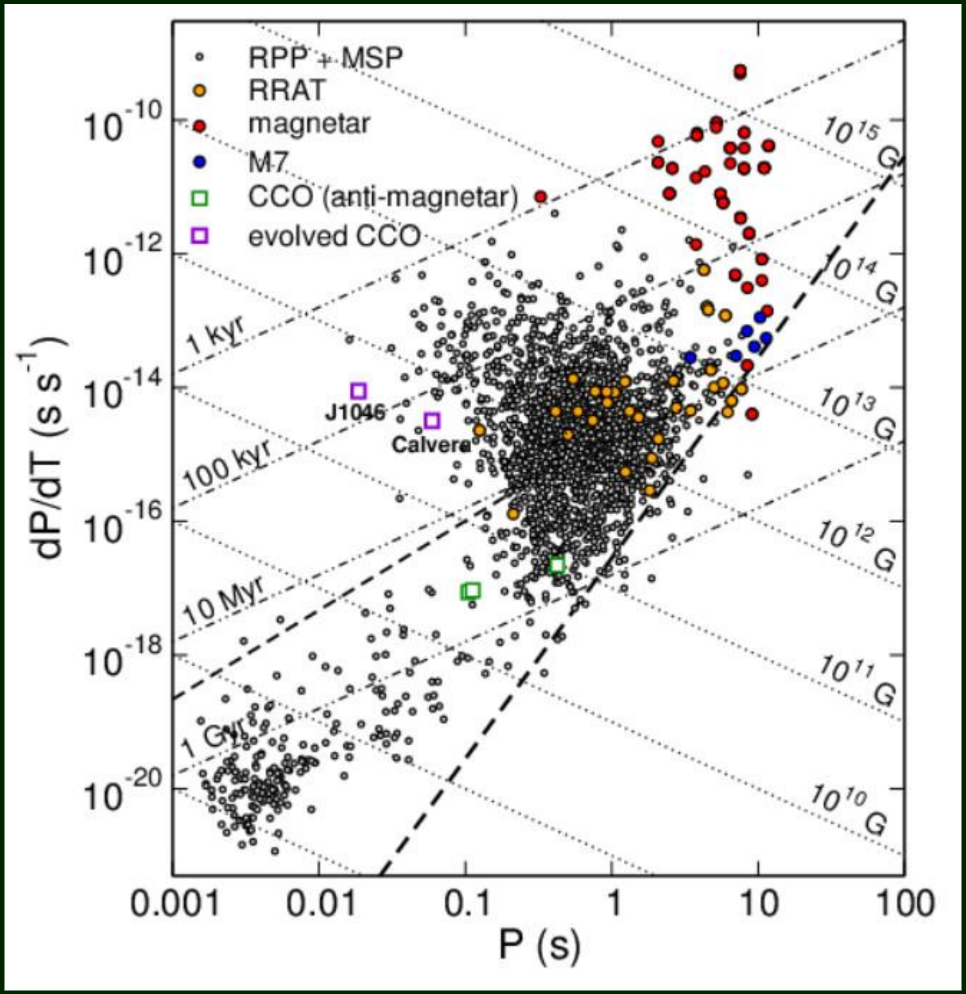
Evolution with decaying and re-emerging magnetic field

Diversity of young neutron stars



Young isolated neutron stars can appear in many flavors:

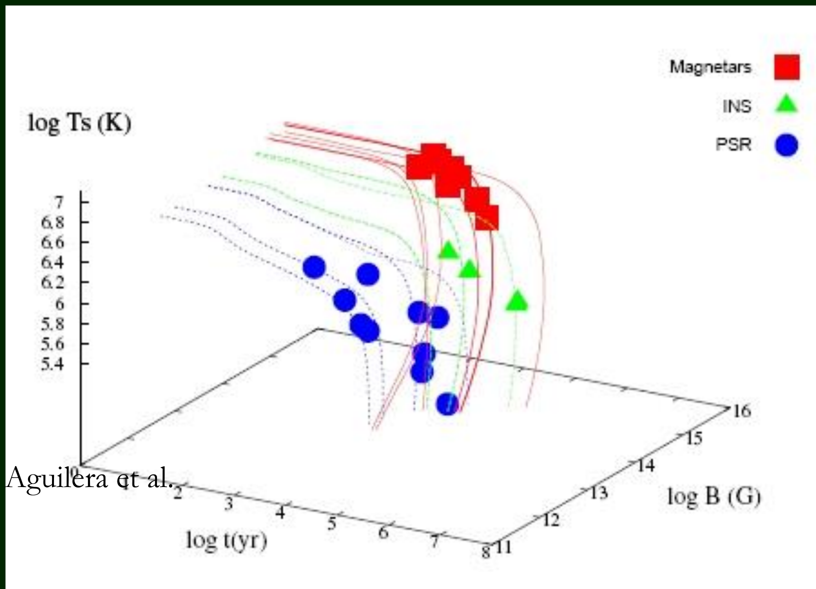
o Radio pulsars
o Compact central X-ray sources in supernova remnants.
o Anomalous X-ray pulsars
o Soft gamma repeaters
o The Magnificent Seven & Co.
o Transient radio sources (RRATs)

The term "GRAND UNIFICATION FOR NEUTRON STARS" was coined by Kaspi (2010)

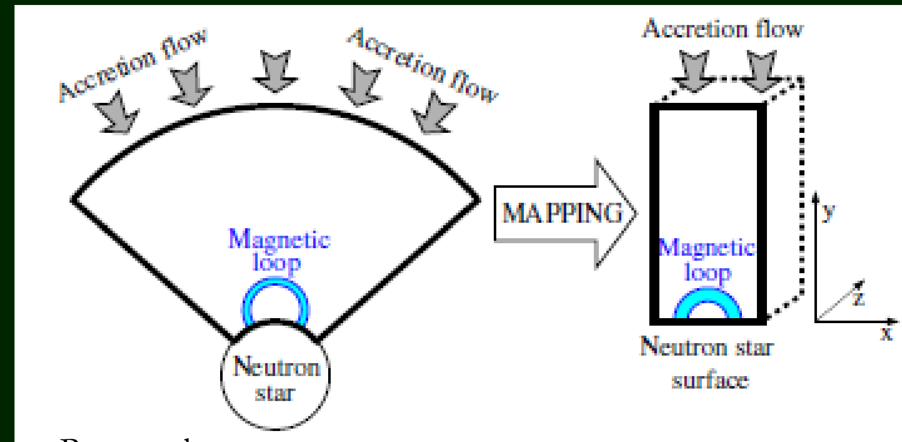
PSRs, magnetars and M7 unified in the model by Popov et al. (2010).

Pires et al. 2015

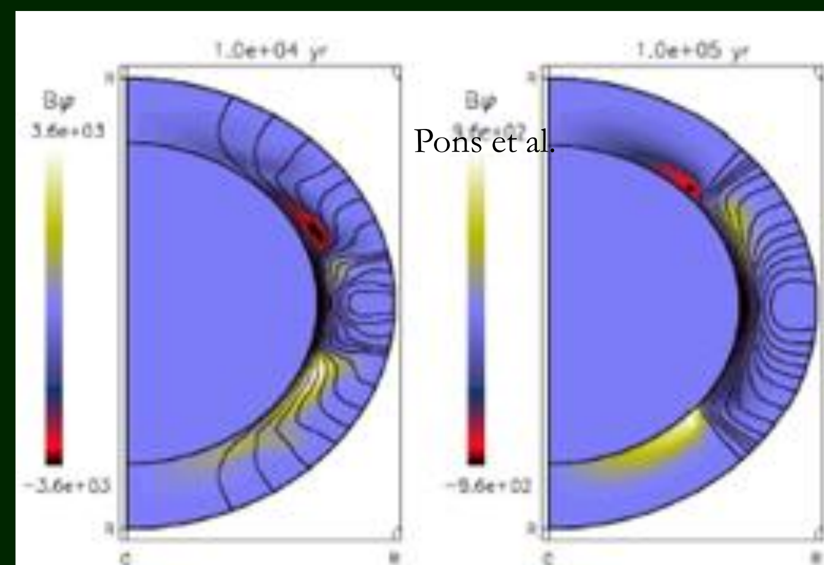
Three main ingredients of a unified model



Aguilera et al.



Page et al.

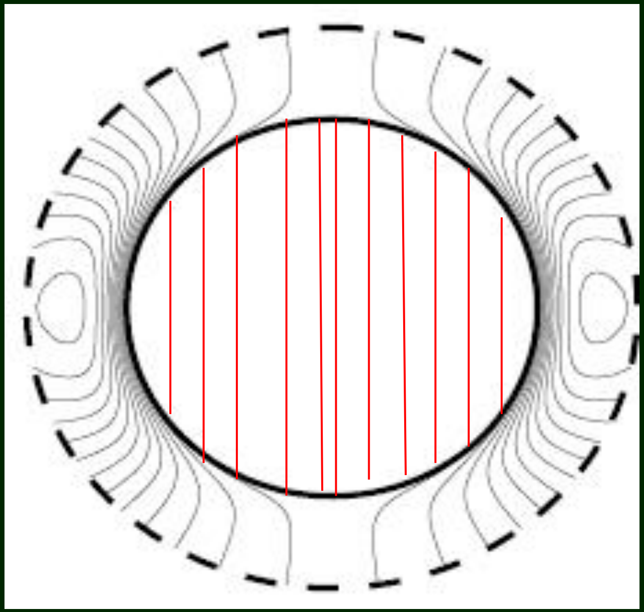


Pons et al.

- Field decay
  - Emerging magnetic field
    - Toroidal magnetic field

Magnetic field decay

Magnetic fields of NSs are expected to decay due to decay of currents which support them.



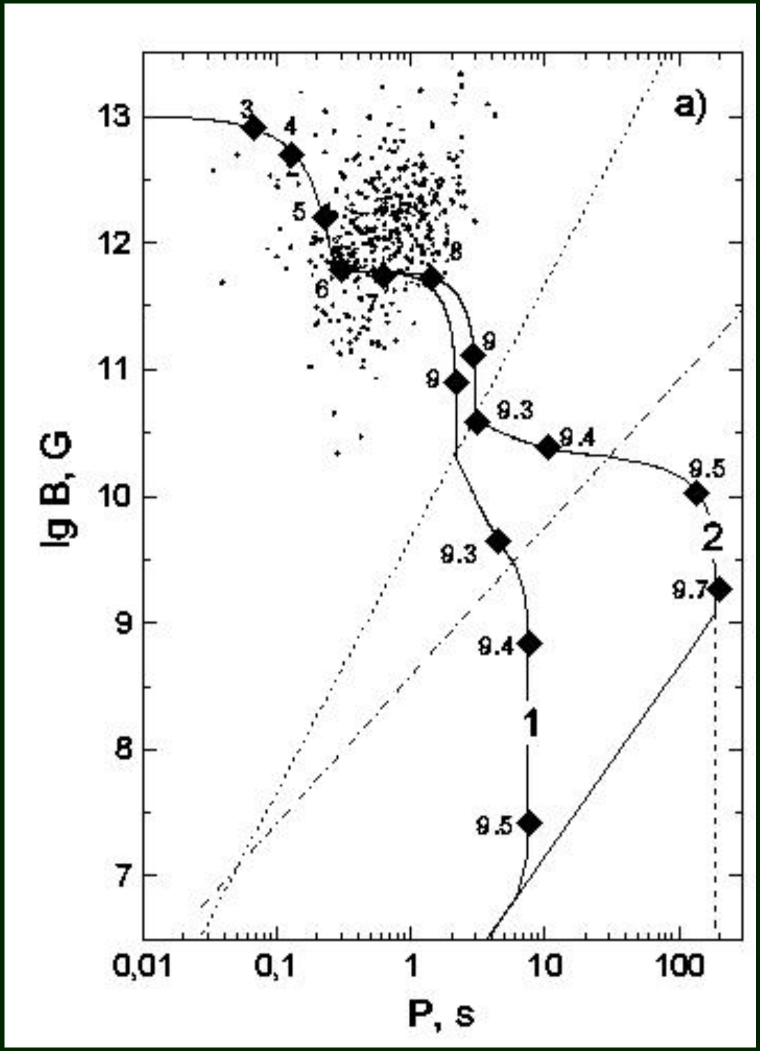
Crustal field of core field?

It is easy to decay in the crust.

In the core the filed is in the form of superconducting vortices.
They can decay only when they are moved into the crust (during spin-down).

Still, in most of models strong fields decay.

Period evolution with field decay



An evolutionary track of a NS is very different in the case of decaying magnetic field.

The most important feature is slow-down of spin-down. Finally, a NS can nearly freeze at some value of spin period.

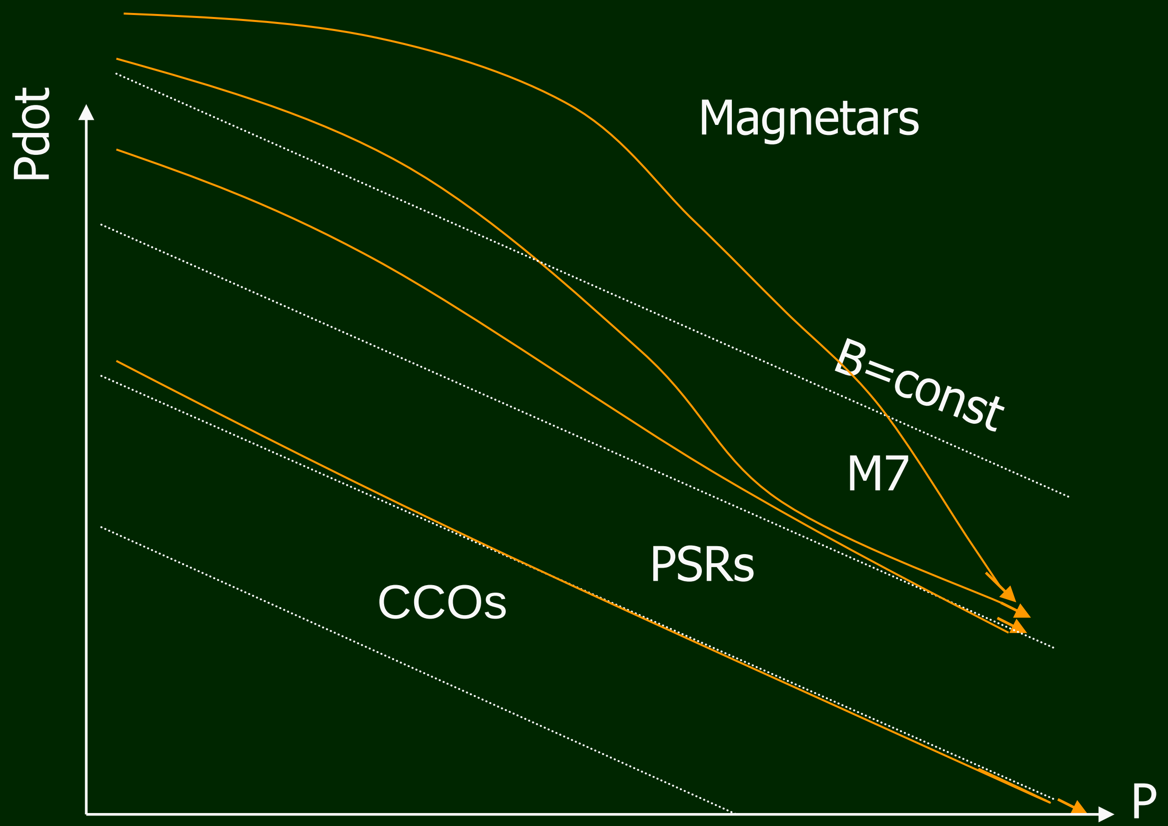
Several episodes of relatively rapid field decay can happen.

Number of isolated accretors can be both decreased or increased in different models of field decay. But in any case their average periods become shorter and temperatures lower.

astro-ph/9707318

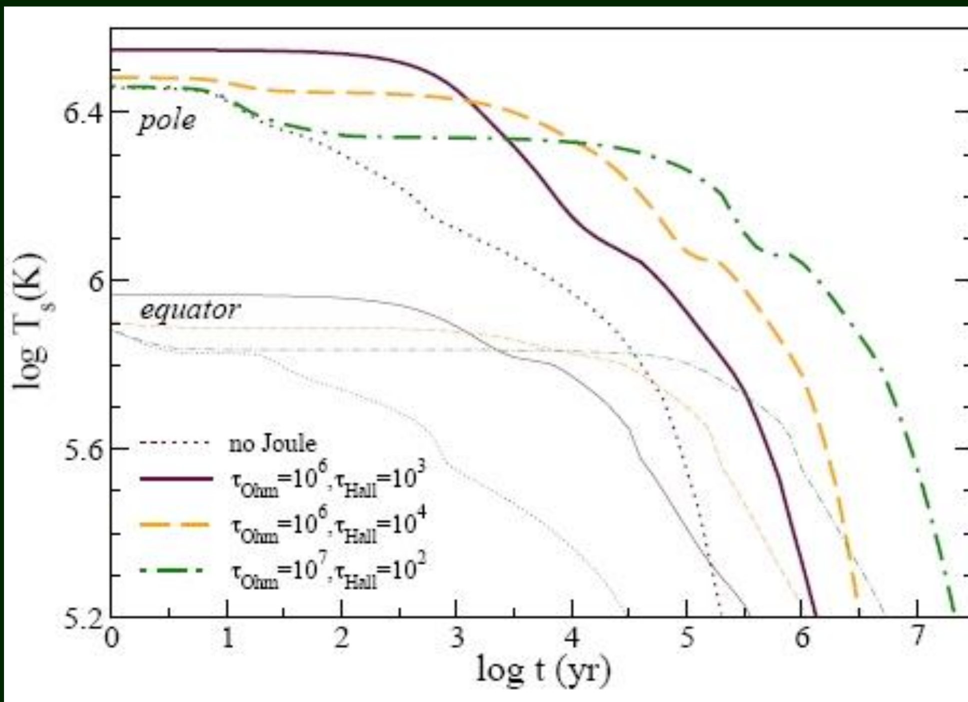
Magnetars, field decay, heating

A model based on field-dependent decay of the magnetic moment of NSs can provide an evolutionary link between different populations (Pons et al.).



Magnetic field decay vs. thermal evolution

Magnetic field decay can be an important source of NS heating.



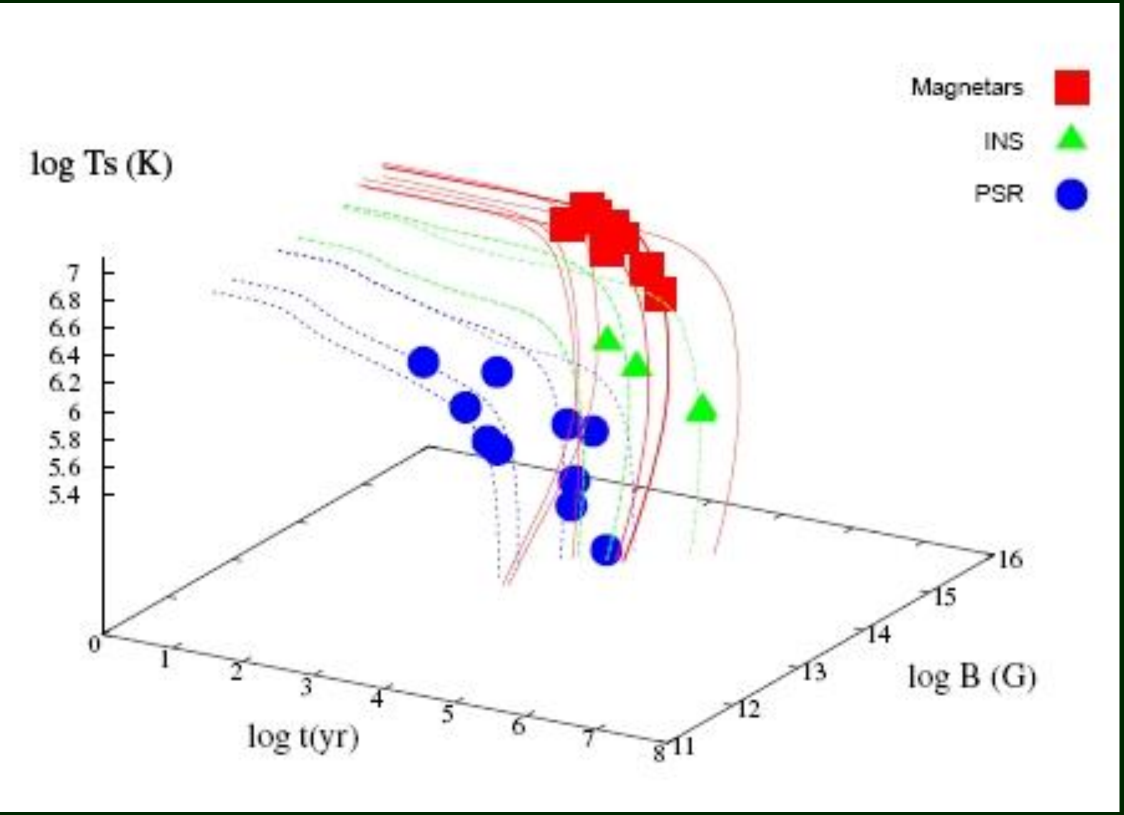
Heat is carried by electrons.
It is easier to transport heat along field lines. So, poles are hotter.
(for light elements envelope the situation can be different).

$$B = B_0 \frac{\exp\left(-t/\tau_{\mathrm{Ohm}}\right)}{1+\frac{\tau_{\mathrm{Ohm}}}{\tau_{\mathrm{Hall}}}\left(1-\exp\left(-t/\tau_{\mathrm{Ohm}}\right)\right)}$$

Ohm and Hall decay

arxiv:0710.0854 (Aguilera et al.)

Joule heating for everybody?



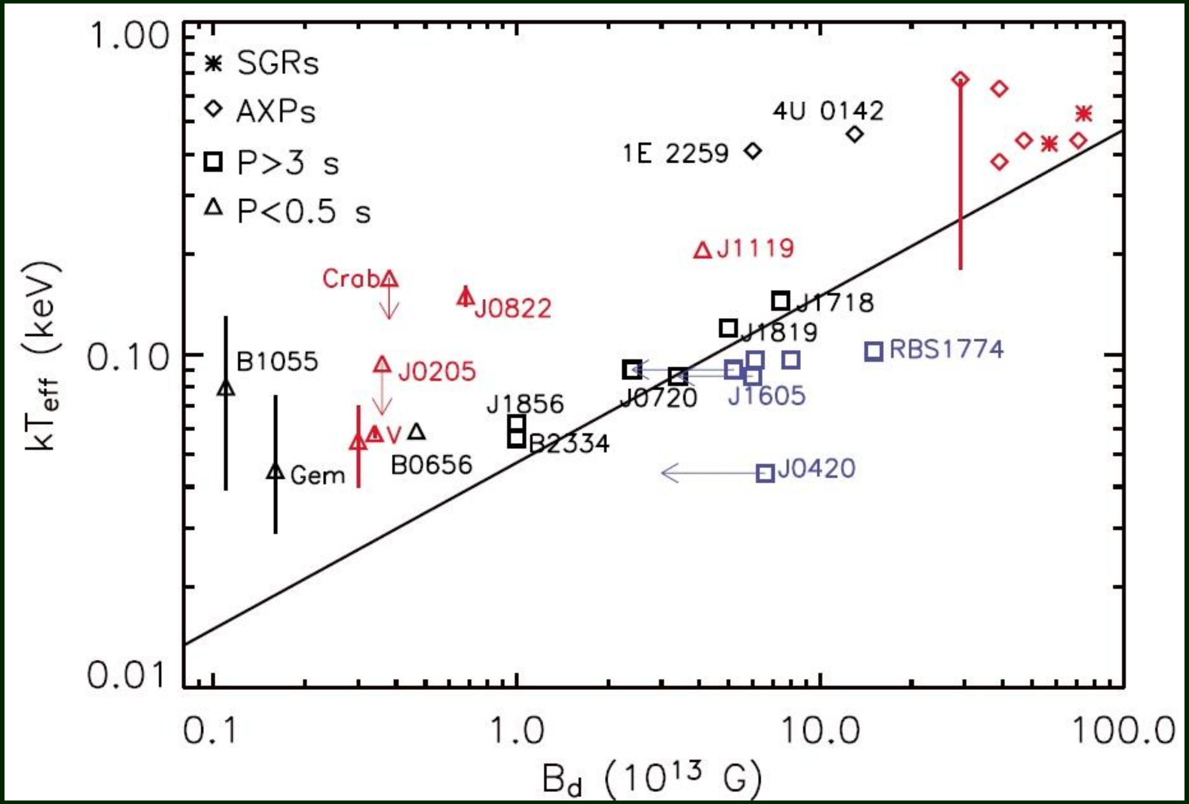
It is important to understand the role of heating by the field decay for different types of INS.

In the model by Pons et al. the effect is more important for NSs with larger initial B.

Note, that the characteristic age estimates (P/2 Pdot) are different in the case of decaying field!

arXiv: 0710.4914 (Aguilera et al.)

Magnetic field vs. temperature



The line marks balance between heating due to the field decay and cooling

It is expected that a NS evolves downwards till it reaches the line, then the evolution proceeds along the line: $T_{eff} \sim B_d^{1/2}$

Selection effects are not well studied here.

A kind of population synthesis modeling is welcomed.

(astro-ph/0607583)

What kind of decay do we see?

Ohmic decay due to phonons

Hall cascade

Both time scales fit, and in both cases we can switch of decay at ~10
either due to cooling, or due to the Hall attractor.

$$B = B_0 \frac{\exp\left(-t/\tau_{\rm Ohm}\right)}{1 + \frac{\tau_{\rm Ohm}}{\tau_{\rm Hall}}\left(1 - \exp\left(-t/\tau_{\rm Ohm}\right)\right)}$$

Hall cascade and field evolution

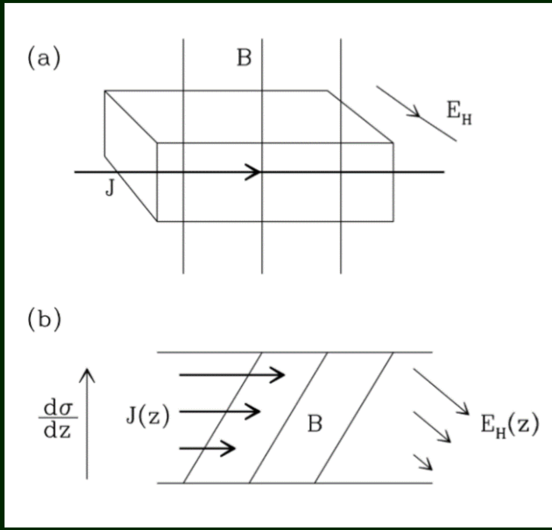
$$\frac{\partial B}{\partial t} = -c\nabla \times E,$$

advection, Ohm, Hall:

$$E = -\frac{1}{c} v \times B + \frac{J}{\sigma} + \frac{J \times B}{n_e e c},$$

$$\tau_{\mathrm{Ohm}} = \frac{4\pi\sigma L^2}{c^2}.$$

$$J = (c/4\pi)(\nabla \times B)$$



With only Hall term we have:

$$\frac{\partial B}{\partial t} = -\nabla \times \left(\frac{J \times B}{n_e e}\right),$$

$$t_{\mathrm{Hall}} = \frac{n_e e L}{J} = \frac{4\pi n_e e L^2}{cB},$$

astro-ph/0402392

Characteristic timescales

$$\tau_{\rm Hall} = \frac{4\pi e n_e L^2}{cB(t)},$$

$$\tau_{\rm Hall} = \tau_{\rm Hall,0}\frac{B_0}{B(t)}.$$

Hall time scale strongly depends on the current value of the field.

$$\tau_{\rm Ohm} = \frac{4\pi\sigma L^2}{c^2},$$

$$\sigma = \frac{\sigma_{\rm Q}\sigma_{\rm ph}}{\sigma_{\rm Q}+\sigma_{\rm ph}}.$$

Ohmic decay depends on the conductivity

$$\tau_{\rm Ohm}^{-1} = \tau_{\rm Ohm,ph}^{-1} + \tau_{\rm Ohm,Q}^{-1}.$$

Resistivity can be due to

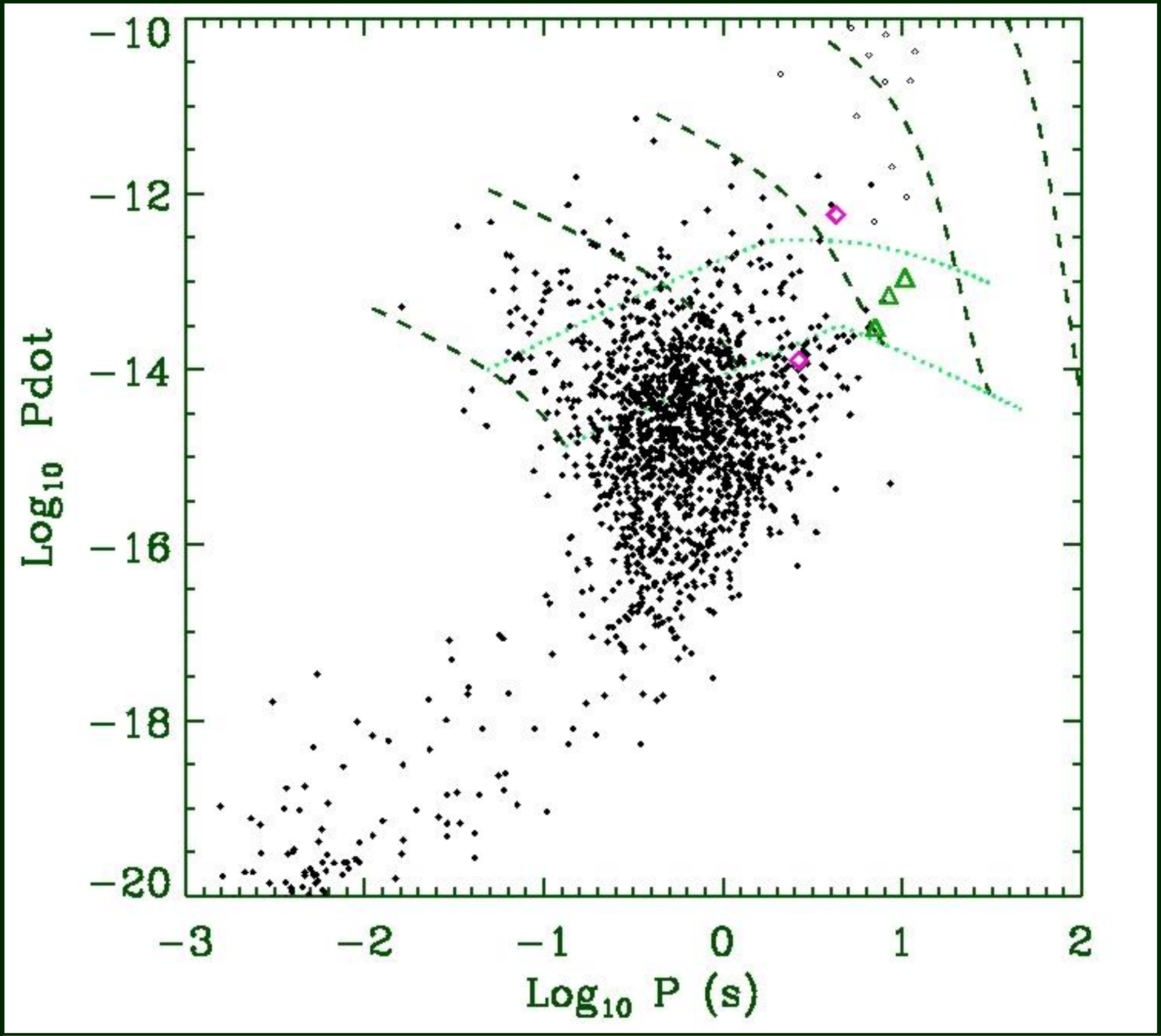
- Phonons
- Impurities

$$\sigma_Q = 4.4\times 10^{25}{\rm s}^{-1}\left(\frac{\rho_{14}^{1/3}}{Q}\right)\left(\frac{Y_e}{0.05}\right)^{1/3}\left(\frac{Z}{30}\right),$$

$$Q = n_{\rm ion}^{-1}\Sigma_i n_i \times (Z^2 - \langle Z\rangle^2).$$

$$\sigma_{\rm ph} = 1.8\times 10^{25}{\rm s}^{-1}\left(\frac{\rho_{14}^{7/6}}{T_8^2}\right)\left(\frac{Y_e}{0.05}\right)^{5/3},$$

P-Pdot diagram and field decay



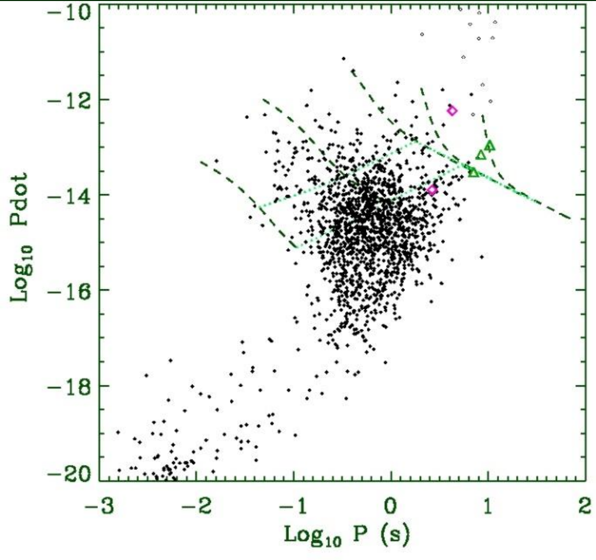
$$B = B_0 \frac{\exp(-t/\tau_{\rm Ohm})}{1 + \frac{\tau_{\rm Ohm}}{\tau_{\rm Hall}}(1 - \exp(-t/\tau_{\rm Ohm}))}$$

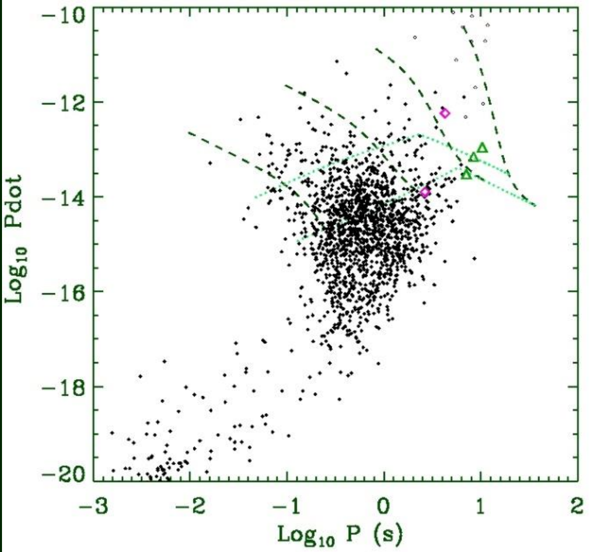
$\tau_{Ohm}=10^6$ yrs

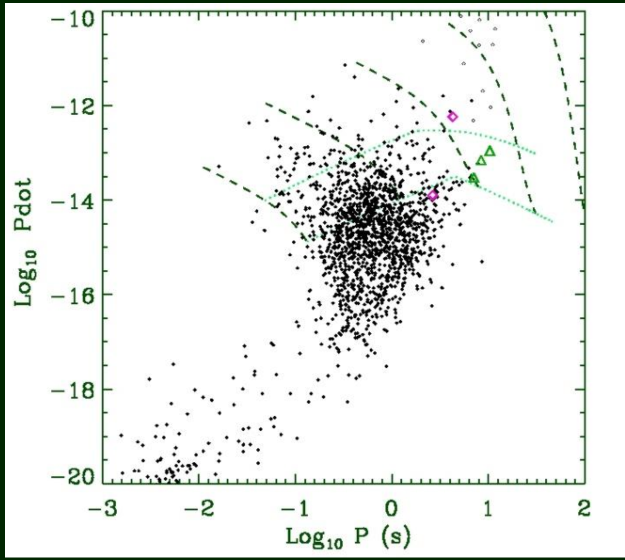
$\tau_{Hall}=10^4/(B_0/10^{15}$ G) yrs

(Popov et al. MNRAS 2010. arXiv: 0910.2190)

Decay parameters and P-Pdot







$\tau_{Ohm}=10^7$ yrs
$\tau_{Hall}=10^2/(B_0/10^{15}\,G)$

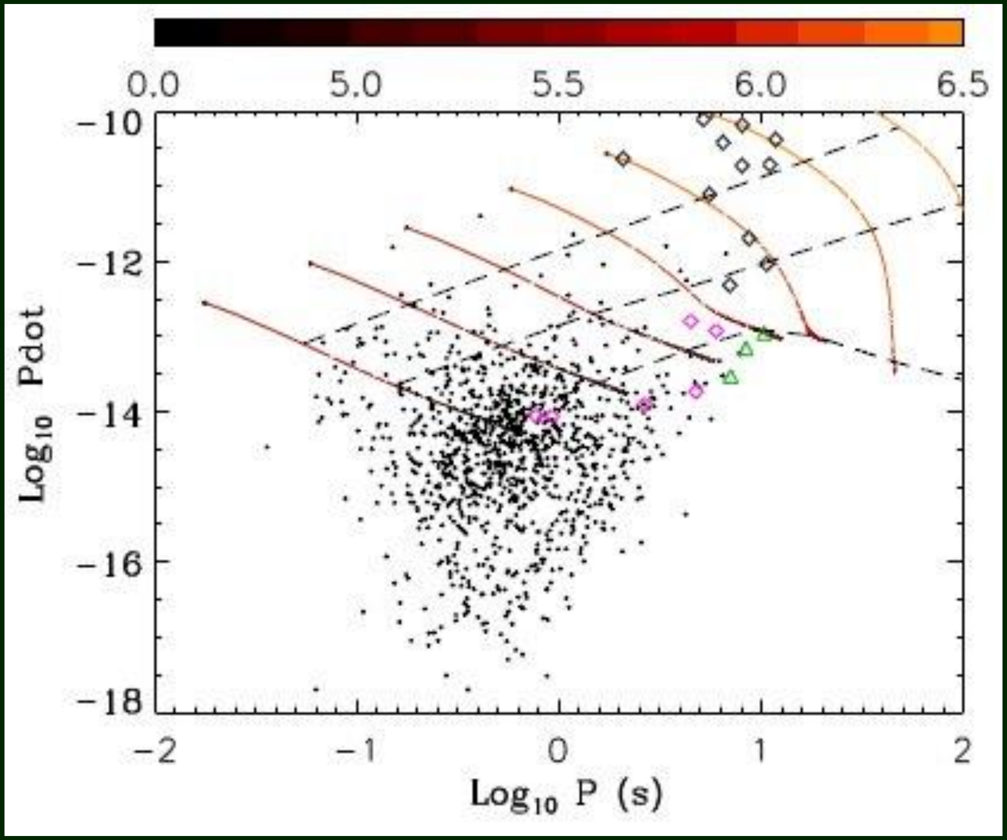
$\tau_{Ohm}=10^6$ yrs
$\tau_{Hall}=10^3/(B_0/10^{15}\,G)$

$\tau_{Ohm}=10^6$ yrs
$\tau_{Hall}=10^4/(B_0/10^{15}\,G)$

Longer time scale for the Hall field decay is favoured.

