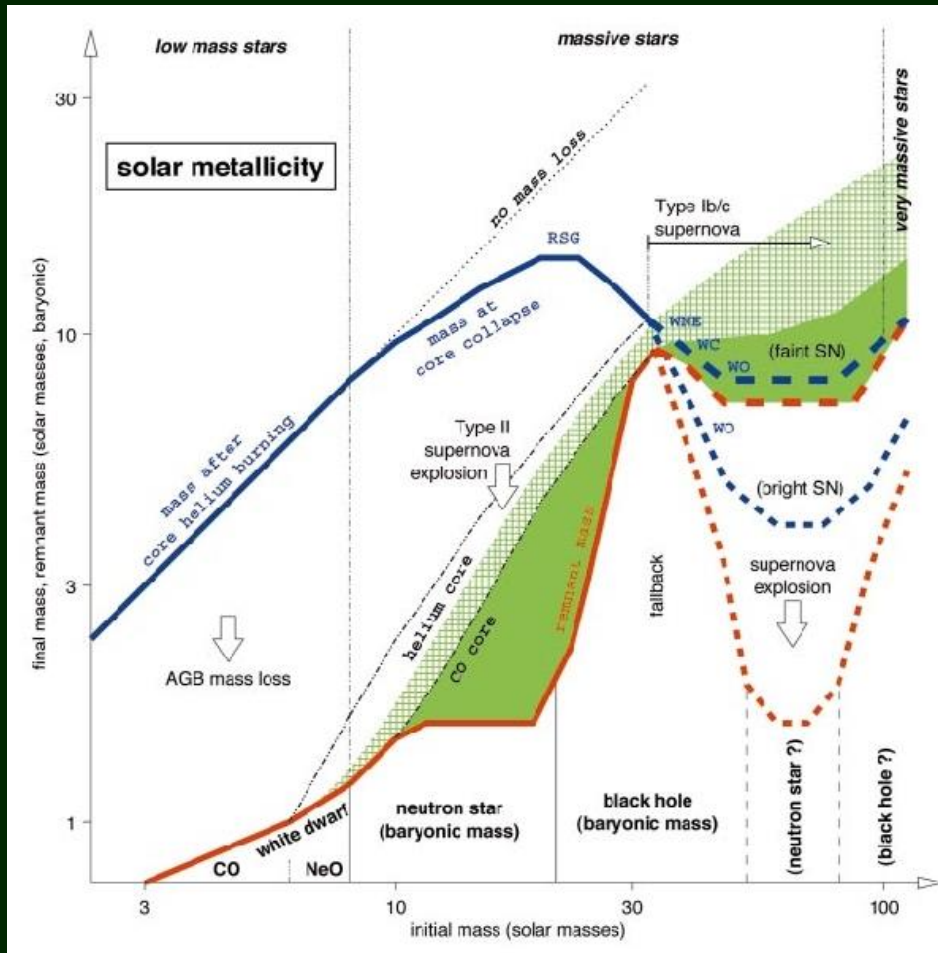
$R_{\text{ns}}$  - neutron star radius;

$R_v$  - neutrinosphere.



# Stellar mass BHs.

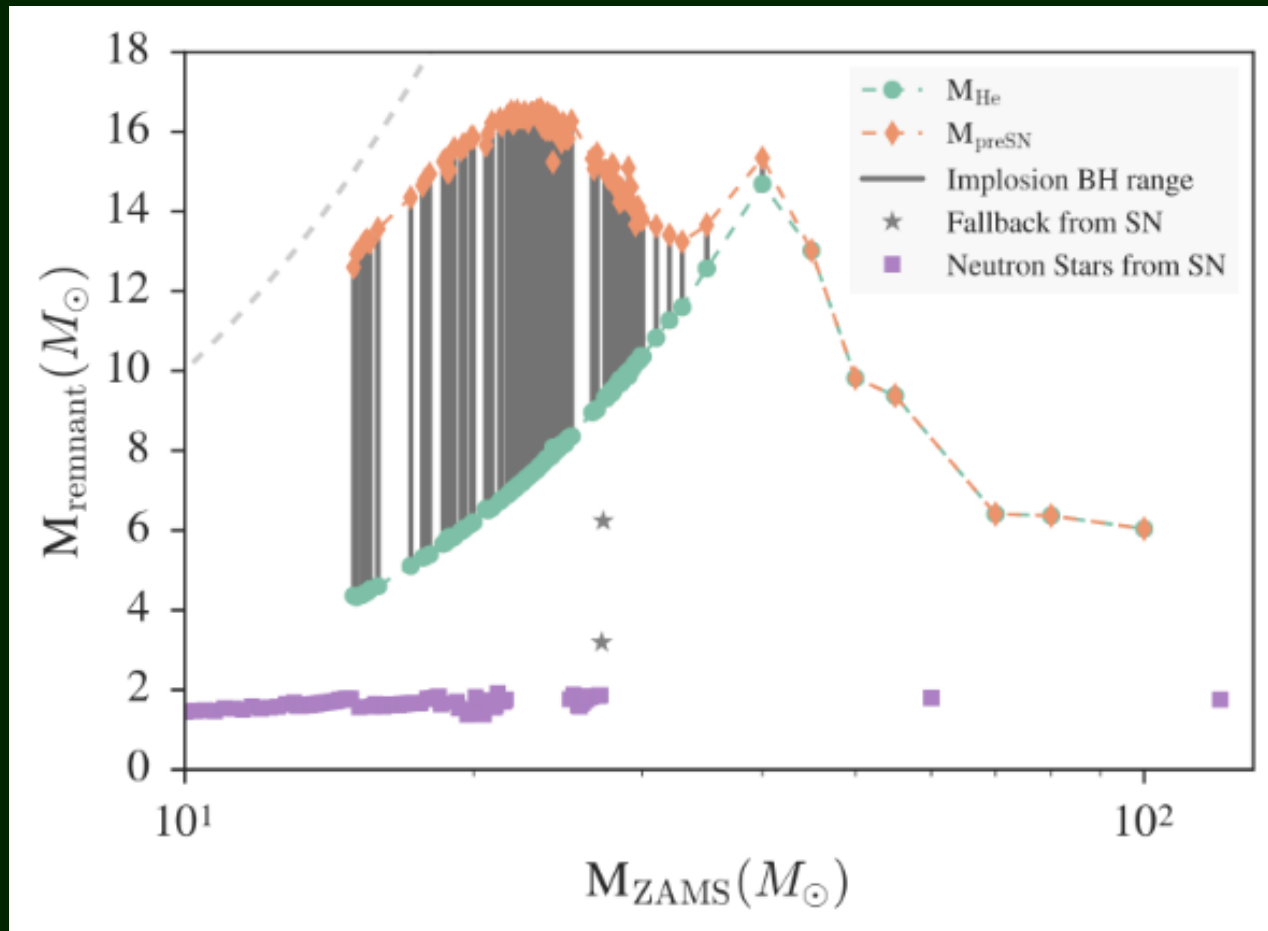
## The case of solar metallicity.



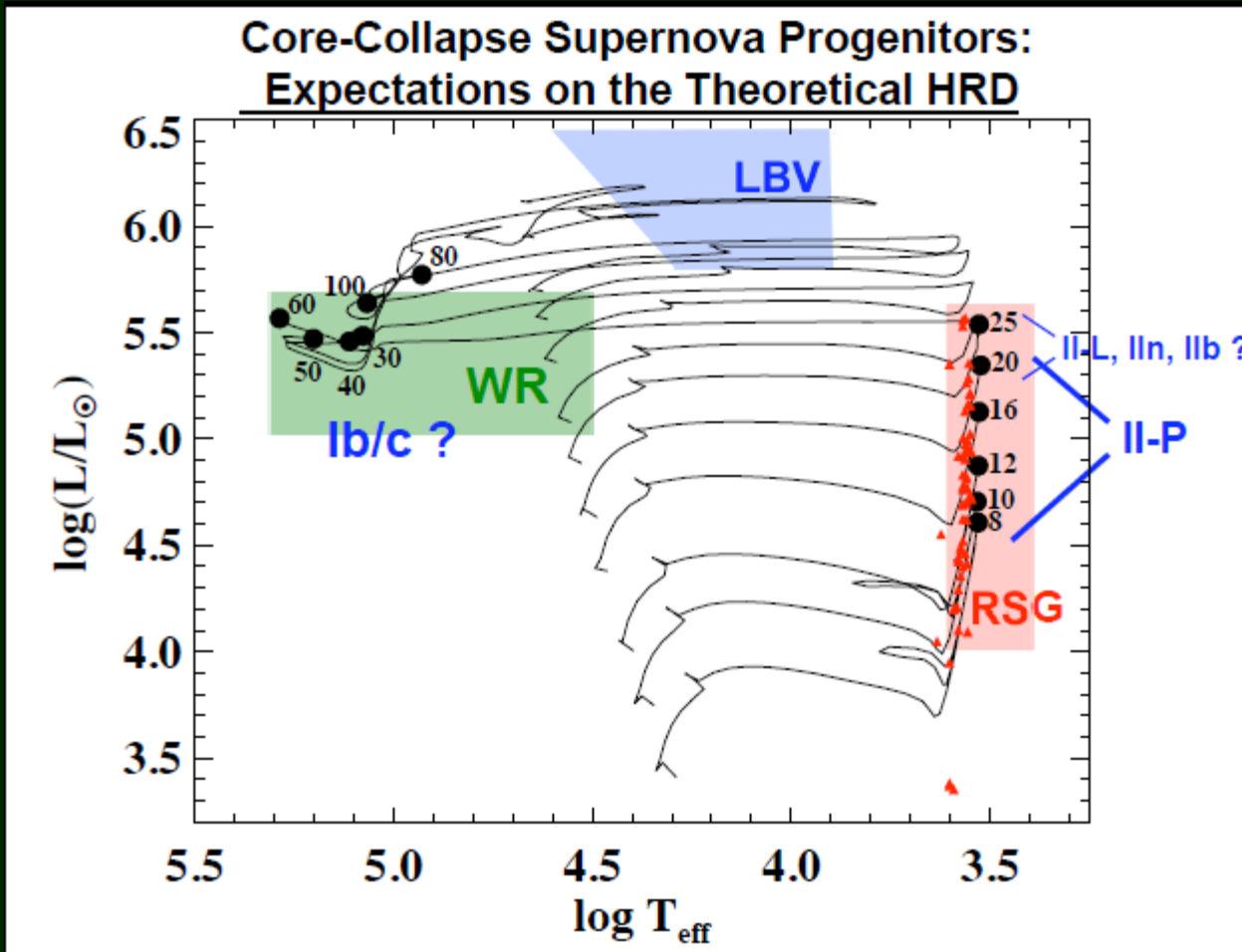
BHs are formed by massive stars. The limiting mass separating BH and NS progenitors is not well known. In addition, there can be a range of masses above this limit in which, again NSs are formed (also, there can be a range in which both types of compact objects form).