It is interesting to look at HMXBs to see if it is possible to derive the effect of field decay and convergence.

Realistic tracks



Using the model by Pons et al. (arXiv: 0812.3018) we plot realistic tracks for NS with masses 1.4 Msolar.

Initial fields are:
3 10^{12}, 10^{13}, 3 10^{13}, 10^{14}, 3 10^{14}, 10^{15}

Color on the track encodes surface temperature.

Tracks start at 10^3 years, and end at 2 10^6 years.

(Popov et al. MNRAS 2010)

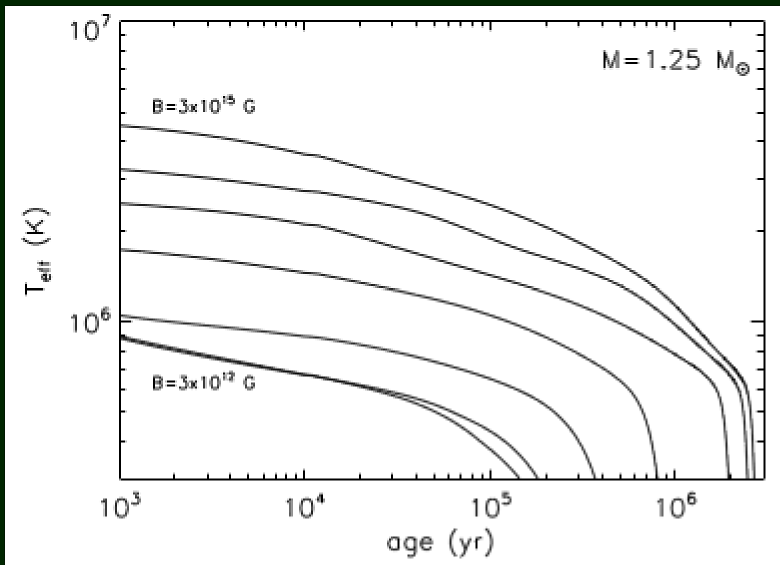
Joint description of NS evolution with decaying magnetic field

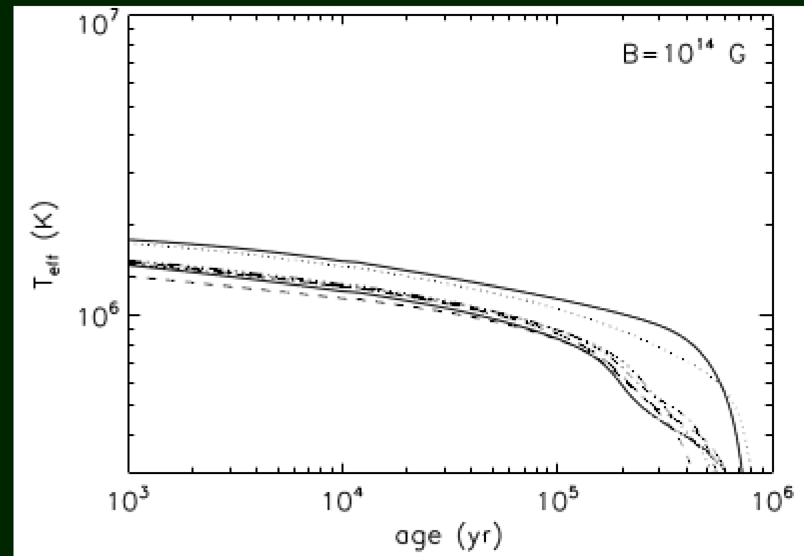
The idea to describe all types of NSs with a unique model using one initial distribution (fields, periods, velocities) and to compare with observational data, i.e. to confront vs. all available observed distributions:
- P-Pdot for PSRs and other isolated NSs
- Log N – Log S for cooling close-by NSs
- Luminosity distribution of magnetars (AXPs, SGRs)
- ……………..

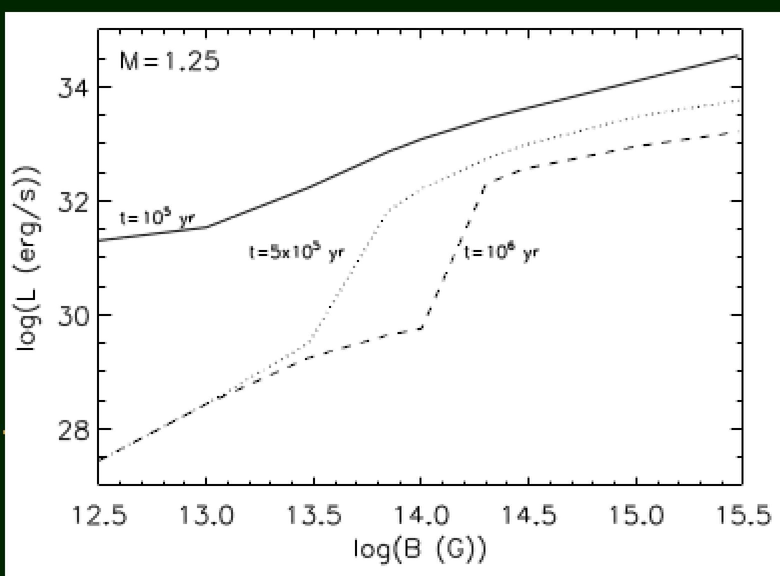
The first step is done in Popov et al. (2010)

The initial magnetic field distribution with $\langle \log B_0 \rangle \sim 13.25$ and $\sigma \sim 0.6$ gives a good fit.
~10% of magnetars.

Cooling curves with decay

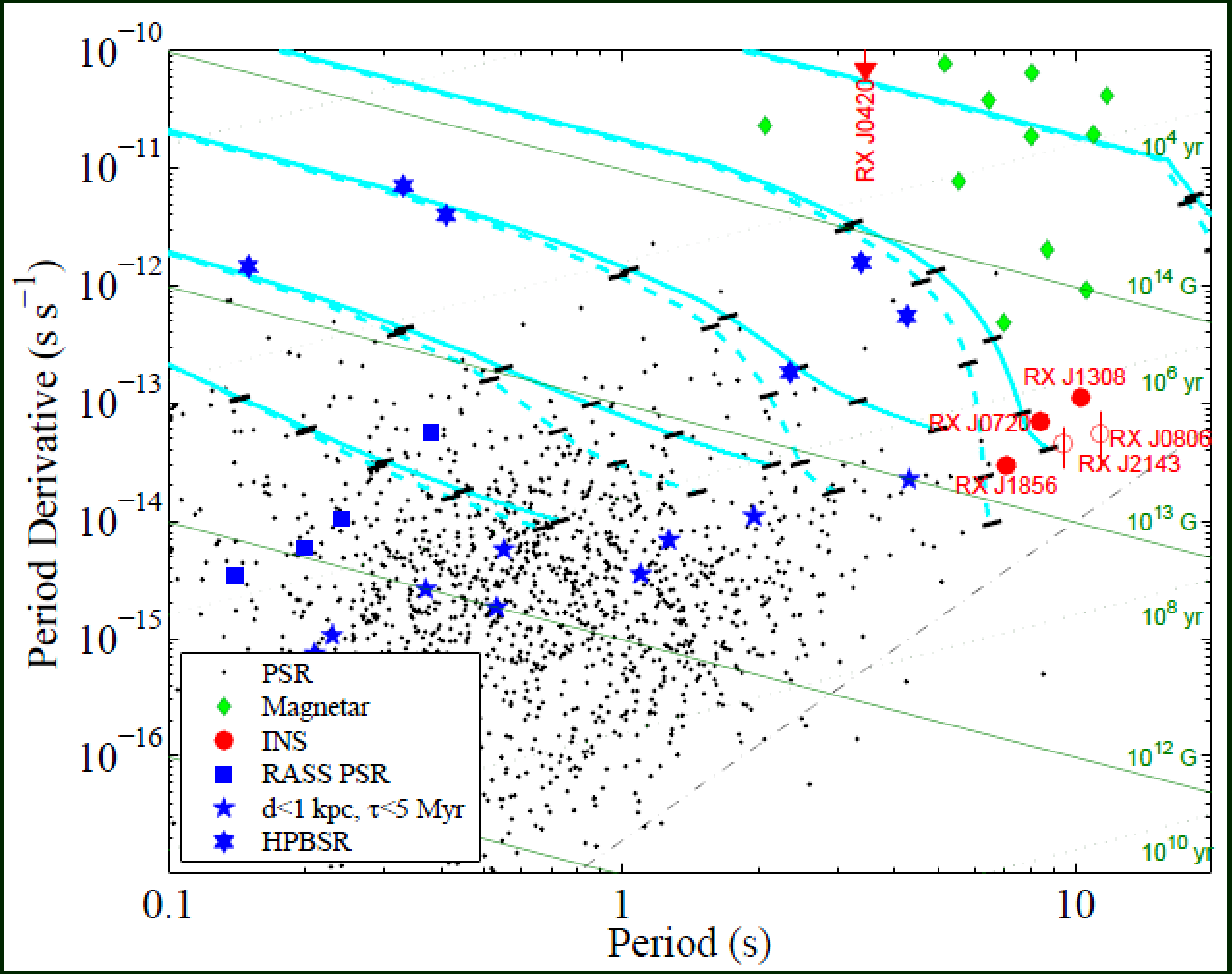






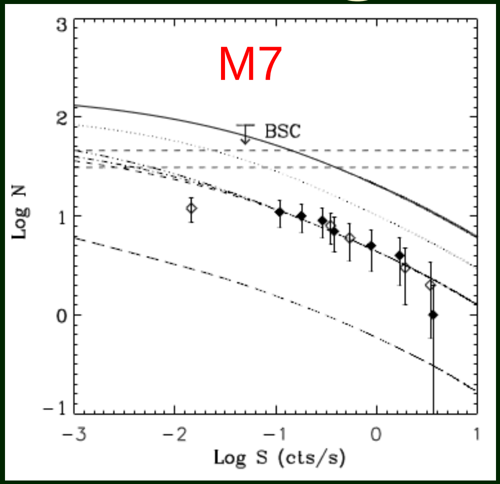
Magnetic field distribution is more important than the mass distribution.

Observational evidence?

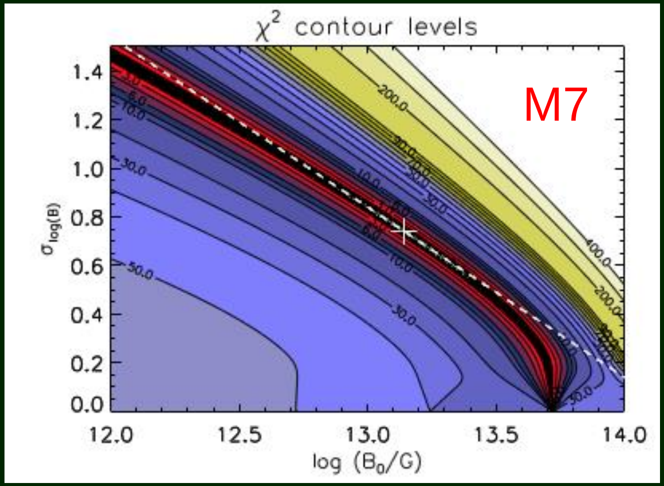


Kaplan & van Kerkwijk arXiv: 0909.5218

Extensive population synthesis: M7, magnetars, PSRs

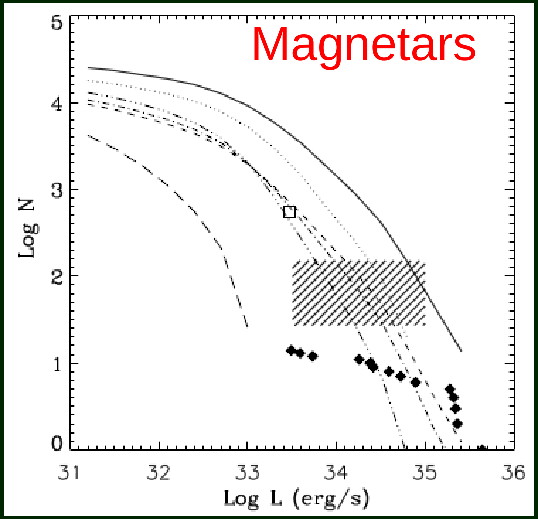


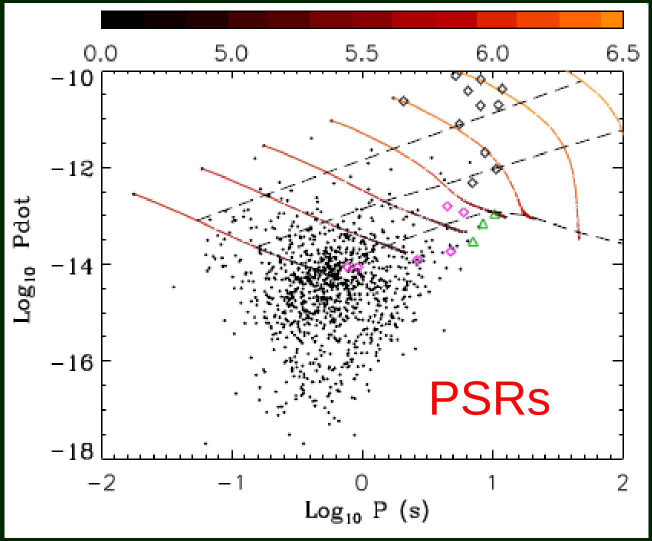
Using one population it is difficult or impossible to find unique initial distribution for the magnetic field



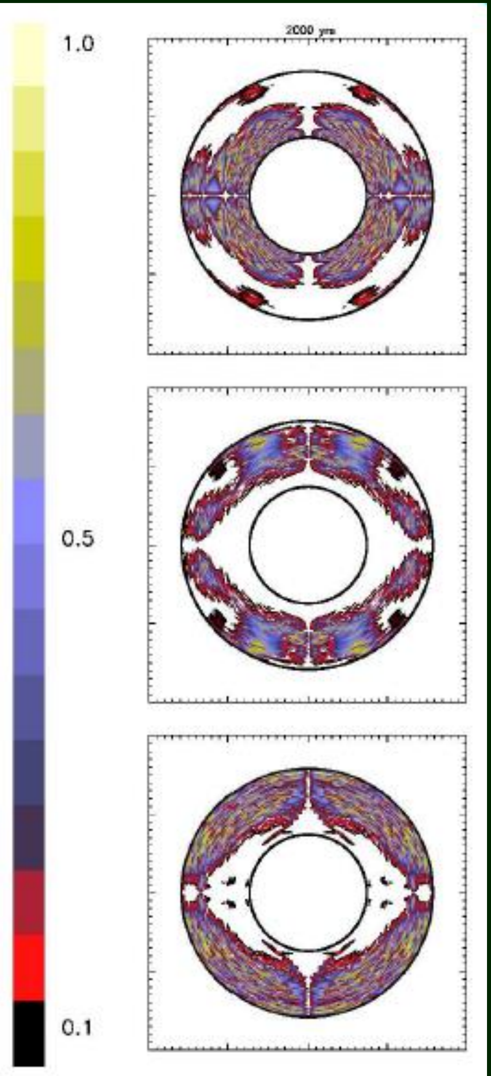
All three populations are compatible with a unique distribution.

Of course, the result is model dependent.



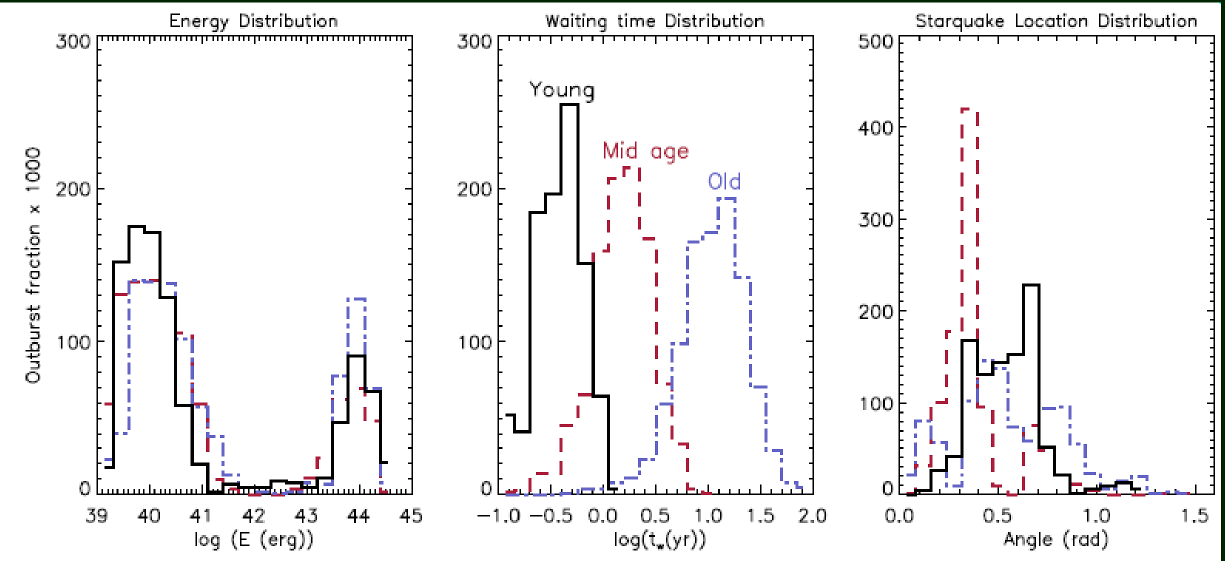


Magnetars bursting activity due to decay



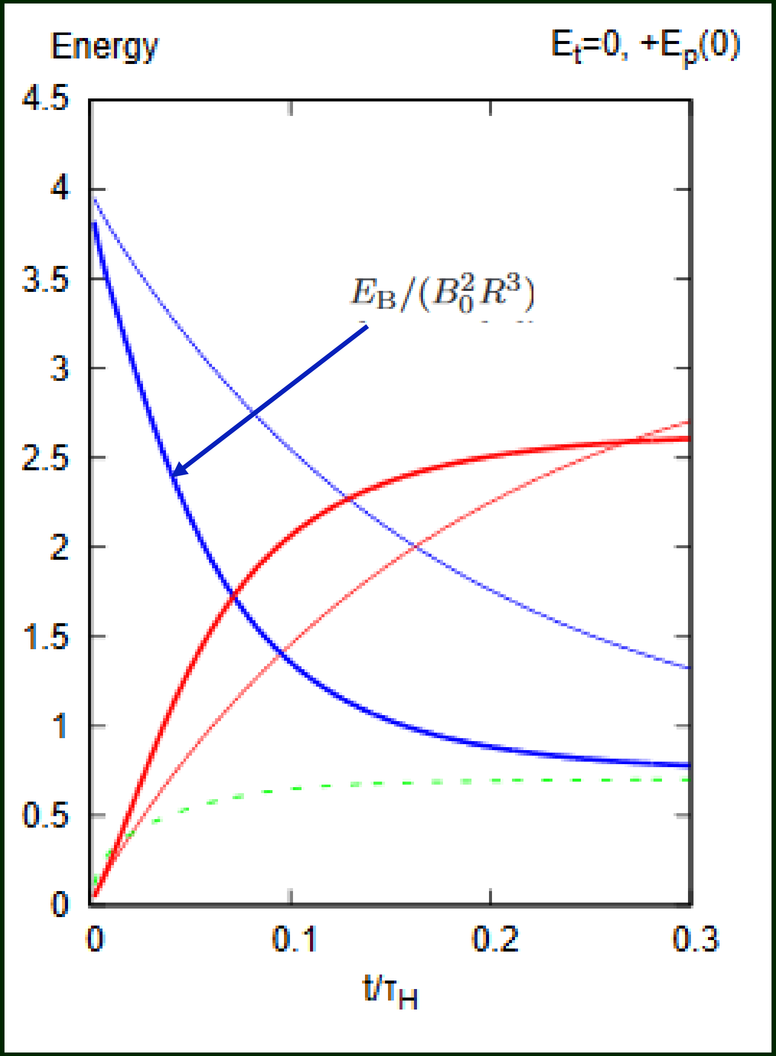
In the field decay model it is possible to study burst activity. Bursts occur due to crust cracking. The decaying field produce stresses in the crust that are not compensated by plastic deformations. When the stress level reaches a critical value the crust cracks, and energy can be released.

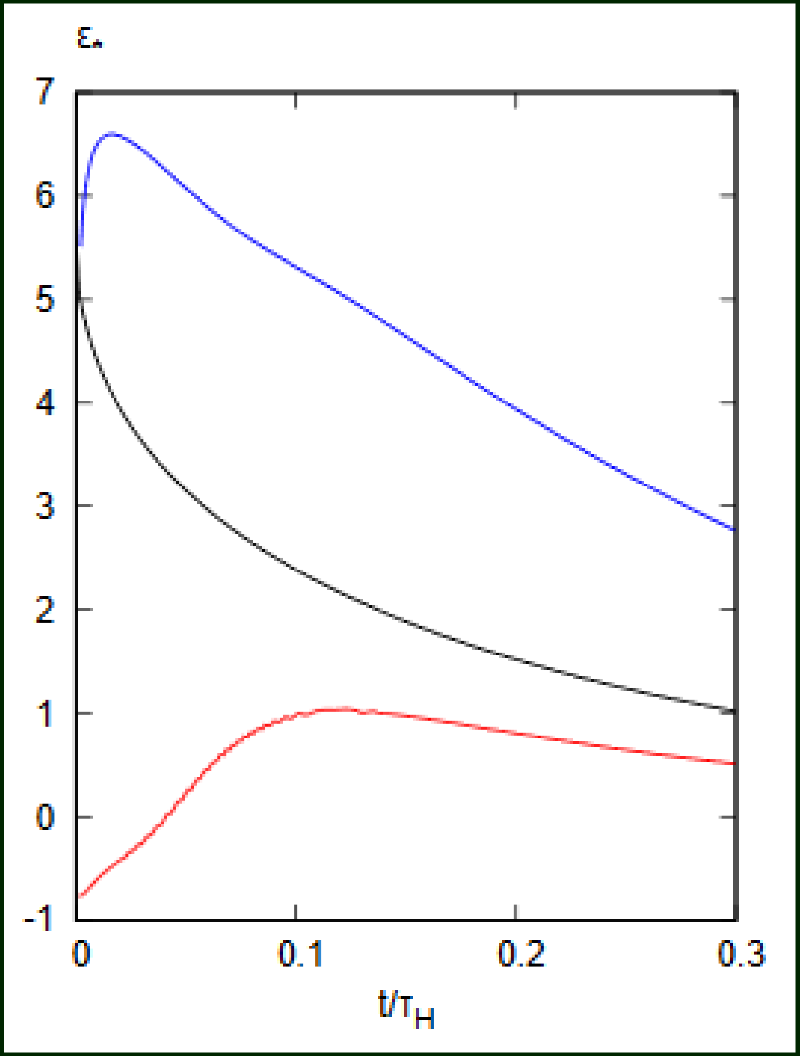
At the moment the model is very simple, but this just the first step.



1101.1098

Field evolution and ellipticity

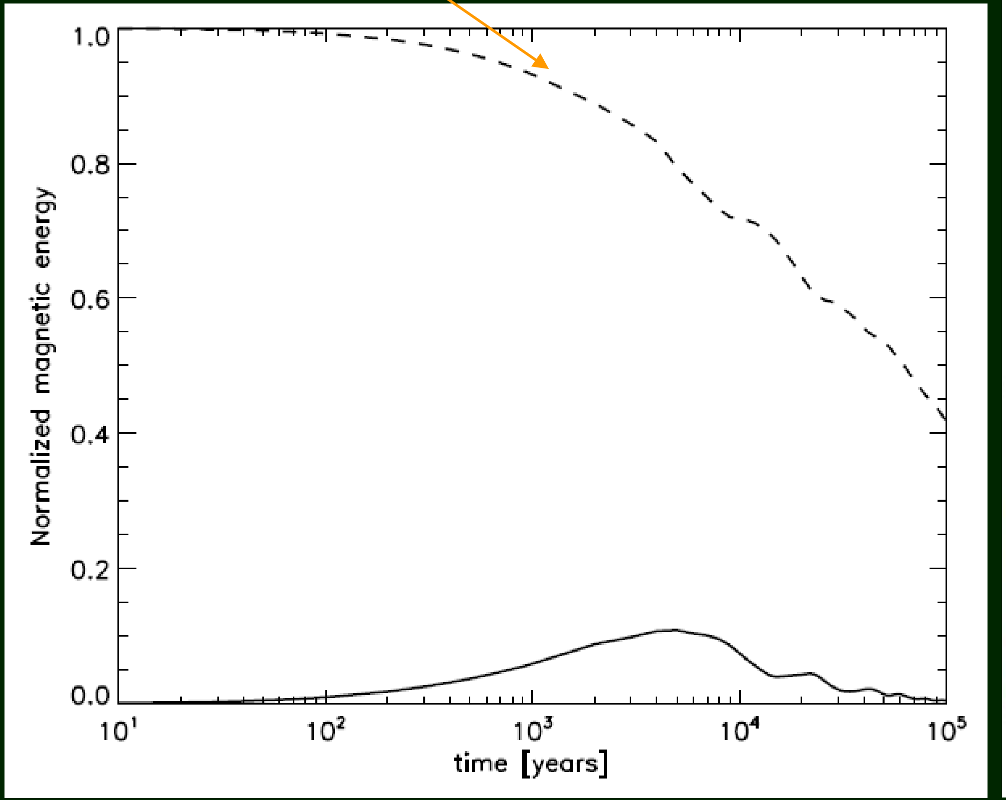


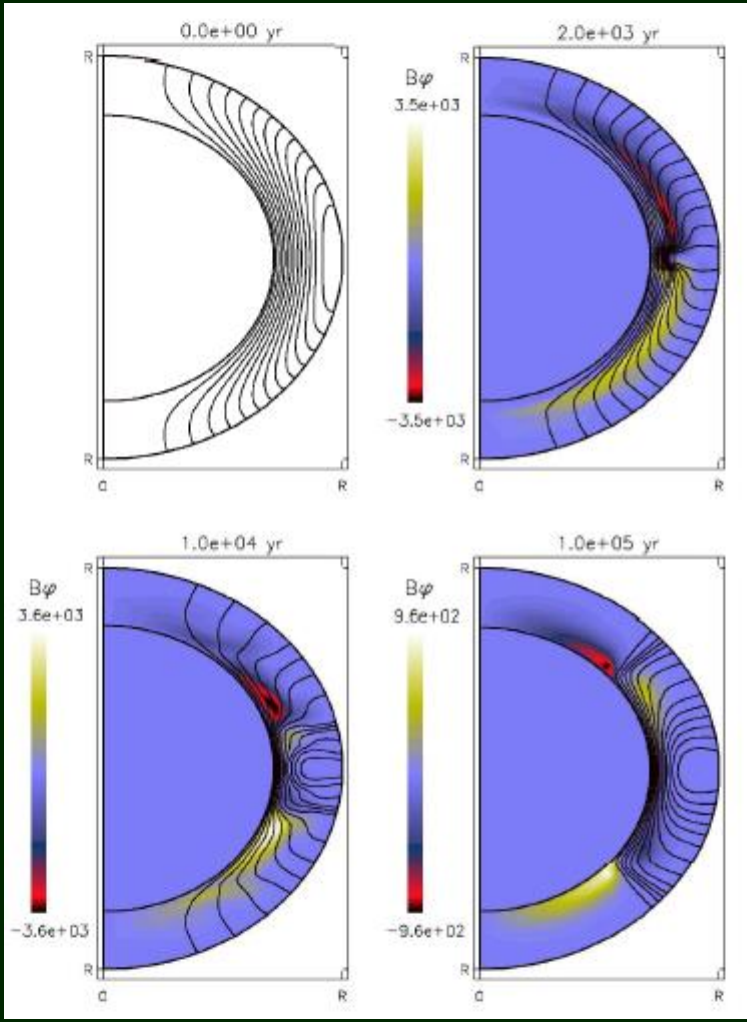


2011.03239

An illustrative model

Poloidal

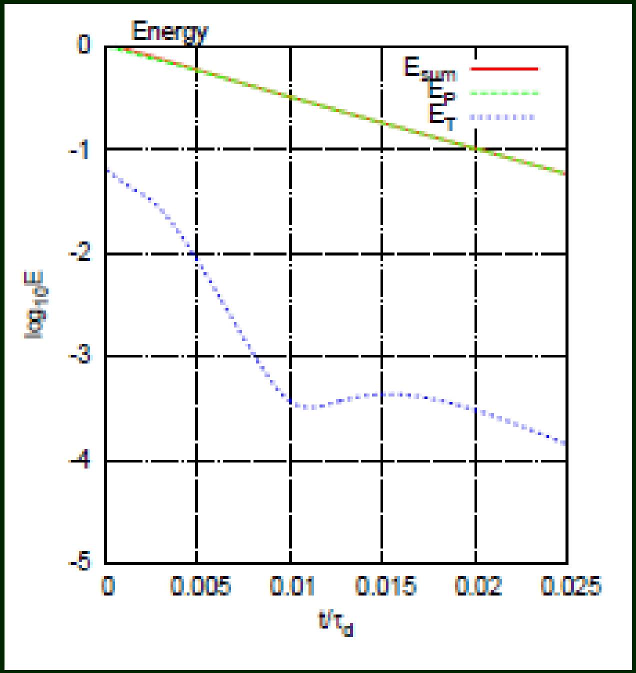




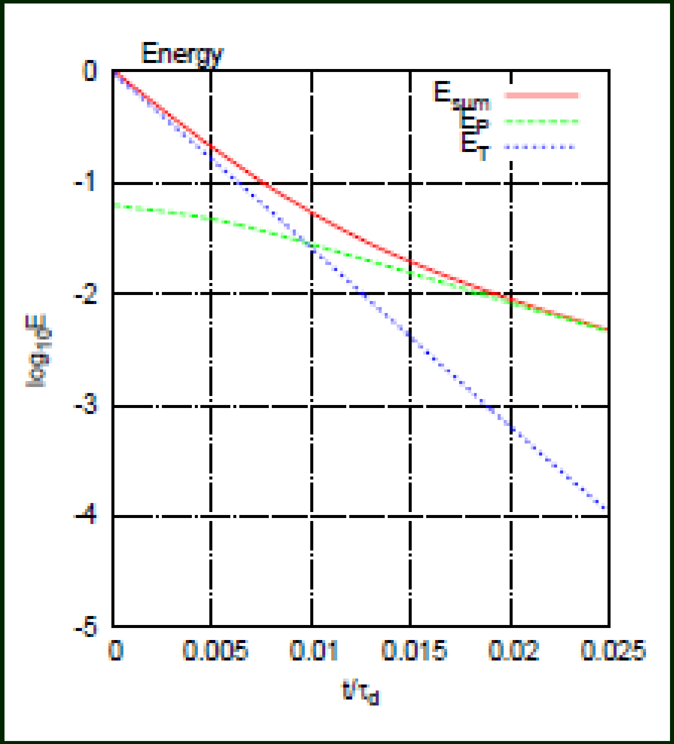
Test illustrates the evolution of initially purely poloidal field

1204.4707

Another model



Initially the poloidal field is large.

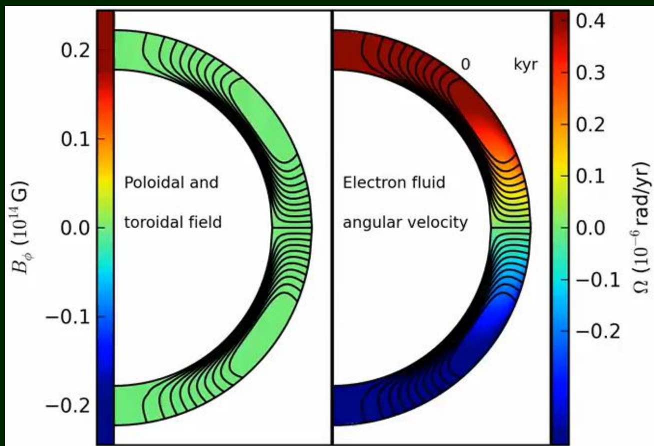


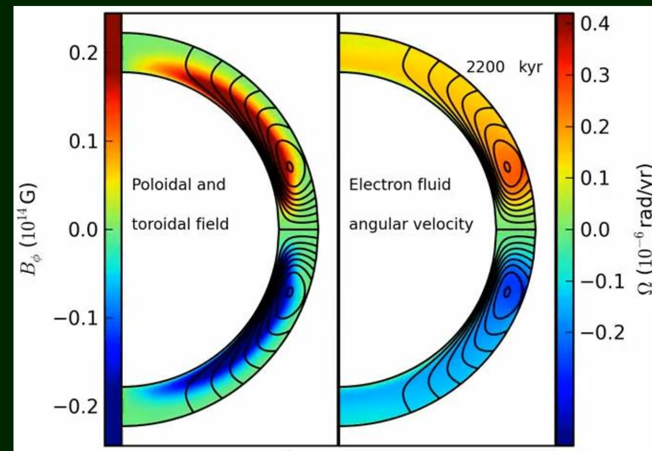
Initially the toroidal field is large.

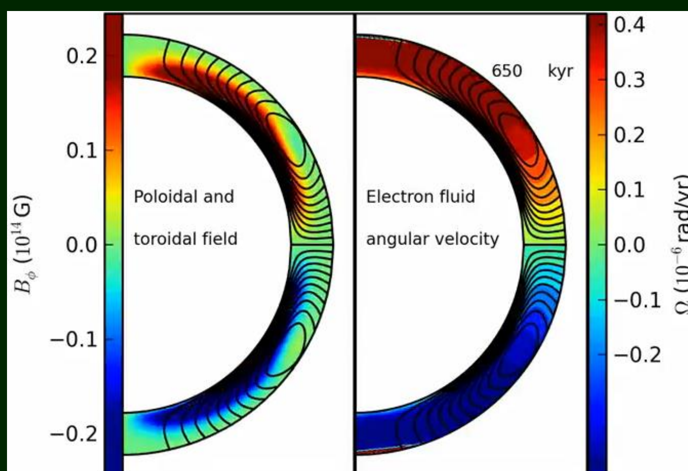
If the toroidal field dominates initially then significant energy is transferred to the poloidal component during evolution.
In the opposite case, when the poloidal component initially dominates, energy is not transferred. The toroidal component decouples.