See 1011.0203 about progenitors

# Initial mass vs. final: ZAMS vs. compact object

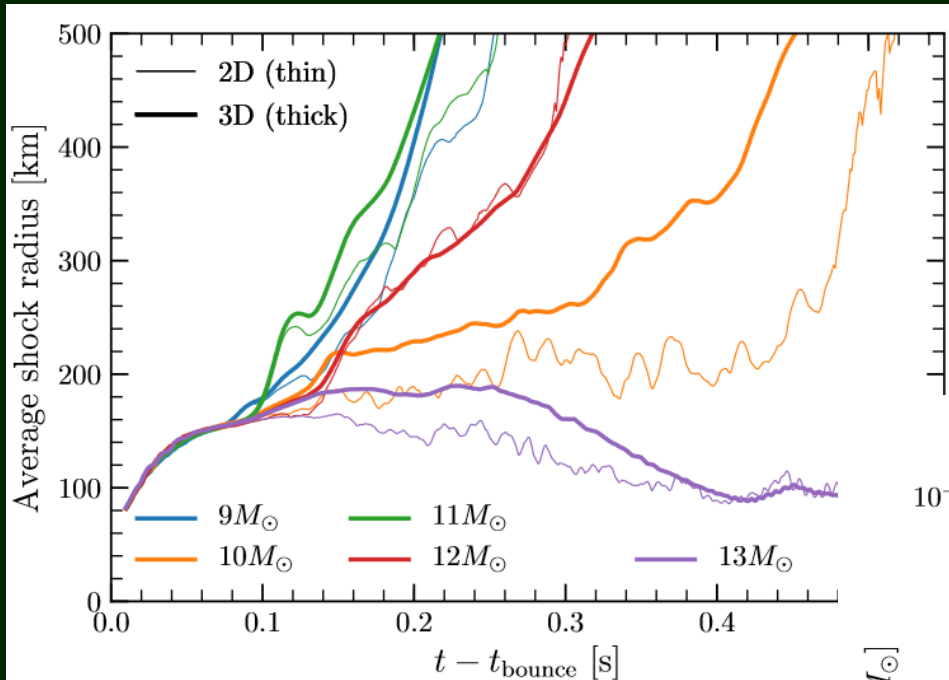


# Supernova progenitors

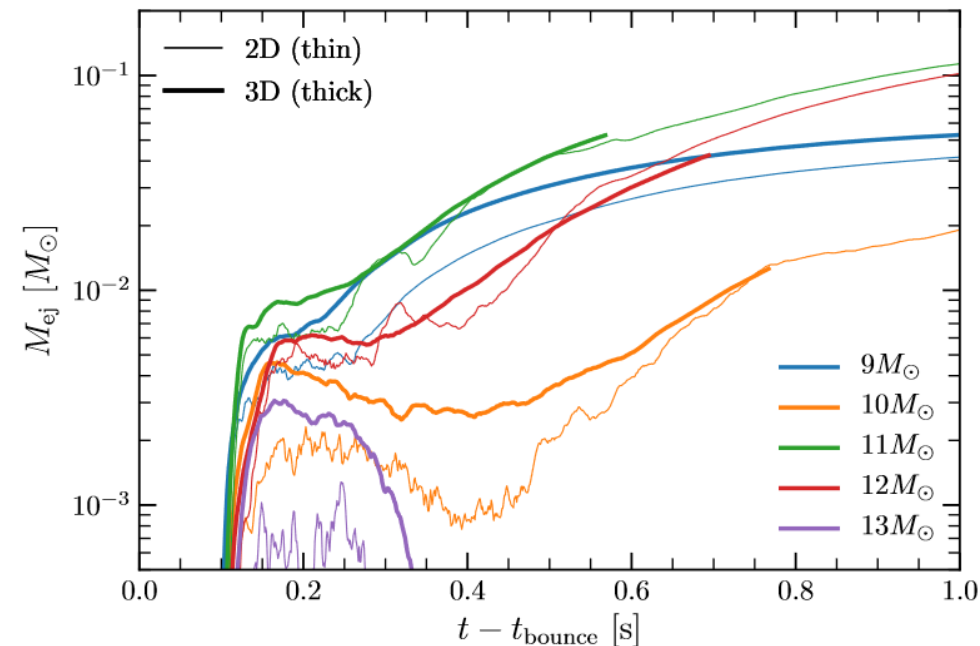


However, there are claims that most of stars  $>18M_{\odot}$  produce BHs (see a review in Smartt arXiv: 1504.02635)

# No explosions for 13 solar mass?

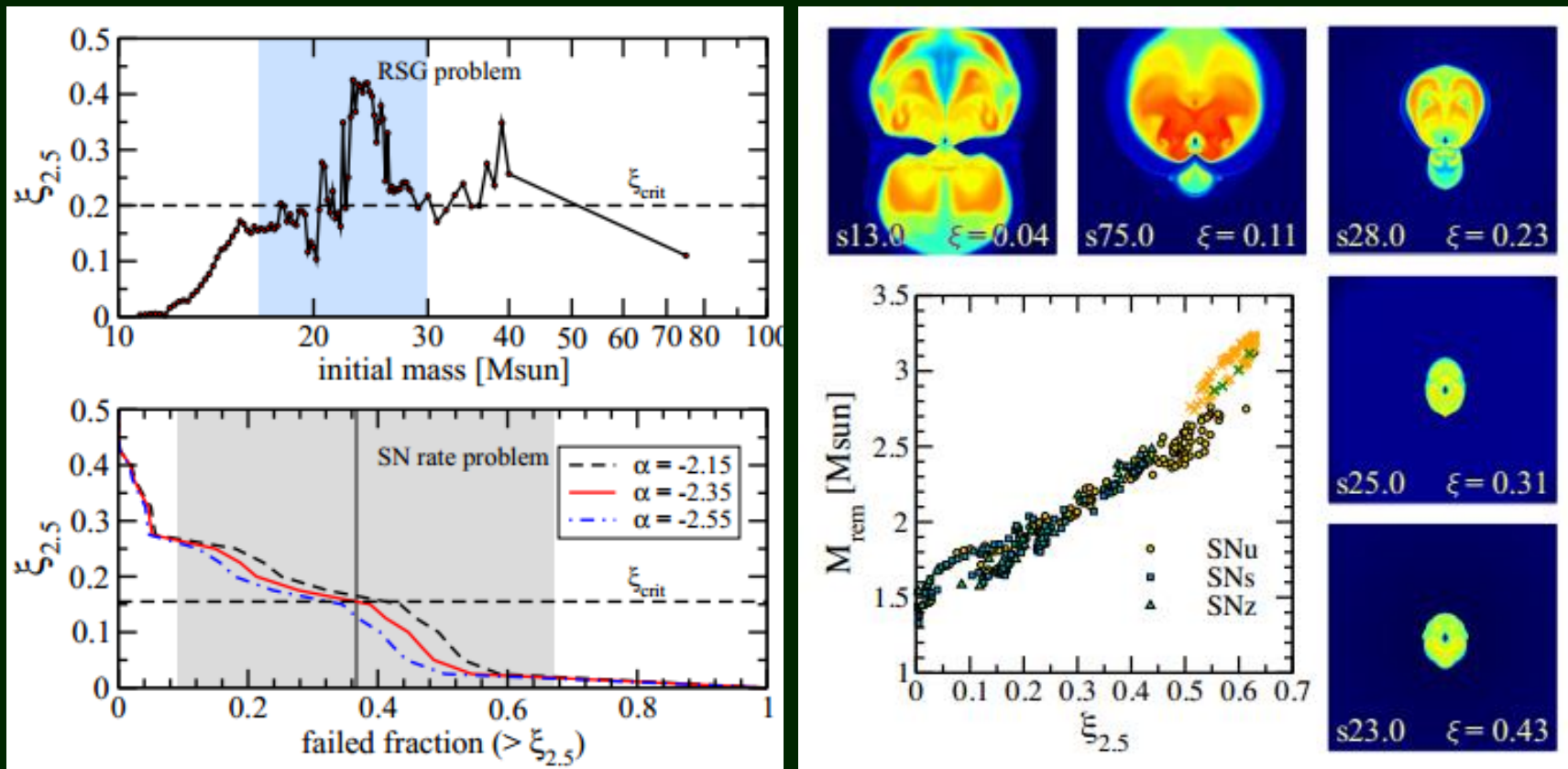


Probably non-rotating progenitors with  $12 < M < 14$  solar mass do not explode.



# Which stars form BHs?

It is proposed that stars with compact internal structure ( $M \sim 20\text{-}30$  Msolar) form BHs not NSs. This explains data on RGs and the SN rate.



# Transients from BH formation

Supergiant progenitors of BHs can have huge convective envelopes.

Convective motions in the outer parts of supergiants generate mean horizontal flows at a given radius with velocities of  $\sim 1$  km/s.

Failed explosions of supergiants - in which the accretion shock onto the neutron star does not revive, leading to black hole formation - may often produce accretion discs that can power day-week (blue supergiants) or week-year (yellow and red supergiants) non-thermal and thermal transients through winds and jets.

These transients will be especially time variable because the angular momentum of the accreting material will vary substantially in time. Observed sources such as Swift J1644+57, iPTF14hls, and SN 2018cow, as well as energetic Type II supernovae (OGLE-2014-SN-073) may be produced by this mechanism.

$$v_h \sim \frac{v_c}{\sqrt{4\pi}} \frac{H}{r}$$

$$j_{\text{rand}} \sim \frac{H v_c}{\sqrt{4\pi}} \sim 6 \times 10^{17} \frac{r v_c}{10^3 R_\odot \text{ km s}^{-1}} \left( \frac{H/r}{0.3} \right) \text{ cm}^2 \text{ s}^{-1}$$