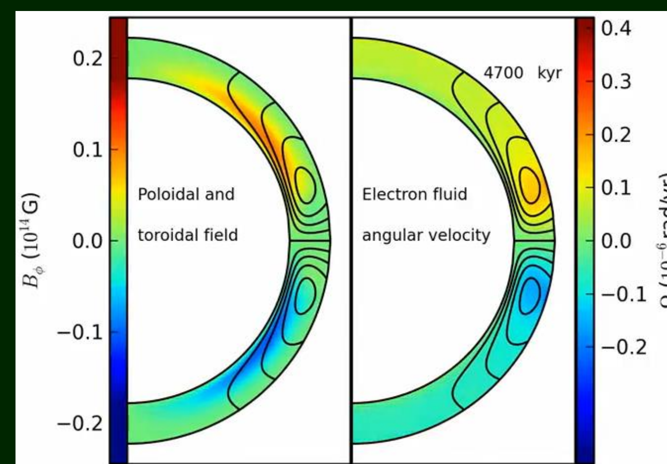
1201.1346

Hall cascade and attractor







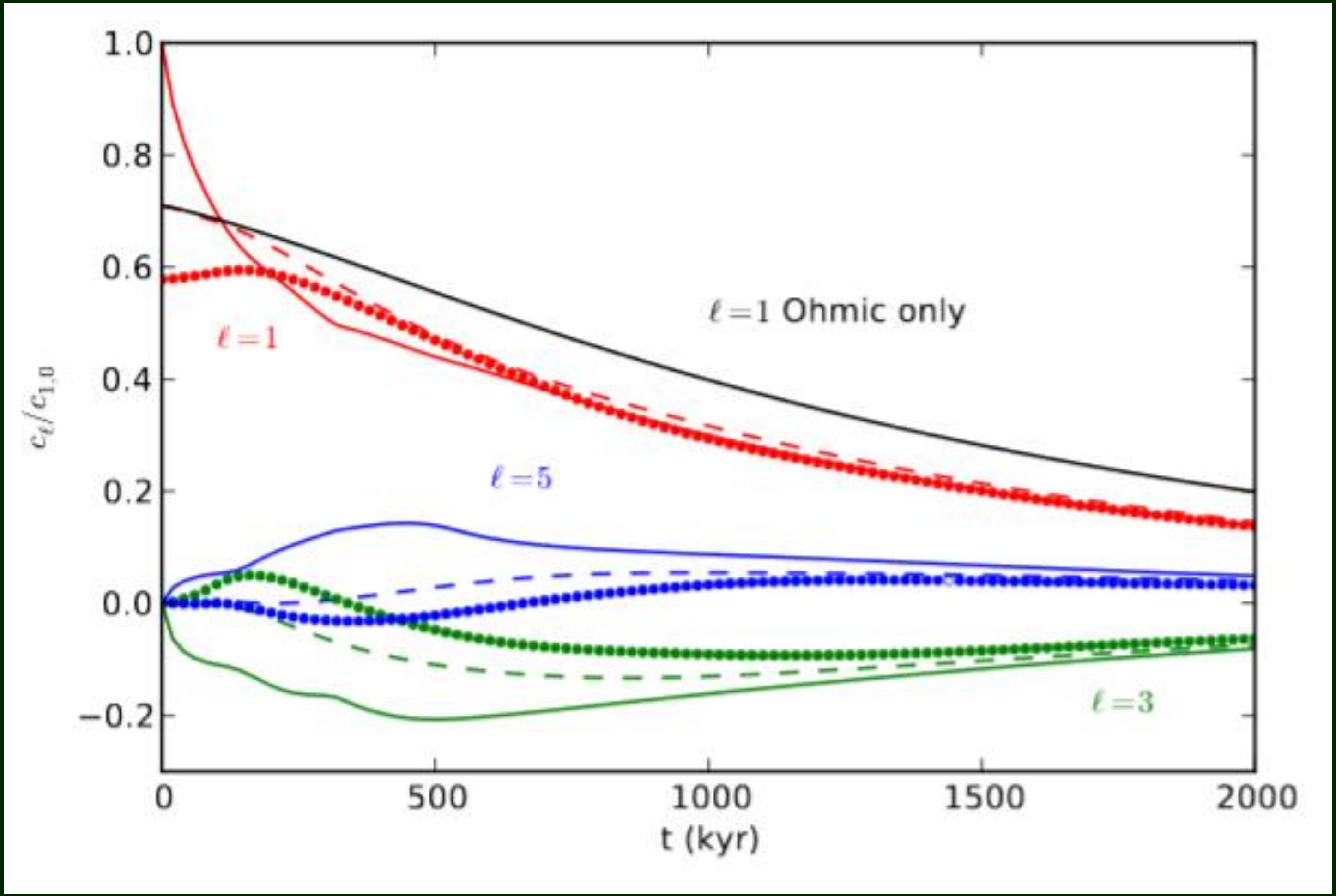


Gourgouliatos and Cumming 2013

Hall cascade can reach the stage of so-called Hall attractor,
where the field decay stalls for some time (Gourgouliatos, Cumming).

http://www.physics.mcgill.ca/~kostasg/research.html

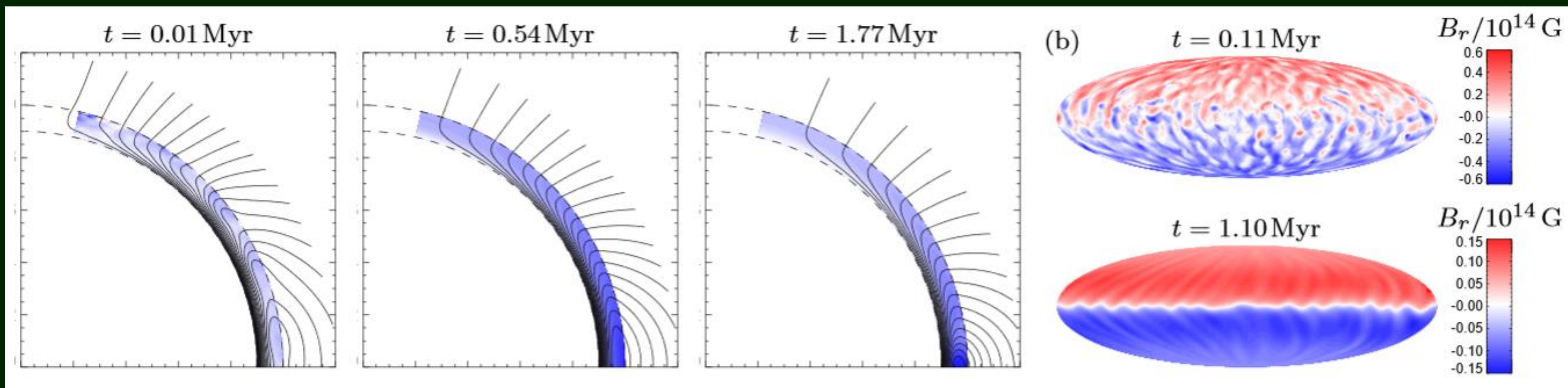
Evolution of different components

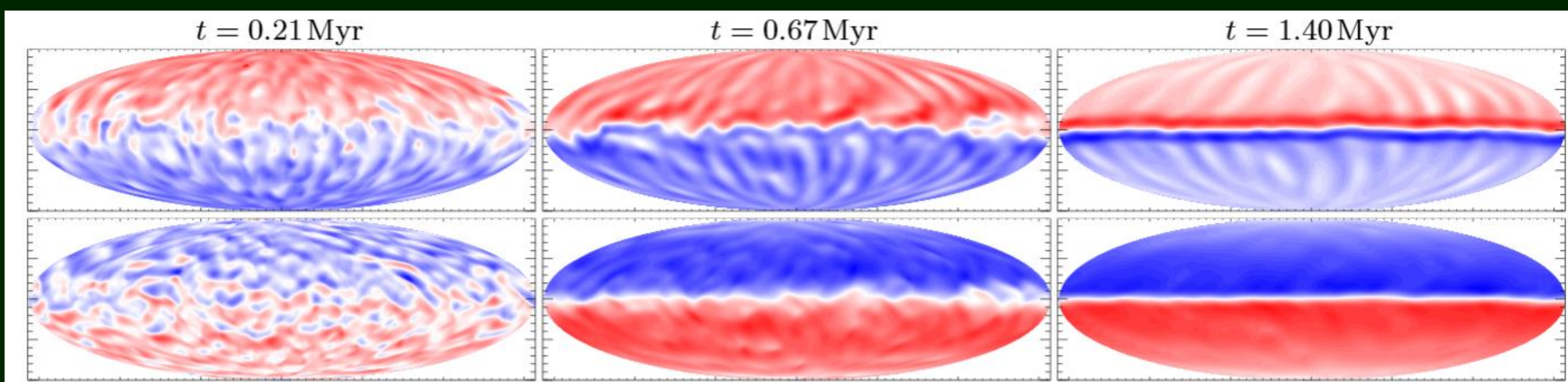


Hall attractor mainly consists of dipole and octupole (+l5)

1311.7004

New studies of the hall cascade



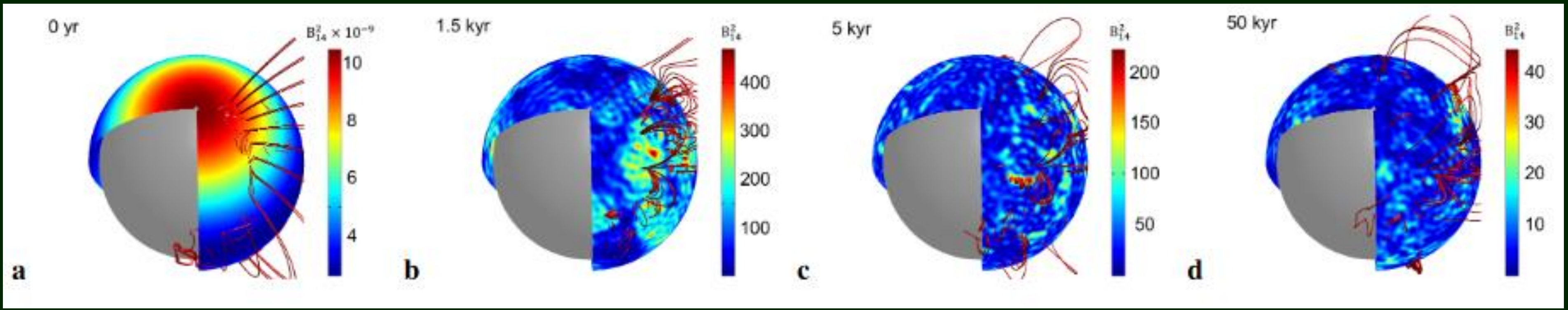


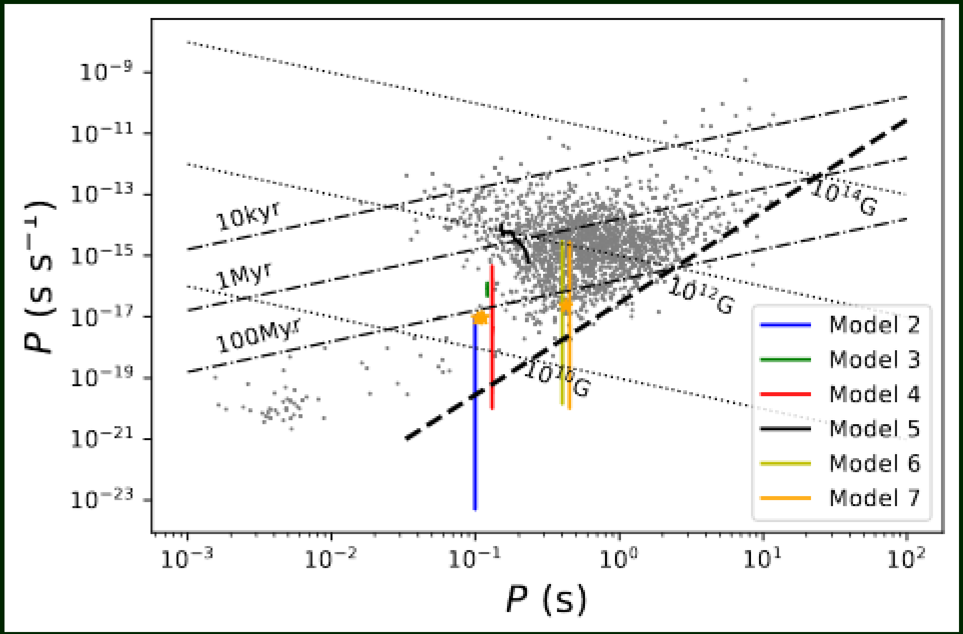
New calculations support the idea of a kind of stable configuration.

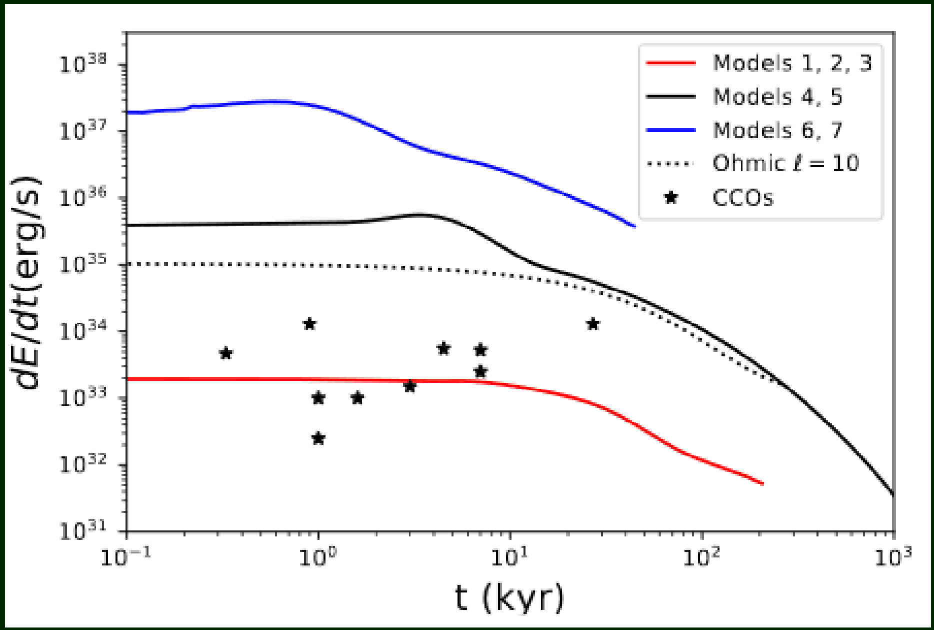
1501.05149

Tangled initial field

Can be important to explain CCOs



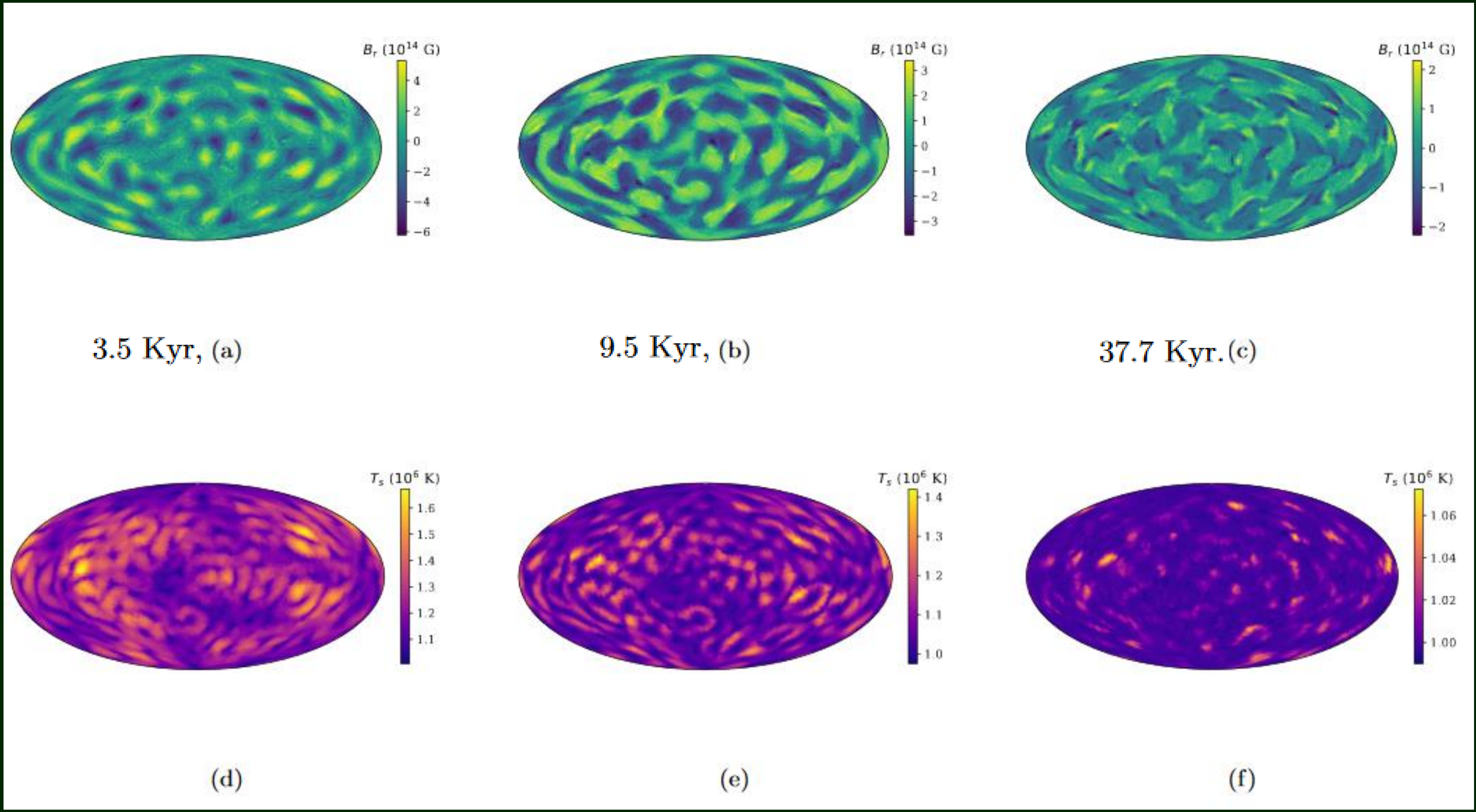




2005.02410

Tangled field - 2

stochastic dynamo scenario



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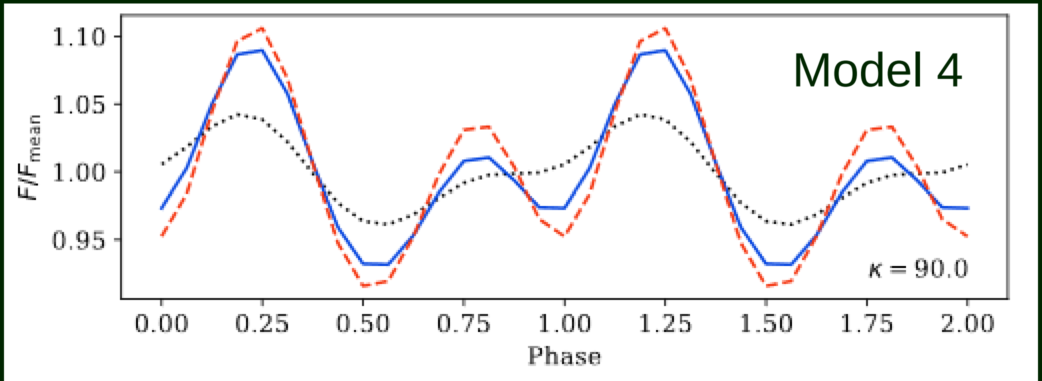
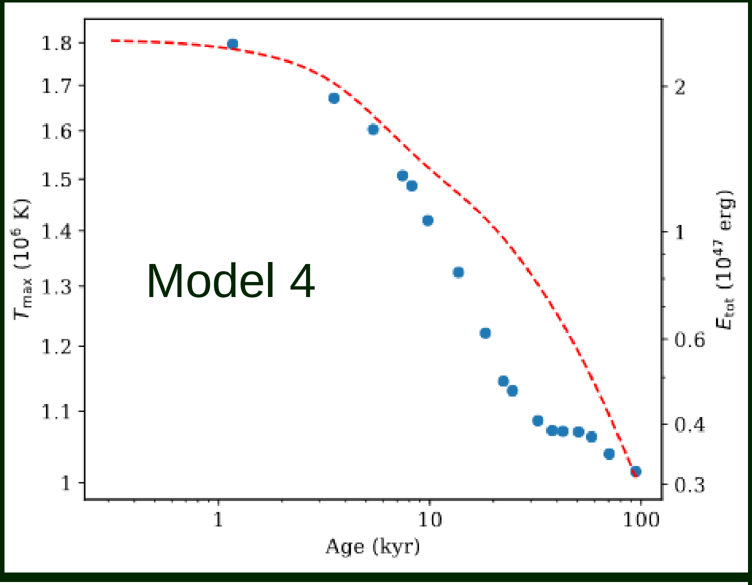
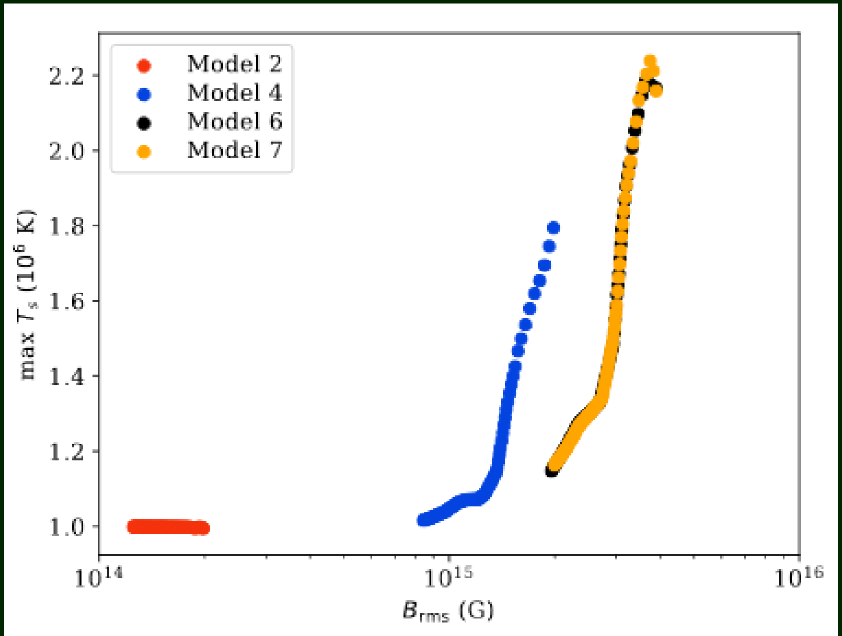
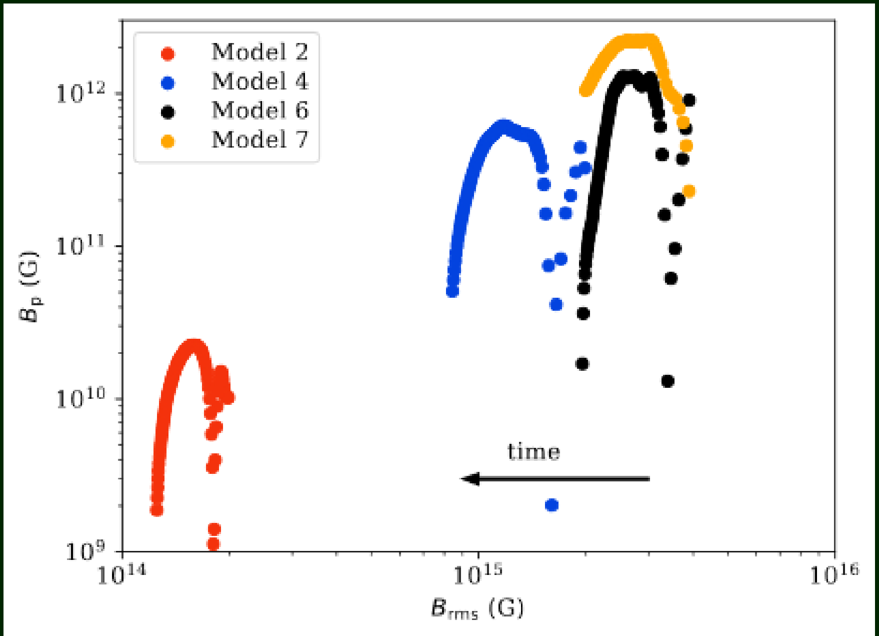
Tangled field - 2









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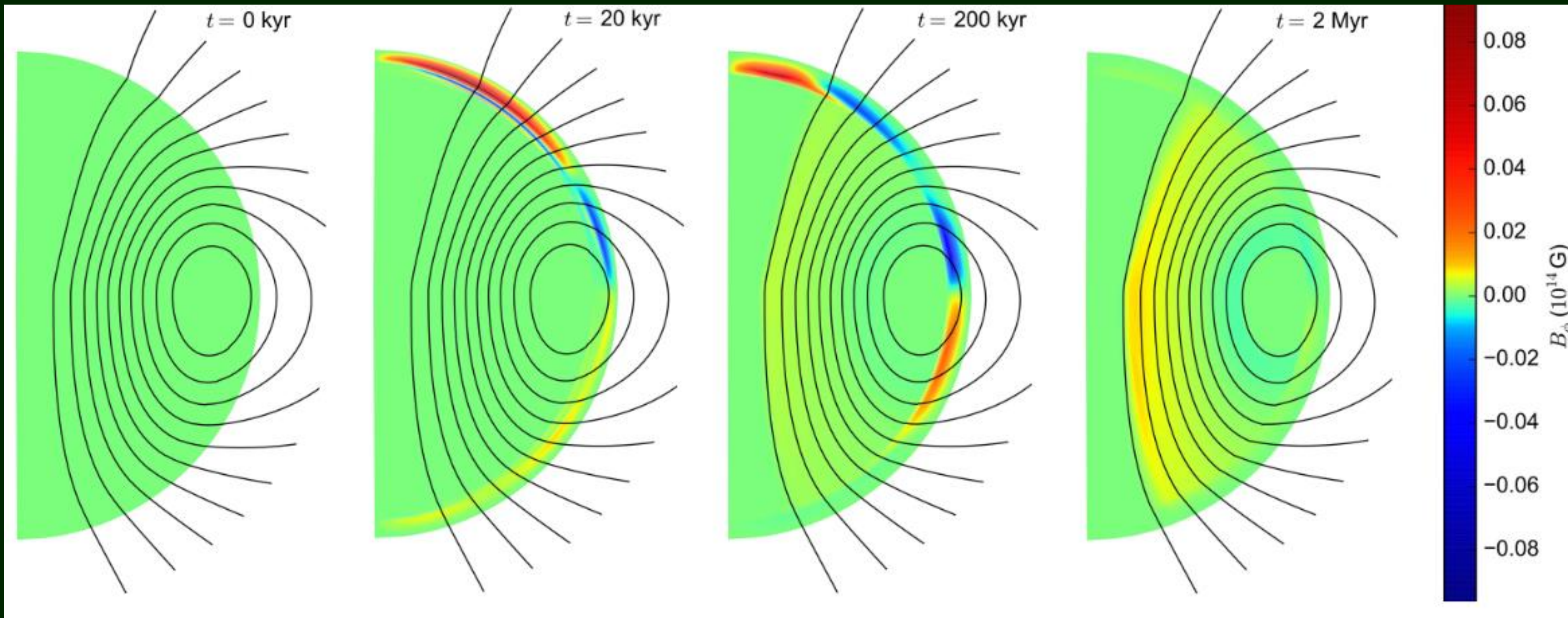
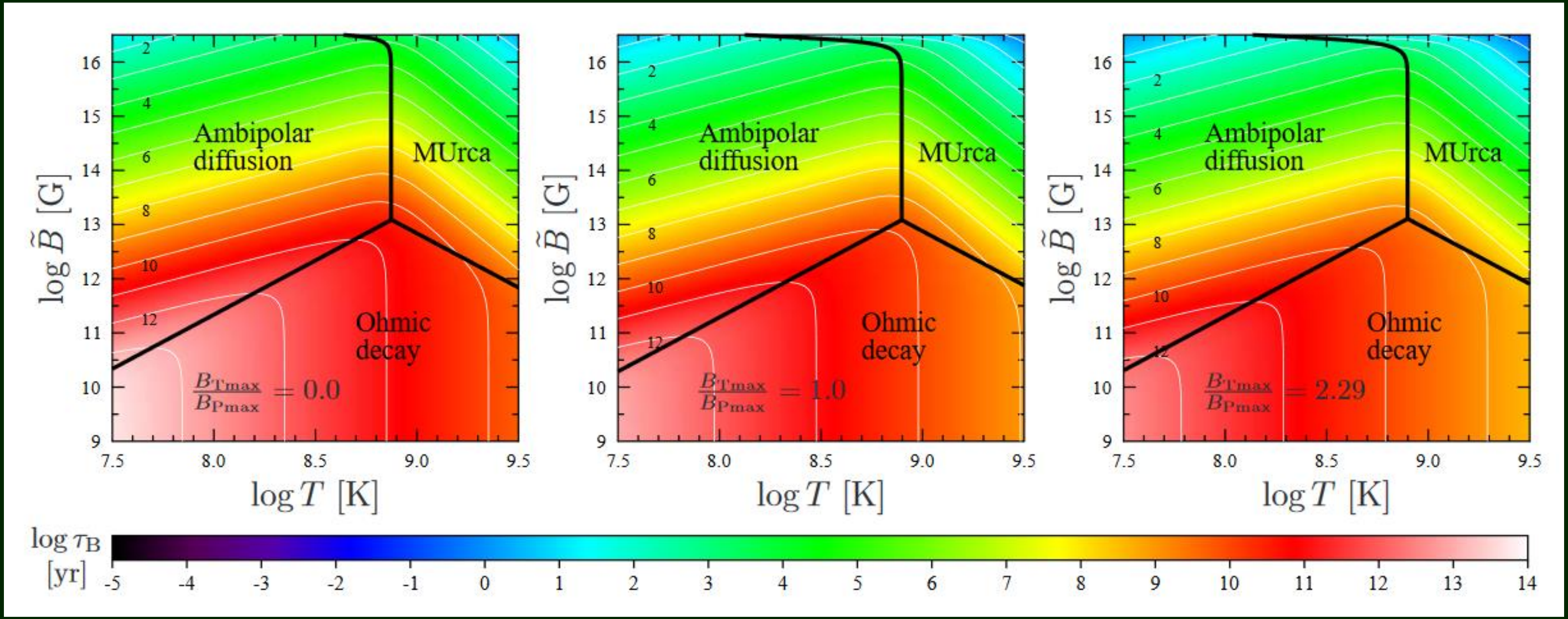
Core and crust field evolution



Hall attractor is confirmed.

1709.09167

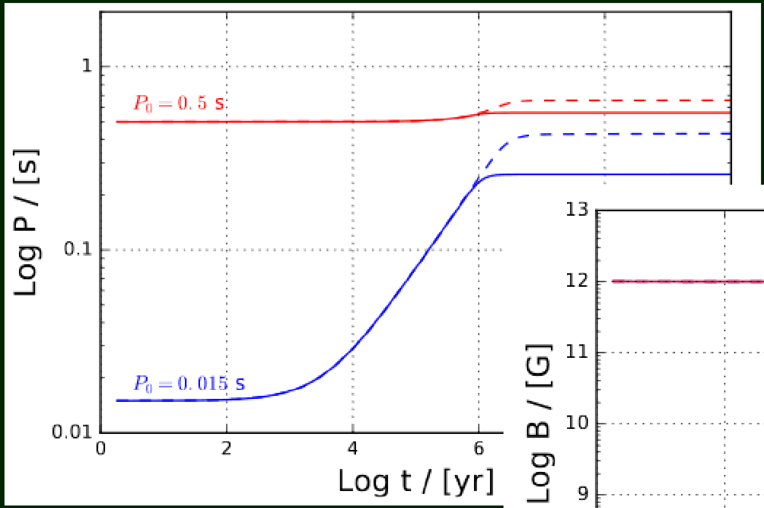
Core field evolution



Typical timescales for the magnetic field dissipation as functions of temperature and the magnetic field strength.

1705.00508, 1805.03956, 2010.07673

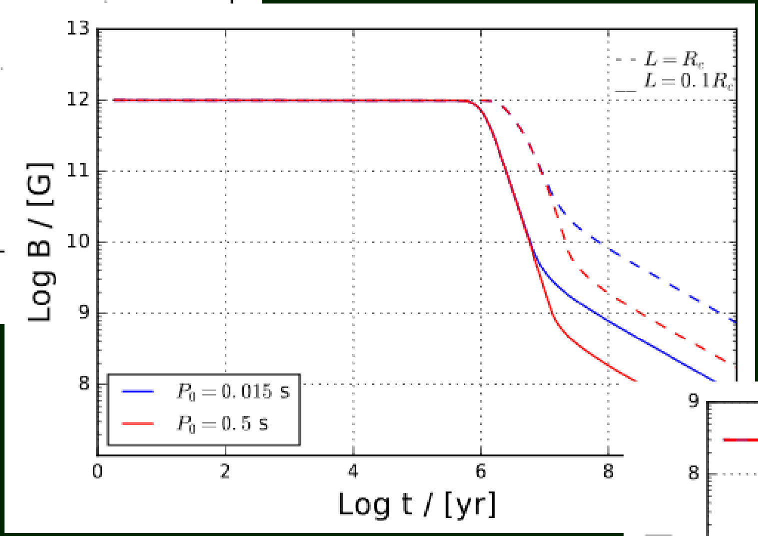
Field evolution due to ambipolar diffusion

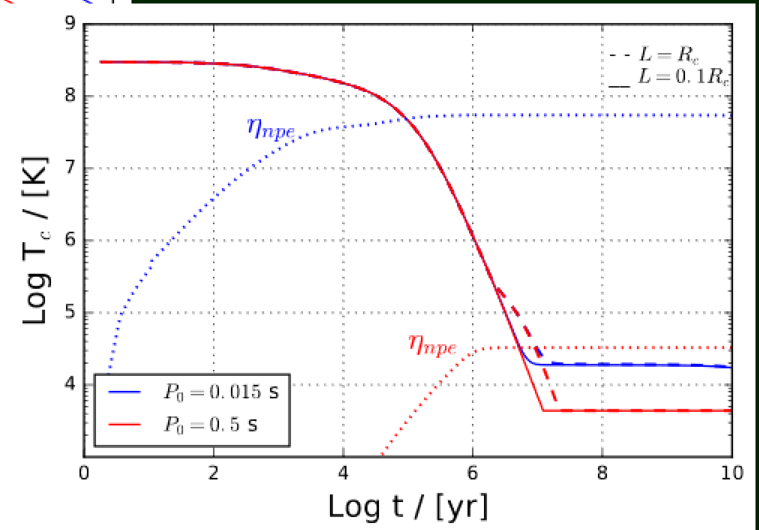


Hypothesis: field decay in MSPs is caused by ambipolar diffusion in the NS core in the non-superfluid/superconductor regime.

$$t_{Ohm} \ll t_{AD}$$

$$\dot{B} \approx -\frac{B}{t_{AD}},$$





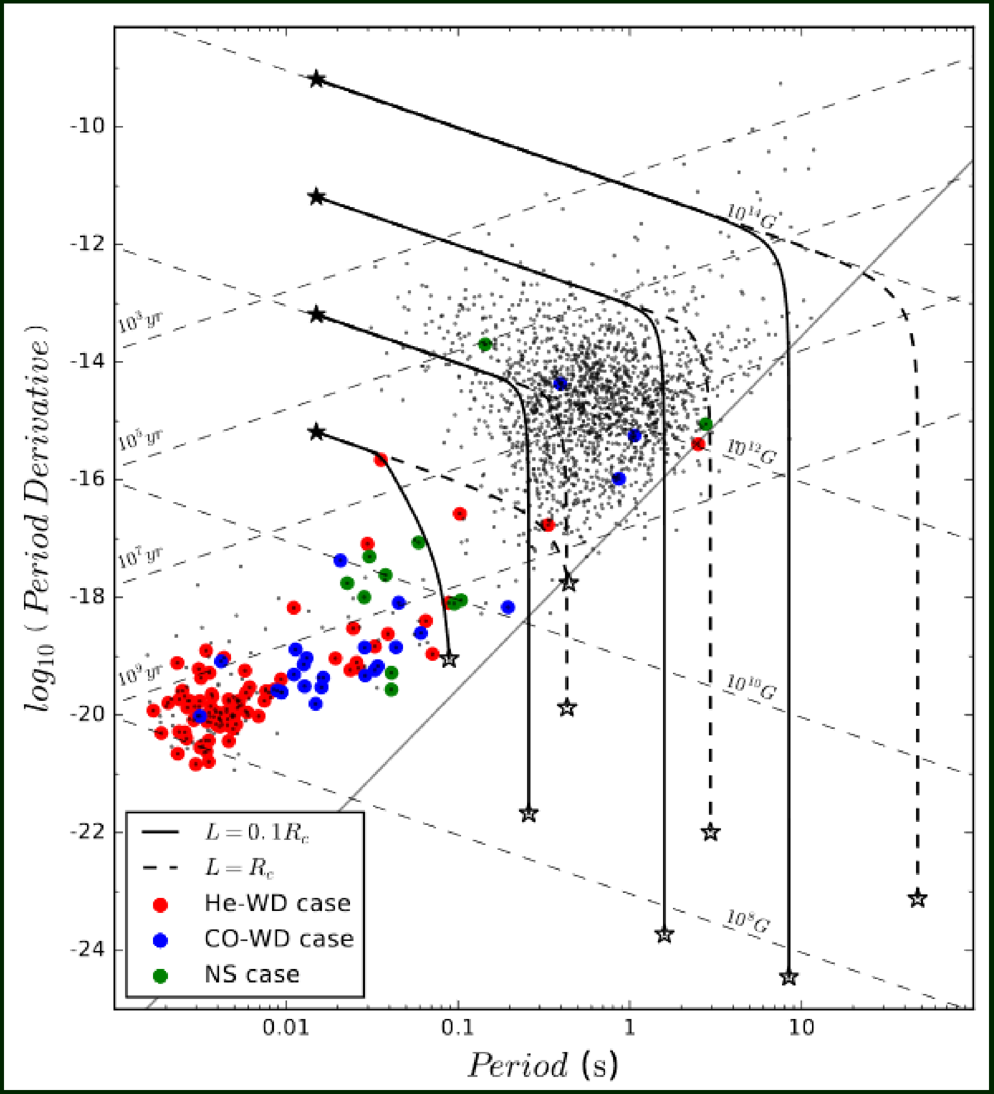
the magnetic field is transported by the charged particles at the ambipolar diffusion velocity

$$\frac{\partial \vec{B}}{\partial t} = \nabla \times \left(\vec{v}_{AD} \times \vec{B}\right)$$

$$t_{AD} = 3 \times 10^9 \frac{T_4^2 L_5^2}{B_8^2} \text{ yr},$$

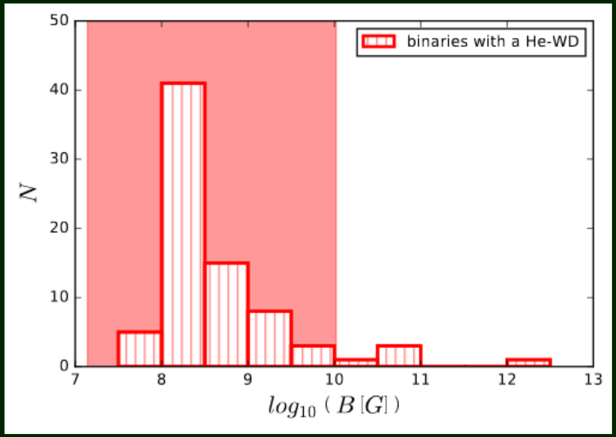
1906.06076

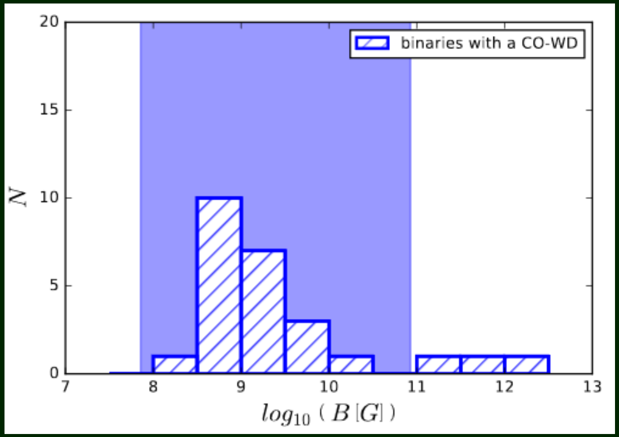
Evolution on the P-Pdot diagram

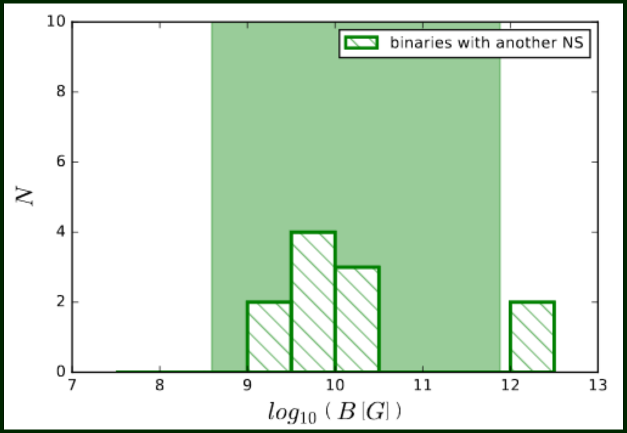


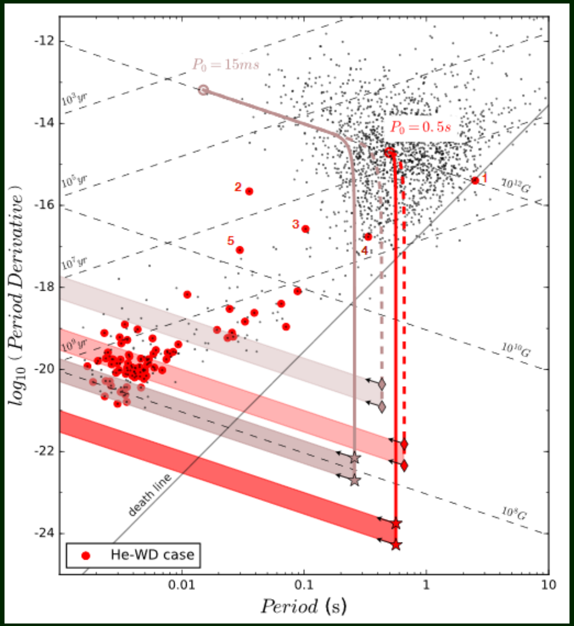
1906.06076

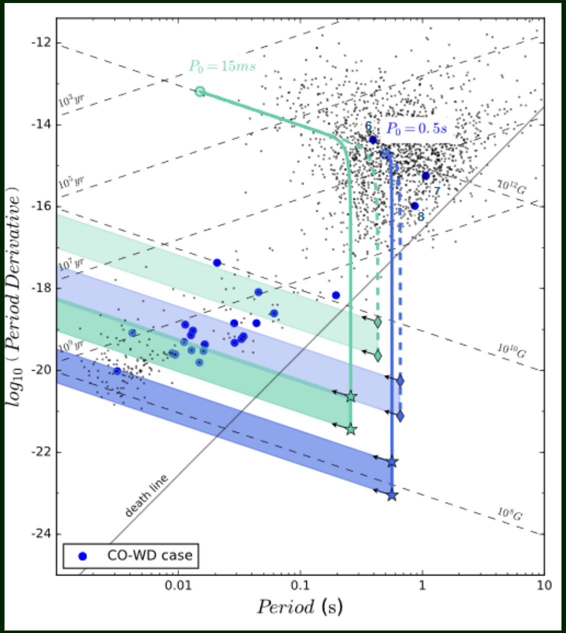
Different types of companions

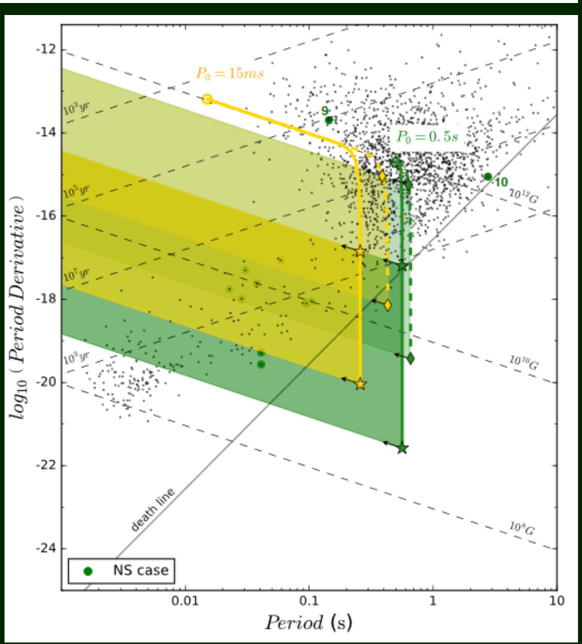






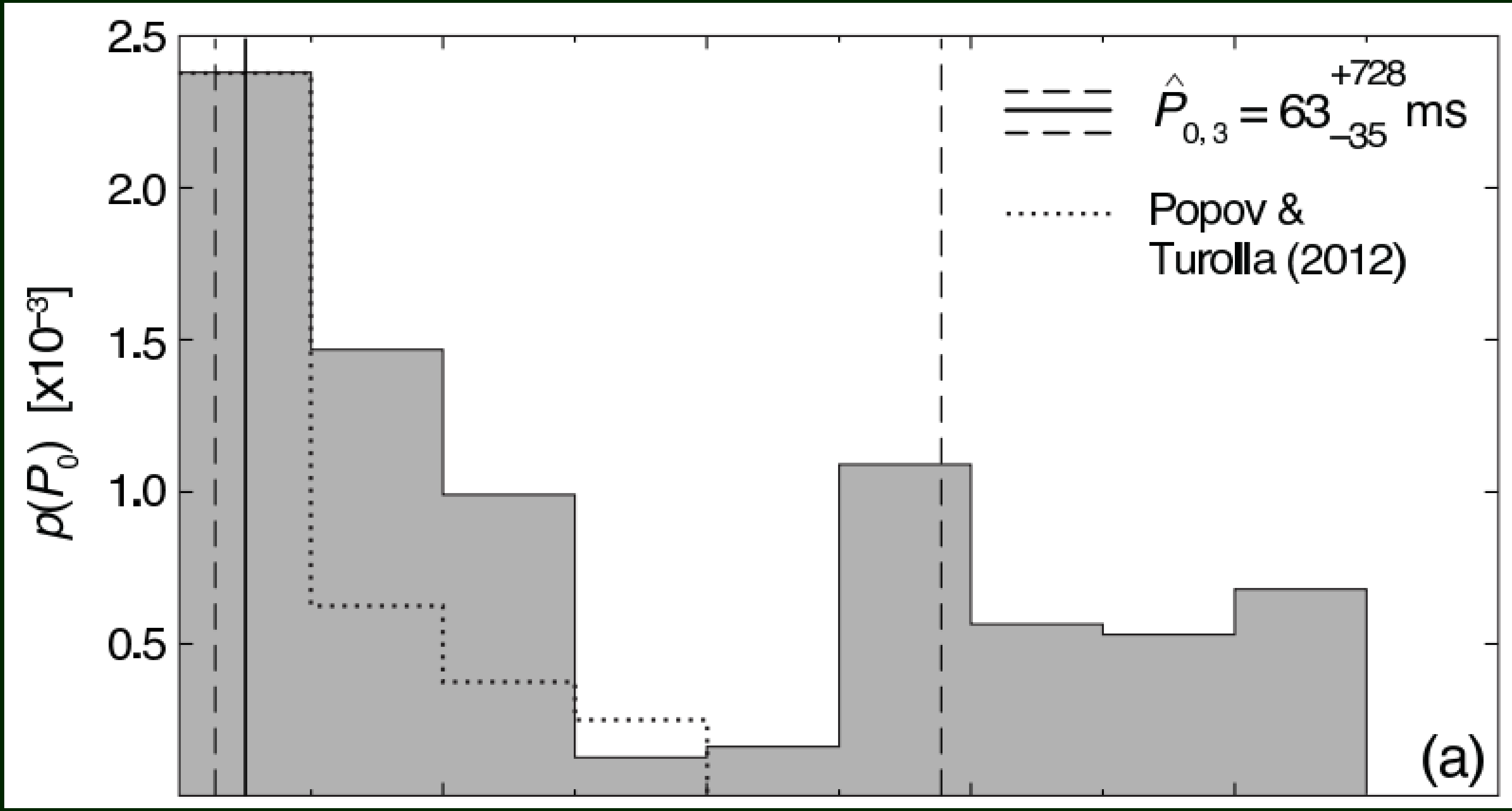






1906.06076

Wide initial spin period distribution



Based on kinematic ages. Mean age – few million years.
Note, that in Popov & Turolla (2012) only NSs in SNRs were used, i.e. the sample is much younger!
Can it explain the difference?

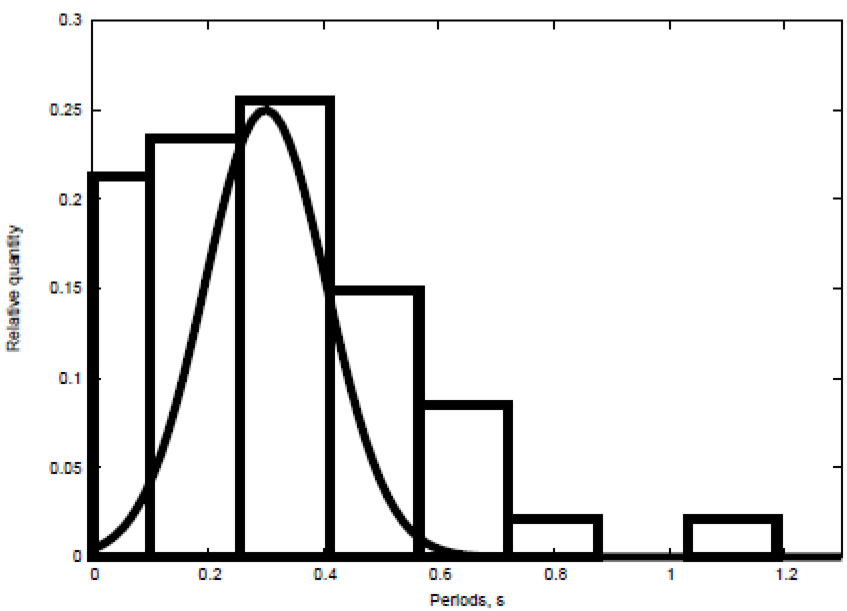
1301.1265

Magnetic field decay and P_0

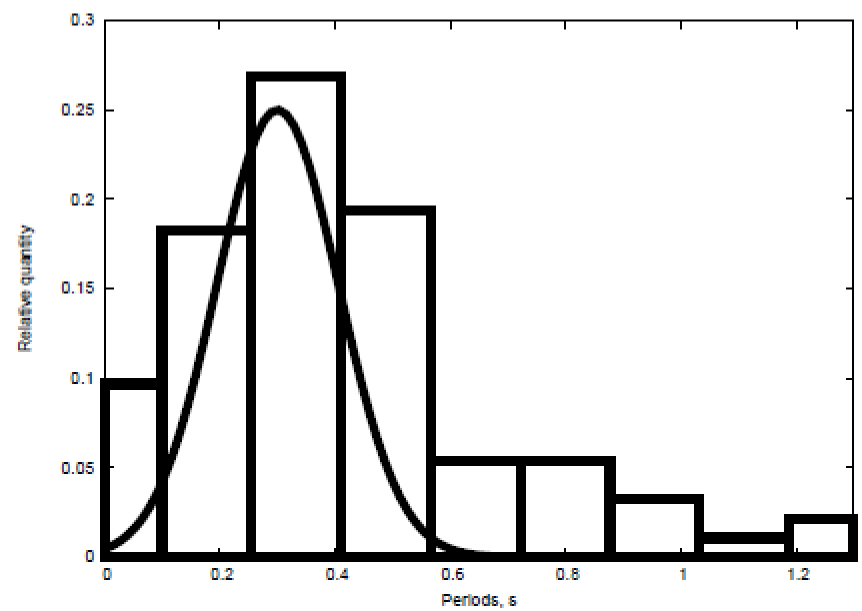
One can suspect that magnetic field decay can influence the reconstruction of the initial spin period distribution.

Exponential field decay with $\tau=5$ Myrs.
$<P_0>=0.3$ s, $\sigma_P=0.15$ s; $<\log B_0/[G]>=12.65$, $\sigma_B=0.55$

$$P_0 = P\sqrt{1-\frac{t}{\tau}}.$$



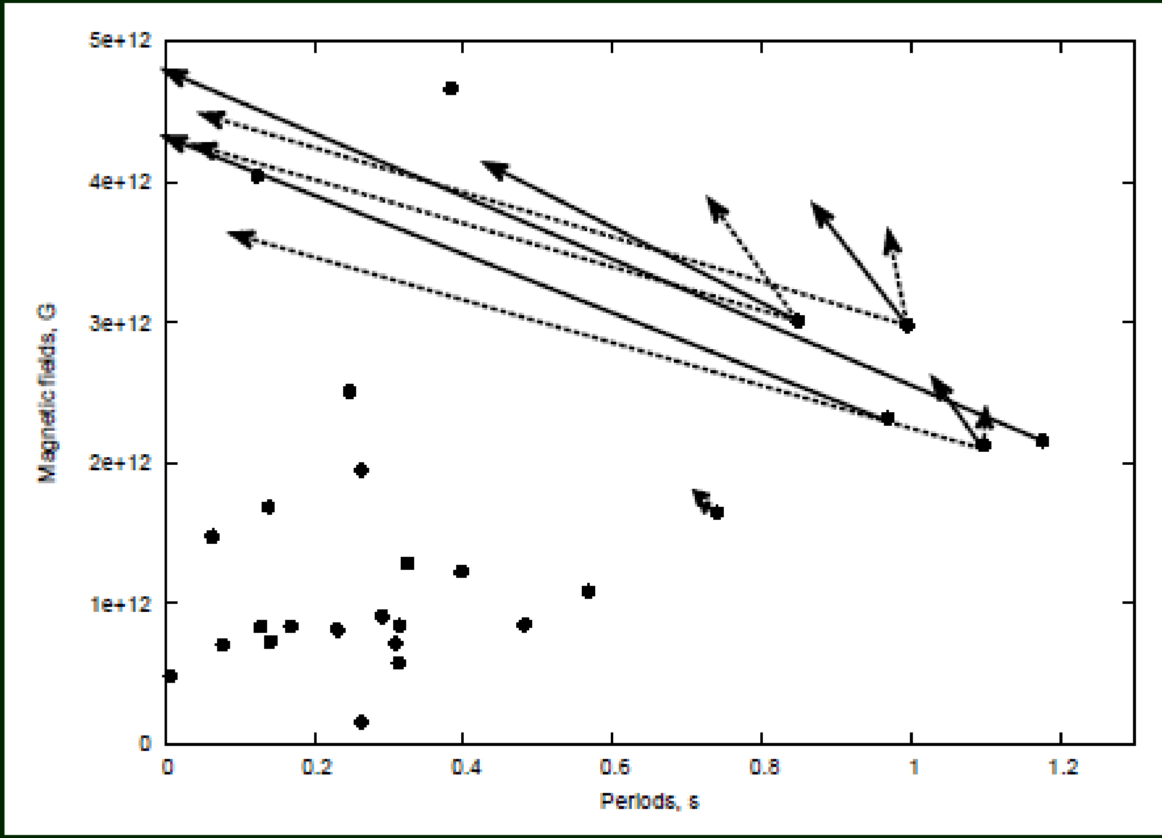
$\tau<10^7$ yrs, $10^5<t$



$10^5<t<10^7$ yrs

Igoshev, Popov 2013

Real vs. reconstructed P_0



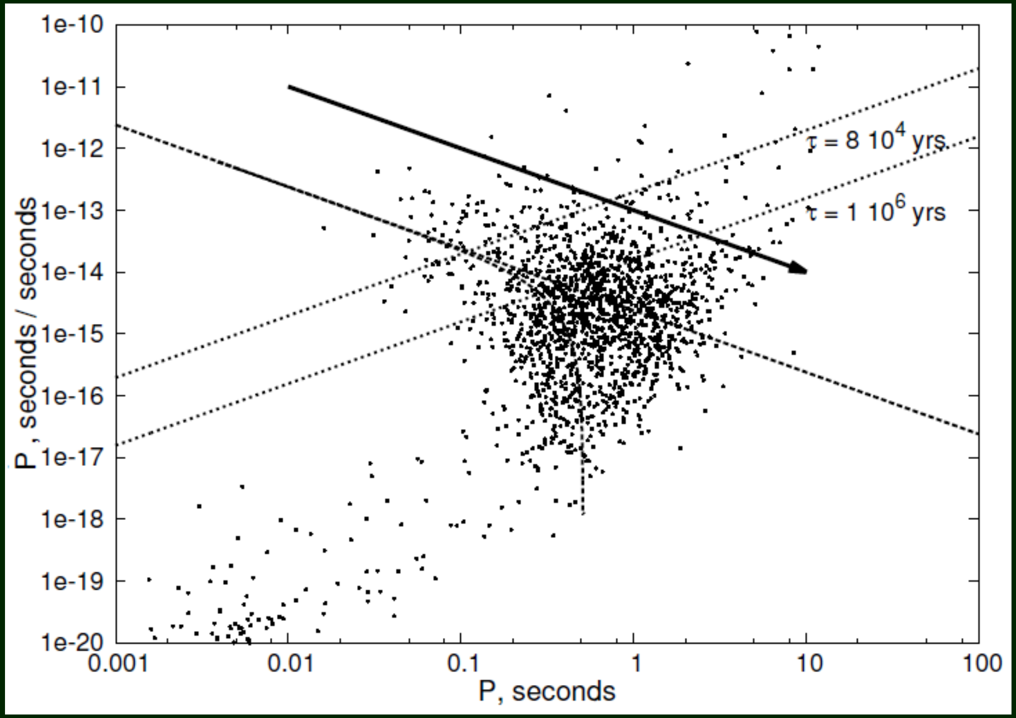
Arrows point to initial parameters of pulsars if the exponential magnetic field decay was operating.

How significantly the reconstructed initial periods changed due to not taking into account the exponential field decay

Igoshev, Popov 2013

Modified pulsar current

We perform a modified pulsar current analysis.
In our approach we analyse the flow not along the spin period axis, as it was done in previous studies, but study the flow along the axis of growing characteristic age.



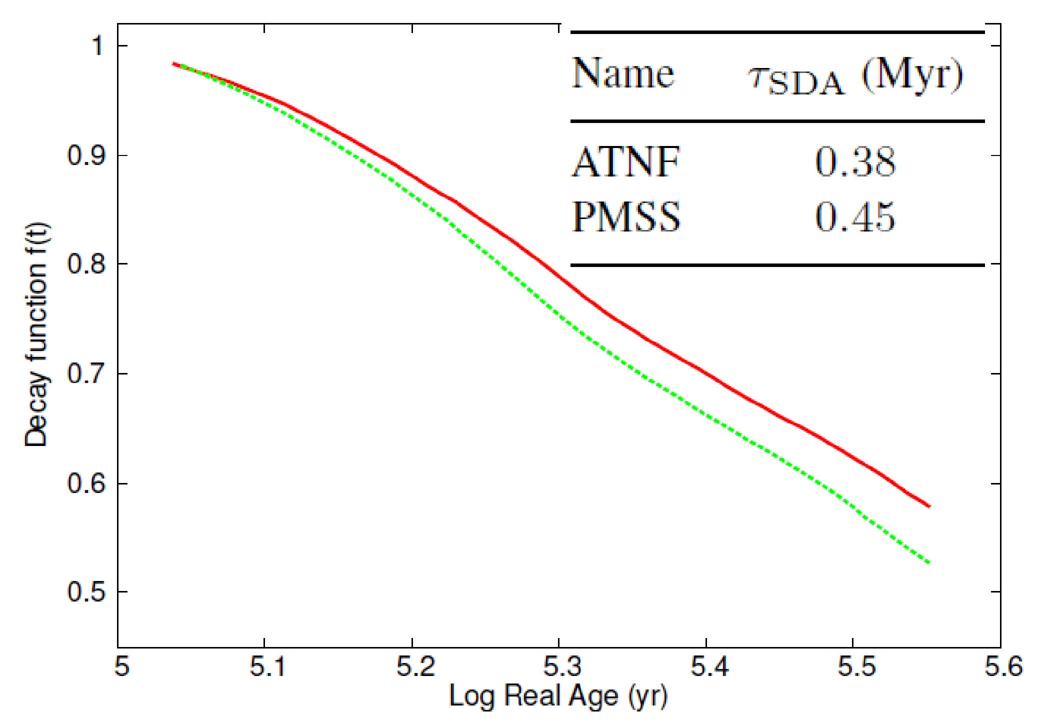
The idea is to probe magnetic field decay.
Our method can be applied only in a limited range of ages.

We use distribution in characteristic ages to reconstruct the field evolution.

Igoshev, Popov (2014). MNRAS

Application to real data

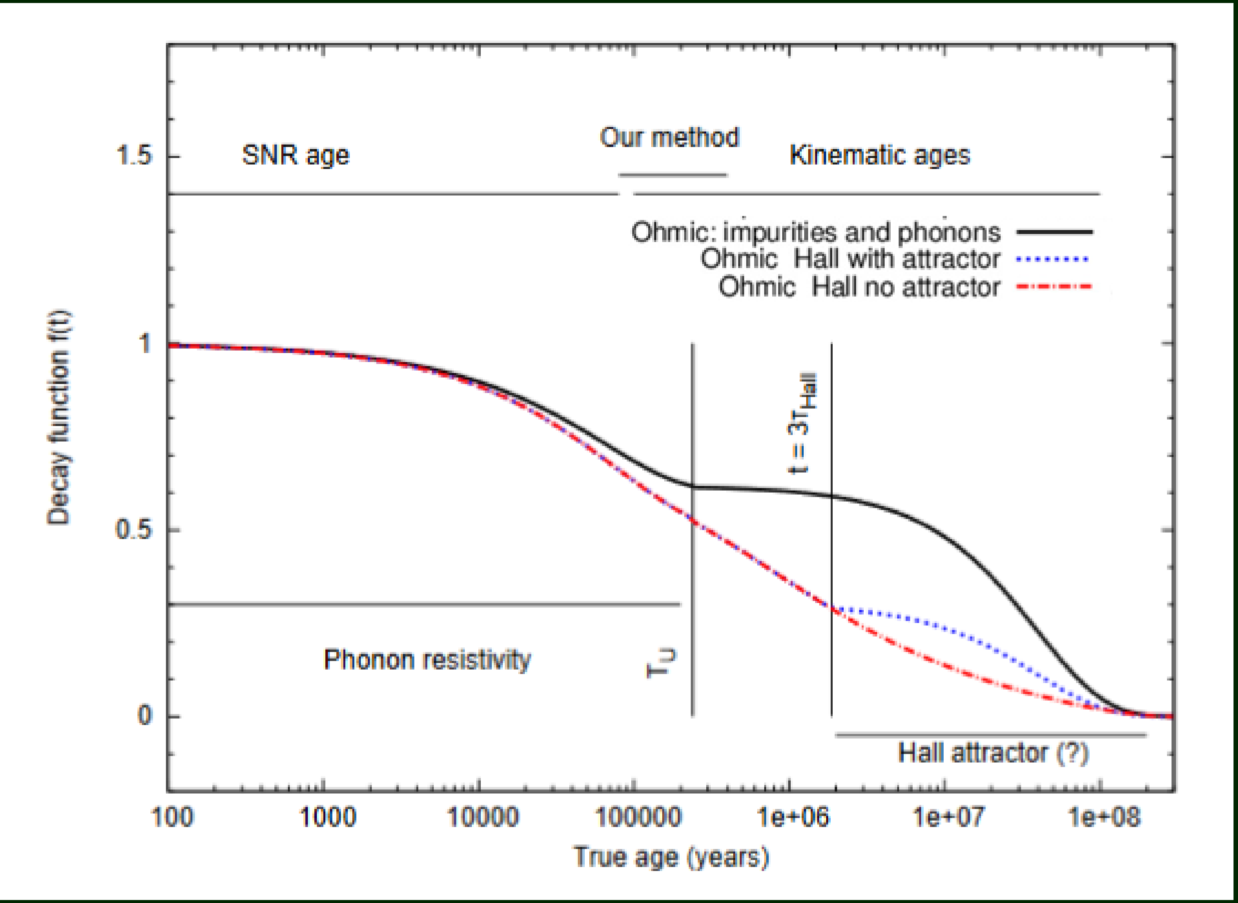
We apply our methods to large observed samples of radio pulsars to study field decay in these objects. As we need to have as large statistics as possible, and also we need uniform samples, in the first place we study sources from the ATNF catalogue (Manchester et al. 2005). Then we apply our methods to the largest uniform subsample of the ATNF — to the PMSS (stands for the Parkes Multibeam and Swinburne surveys) (Manchester et al. 2001).



We reconstruct the magnetic field decay in the range of true (statistical) ages:
$8\ 10^4 < t < 3.5\ 10^5$ yrs
which corresponds to characteristic ages
$8\ 10^4 < \tau < 10^6$ yrs.
In this range, the field decays roughly by a factor of two.
With an exponential fit this corresponds to the decay time scale ~4 10^5 yrs.
Note, this decay is limited in time.

Igoshev, Popov (2014)

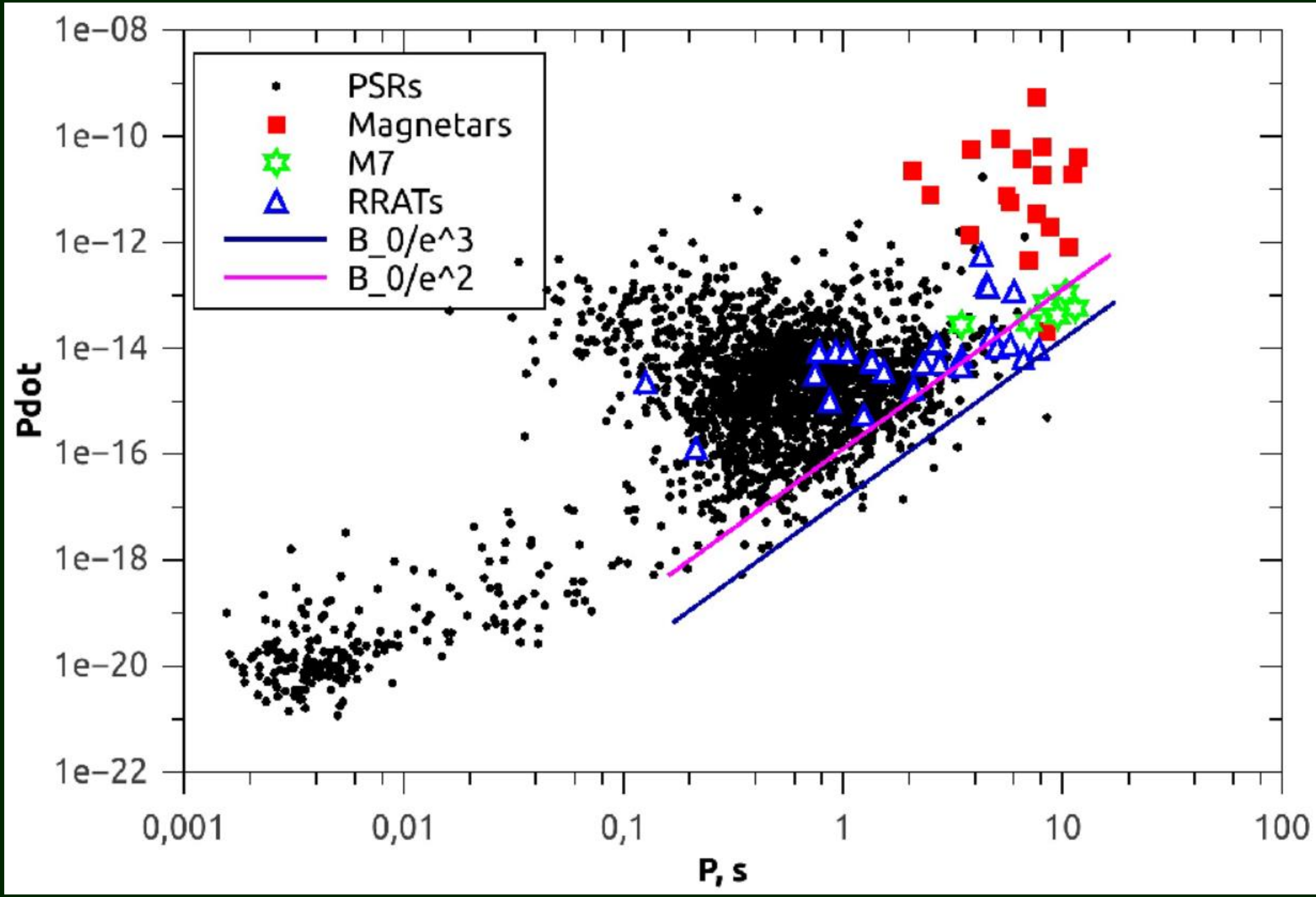
Comparison of different options



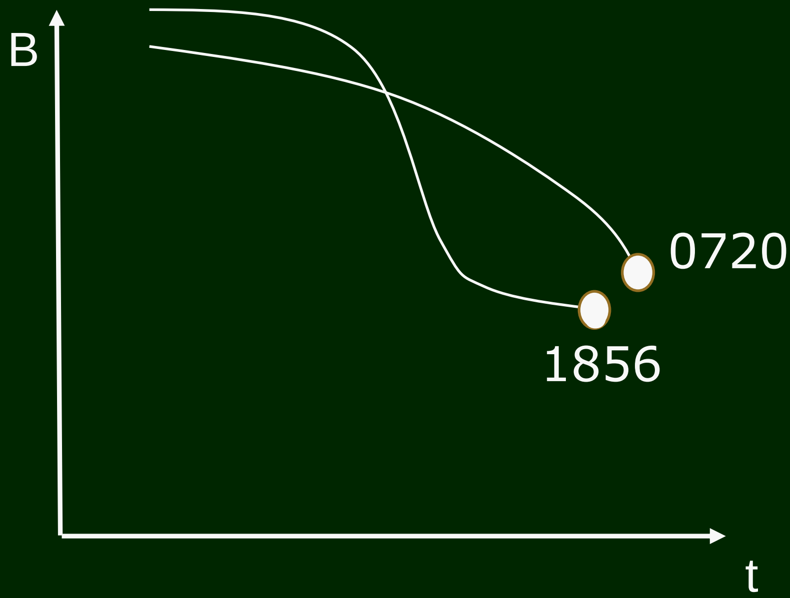
We think that at the ages ~10^5 yrs and below for normal pulsars we see mostly Ohmic decay, which then disappears as NSs cool down below the critical T.

Igoshev, Popov (2015)

Getting close to the attractor



Tracks on the P-pdot diagram



Kinematic age is larger for 0720, but characteristic age – for 1856.

It seems that 1856 is now on a more relaxed stage of the magneto-rotational evolution.

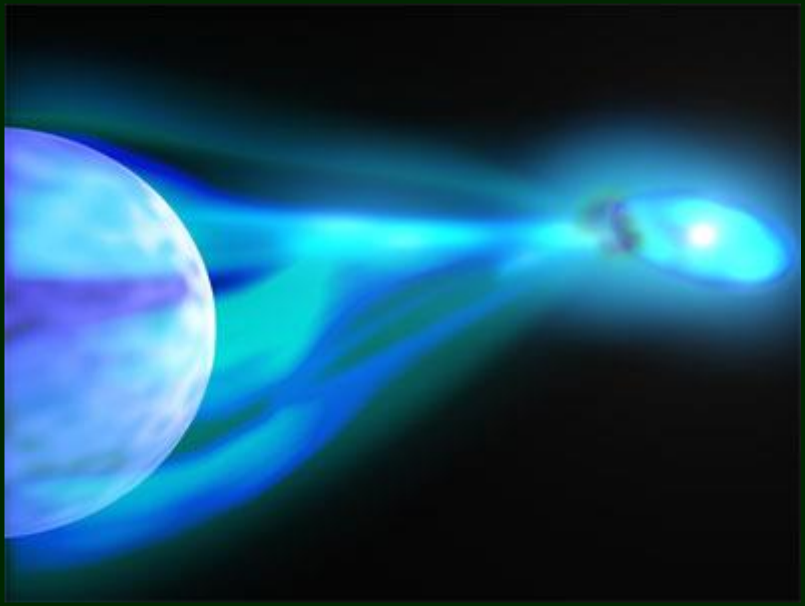
RX J0720 shows several types of activity, but RX J1856 is a very quiet source.

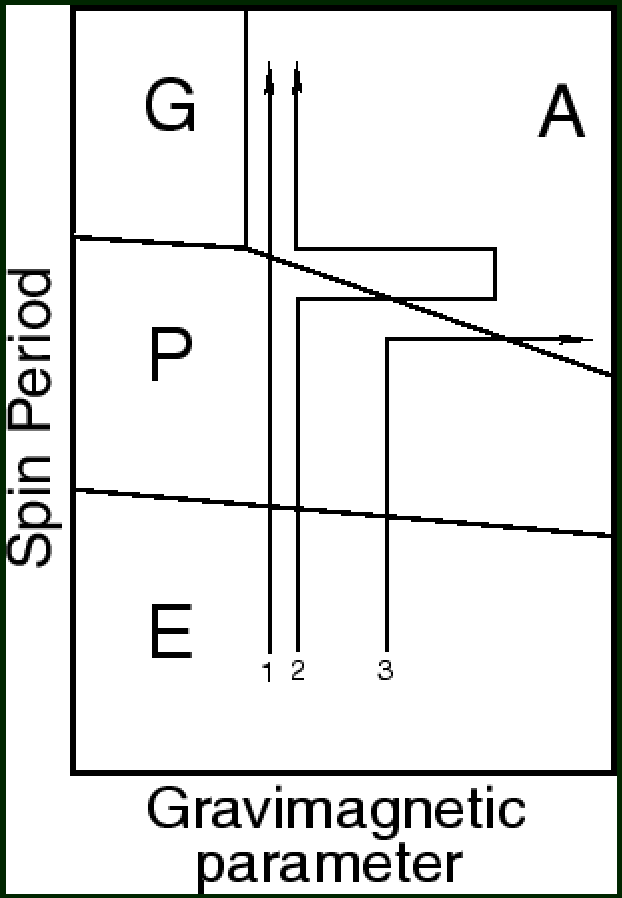
Non-monotonic evolution?

SXP 1062

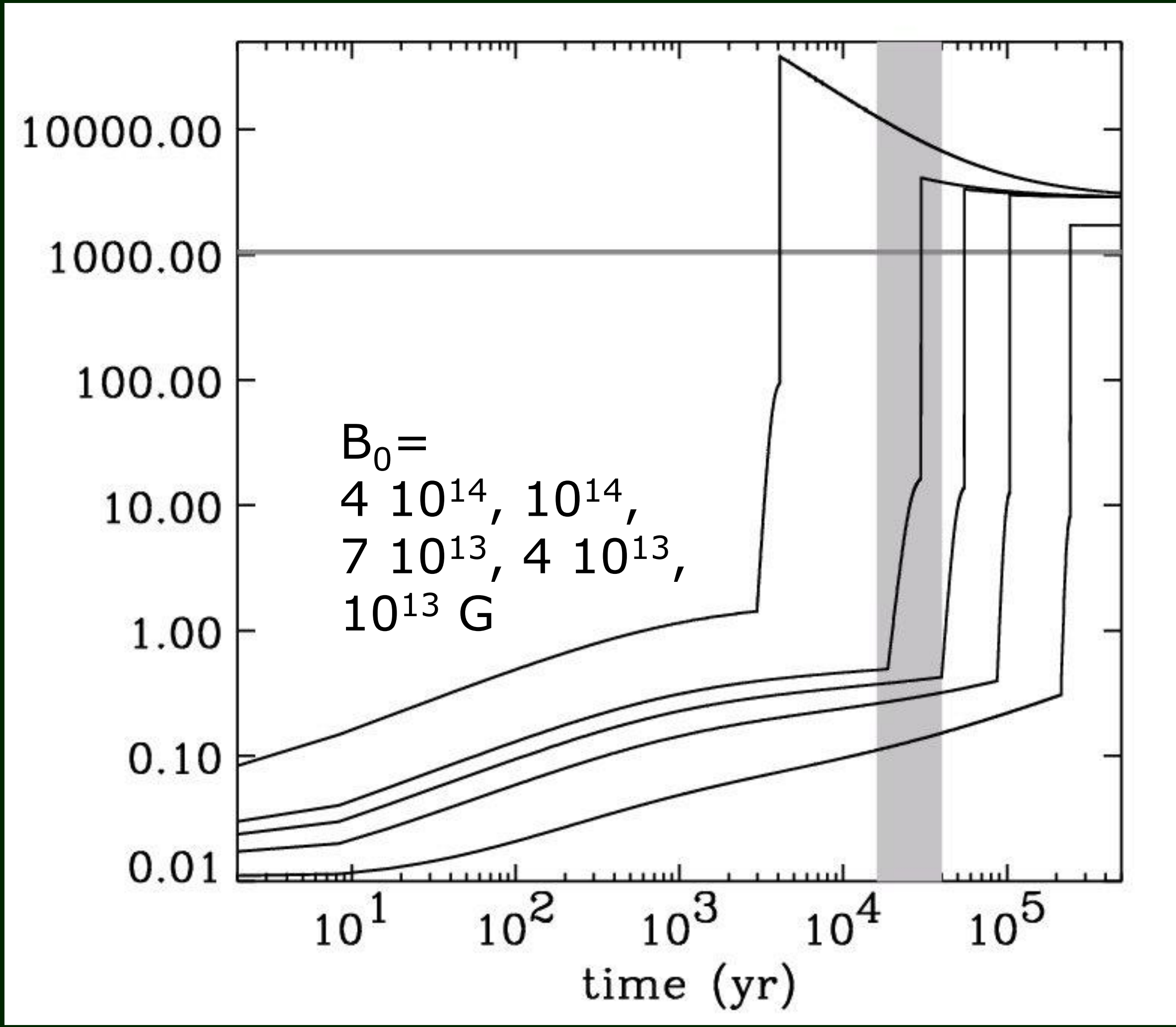
A peculiar source was discovered in SMC.
Be/Xray binary, P=1062 sec.
A SNR is found. Age ~10^4 yrs.
(1110.6404; 1112.0491)

Typically, it can take ~1 Myr for a NS with B~10^{12} G to start accretion.





Evolution of SXP 1062



A model of a NS with initial field ~10^{14} G which decayed down to ~10^{13} G can explain the data on SXP 1062.

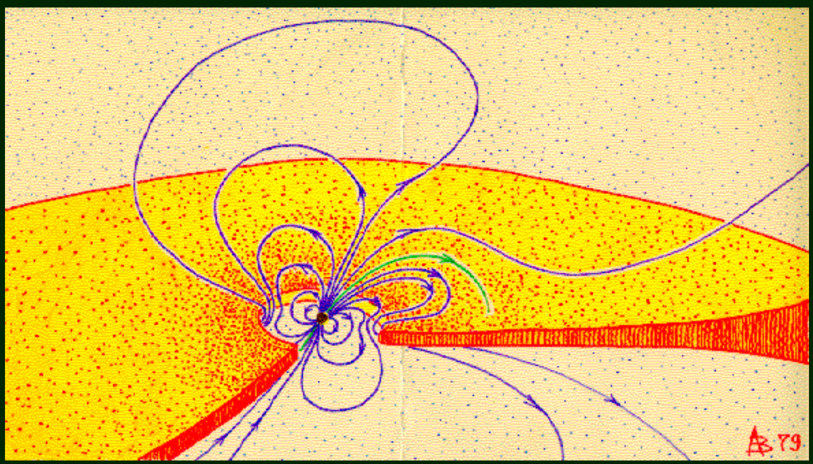
1112.2507

Some new data in 1304.6022

Many other scenarios have been proposed.
We need new observational data.