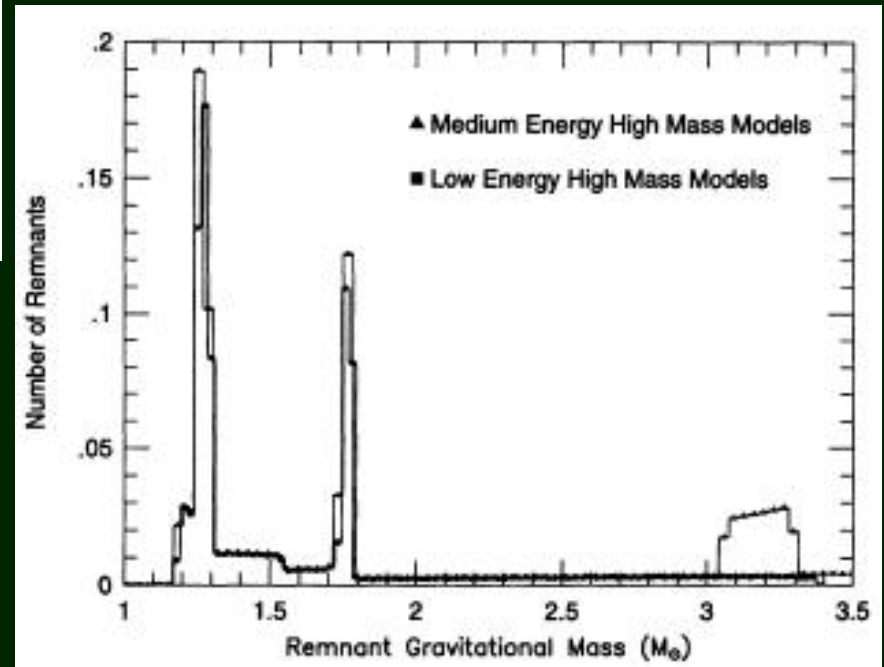
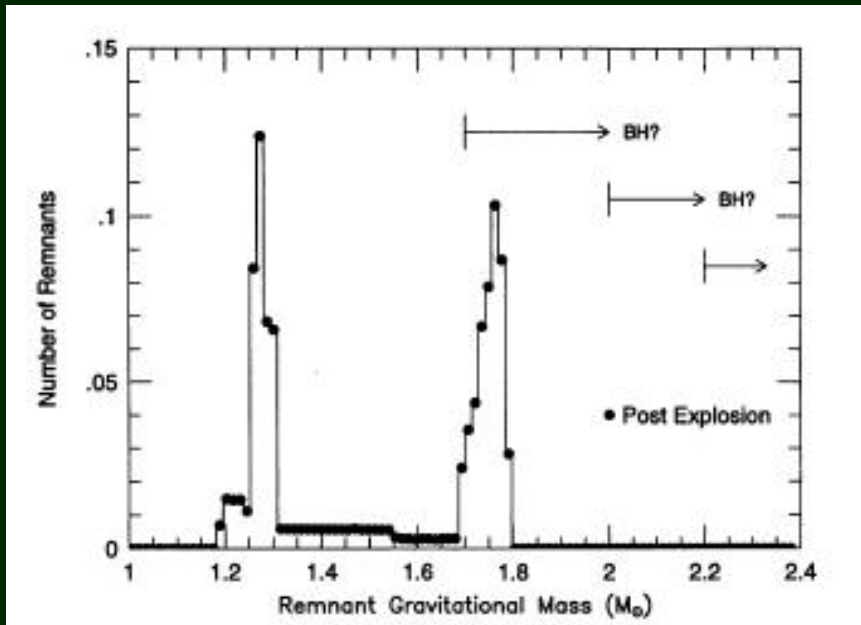
$v_c$  ( $\sim 10$  km/s) –  
convective velocity

$v_h$  – horizontal velocity

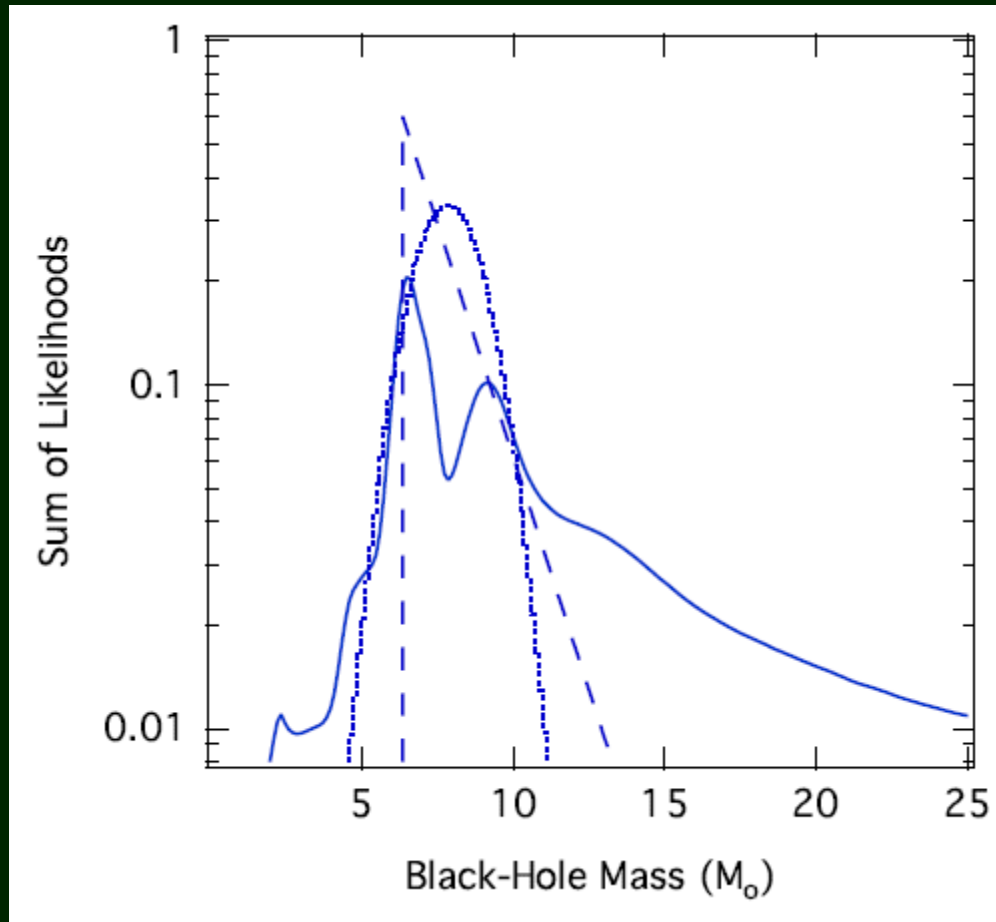
$$j_{\text{ISCO}} = 1.15 - 3.5 \frac{GM}{c} \sim 0.5 - 1.5 \times 10^{17} \left( \frac{M}{10 M_\odot} \right) \text{ cm}^2 \text{ s}^{-1}$$