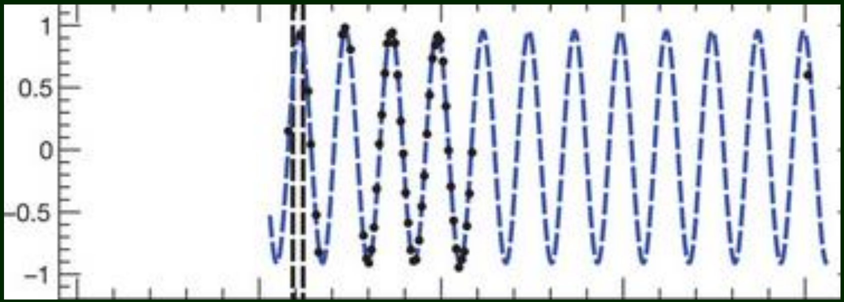
Accreting magnetars

Typically magnetic fields of neutron stars in accreting X-ray binaries are estimated with indirect methods.



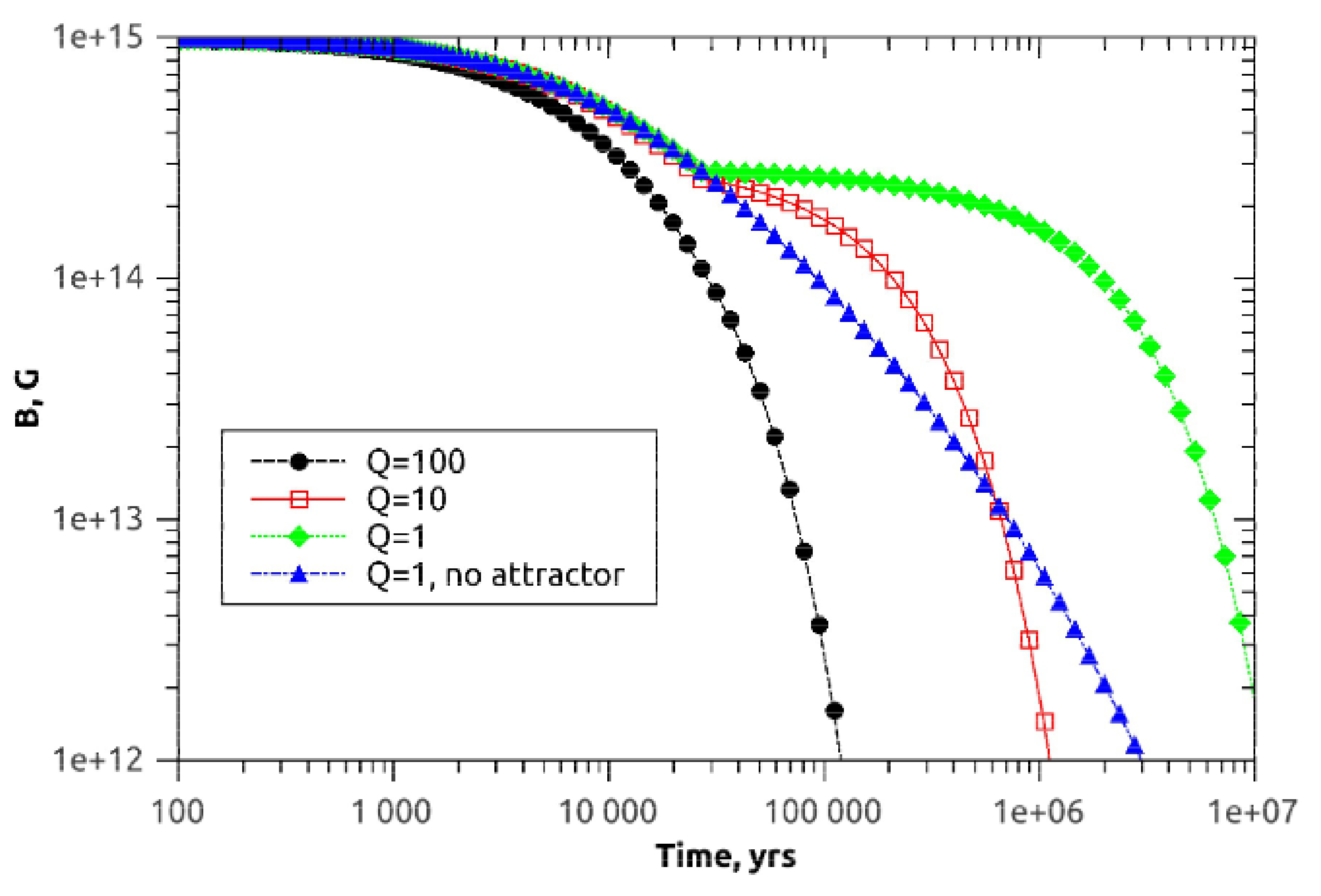
- Spin-up
- Spin-down
- Equilibrium period
- Accretion model
- …….

- ULX. NuSTAR J095551+6940.8 (M82 X-2). Ekşi et al. (2015).
- ULX. NGC 5907. Israel et al. (2017a)
- ULX. NGC 7793 P13. Israel et al. (2017b).
- 4U0114+65. Sanjurjo et al. (2017).
- 4U 2206+54. Ikhsanov & Beskrovnaya (2010).
- SXP1062. Fu & Li (2012)
- Swift J045106.8-694803. Klus et al. (2013).



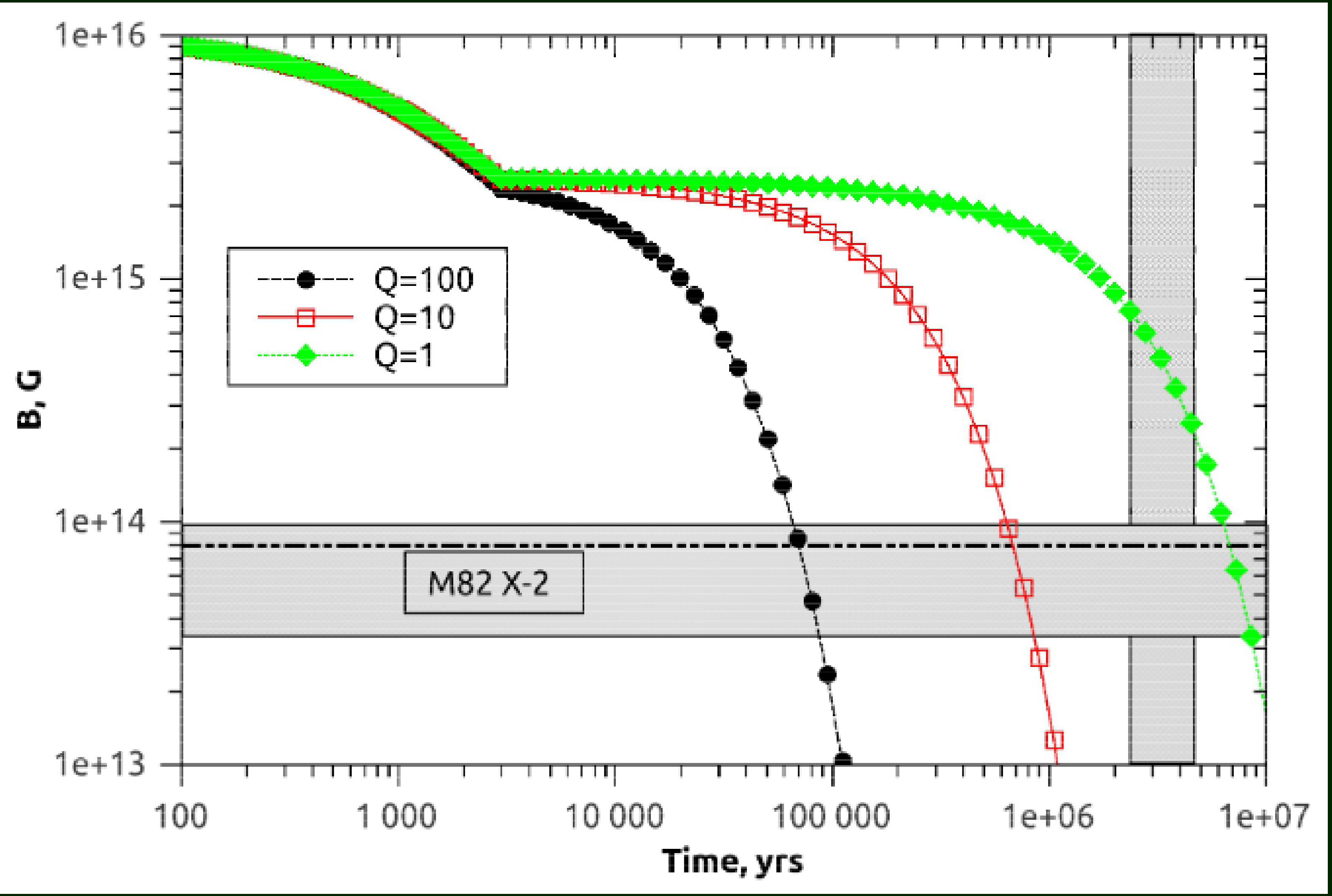
1709.10385

Field evolution in a magnetar



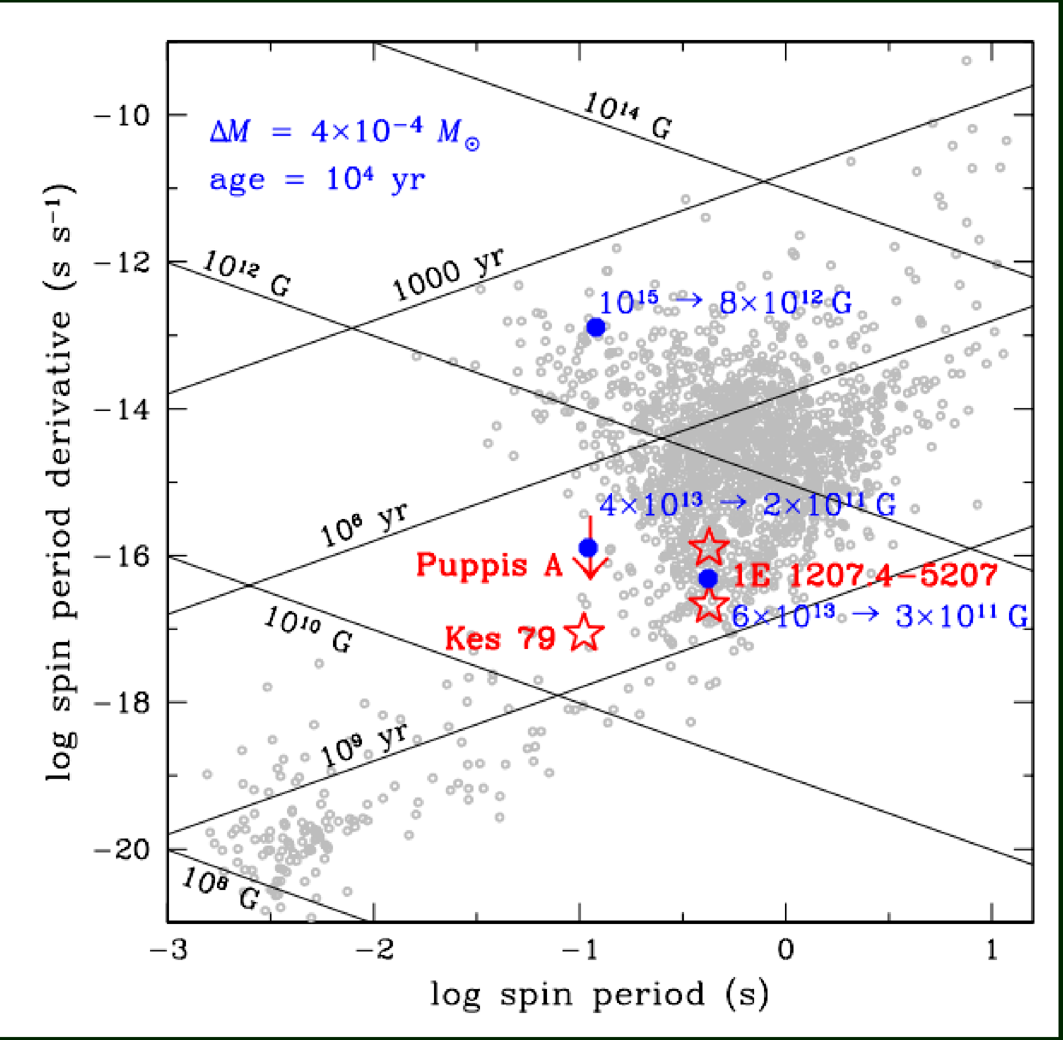
1709.10385

Parameters of ULX M82 X-2



1709.10385

Anti-magnetars

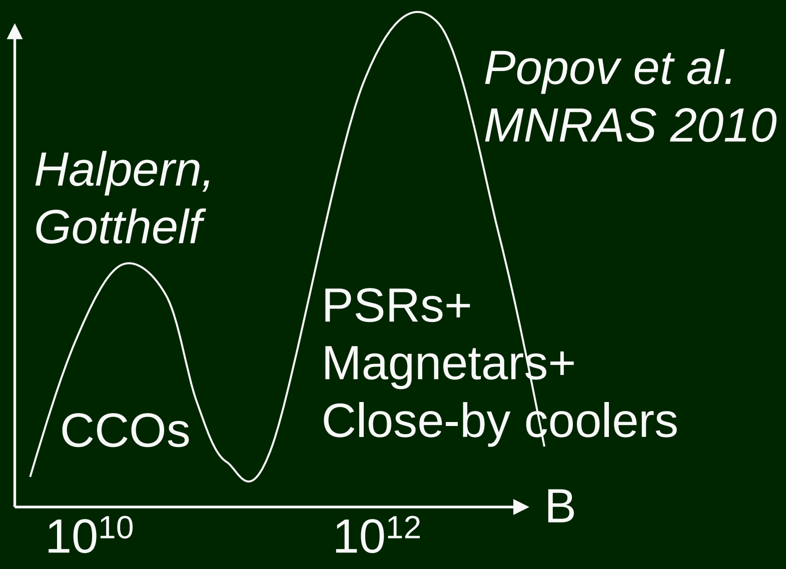


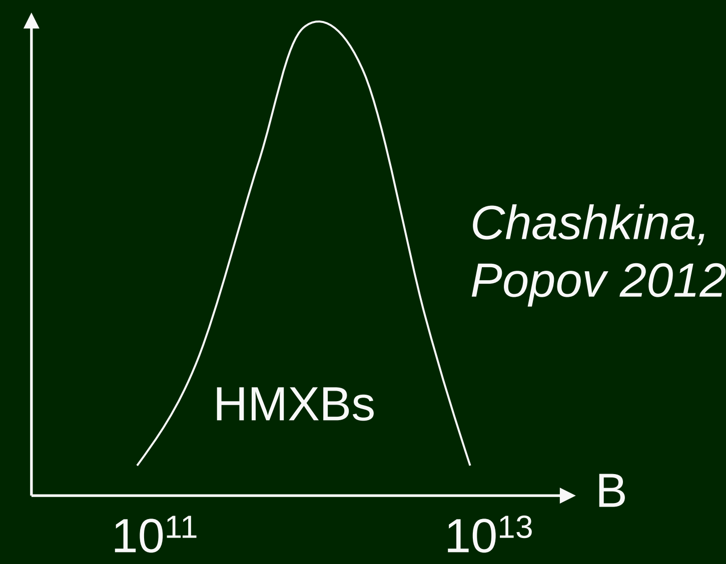
Note, that there is no room for antimagnetars from the point of view of birthrate in many studies of different NS populations.

New results 1301.2717
Spins and derivative are measured for
PSR J0821-4300 and
PSR J1210-5226

Ho 1210.7112

Evolution of CCOs



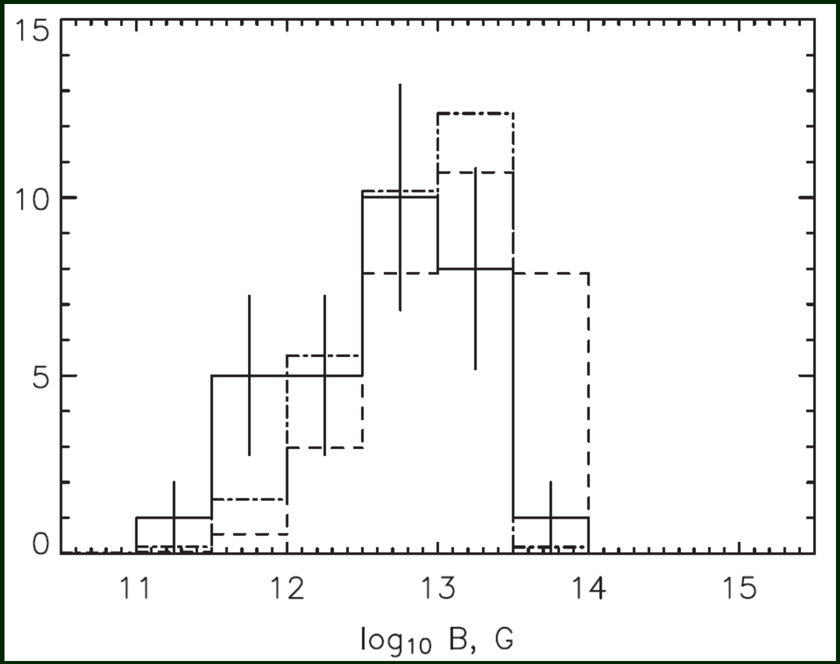


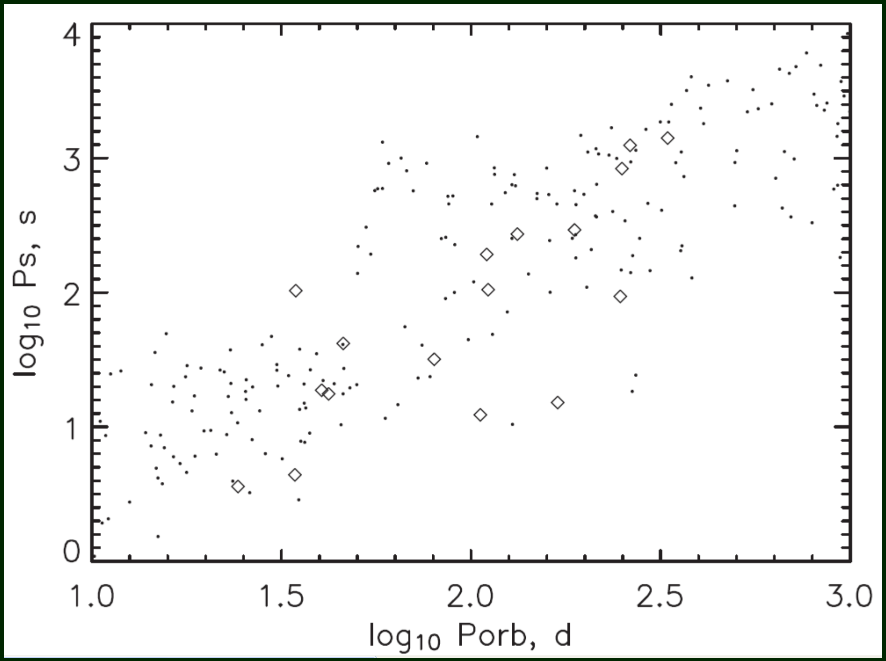
Among young isolated NSs about 1/3 can be related to CCOs.
If they are anti-magnetars, then we can expect that 1/3 of NSs in HMXBs are also low-magnetized objects.
They are expected to have short spin periods <1 sec.
However, there are no many sources with such properties.
The only good example - SAX J0635+0533. An old CCO?

Possible solution: emergence of magnetic field
(see physics in Ho 2011, Vigano, Pons 2012).

Observations vs. theory

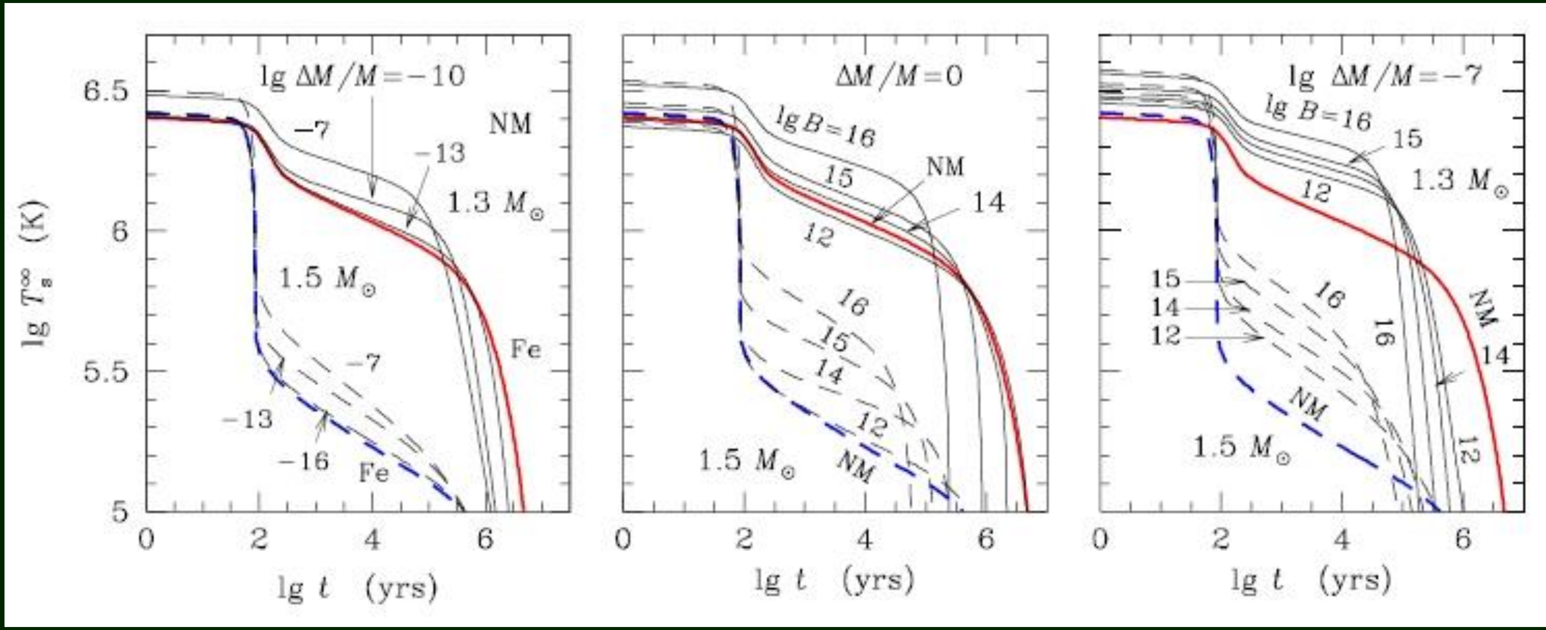
We use observations of Be/X-ray binaries in SMC to derive magnetic field estimates, and compare them with prediction of the Pons et al. model.





Chashkina, Popov (2012)

Where are old CCOs?



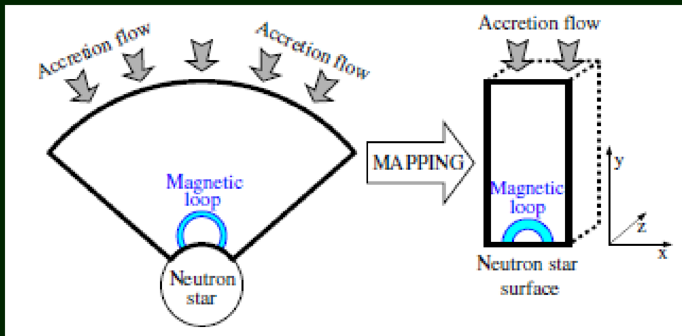
Yakovlev, Pethick 2004

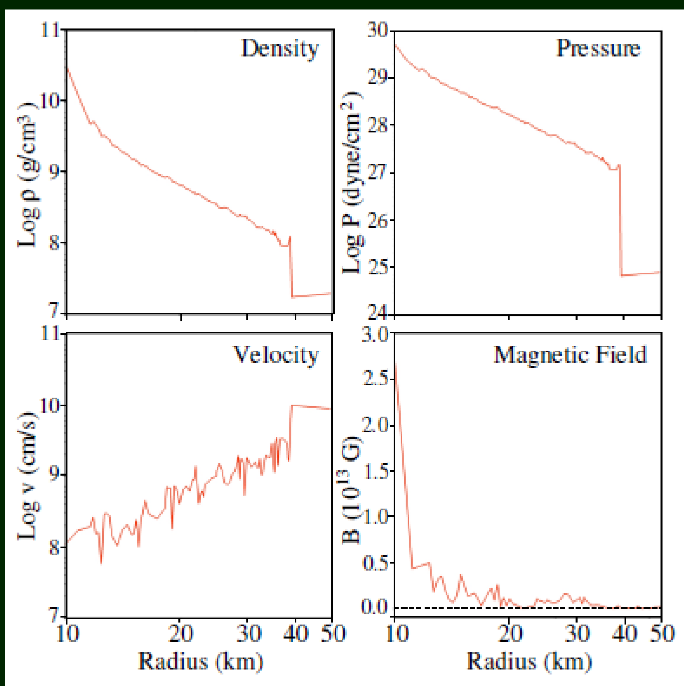
According to cooling studies they have to be bright till at least 10^5 years.
But only one candidate (2XMM J104608.7-594306 Pires et al.) to be a low-B cooling NS is known (Calvera is also a possible candidate).

We propose that a large set of data on HMXBs and cooling NSs is in favour of field emergence on the time scale $10^4 \le \tau \le 10^5$ years (arXiv:1206.2819).

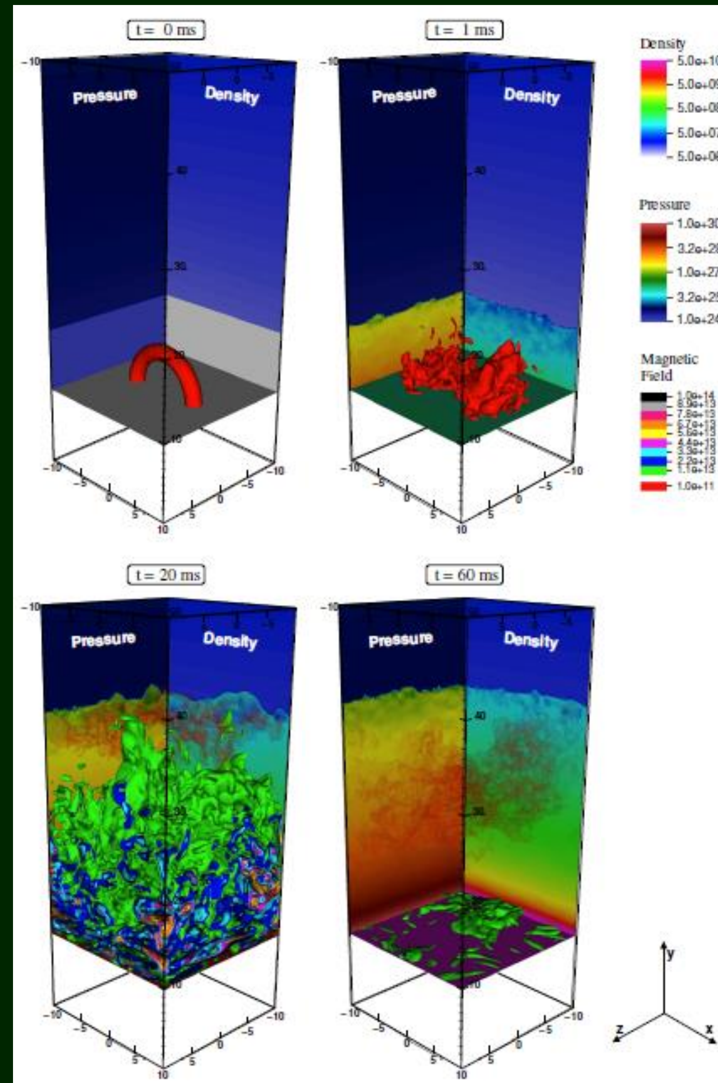
Some PSRs with thermal emission for which additional heating was proposed can be descendants of CCOs with emerged field.

How the field is buried





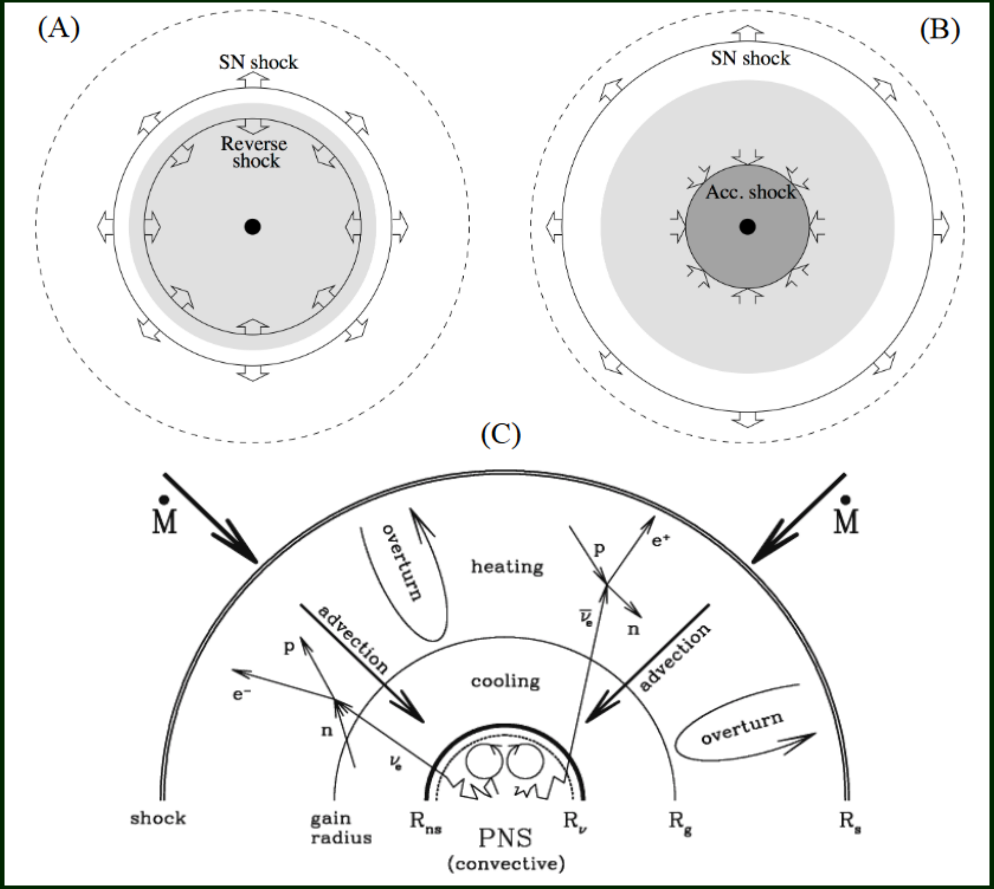
For t=60 msec

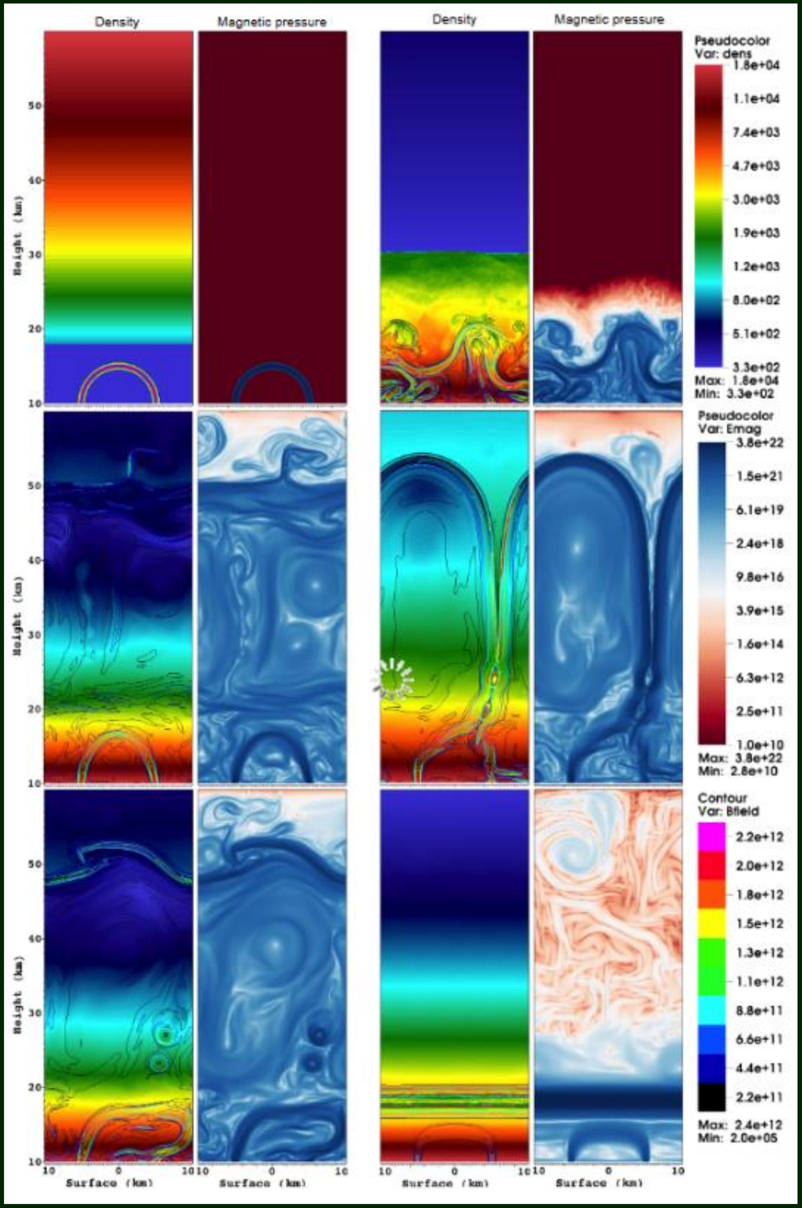


1212.0464

See 1210.7112 for a review of CCOs magnetic fields

More advanced model



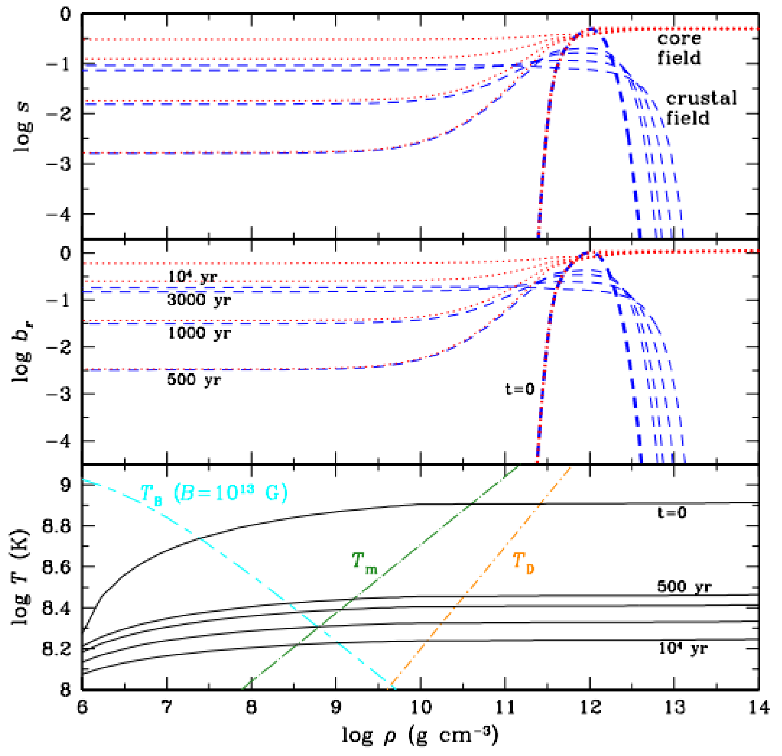


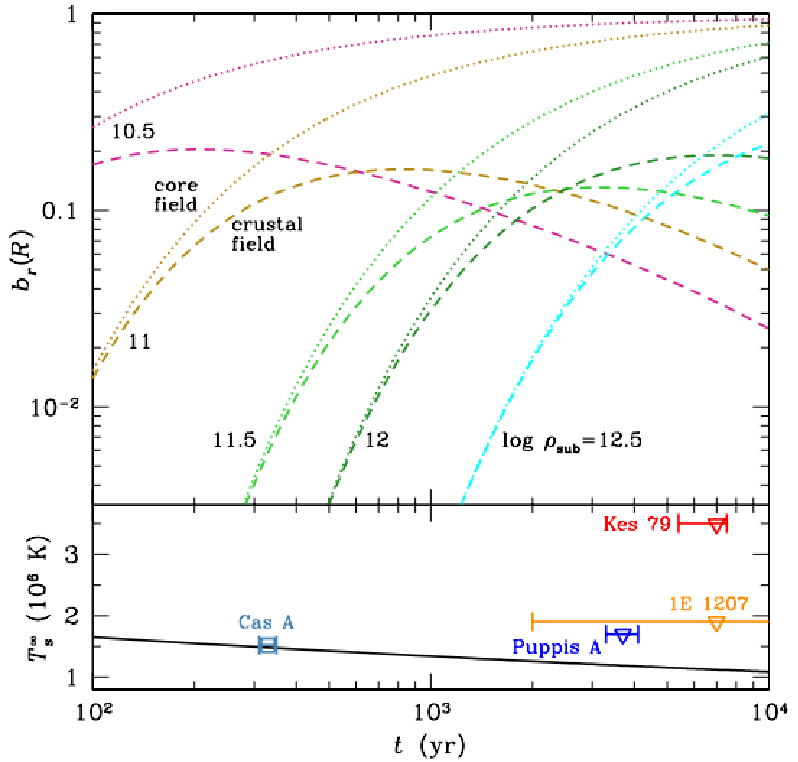
1809.07057

Emerging field: modeling

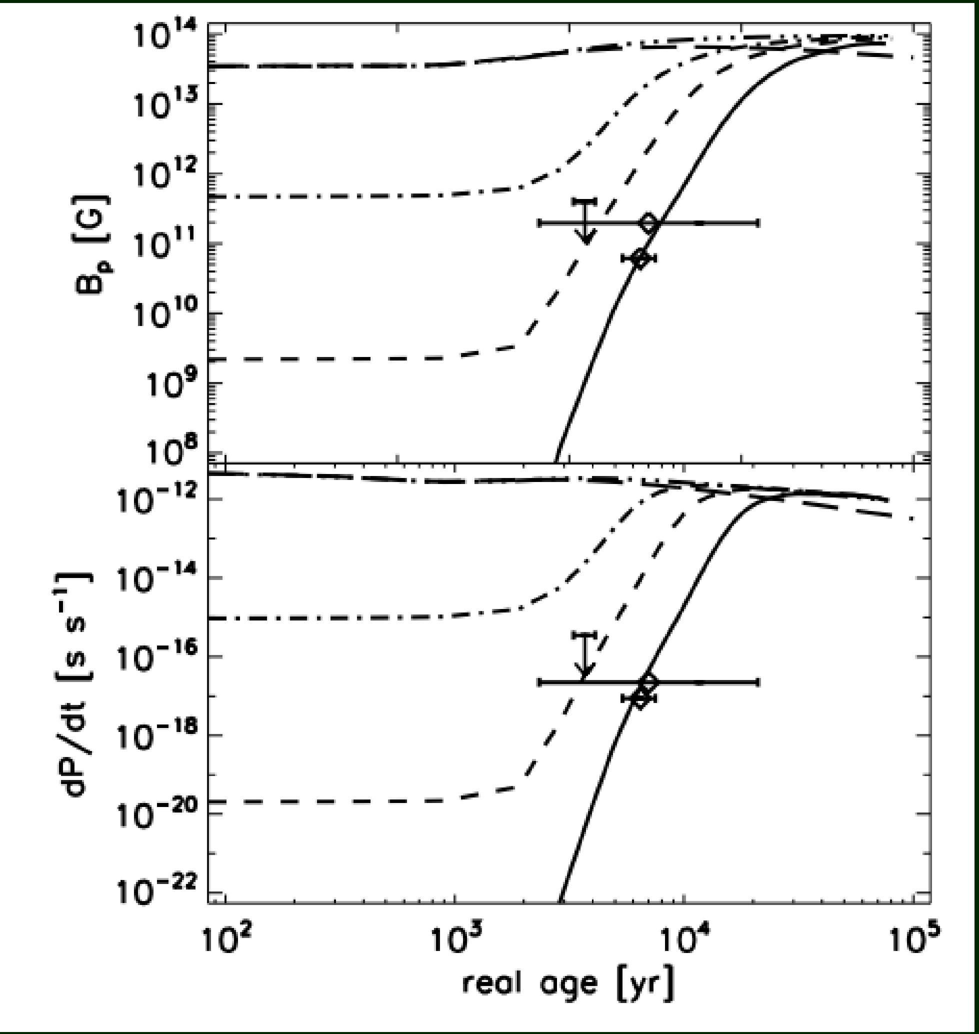
1D model of field emergence

Dashed – crustal, dotted – core field





Another model



2D model with field decay

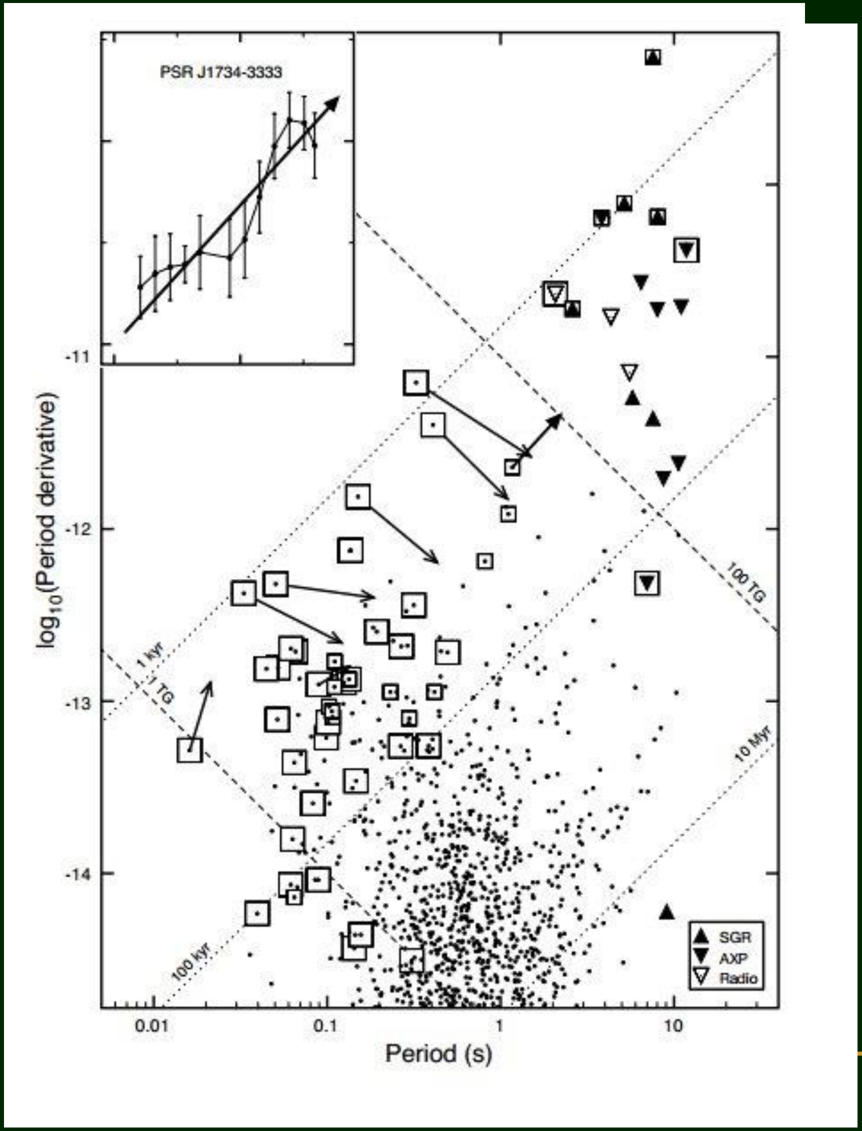
Ohmic diffusion dominates in field emergence, but Hall term also can be important.

Calculations confirm that emergence on the time scale 10^3-10^5 years is possible.

$B_{0p}=10^{14}$ G

Vigano, Pons 2012 1206.2014

Emerged pulsars in the P-Pdot diagram



Emerged pulsars are expected to have
P~0.1-0.5 sec
B~10^{11}-10^{12} G
Negative braking indices or at least n<2.
About 20-40 of such objects are known.

Parameters of emerged PSRs:
similar to “injected” PSRs
(Vivekanand, Narayan, Ostriker).

The existence of significant fraction
of “injected” pulsars formally
do not contradict recent pulsar current studies
(Vranesevic, Melrose 2011).

Part of PSRs supposed to be born with
long (0.1-0.5 s) spin periods can be
matured CCOs.

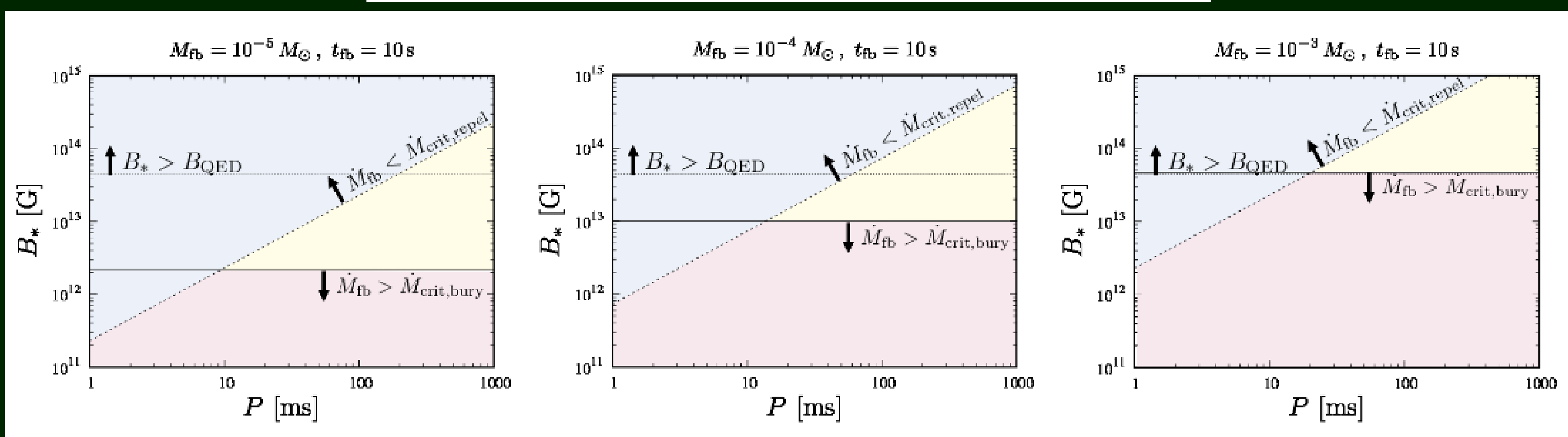
Espinoza et al. arXiv: 1109.2740, 1211.5276

Field, rotation, fallback

$$\dot{M}_{\rm crit,repul} \sim 3\times10^{-5}\,M_\odot\,{\rm s}^{-1}\,\frac{\xi_{\rm s,crit}}{0.2}\,\frac{(4\pi D_{\rm fb}\sqrt{\xi_{\rm s}})_{\rm crit}}{5.3}\left(\frac{B_*}{10^{13}\,{\rm G}}\right)^2\left(\frac{P}{10\,{\rm ms}}\right)^{-2}\left(\frac{t_{\rm fb}}{20\,{\rm s}}\right)^{2/3}$$

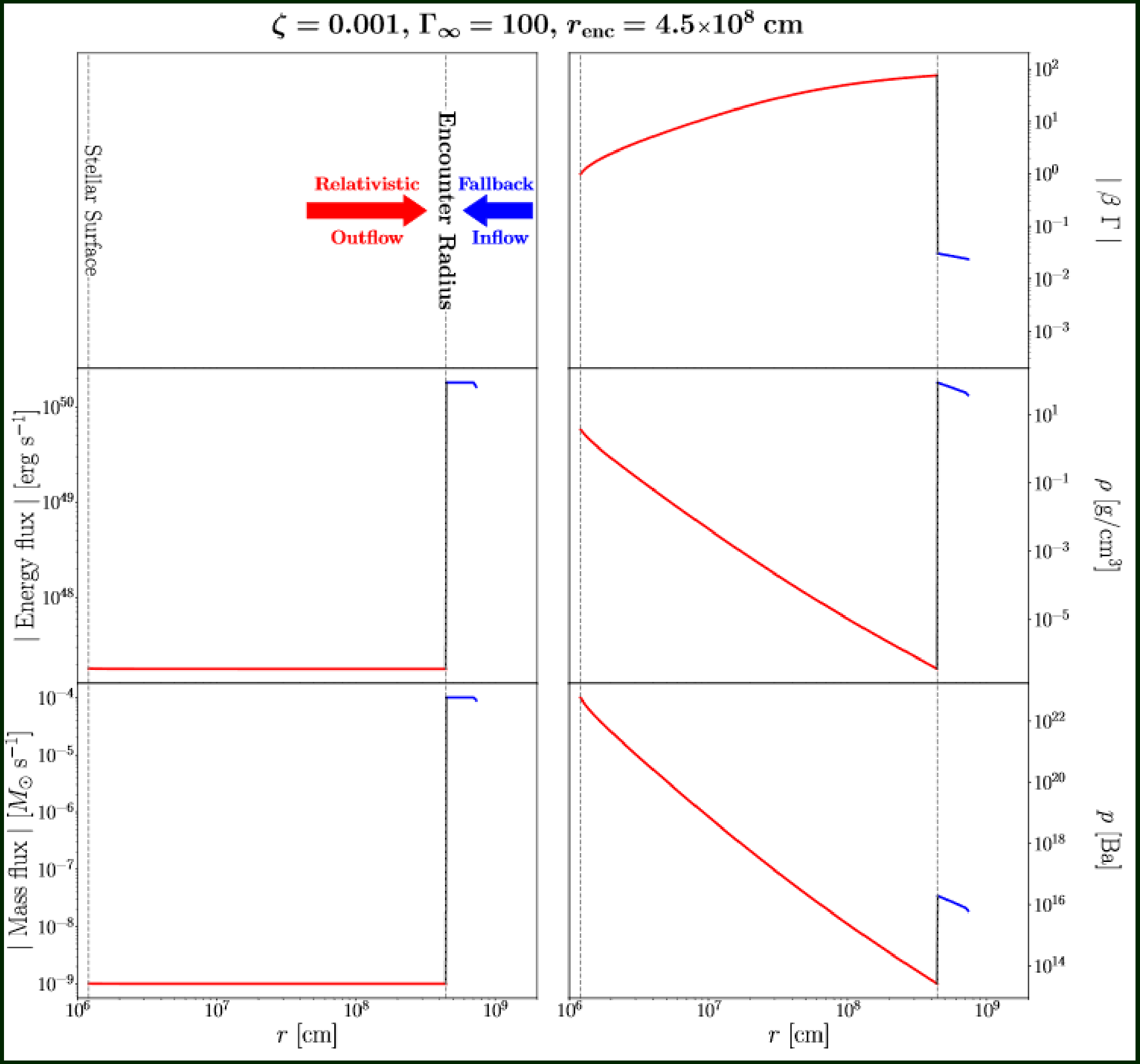
$$\frac{{B_*}^2}{8\pi} \lesssim \rho v^2 \sim \frac{\dot{M}}{4\pi R_*^2}\sqrt{\frac{GM_{\rm c}}{R_*}}$$

$$\dot{M}_{\rm crit,bury} \sim 3\times10^{-6}\,M_\odot\,{\rm s}^{-1}\left(\frac{B_*}{10^{13}\,{\rm G}}\right)^2$$



1809.00487

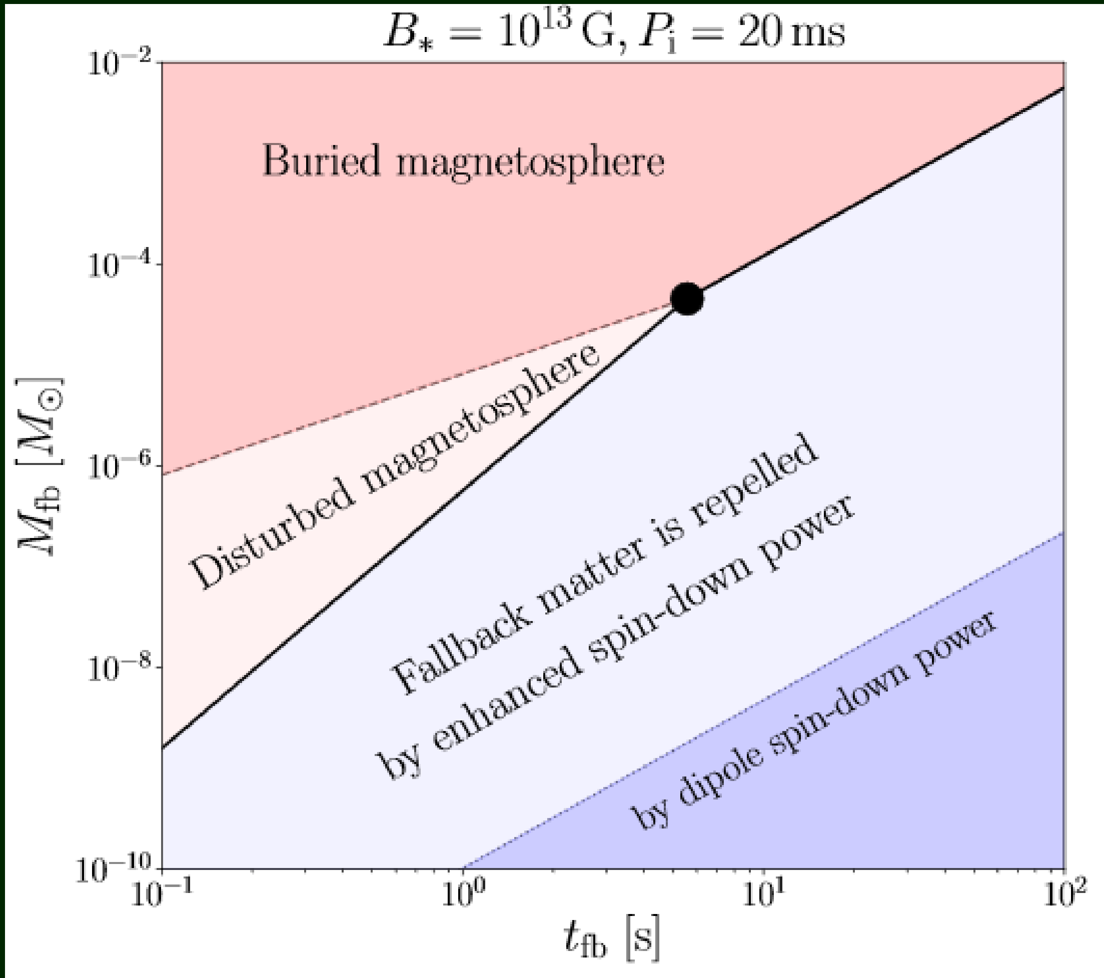
Fallback matter interacting with a NS

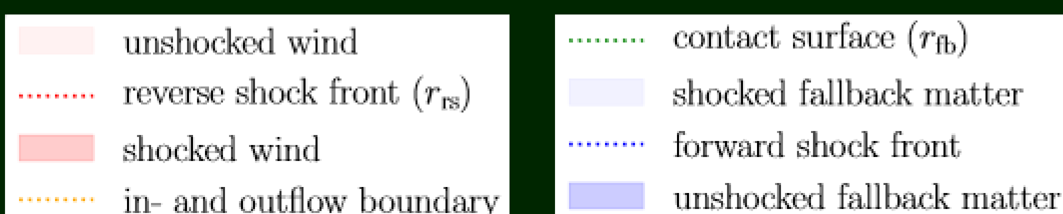


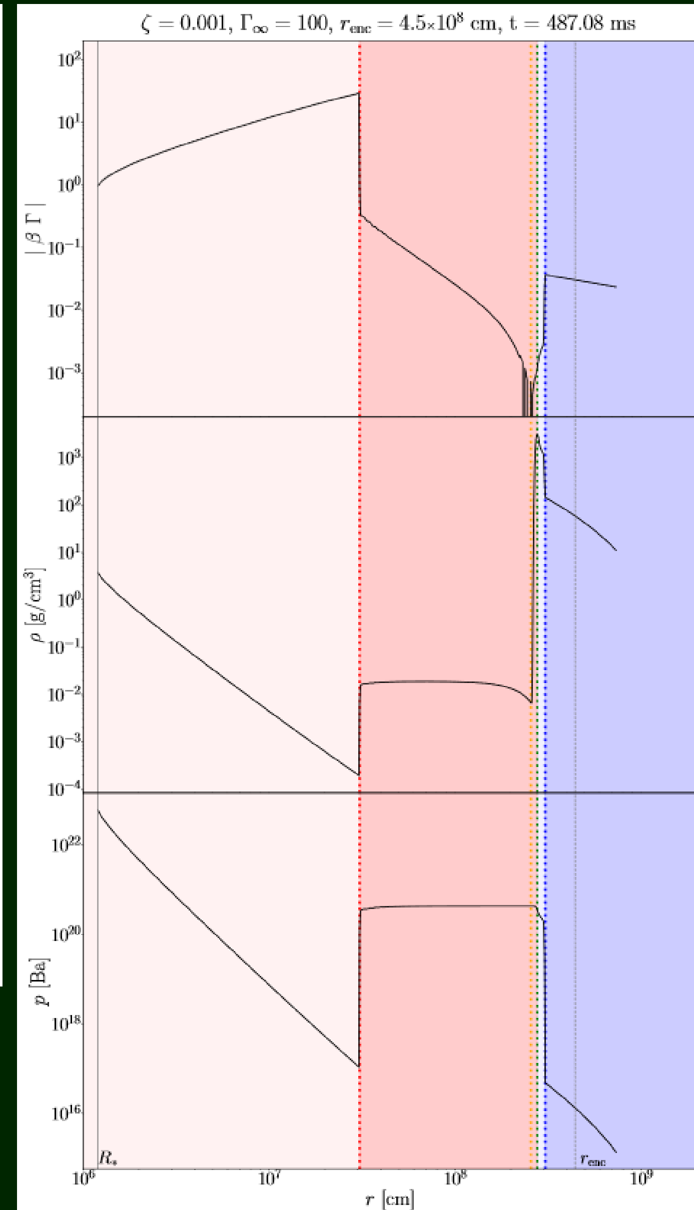
Fallback matter interacts with relativistic wind, magnetosphere, and finally – with the NS surface.

Depending on parameters (ΔM, field, spin, etc.) different regimes can appear. Thus, NSs can appear at different stages and can be observed as sources with different properties.

2103.09461

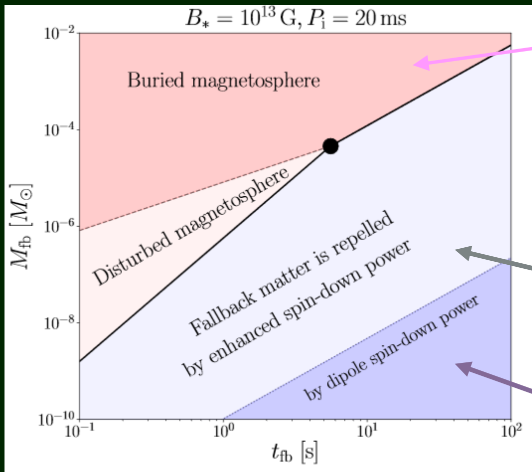

$B_* = 10^{13}$ G, $P_i = 20$ ms
Buried magnetosphere
Disturbed magnetosphere
Fallback matter is repelled
by enhanced spin-down power
by dipole spin-down power
M_{fb} [$M_\odot$]
t_{fb} [s]
10^{-2}
10^{-4}
10^{-6}
10^{-8}
10^{-10}
10^{-1}
10^{0}
10^{1}
10^{2}


unshocked wind
reverse shock front (r_{rs})
shocked wind
in- and outflow boundary
contact surface (r_{fb})
shocked fallback matter
forward shock front
unshocked fallback matter


$\zeta = 0.001$, $\Gamma_\infty = 100$, $r_{enc} = 4.5\times10^8$ cm, t = 487.08 ms
$|\beta\,\Gamma|$
ρ [g/cm^3]
p [Ba]
R_*
r_{enc}
r [cm]

2103.09461

Variants of young NS properties



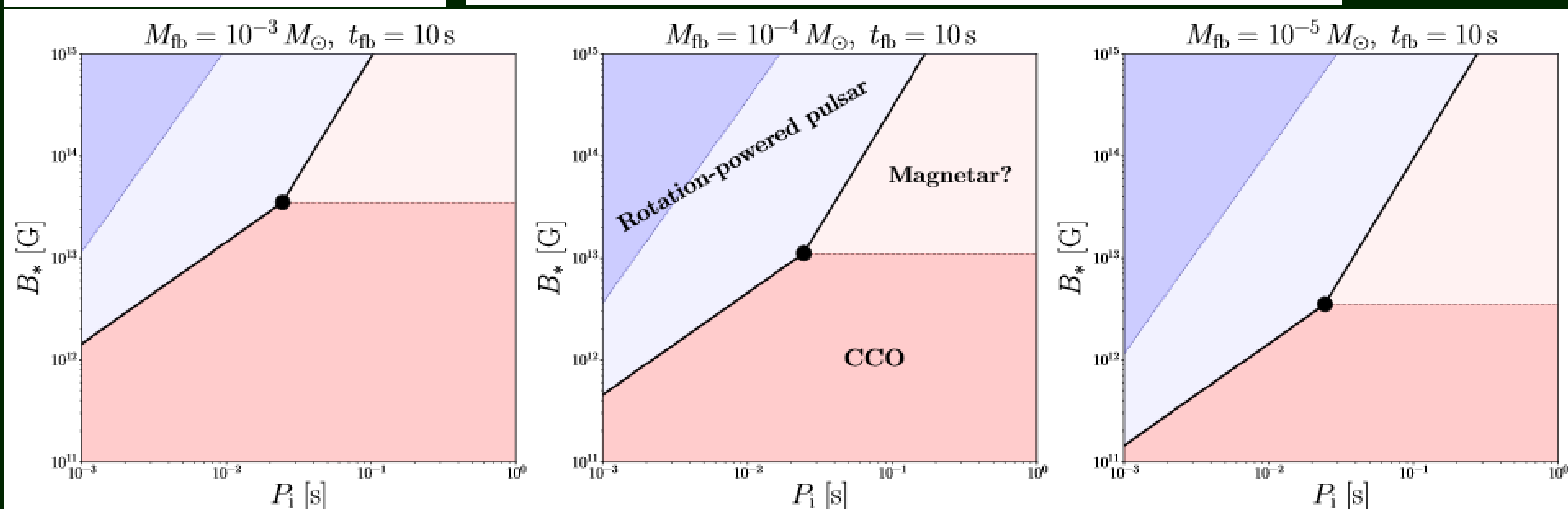
$$M_{\rm fb} > 8.2\times10^{-5}\,M_\odot\,{B_{*,13}}^2 t_{\rm fb,1},$$

$$\dot{M}_{\rm fb} = \dot{M}_{\rm fb,ini}\times\begin{cases}1 & t\le t_{\rm fb}\\ (t/t_{\rm fb})^{-l} & t>t_{\rm fb}\end{cases},$$

$$\dot{M}_{\rm fb,ini} = \frac{l-1}{l}\frac{M_{\rm fb}}{t_{\rm fb}} \sim 1\times10^{-5}\,M_\odot\,{\rm s}^{-1}\left(\frac{l-1}{l}\right)M_{\rm fb,-4}{t_{\rm fb,1}}^{-1},$$

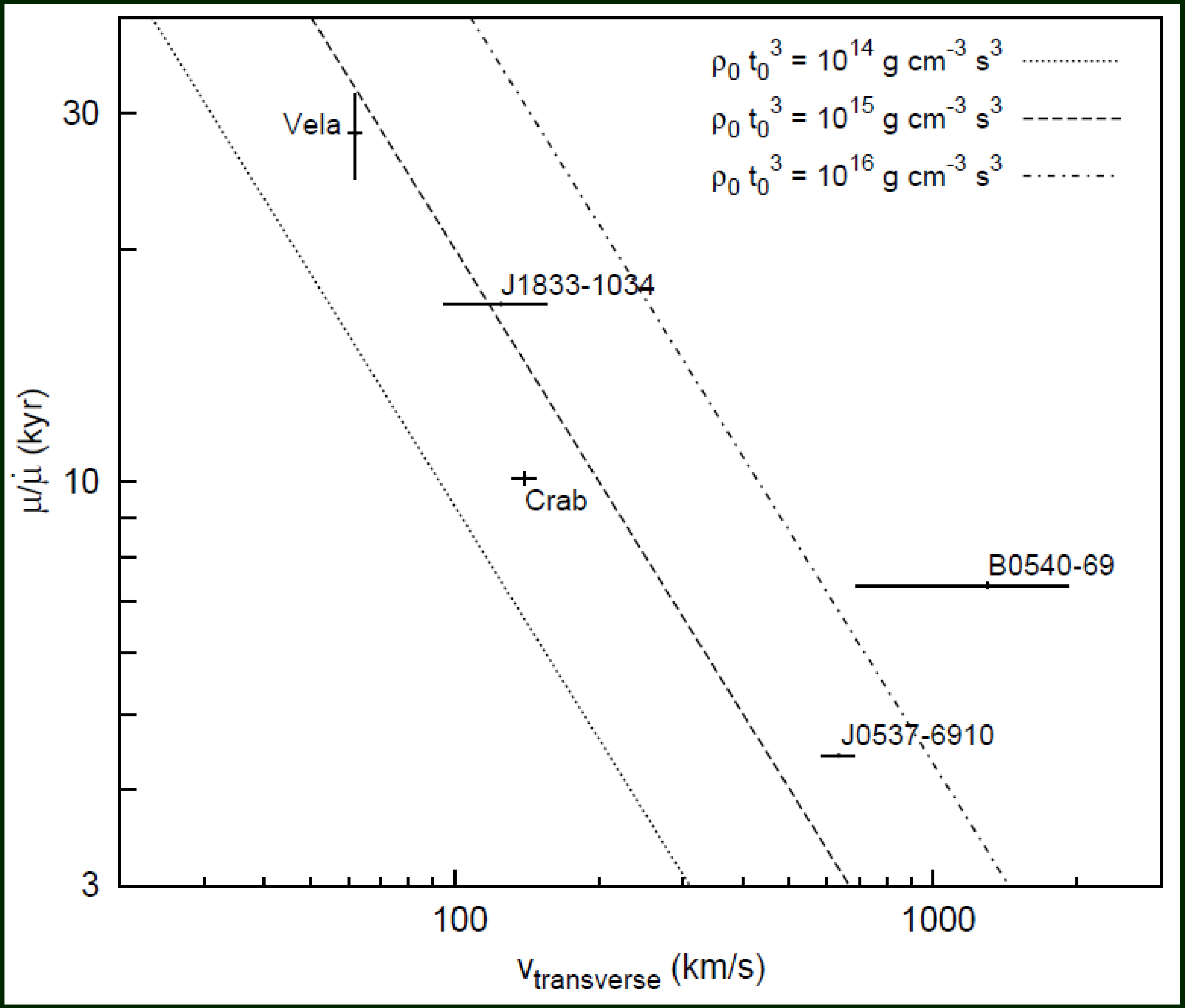
$$M_{\rm fb,crit} \approx \begin{cases}5.2\times10^{-3}\,M_\odot\,{B_{*,13}}^2{P_{\rm i,-2}}^{-14/3}{t_{\rm fb,1}}^{23/9} & r_\Lambda > R_*\\ 4.8\times10^{-4}\,M_\odot\,{B_{*,13}}^2{P_{\rm i,-2}}^{-2}{t_{\rm fb,1}}^{5/3} & r_\Lambda \le R_*\end{cases}.$$

$$M_{\rm fb,crit} \approx 7.7\times10^{-8}\,M_\odot\,(1+\sin\chi^2)\,{B_{*,13}}^2{P_{\rm i,-2}}^{-4}{t_{\rm fb,1}}^{5/3}\quad \text{(dipole)}.$$



2103.09461

Growing field and kick velocities?

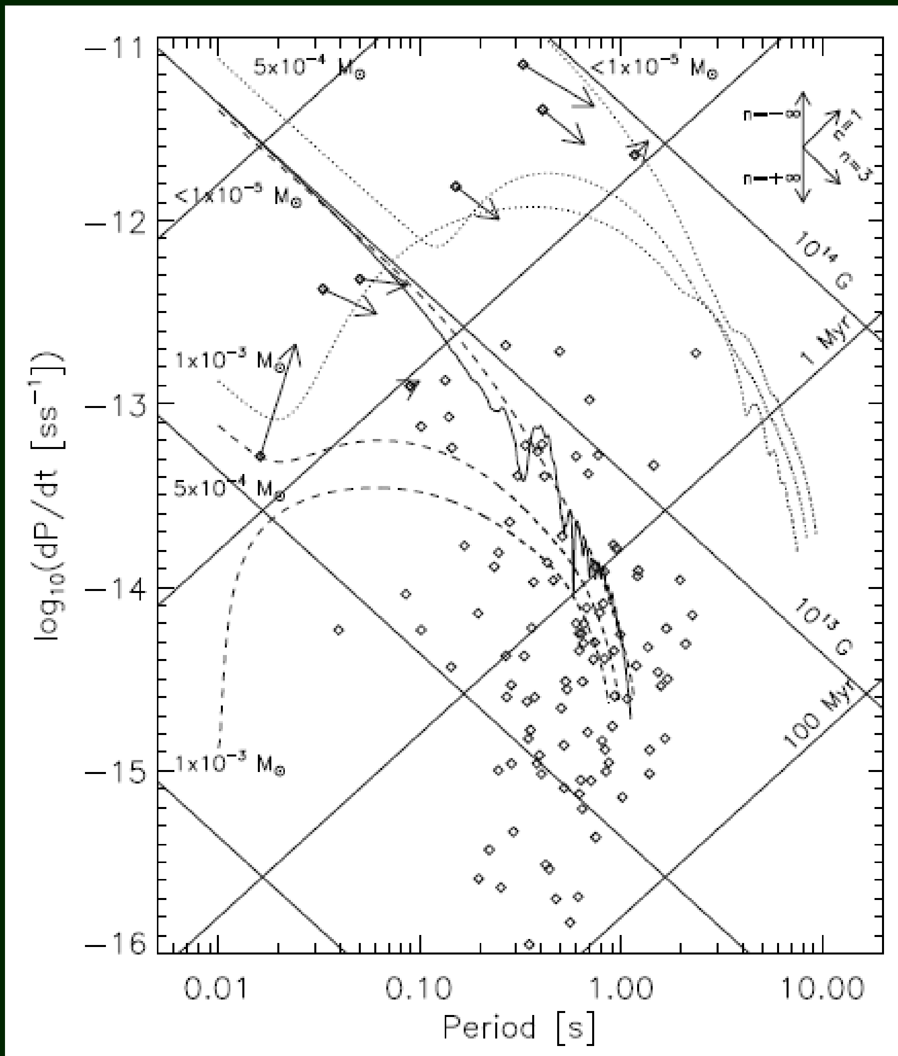


The idea is that n<3 are explained as due to growing field. Then it is possible to estimate the timescale for growing and plot it vs. velocity.

Larger kick –
- smaller fallback –
  - faster field growing

1207.1219

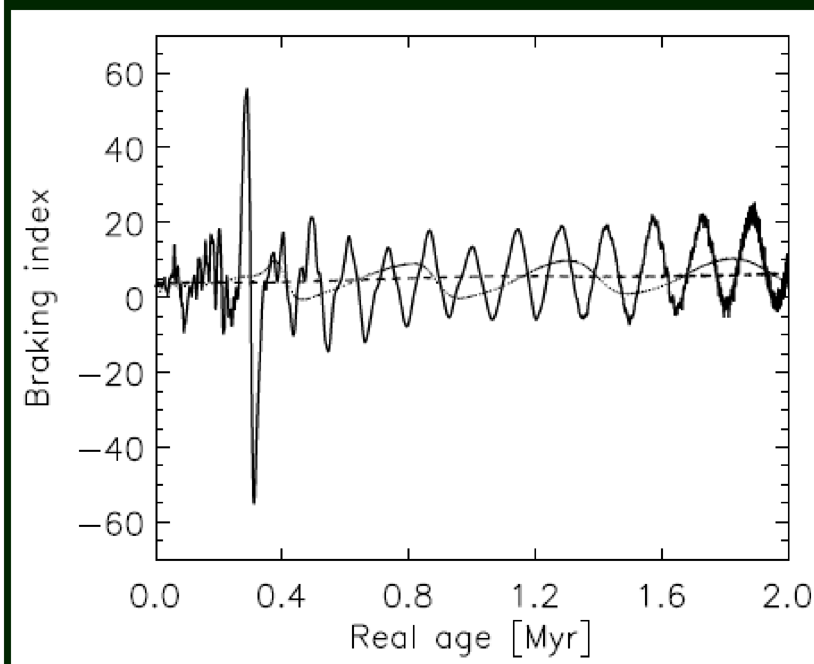
Evolution of PSRs with evolving field



Three stages:

1. n<=3 Standard + emerging field
2. n>3 Orhmic field decay
3. oscillating and large n – Hall drift

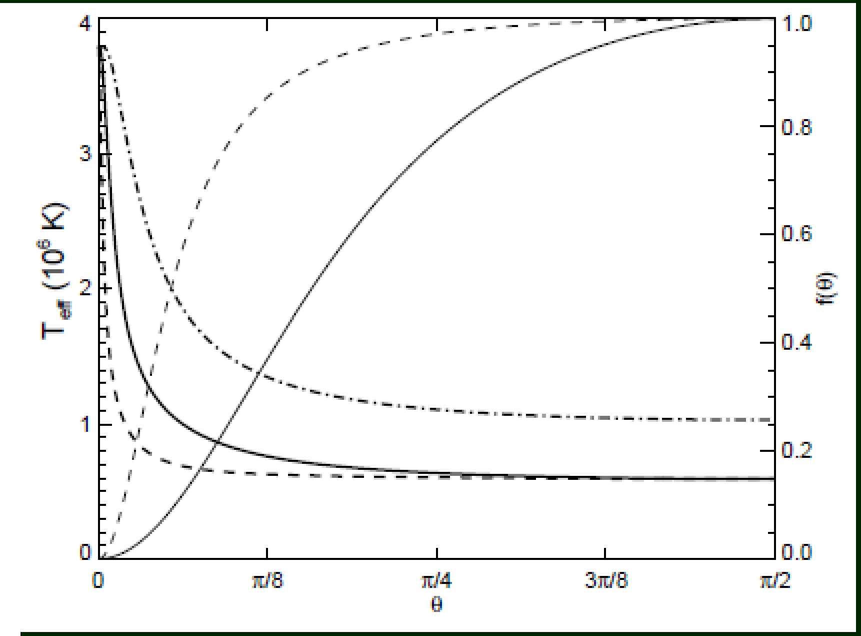
$$n = 3 - 4\frac{\dot{B}_0}{B_0}\tau_c \equiv 3 - 4\frac{\tau_c}{\tau_B},$$

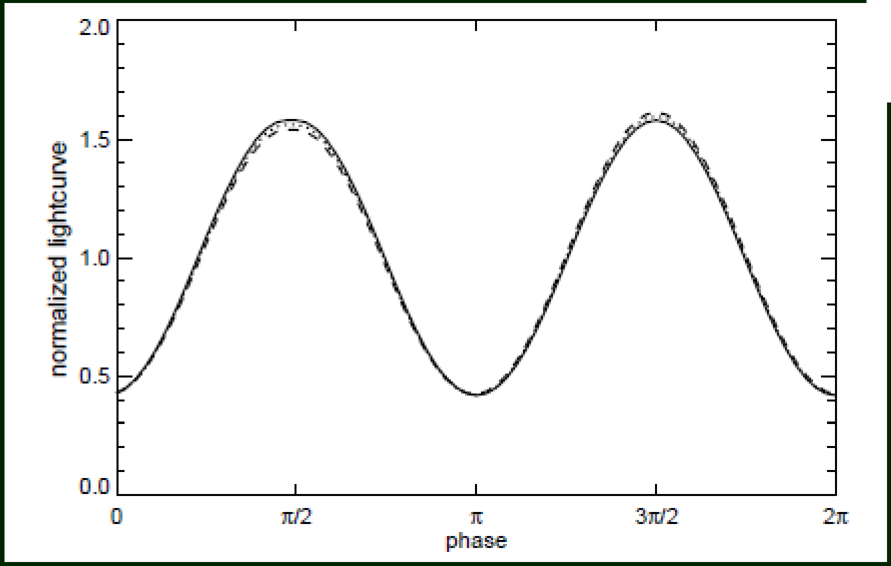


1209.2273

Buried field in Kes79?

The idea is to reconstruct surface temperature distribution, and then calculate which field configuration can produce it.