# Mass spectrum of compact objects

Results of numerical models

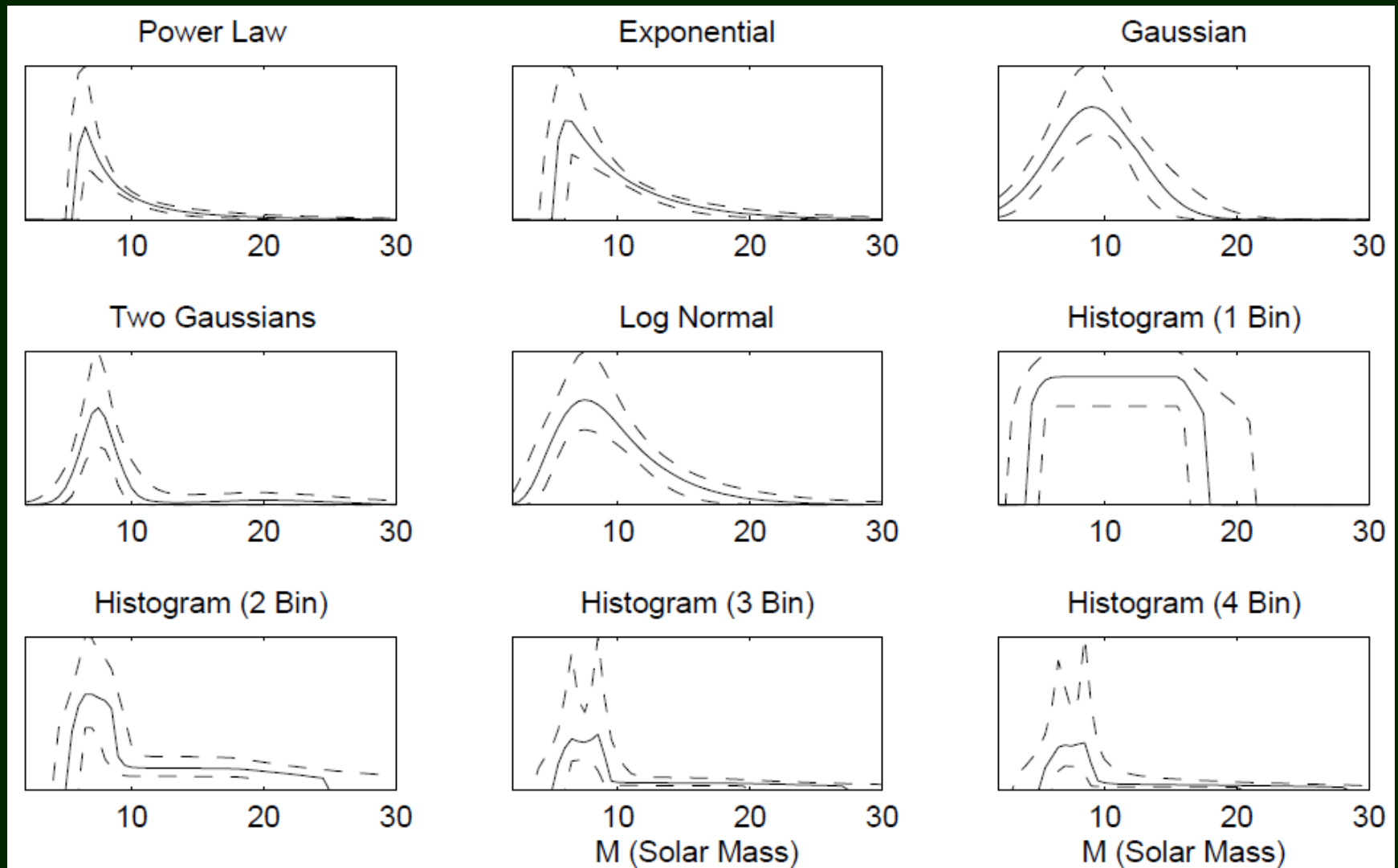


# BH mass function

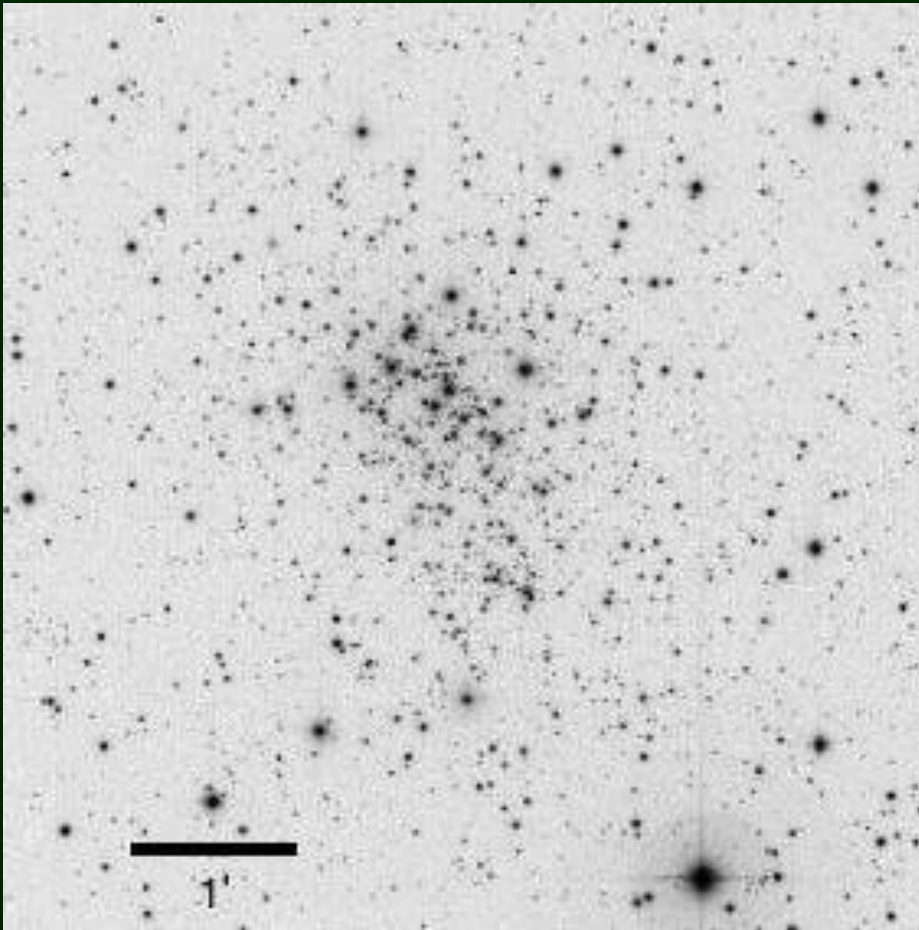


Likelihood based on 16 systems

# BH mass distribution



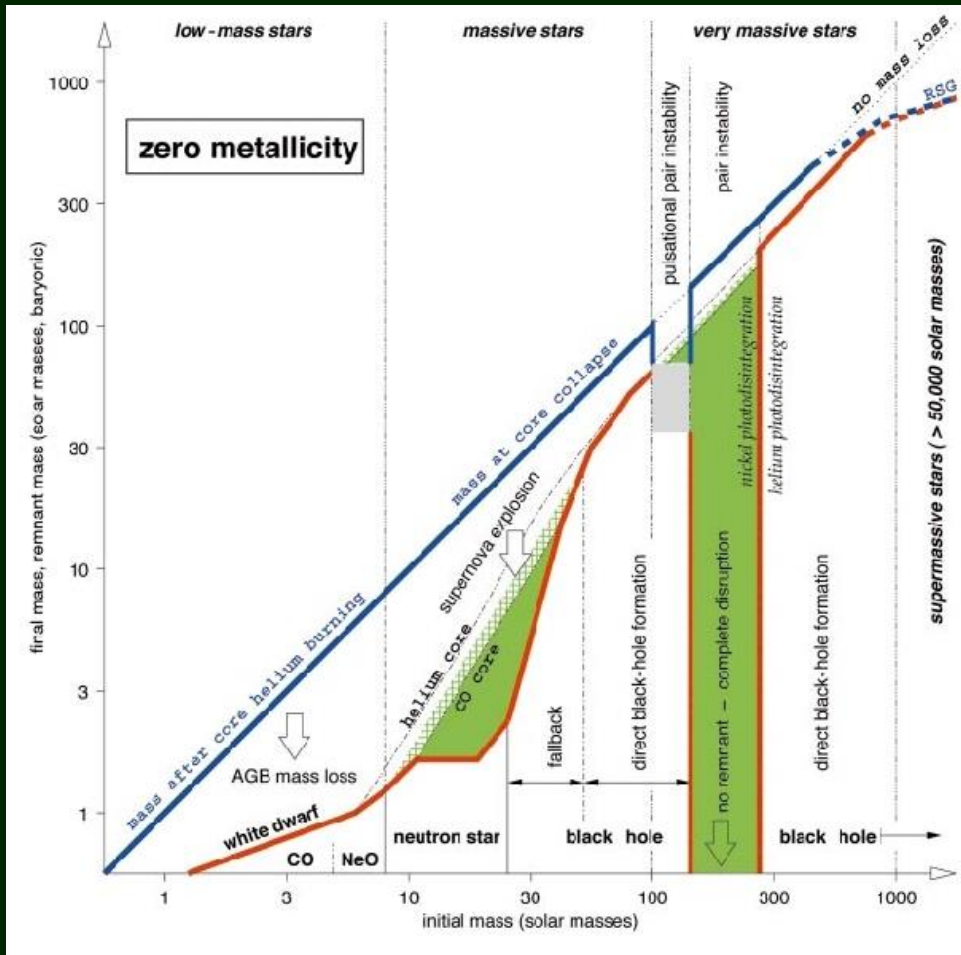
# A NS from a massive progenitor



Anomalous X-ray pulsar in the cluster Westerlund1 most probably has a very massive progenitor,  $>40 M_{\odot}$ .

# Stellar mass BHs.

## The case of zero metallicity



Pop III massive stars could produce very massive BHs which became seeds for formation of supermassive BHs.

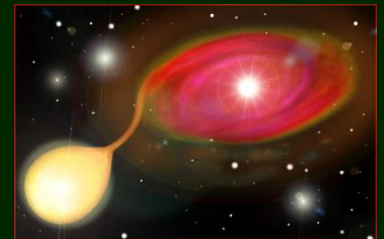
# BHs and NSs in close binary systems

Studying close binaries with compact objects we can obtain mass estimates for progenitors of NSs and BHs (see, for example, Ergma, van den Heuvel 1998 A&A 331, L29).

An interesting result was obtained for the NS system GX 301-2. The progenitor mass was found to be equal to 50 solar masses or more. On the other hand, for many systems with BHs estimates of progenitor masses are lower: 20-50 solar masses.

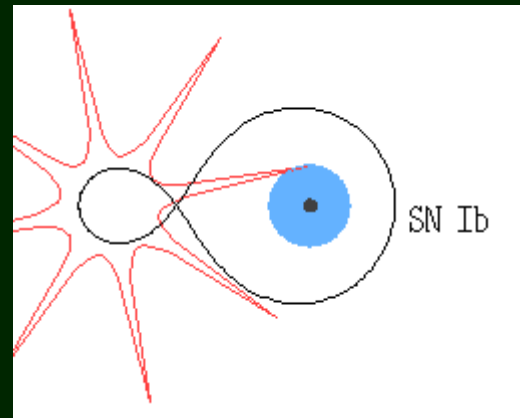
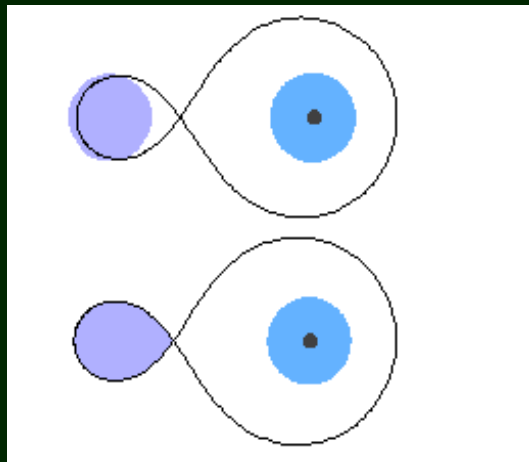
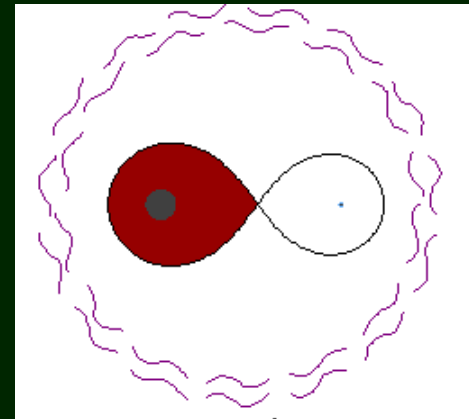
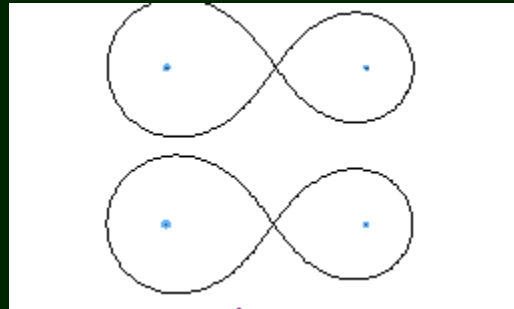
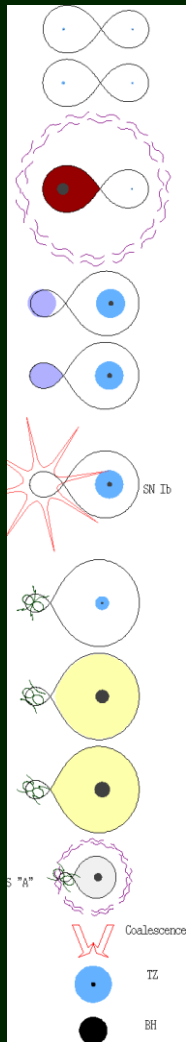
Finally, for the BH system LMC X-3 the mass of the progenitor is estimated as  $>60$  solar masses.

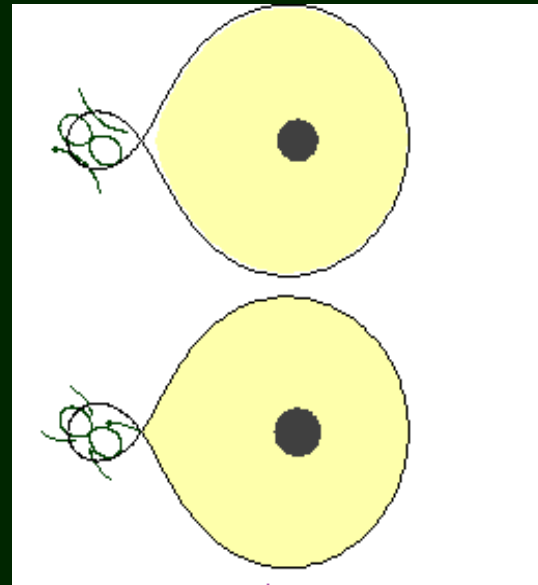
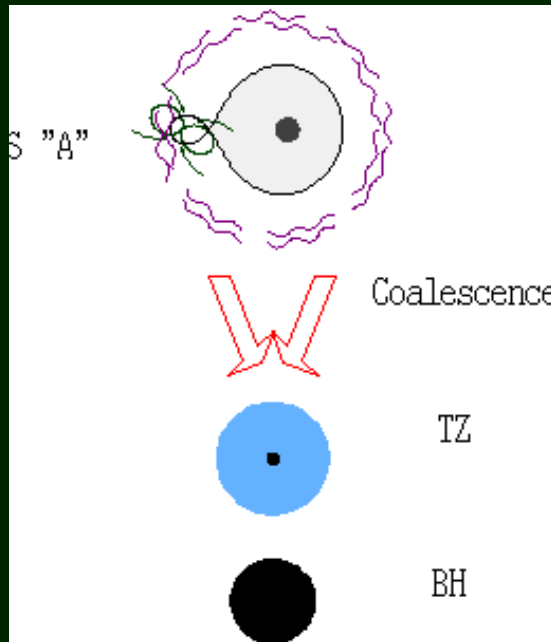
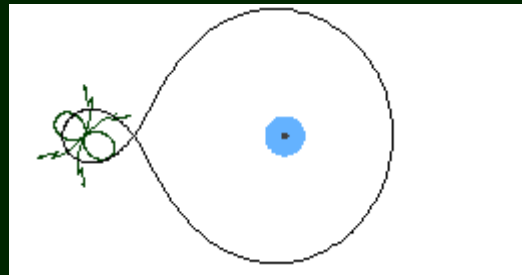
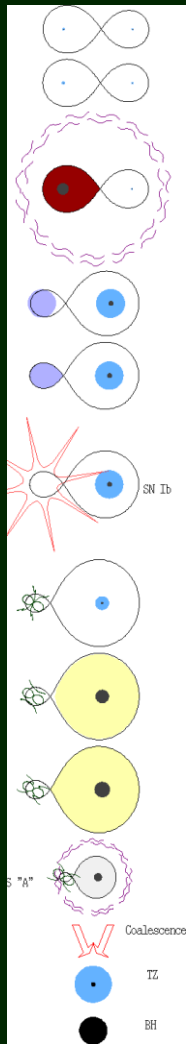
So, the situation is rather complicated. Most probably, in some range of masses, at least in binary systems, both variants are possible.



# Binary evolution

A BH can be formed even from stars each below the limit.





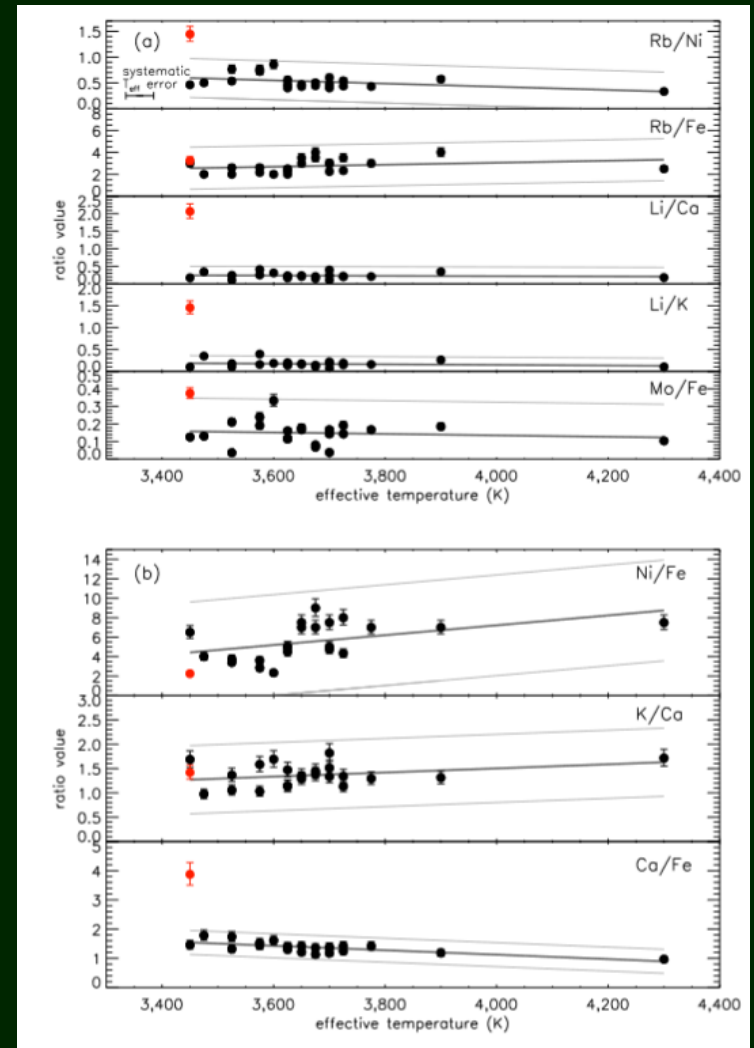
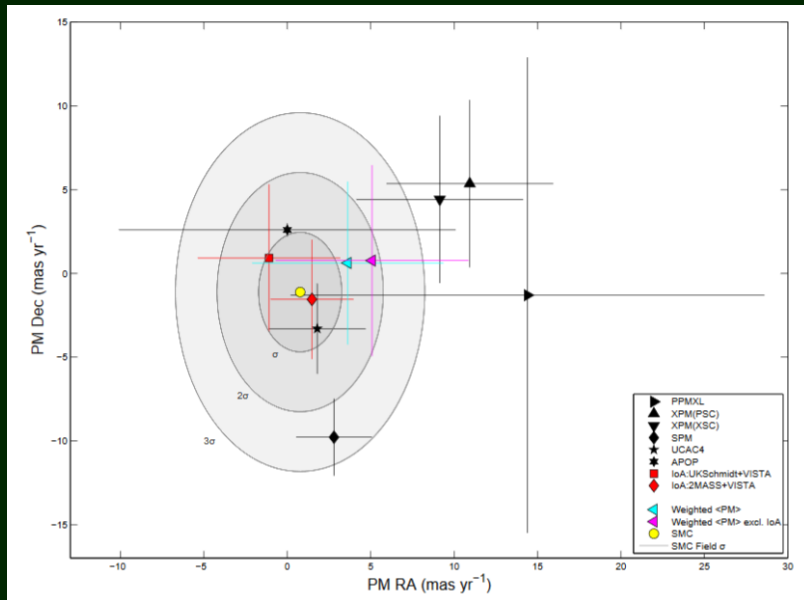
"Scenario machine" calculations

# Thorne-Zytzkow candidates

Chemical composition anomalies.

Discussion:

1. Large proper motion – not in SMC  
1601.05455
2. In SMC 1602.08479  
Gaia DR2 confirms 1804.10192



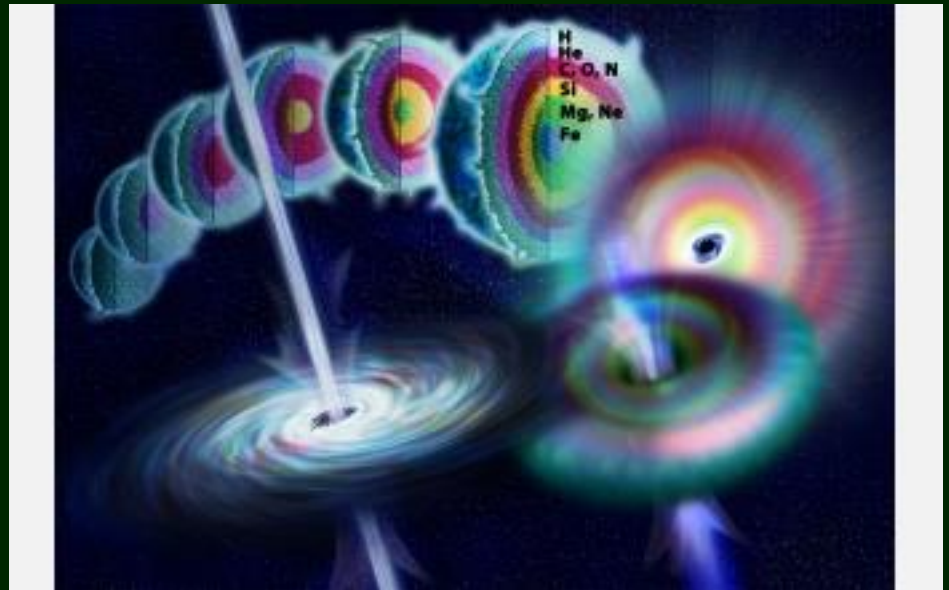
1406.0001

# GRBs and BHs

According to the standard modern model of long GRBs, a BH is the main element of the “central engine”.

So, studying GRBs we can hope to get important information about the first moments of BH's life.

See a very brief review in  
[arXiv:1302.6461](https://arxiv.org/abs/1302.6461)