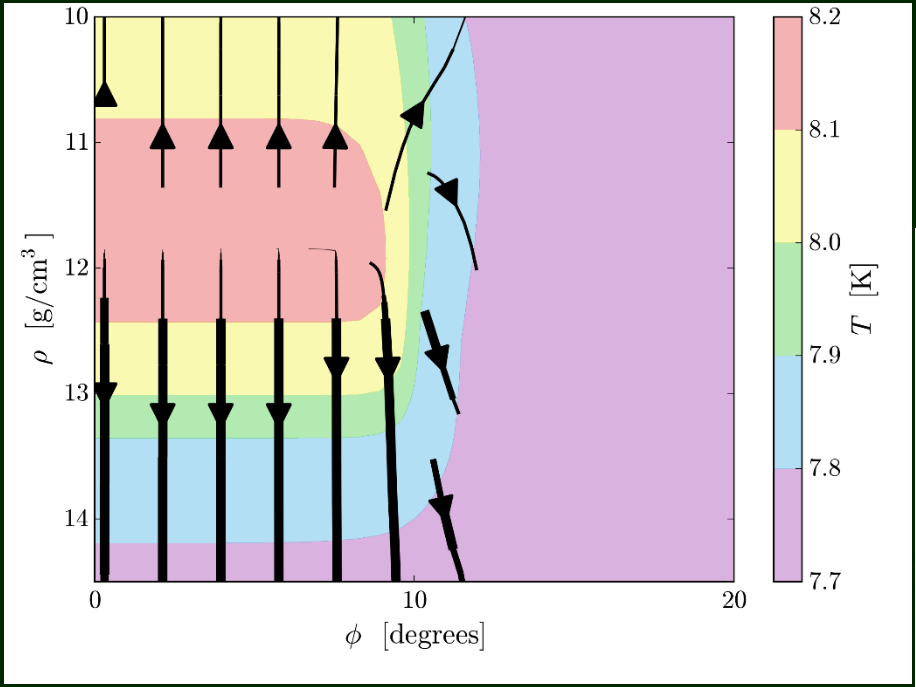


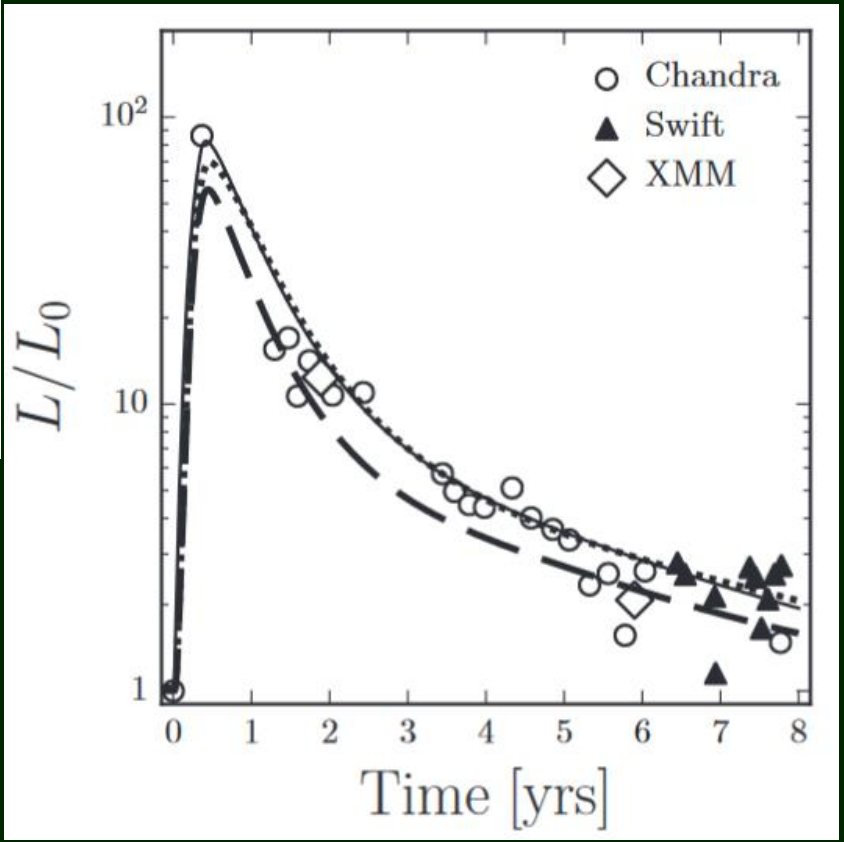


Very large pulse fraction (64%) in the anti-magnetar Kes 79.
Large sub-surface magnetic field can explain the existence of compact hot spots.
Then the field must have been buried in a fall-back episode.

1110.3129

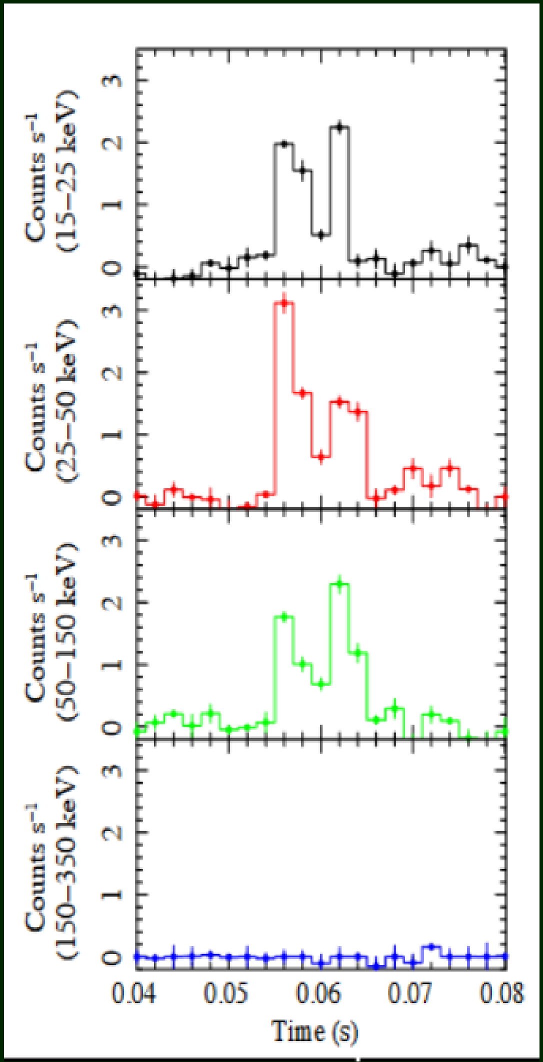
Hidden magnetar in RCW103

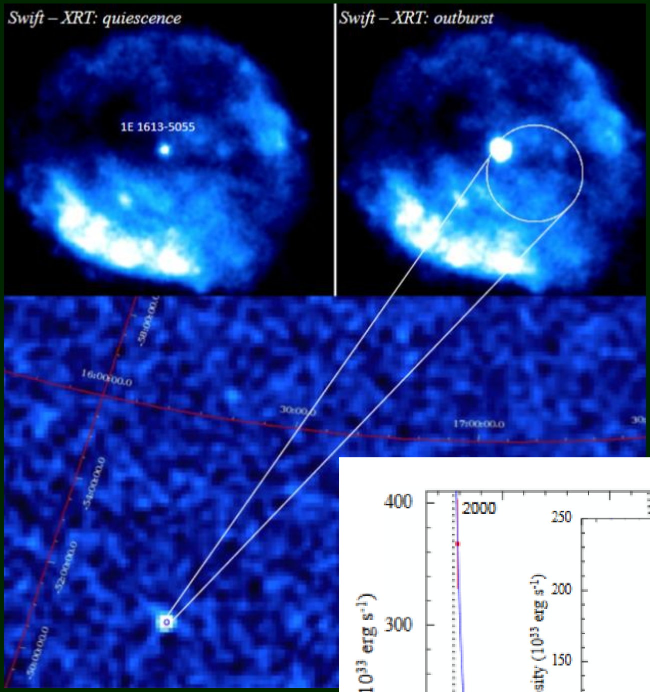


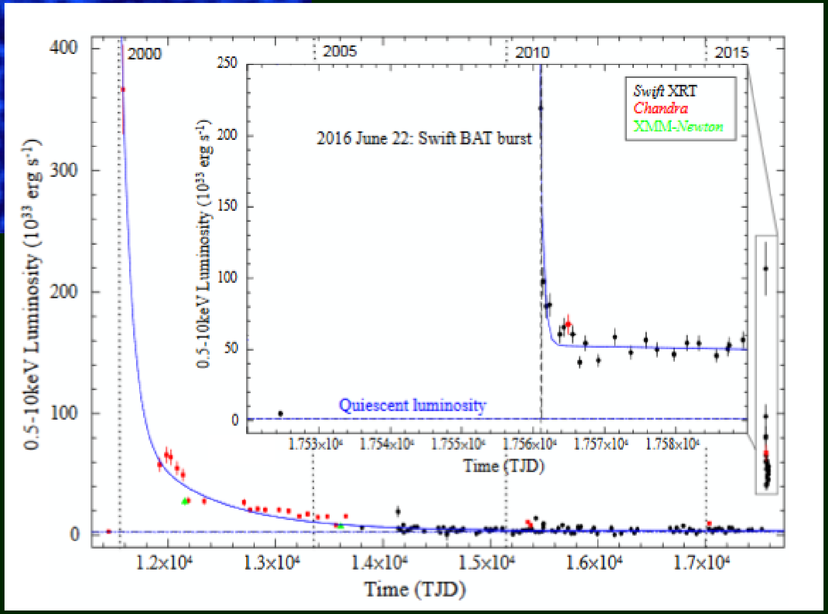


1504.03279

Not so hidden!



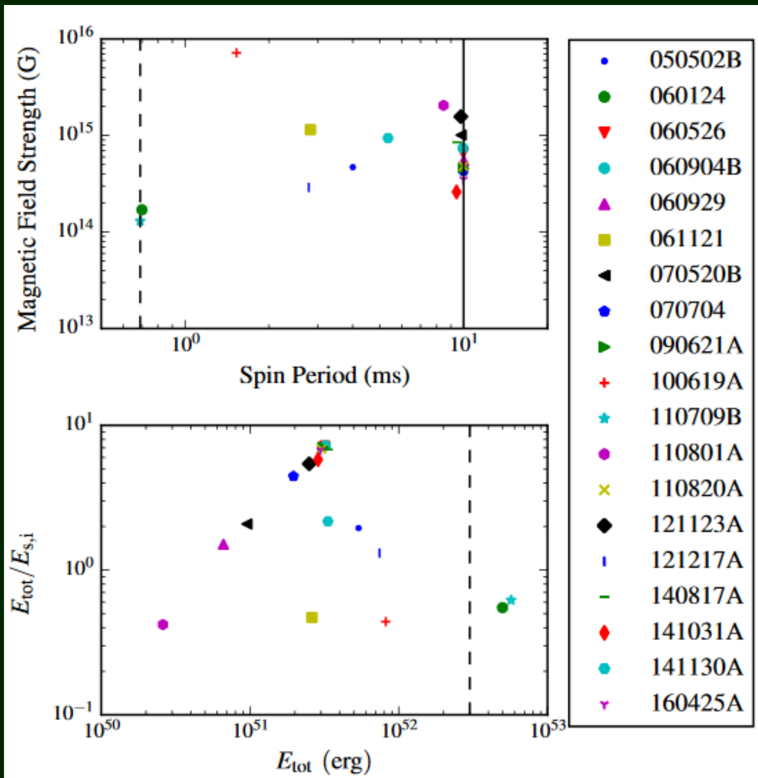




1607.04107

1607.04264

GRBs and fallback onto magnetars



Giant X-ray flares in GRB happen after ~30-10^5 s.

Rotational energy ~2 10^{52} erg P_{ms}^{-2}

$$\dot{M}_D(t) = \dot{M}_{fb} - \dot{M}_{acc} - \dot{M}_{prop},$$

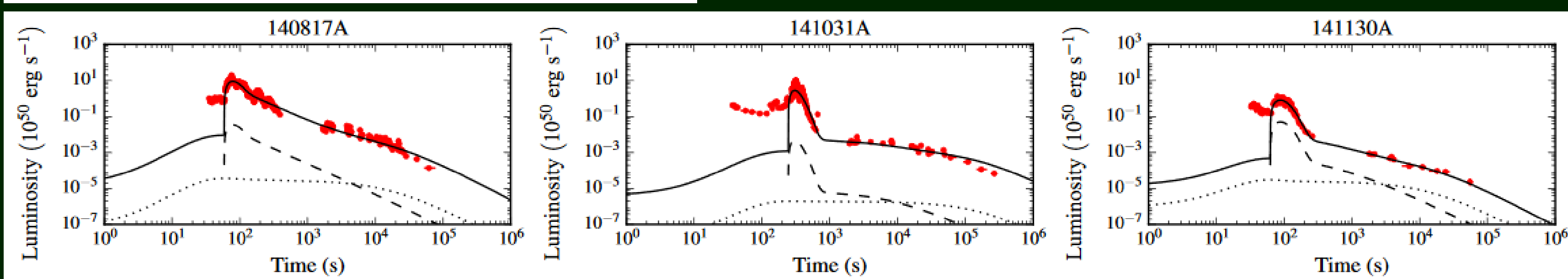
$$\dot{M}_{fb}(t) = \frac{M_{fb}}{t_{fb}}\left(\frac{t+t_{fb}}{t_{fb}}\right)^{-\frac{5}{3}}$$

$$L_{dip} = -\tau_{dip}\omega$$

$$L_{prop} = -\tau_{acc}\omega$$

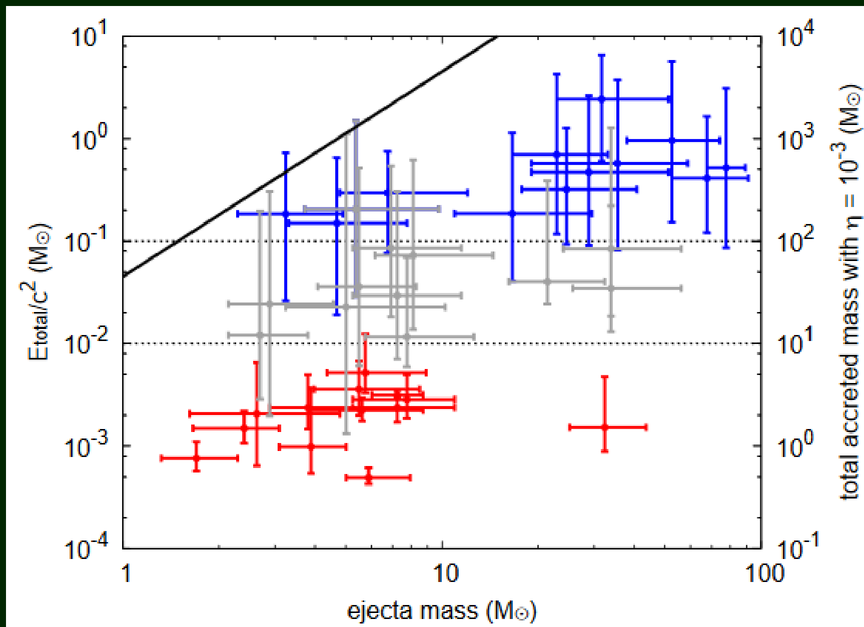
τ_{acc} and τ_{dip} are the accretion and dipole torques

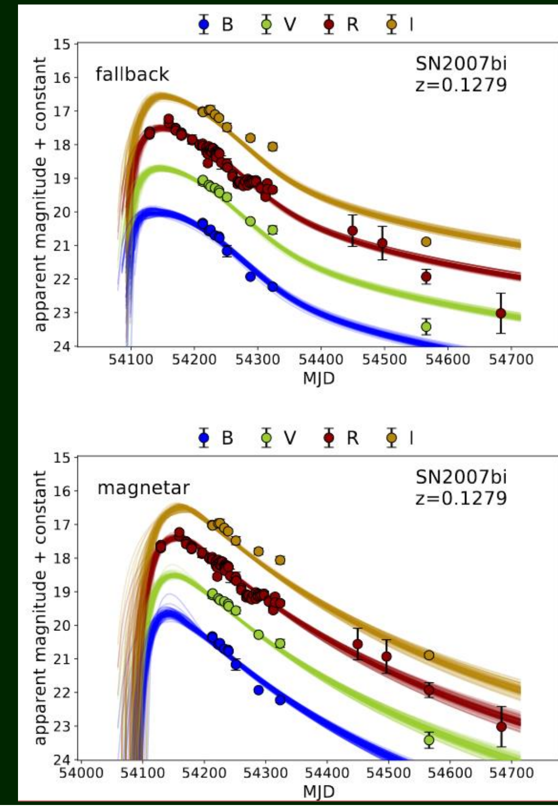
$$L_{tot} = \frac{1}{f_B}\left(\eta_{prop}L_{prop} + \eta_{dip}L_{dip}\right)$$

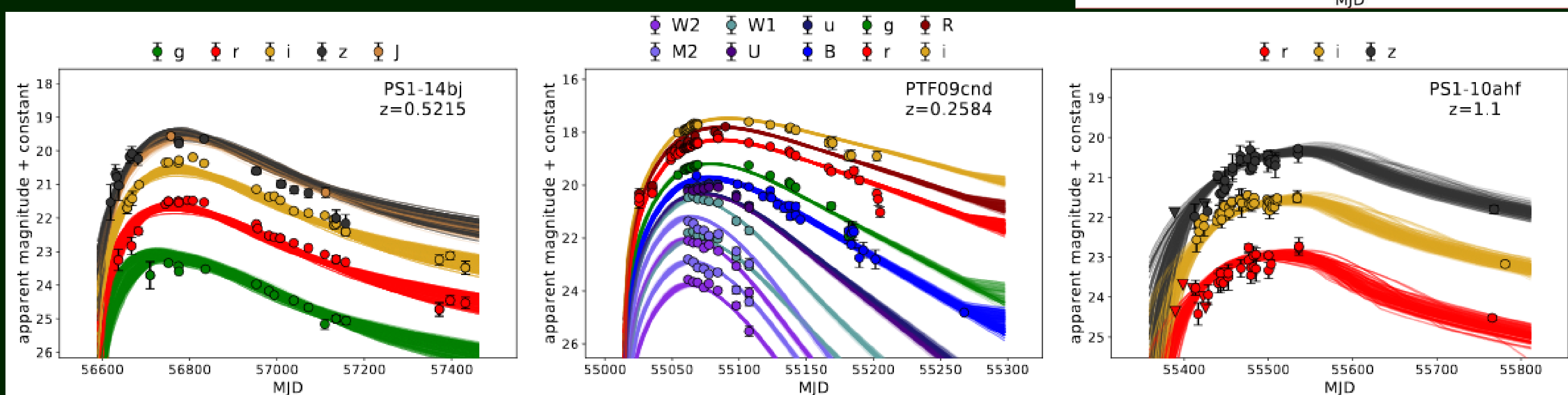


1805.09022

Fallback to power SN







1806.00090

$$L_{\rm fallback}(t) = \begin{cases} L_1 \left(\frac{t_{\rm tr}}{1\ {\rm sec}}\right)^{-\frac{5}{3}} \equiv L_{\rm flat} & (t < t_{\rm tr}) \\ L_1 \left(\frac{t}{1\ {\rm sec}}\right)^{-\frac{5}{3}} & (t \geq t_{\rm tr}) \end{cases}$$

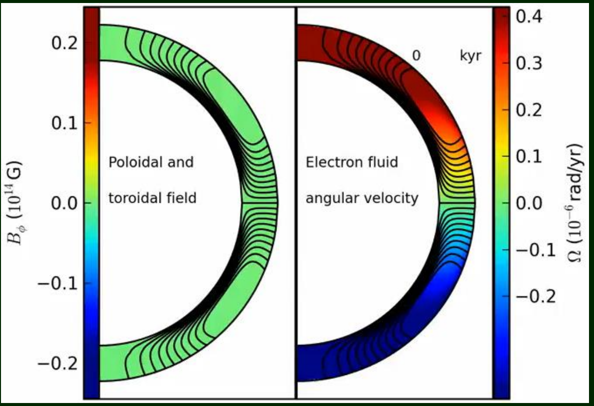
Conclusions

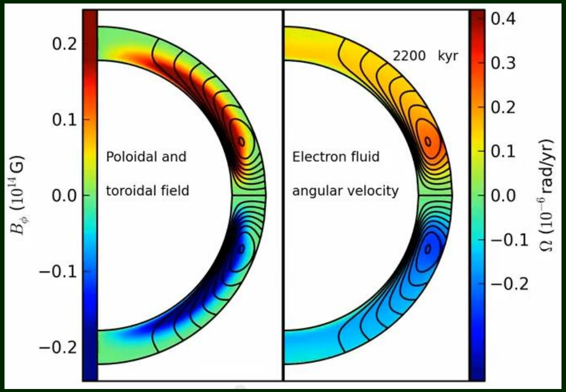
- Decaying magnetic field results in additional heating of a NS and decreasing its spin-down rate
- Field decay can be more important for large initial fields, for "standard" fields (~10^{12} G) it is not important
- It is possible to describe different types of young NSs (PSRs, magnetars, M7 etc.) in the model with decaying magnetic field
- Re-merging magnetic field can be an important ingredient
- With re-emerging field we can add to the general picture also CCOs.
- Recent studies indicate that in the life of normal radio pulsars there is a period when their magnetic field decay
- Hall cascade (and attractor) can be an important ingredient of the field evolution.
- At the moment we cannot state that we see the Hall attractor in the population of normal radio pulsars
- Also, we do not see that any of the M7 NSs are at the attractor stage, as its properties are predicted by GC2013
- Probably, the attractor stage is reached later, or its properties are different form the predicted ones.

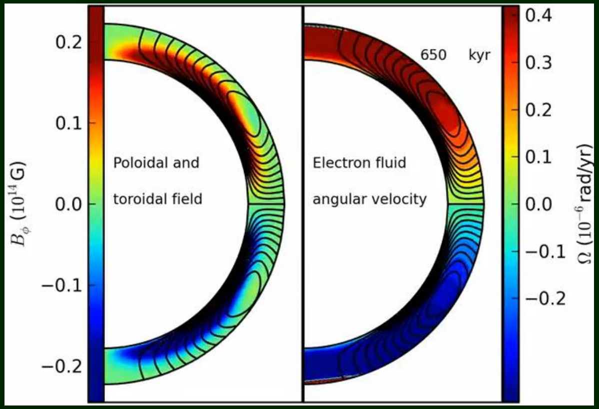
Papers to read

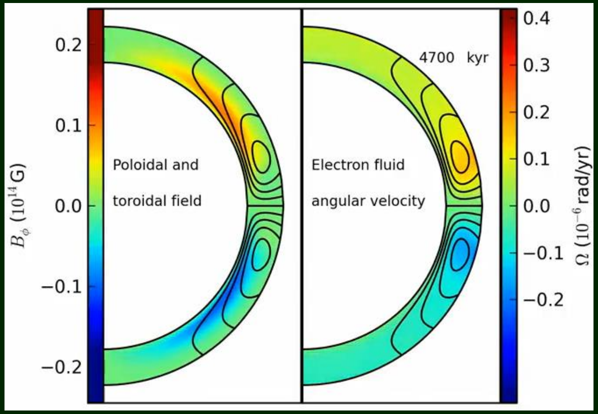
- Pons, Geppert "Magnetic field dissipation in neutron star crusts: from magnetars to isolated neutron stars " astro-ph/0703267
- Popov et al. "Population synthesis studies of isolated neutron stars with magnetic field decay" MNRAS (2009) arXiv: 0910.2190
- Ho ``Evolution of a buried magnetic field in the central compact object neutron stars ‘’ arXiv:1102.4870
- Pons et al. "Pulsar timing irregularities and the imprint of magnetic field Evolution" arXiv: 1209.2273
- Cumming et al. "MAGNETIC FIELD EVOLUTION IN NEUTRON STAR CRUSTS DUE TO THE HALL EFFECT AND OHMIC DECAY" astro-ph/0402392

Hall cascade and attractor







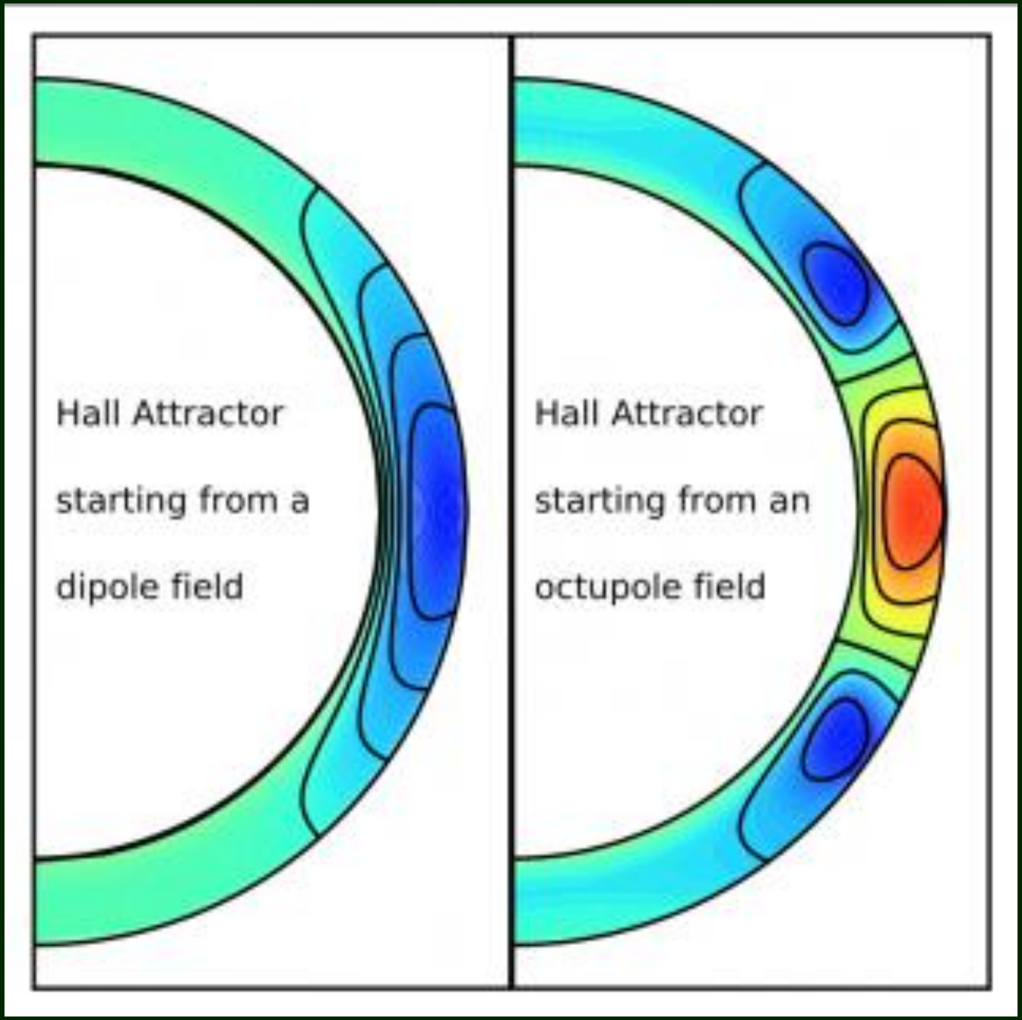


Hall cascade can reach the stage of so-called Hall attractor, where the field decay stalls for some time (Gourgouliatos, Cumming).

Gourgouliatos and Cumming 2013

http://www.physics.mcgill.ca/~kostasg/research.html

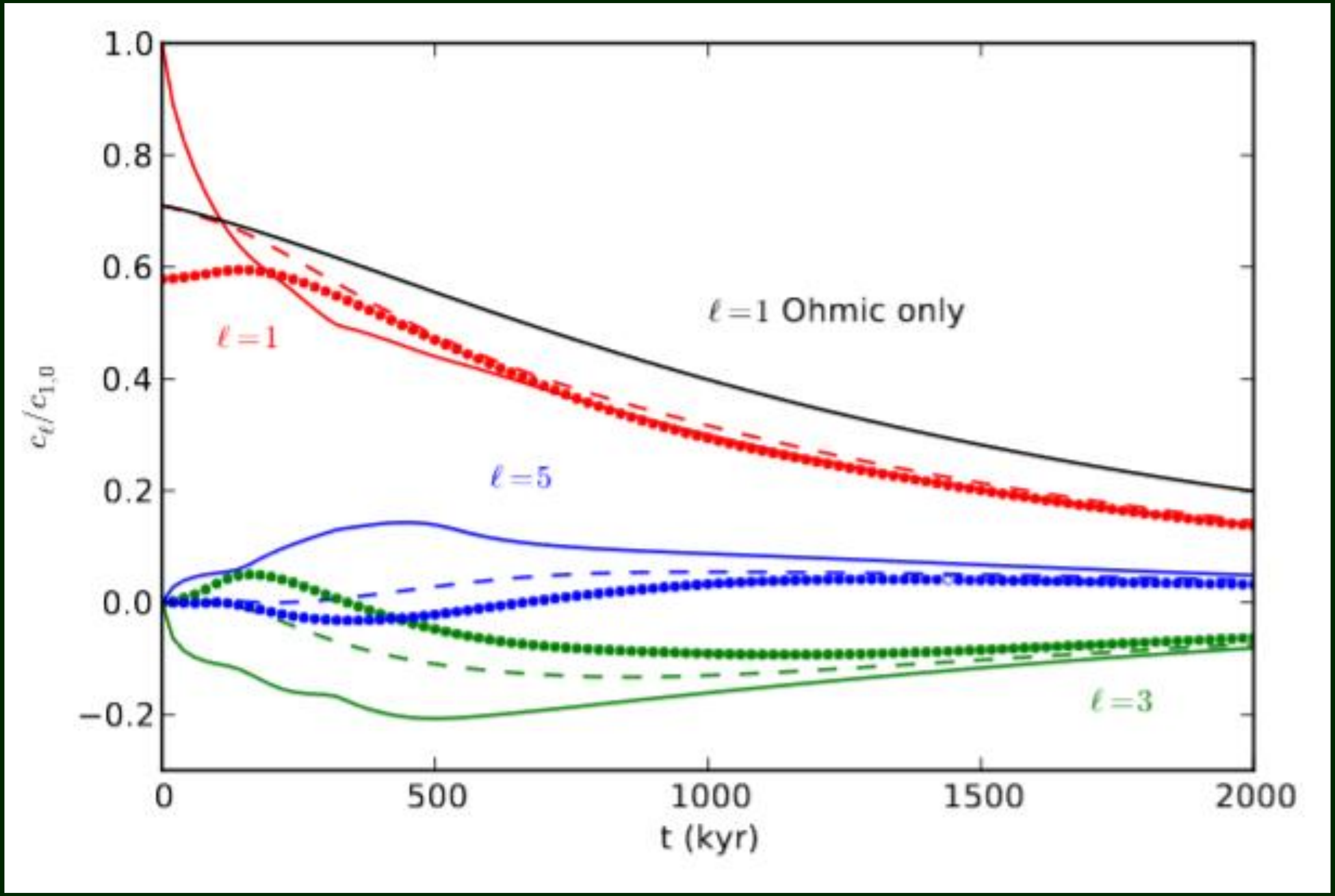
Hall attractor



After some time the Hall cascade decays as the field finds a new| stable configuration.

1311.7345

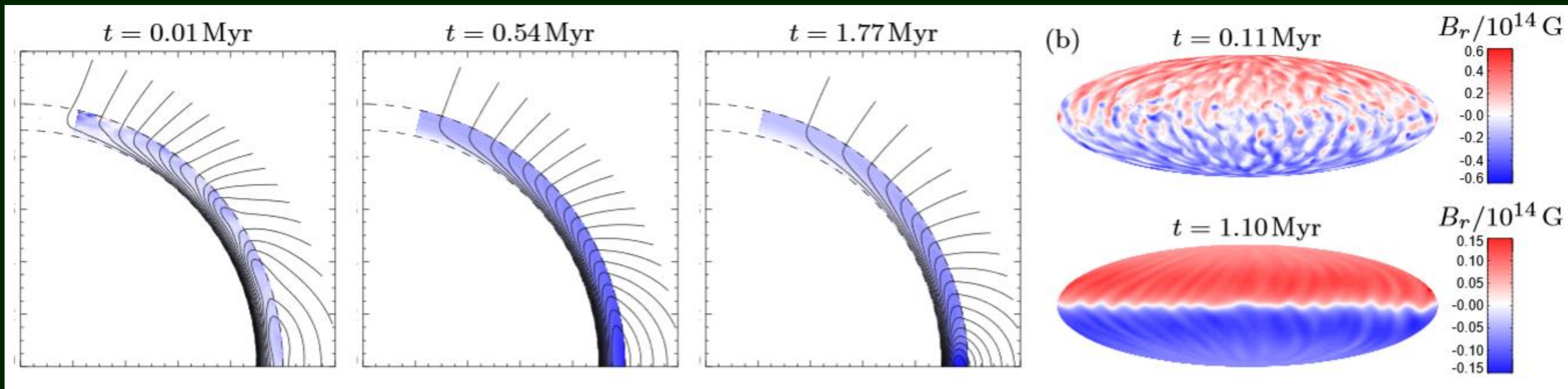
Evolution of different components

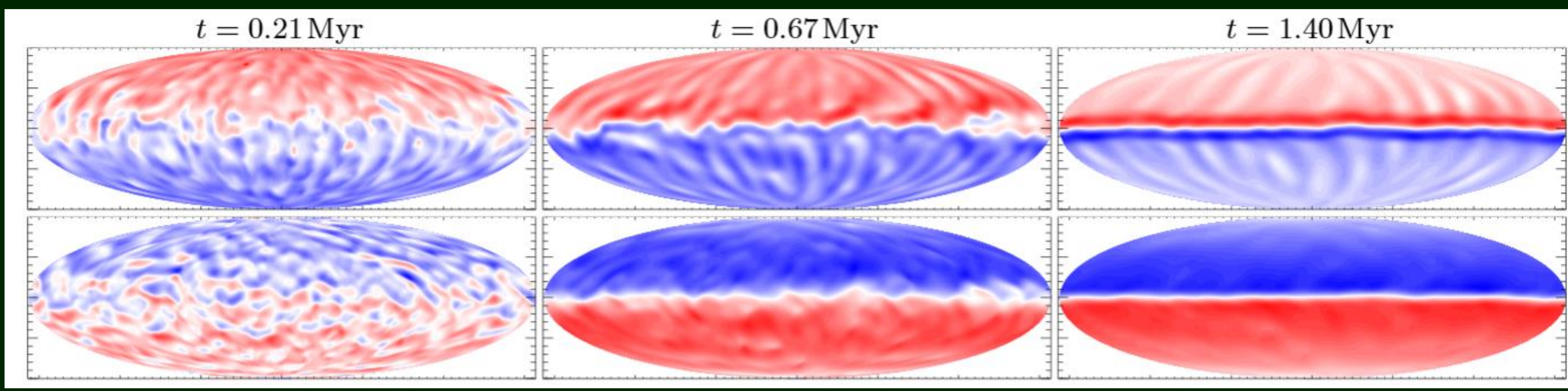


Hall attractor mainly consists of dipole and octupole

1311.7004

New studies of the hall cascade

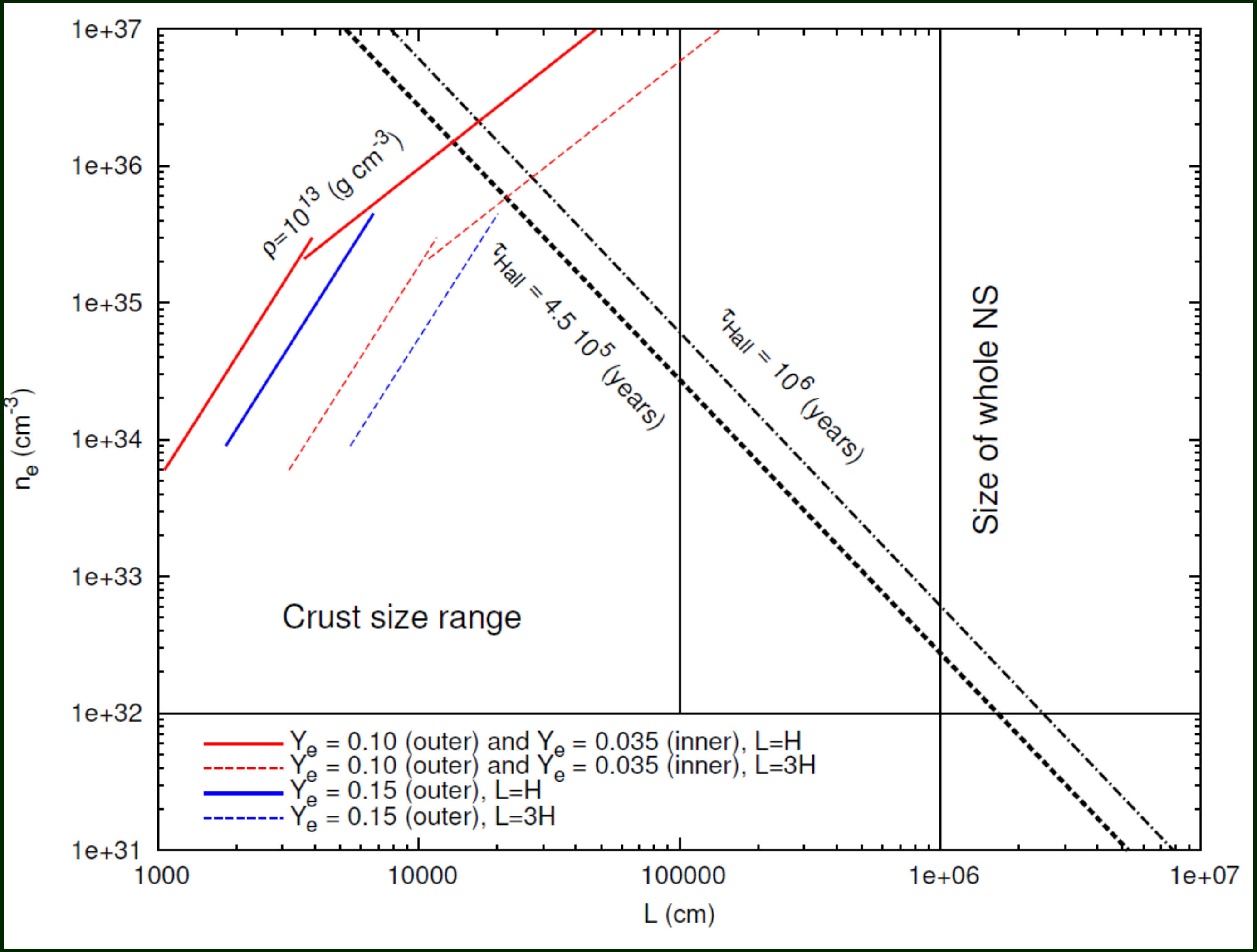




New calculations support the idea of a kind of stable configuration.