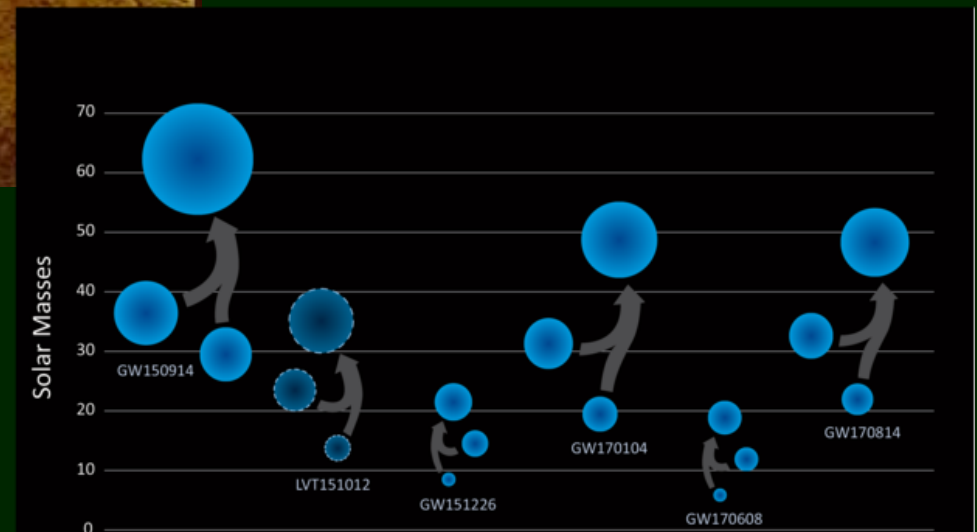


# BHs from GW signals



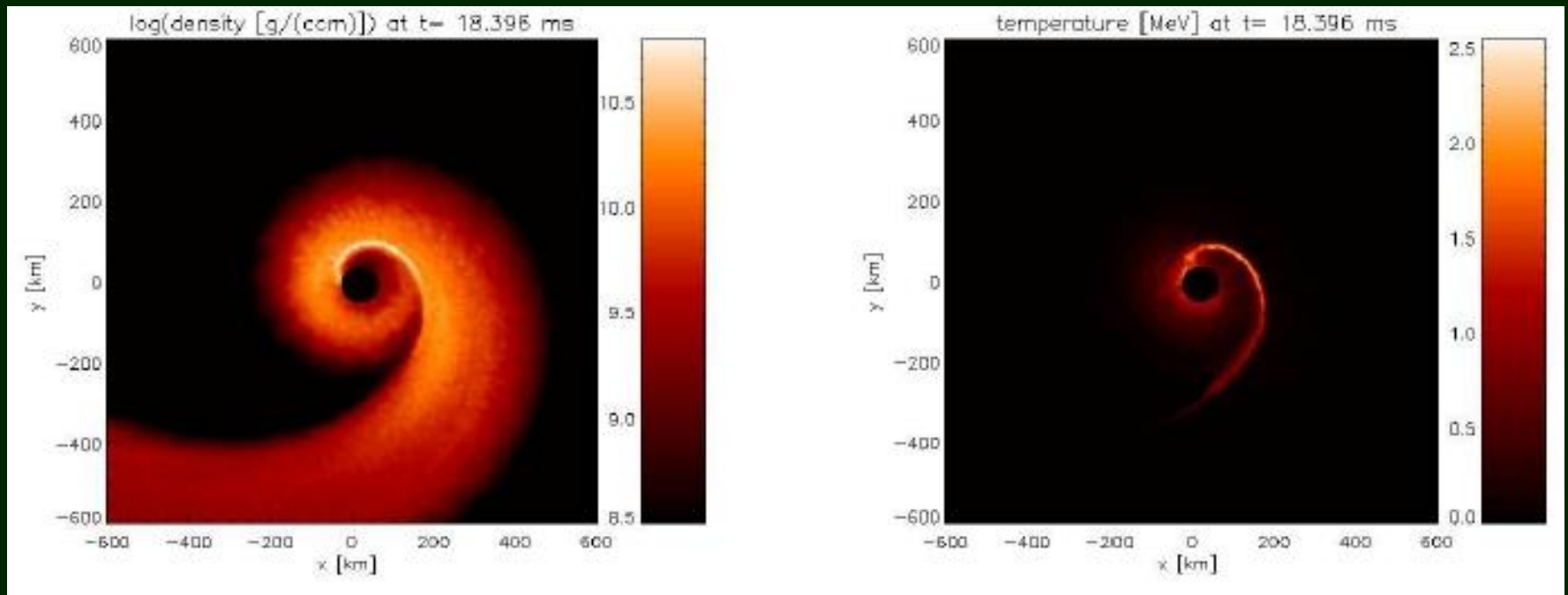
*LIGO* measure signals from compact object mergers.

These signals are more powerful for larger masses. So, even being rarer per unit volume, BH+BH mergers are more frequent in the data.



# NS and BH coalescence

Some numerical models show (astro-ph/0505007, 0505094) that such events do not produce GRBs. Some show that they do.

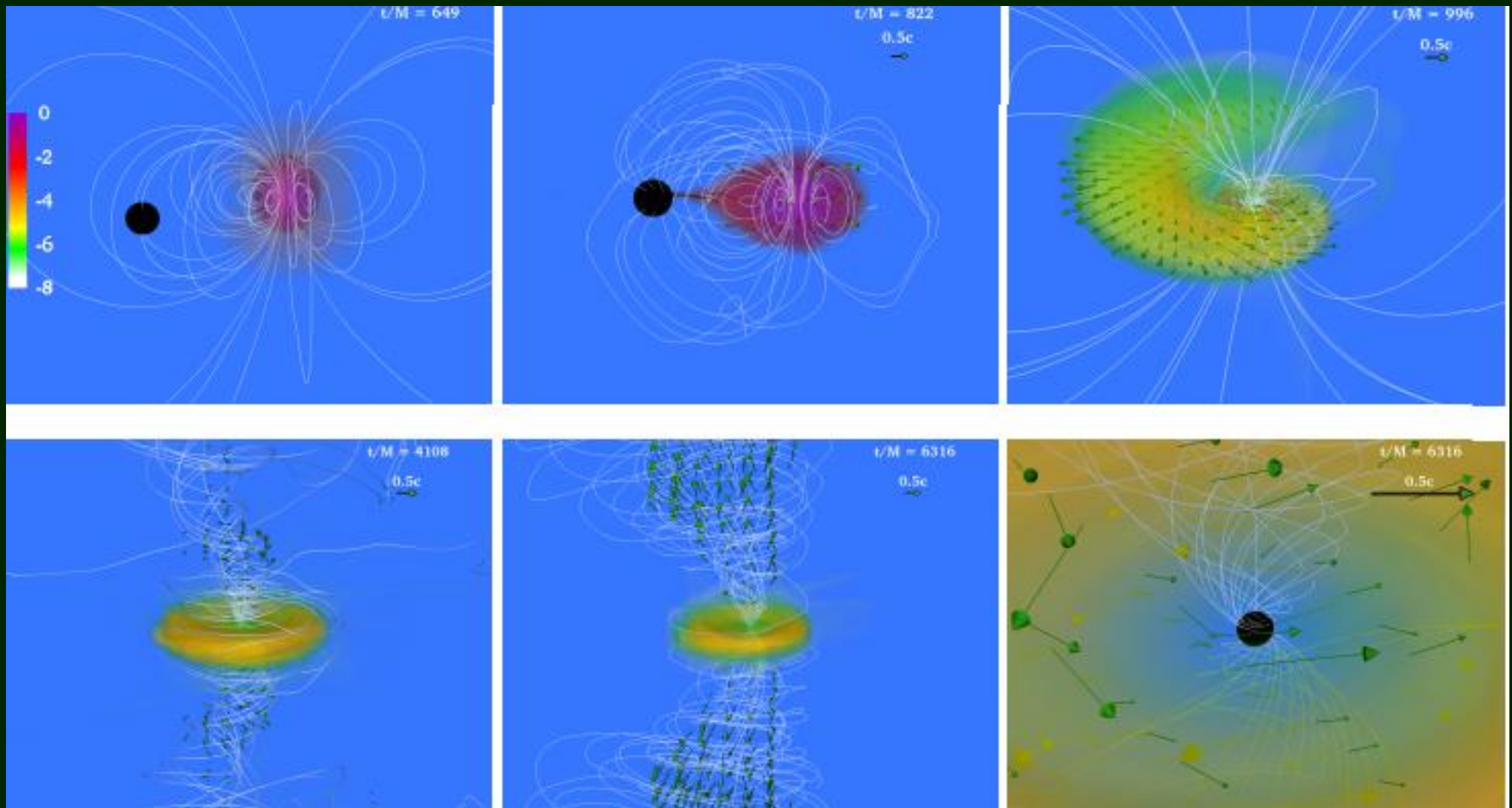


BH-NS mergers are still a popular subject of studies:

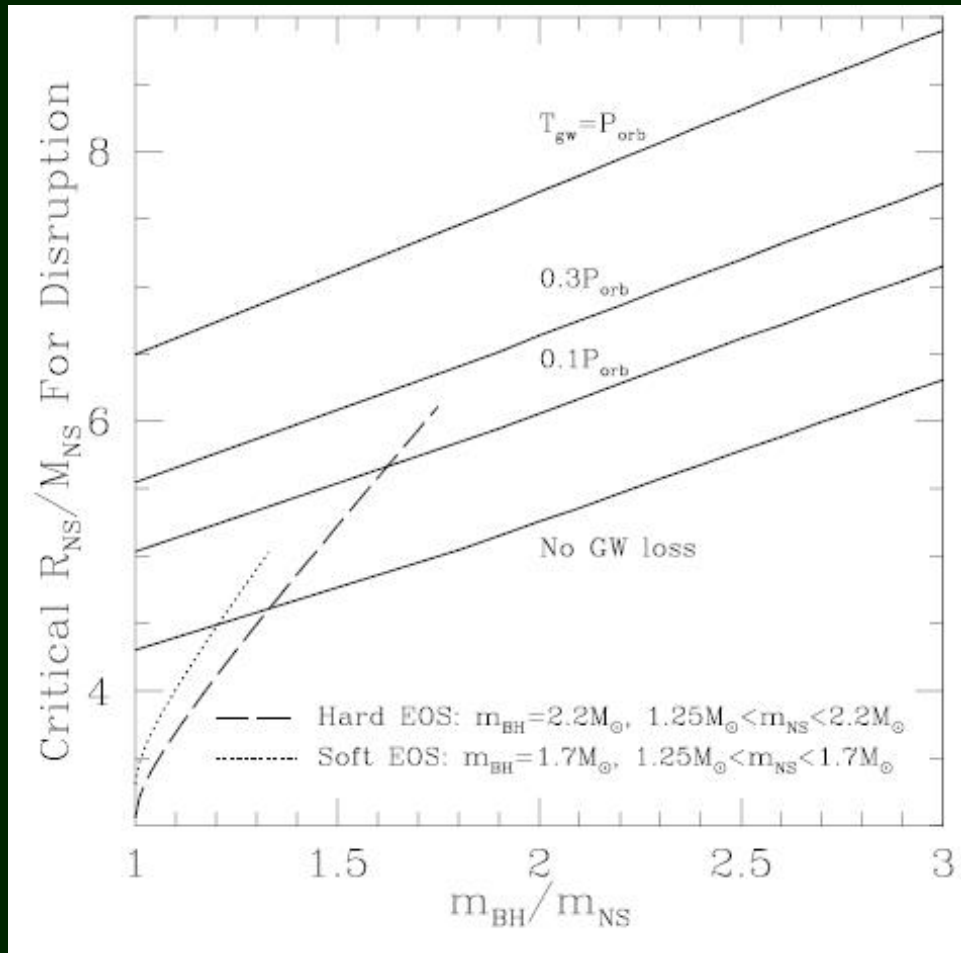
1105.3175, 1103.3526, 1210.8153, 1302.6297, 1301.5616, 1304.3384.

# Magnetic field jet launch

Neutron star magnetic field helps to launch the jet. But disc is still necessary!



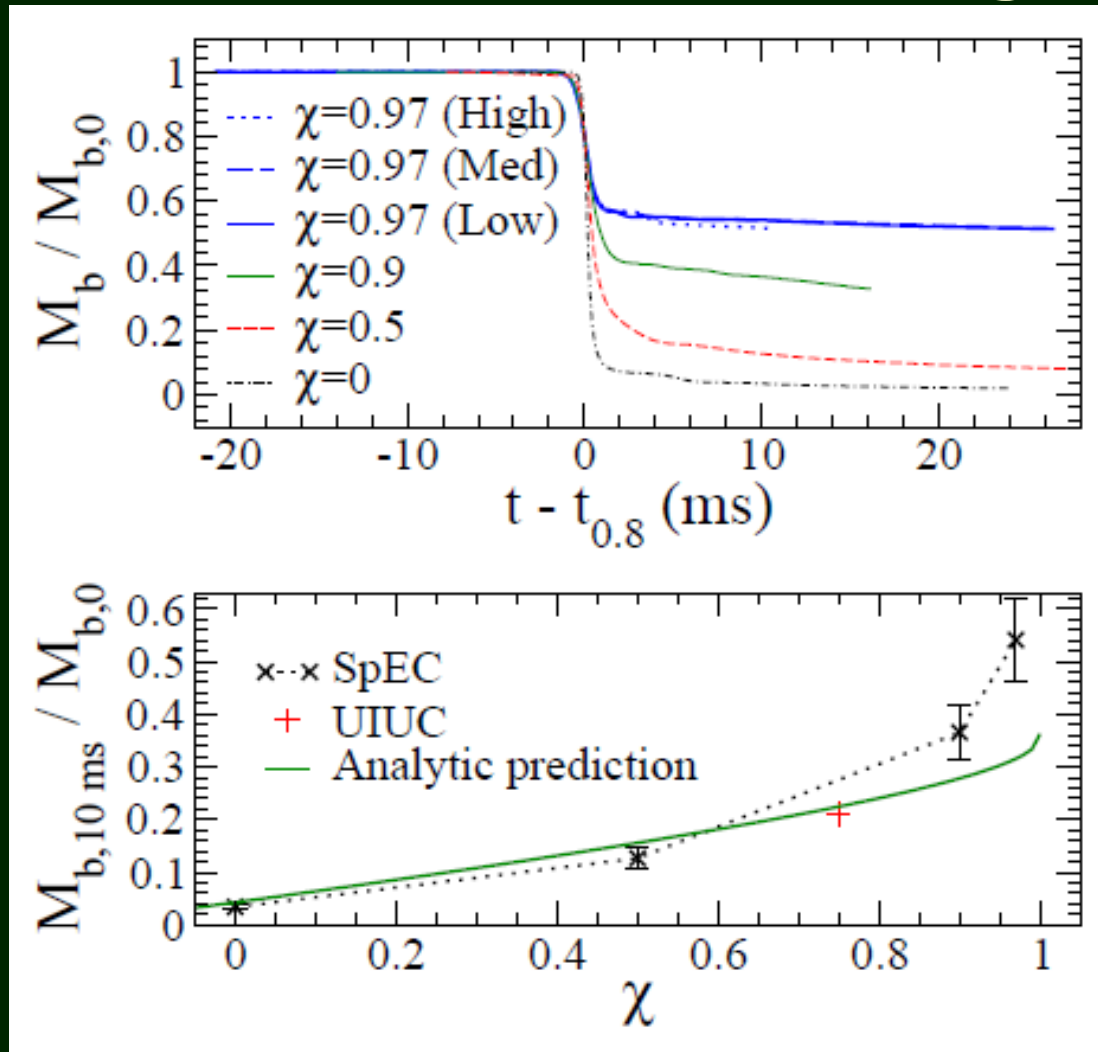
# Prompt mergers of NSs with BHs



Coleman Miller demonstrated that in NS-BH coalescence most probably there is no stable mass transfer and an accretion disc is not formed. This means – **no GRB!**

The top solid line is constructed by assuming that the neutron star will plunge when, in one full orbit, it can reduce its angular momentum below the ISCO value via emission of gravitational radiation. The next two solid lines reduce the allowed time to 30 and 10% of an orbit. The bottom line ignores gravitational radiation losses entirely.

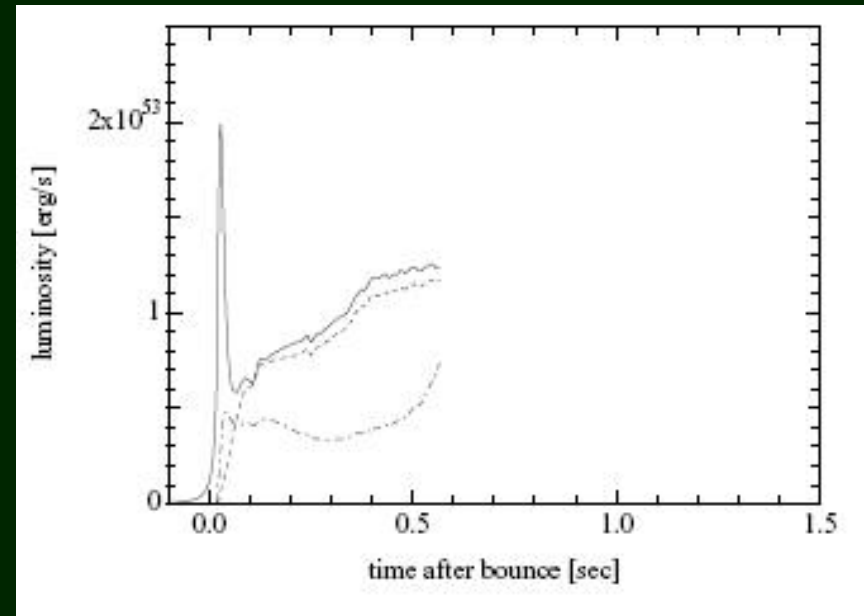
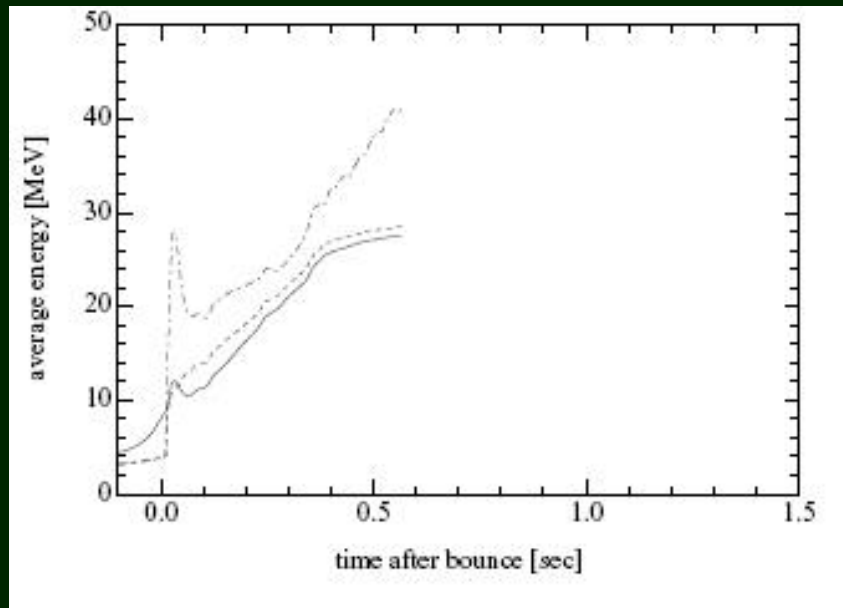
# Extremal BH-NS mergers



It is possible to form a massive disc around a BH during BH-NS merger. However, not for non-rotating BHs.

# Supernovae

The neutrino signal during a (direct) BH formation must be significantly different from the signal emitted during a NS formation. (arXiv: 0706.3762)



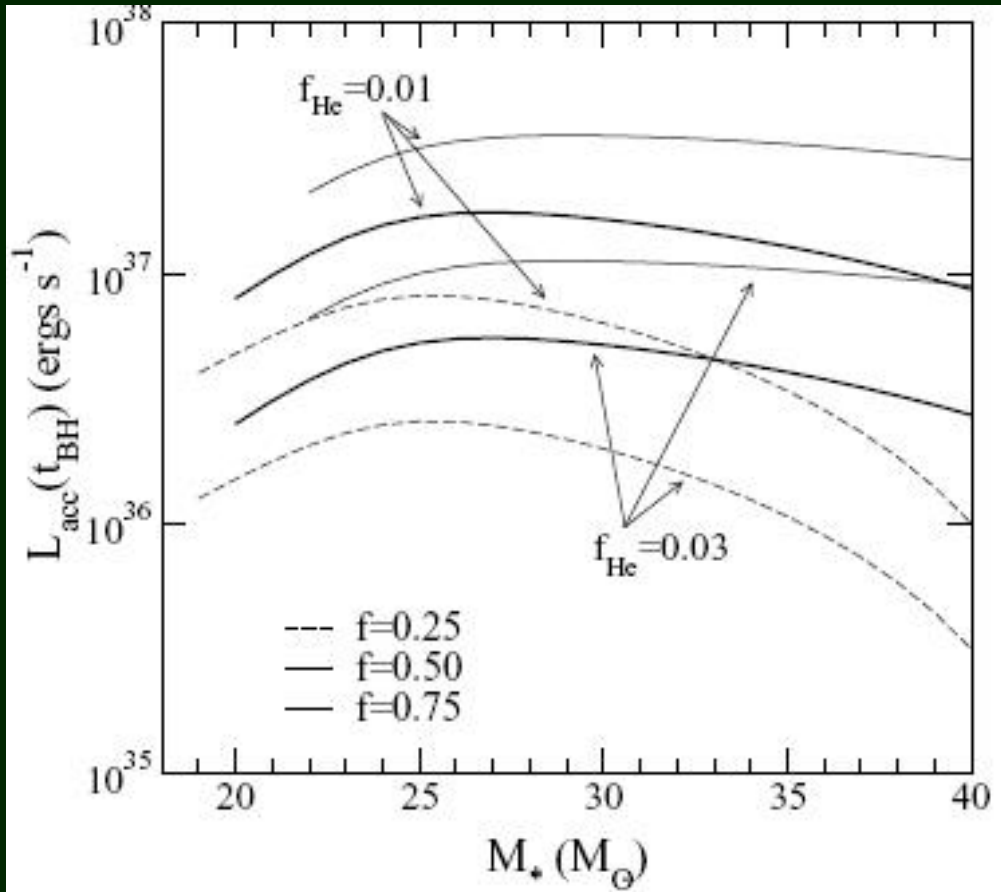
*Different curves are plotted for different types of neutrino:  
electron – solid, electron anti-neutrino – dashed, mu and tau-neutrinos – dot dashed.*

Constant growth of neutrino energy and a sharp cut-off indicate a BH formation. Result depends on the EoS.

See some new results in: arXiv:0809.5129

BH formation in a PNS collapse and neutrino spectra

# BH signatures in SN light curves



$$\dot{M} \propto t^{-5/3}$$

$$L_{\text{acc}}(t) = L_{\text{Edd}} \left( \frac{t}{t_{\text{dust}}} \right)^{-25/18}$$

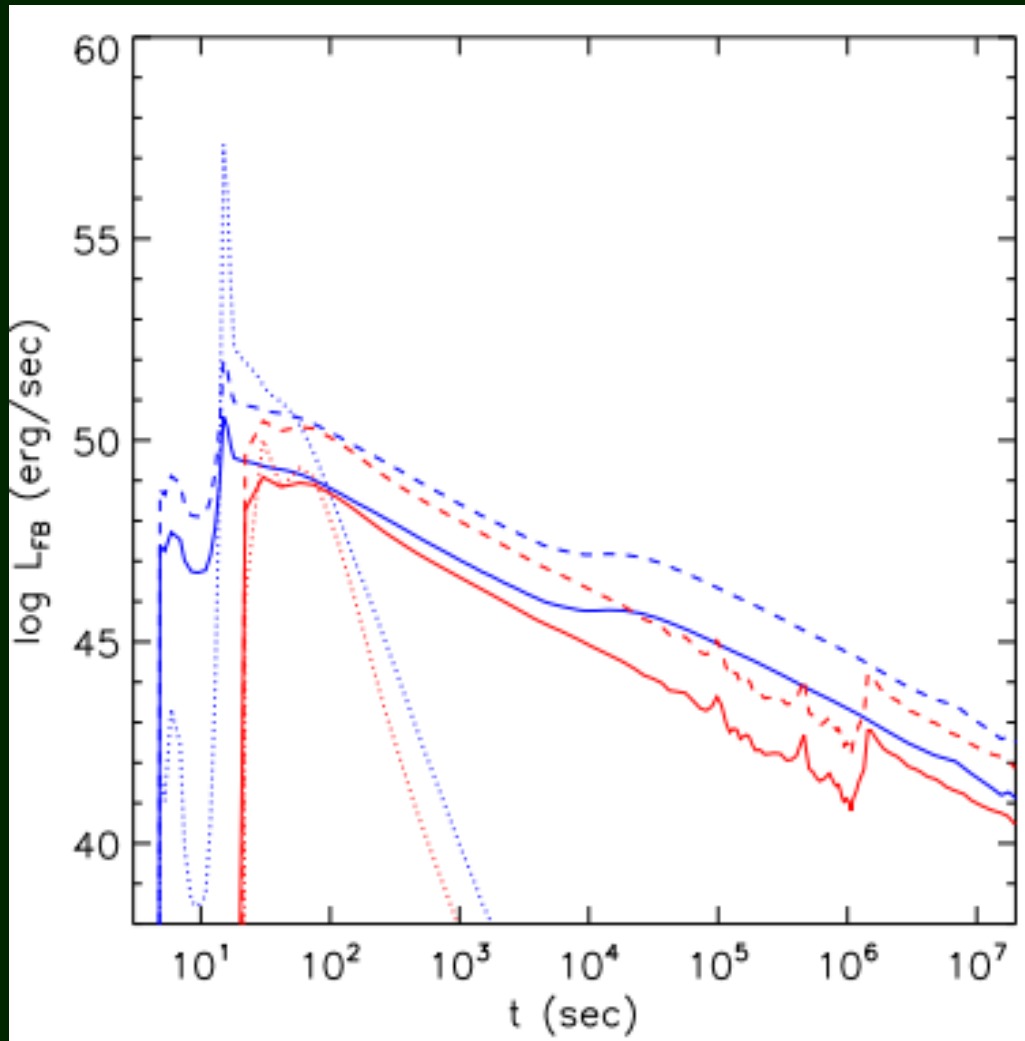
For this plot no radioactive heating is taken into account.

An accreting BH can “emerge” after ~few months-years.

Balberg, Shapiro astro-ph/0104215

(see also Zampieri et al., 1998, ApJ 505, 876)

# New calculations



Several mechanisms of energy release in a fall-back are calculated:

- “accretion heating” (solid line)
- neutrino annihilation (dotted line)
- Blandford-Znajek emission (dashed line).

Estimates show that fallback can potentially lead to large amount of energy deposition to the ejecta, powering super-luminous supernovae.

# A BH birth???

## EVIDENCE FOR A BLACK HOLE REMNANT IN THE TYPE IIL SUPERNOVA 1979C

D. J. PATNAUDE, A. LOEB, & C. JONES

Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138, USA

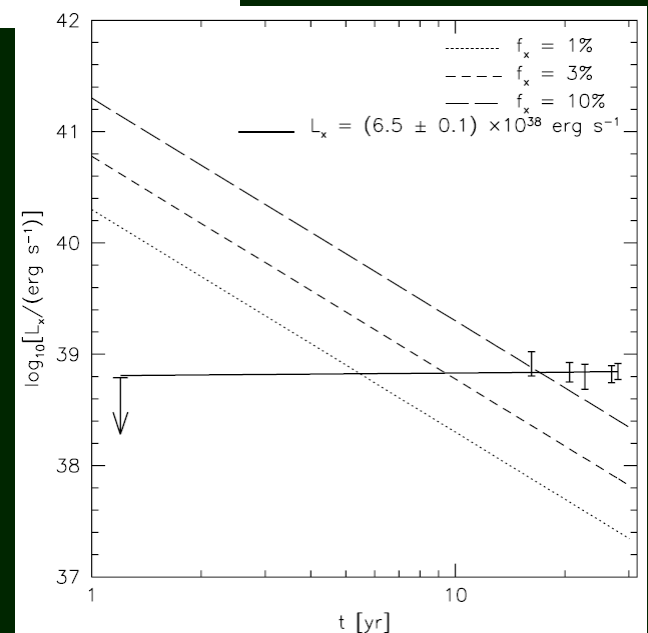
Draft version December 8, 2009

### ABSTRACT

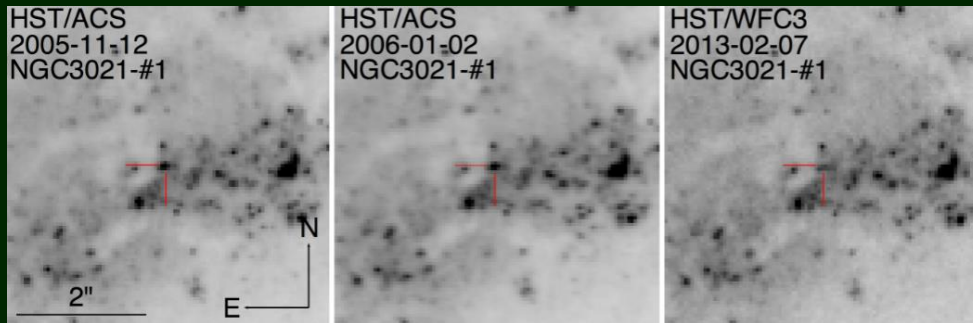
We present an analysis of archival X-ray observations of the Type IIL supernova SN 1979C. We find that its X-ray luminosity is remarkably constant at  $(6.5 \pm 0.1) \times 10^{38} \text{ erg s}^{-1}$ . The high and steady luminosity is evidence for a stellar-mass ( $\sim 5\text{--}10M_{\odot}$ ) black hole accreting material from either a supernova fallback disk or possibly from a binary companion. We find that the bright and steady X-ray light curve is not consistent with either a model for a supernova powered by magnetic braking of a rapidly rotating magnetar, or a model where the blast wave is expanding into a dense circumstellar wind.

TABLE 1  
X-RAY OBSERVATIONS OF SN 1979C

$\Delta t$ yr	Count Rate $10^{-4}$ cps	$F_X^a$ $10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1}$	$L_X^b$ $10^{38} \text{ erg s}^{-1}$	Mission
0.7	< 3.0	< 2.3	< 6.3	<i>Einstein</i> (HRI)
16.2	$6.7 \pm 0.7$	$3.0 \pm 0.3$	$8.2 \pm 0.9$	<i>ROSAT</i> (HRI)
20.6	$42. \pm 2.0$	$2.5 \pm 0.2$	$6.9 \pm 0.6$	<i>Chandra</i> (ACIS-S)
22.7 <sup>c</sup>	...	$2.3 \pm 0.3$	$6.3 \pm 0.7$	<i>XMM-Newton</i> (MOS)
26.9	$40. \pm 0.8$	$2.4 \pm 0.2$	$6.6 \pm 0.5$	<i>Chandra</i> (ACIS-S)
28.0	$43. \pm 0.3$	$2.6 \pm 0.2$	$7.0 \pm 0.5$	<i>Chandra</i> (ACIS-S)



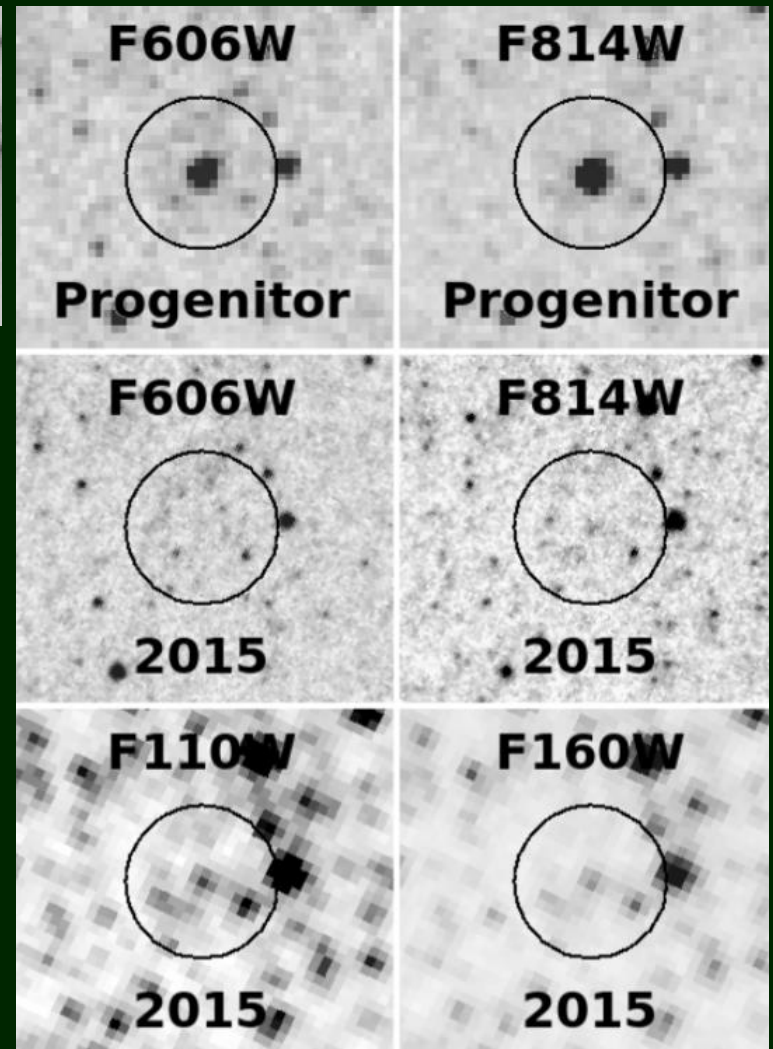
# Disappearance of stars



The event is consistent with the ejection of the envelope of a red supergiant in a failed supernova and the late-time emission could be powered by fallback accretion onto a newly-formed black hole.

Progenitor mass  $\sim 23$ -28 solar.  
Consistent with the missing RSG problem.

In 2018  $\sim 30$  examples of identified SN progenitors are known, and there  $\sim 40$  upper limits (1802.07870).



1609.01283

# Conclusions

- There can be different kinds of BHs: PBH, stellar, IMBH, SMBH
- Stellar mass BHs can be observed due to
  - accretion in binaries
  - GRBs
  - GW
  - in SN
- Mass interval for stellar mass BH formation is not certain