1501.05149

Where the currents are located?



$$\tau_{\rm Hall} \approx \frac{4\pi e L^2 n_e}{cB}$$

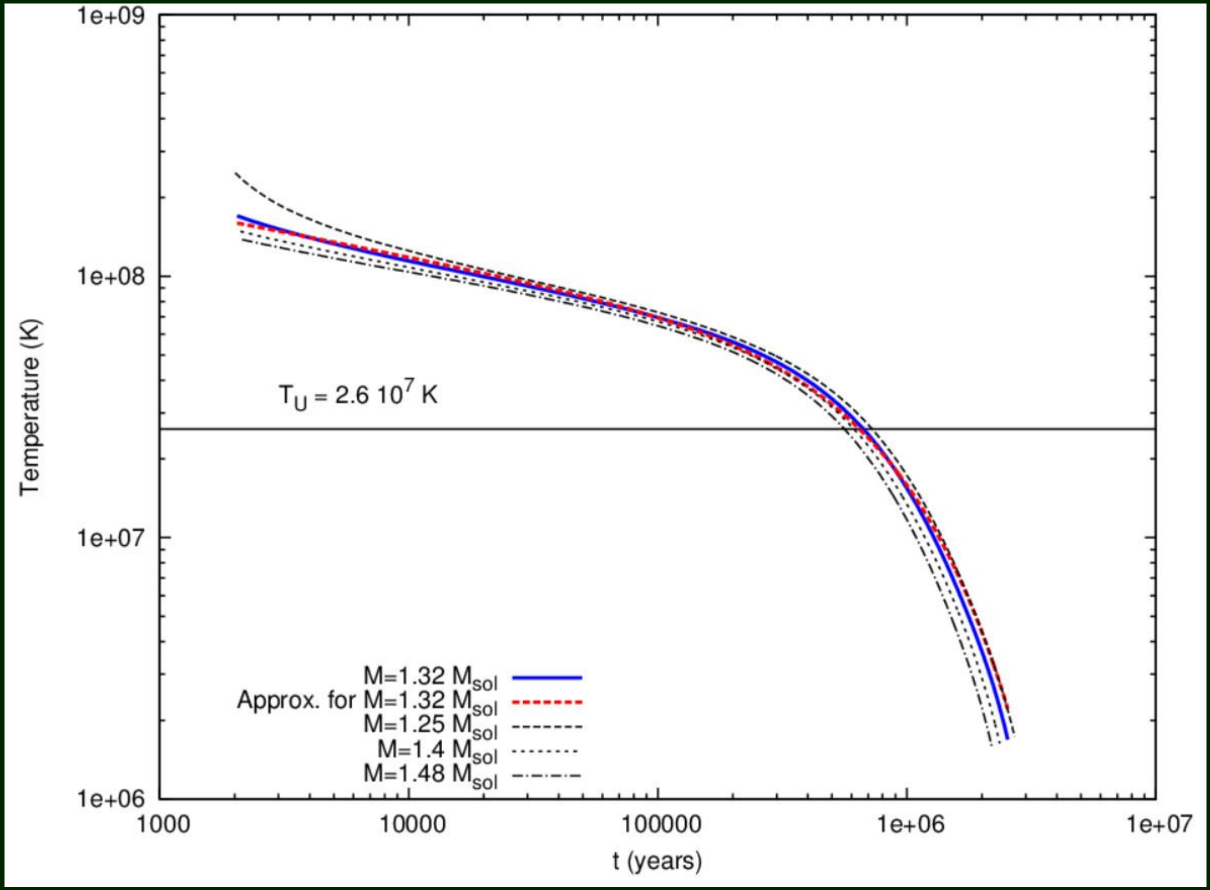
$$L \approx H = P(\rho)/(\rho g)$$

Igoshev, Popov (2015)

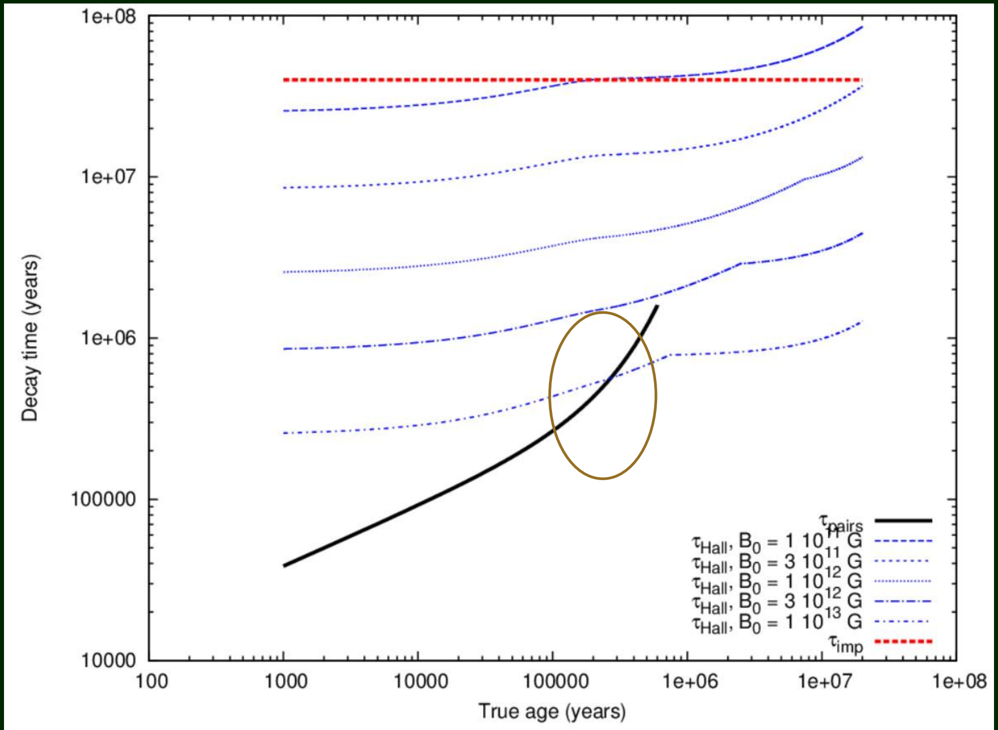
Thermal evolution

Calculations are made by Shternin et al.

We fit the numerical results to perform a population synthesis of radio pulsars with decaying field.



Different decay time scales



In the range of ages interesting for us the Hall rate is about the same value as the rate of Ohmic dissipation due to phonons.

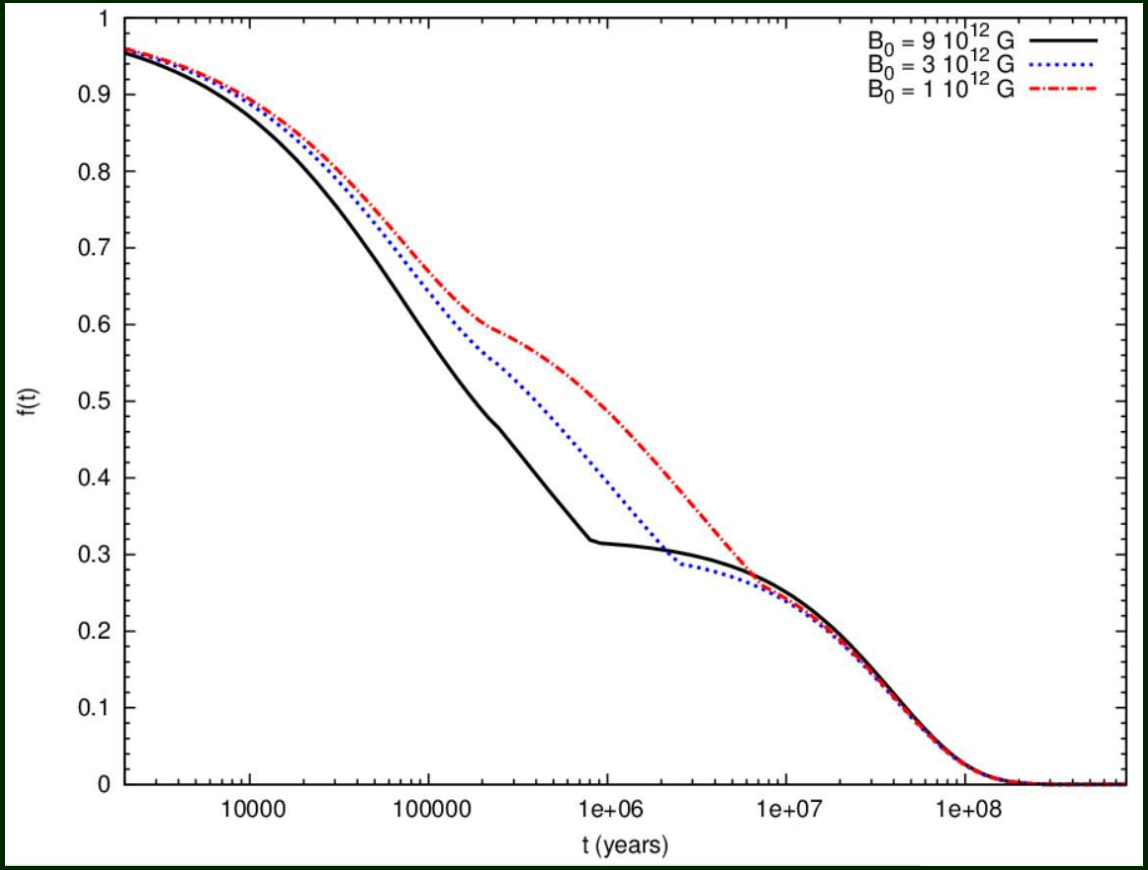
$$B = B_0 \frac{\exp(-t/\tau_{\rm Ohm})}{1 + \frac{\tau_{\rm Ohm}}{\tau_{\rm Hall}}(1 - \exp(-t/\tau_{\rm Ohm}))}$$

$$\tau_{\rm imp} = 5.7 \frac{\rho_{14}^{5/3}}{Q}\left(\frac{Z}{30}\right)\left(\frac{Y_e}{0.05}\right)^{1/3}\left(\frac{Y_n}{0.8}\right)^{10/3} \times$$
$$\times \left(\frac{f}{0.5}\right)^2 \left(\frac{g_{14}}{2.45}\right) \text{ Myrs},$$

$$\tau_{\rm phonon} = 2.2 \frac{\rho_{14}^{15/6}}{T_8^2}\left(\frac{Y_e}{0.05}\right)^{5/3}\left(\frac{Y_n}{0.8}\right)^{10/3} \times$$
$$\times \left(\frac{f}{0.5}\right)^2 \left(\frac{g_{14}}{2.45}\right)^{-2} \text{ Myrs},$$

Magnetic field evolution

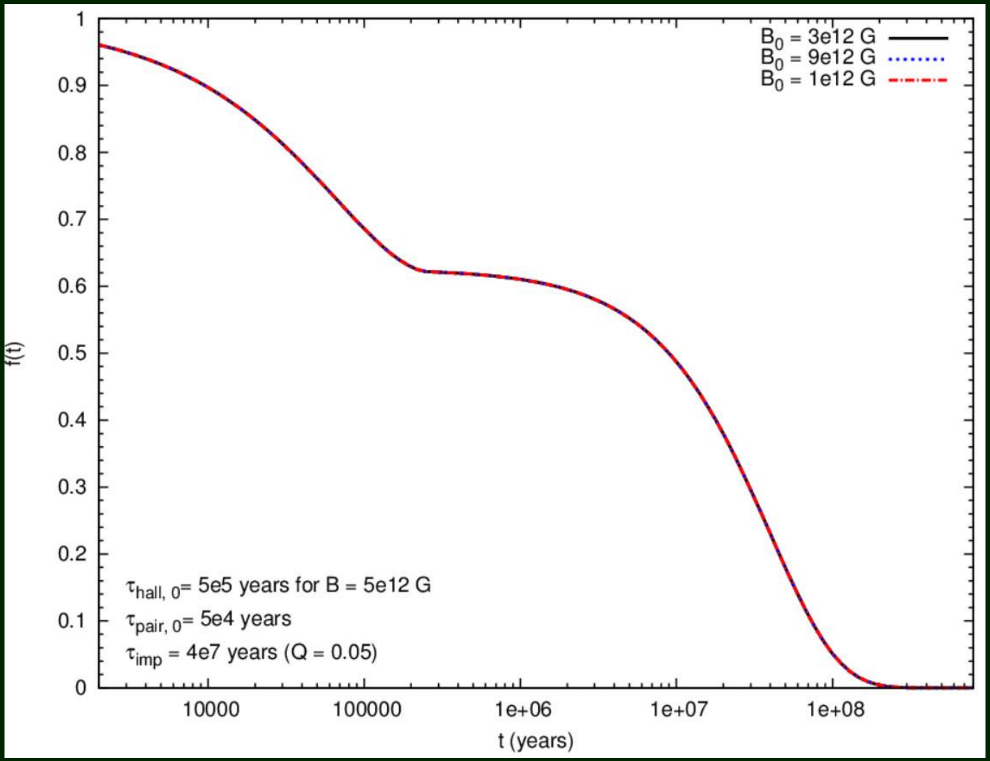
$$B = B_0 \frac{\exp\left(-t/\tau_{\rm Ohm}\right)}{1 + \frac{\tau_{\rm Ohm}}{\tau_{\rm Hall}}\left(1 - \exp\left(-t/\tau_{\rm Ohm}\right)\right)}$$



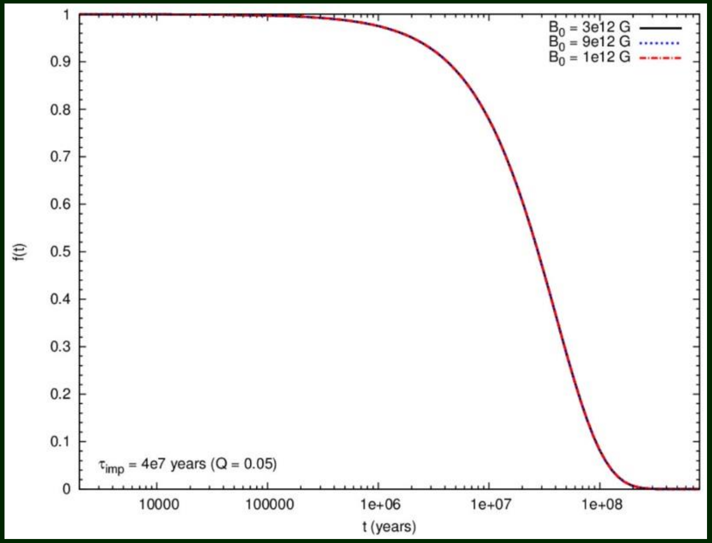
All inclusive:
- Hall
- Phonons
- Impurities

Igoshev, Popov (2015)

Only Ohmic decay

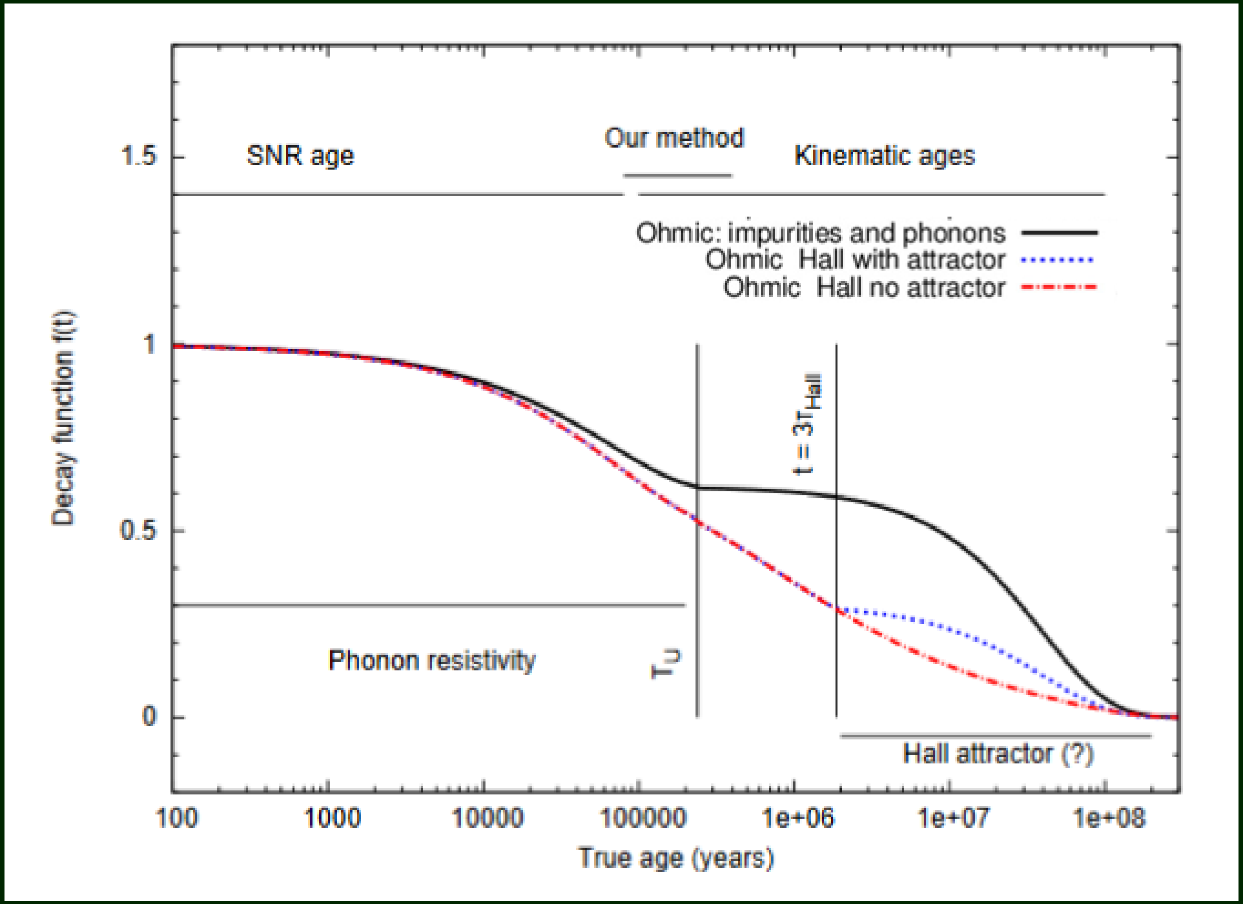


Here the Hall cascade is switched off



In one figure we have Ohmic decay only due to impurities, on another one – phonons are added.

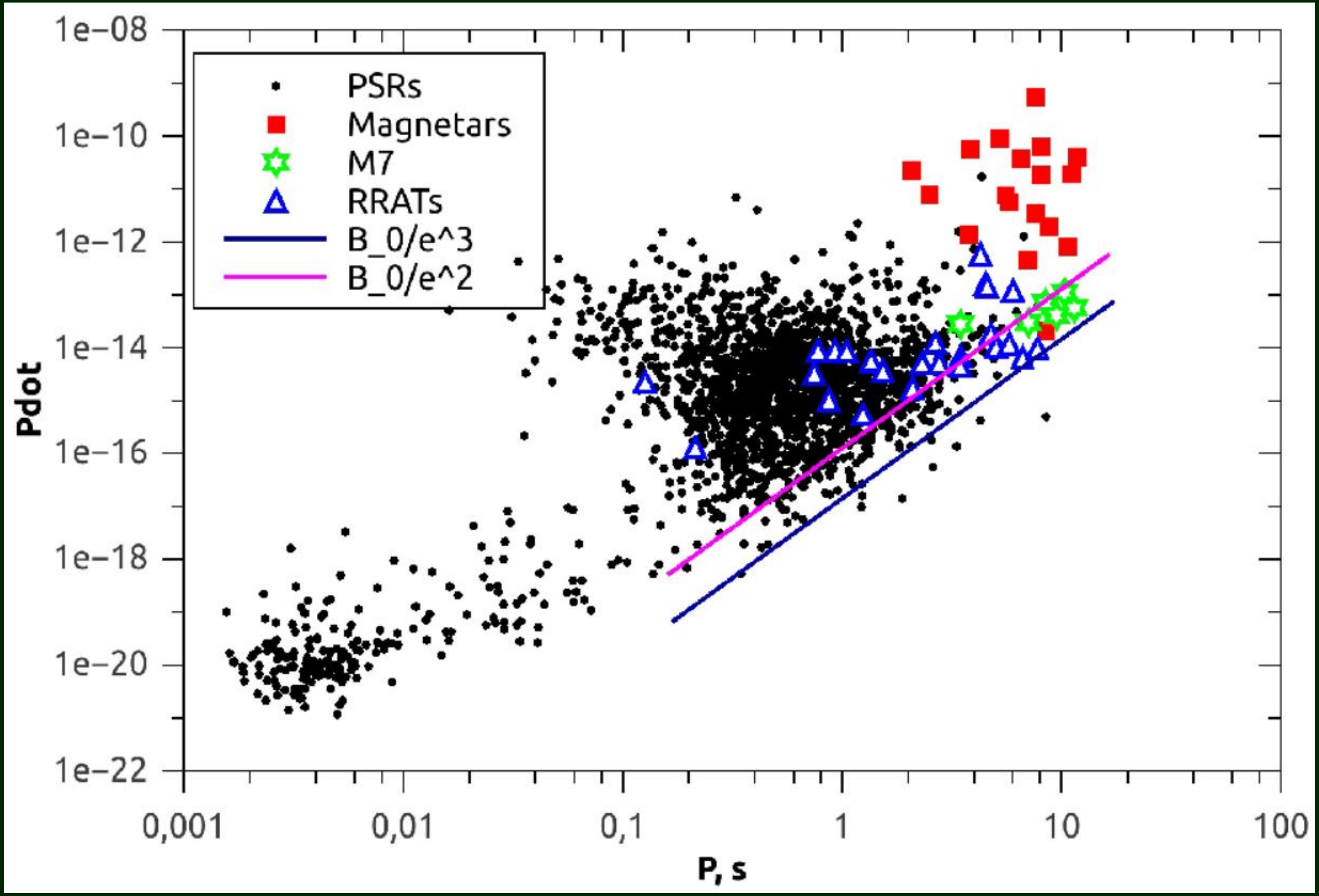
Comparison of different options

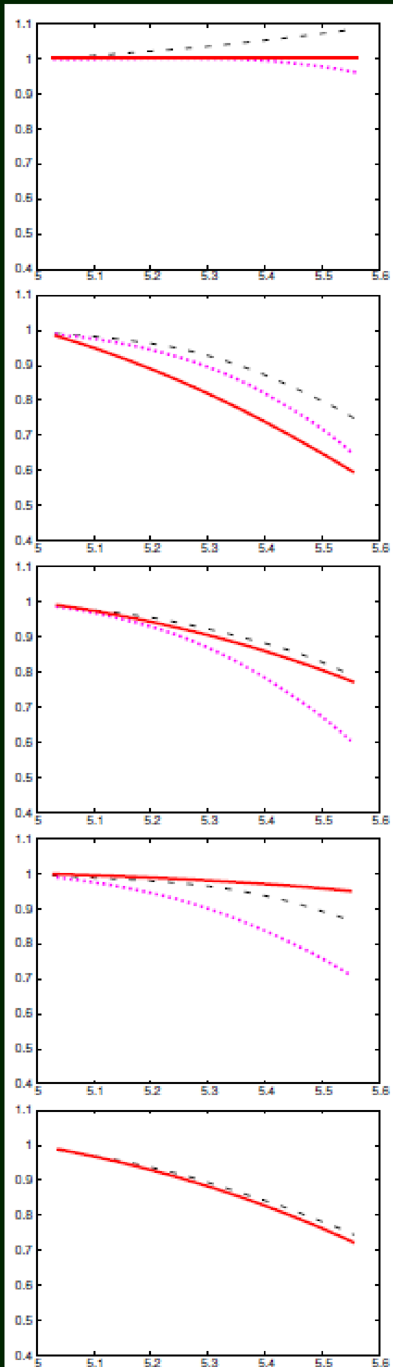


We think that at the ages ~10^5 yrs and below for normal pulsars we see mostly Ohmic decay, which then disappears as NSs cool down below the critical T.

Igoshev, Popov (2015)

Getting close to the attractor





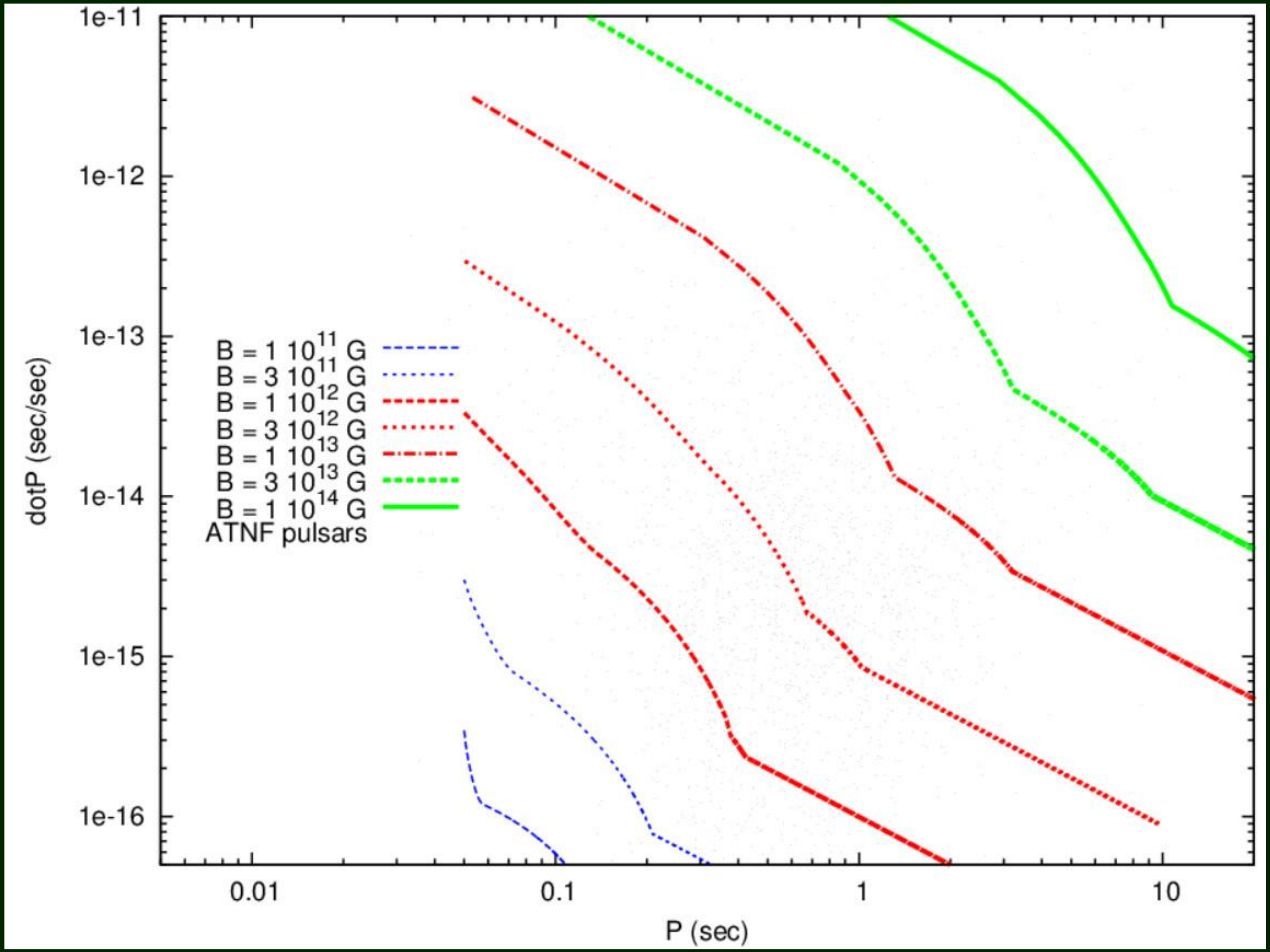
Tests

We make extensive tests of the method and obtain that in most of the cases it is able to uncover non-negligible magnetic field decay (more than a few tens of per cent during the studied range of ages) in normal radio pulsars for realistic initial properties of neutron stars.

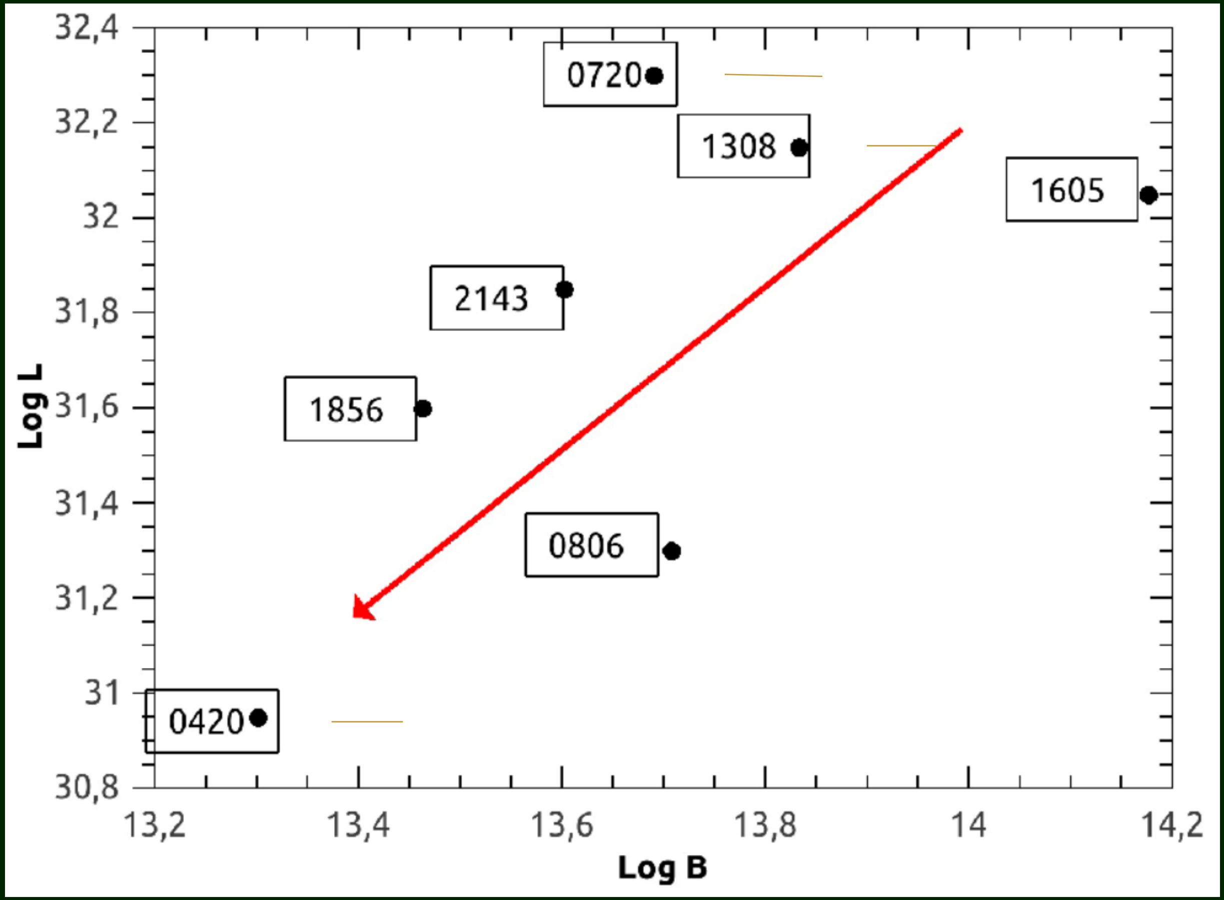
Name	$\log \mu_{B_0}$ [G]	$\log \sigma_{B_0}$ [G]	μ_{P_0} [s]	σ_{P_0} [s]	α	$\tau_{\rm D}$ [Myr]	$\tau_{\rm SDA}$ [Myr]
A1	12.60	0.47	0.33	0.23	0.50	∞	∞
A2	12.95	0.55	0.30	0.15	0.50	∞	10
B1	12.60	0.47	0.33	0.23	0.50	0.5	1.00
B2	12.95	0.55	0.30	0.15	0.50	0.5	0.690
C1	12.60	0.47	0.33	0.23	0.50	1	1.15
C2	12.95	0.55	0.30	0.15	0.50	1	0.560
D1	12.60	0.47	0.33	0.23	0.50	5	2.00
D2	12.95	0.55	0.30	0.15	0.50	5	0.80
E	13.04	0.55	0.22	0.32	0.44	~ 0.8	0.880

(Synthetic samples are calculated by Gullon, Pons, Miralles)

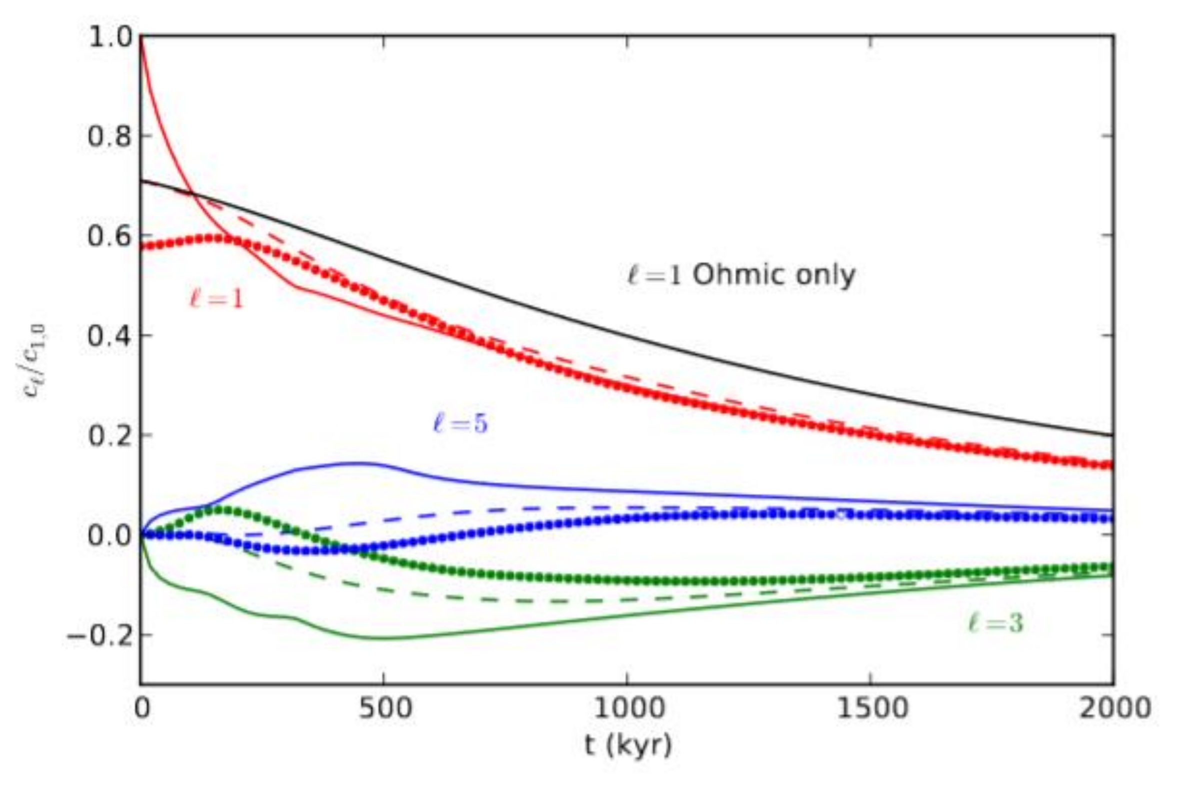
Evolution with field decay

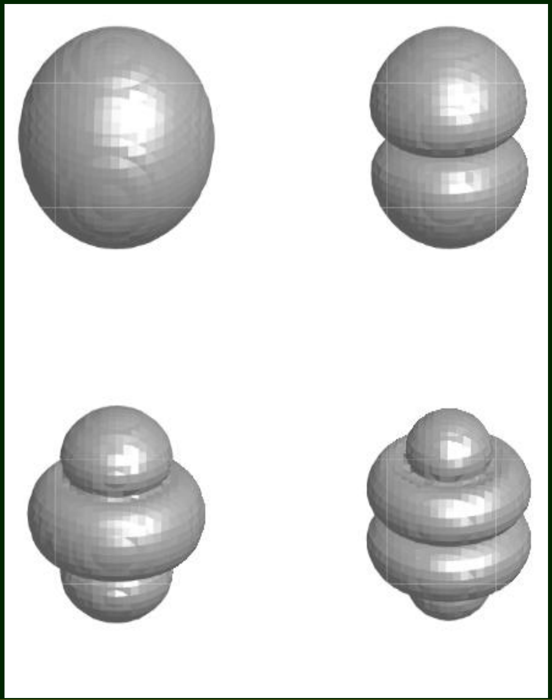


Who is closer to the attractor stage?



Evolution of different components





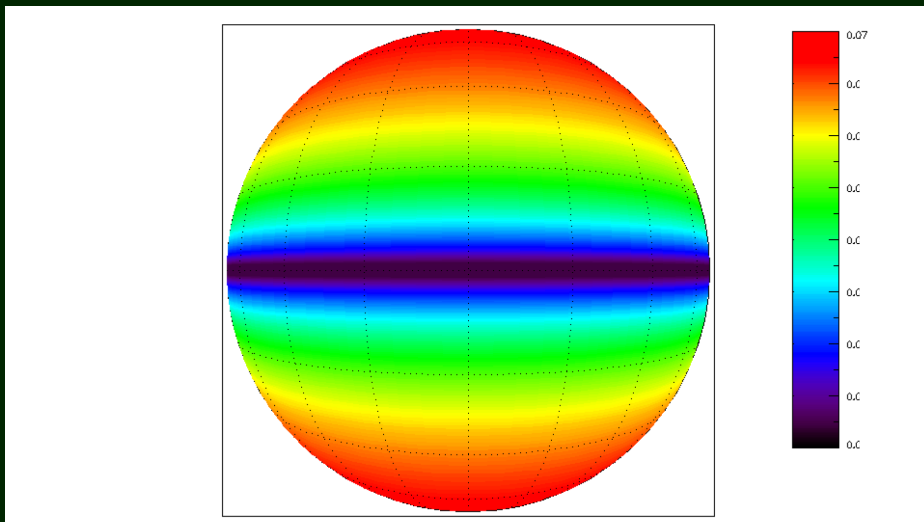
Hall attractor mainly consists of dipole and octupole

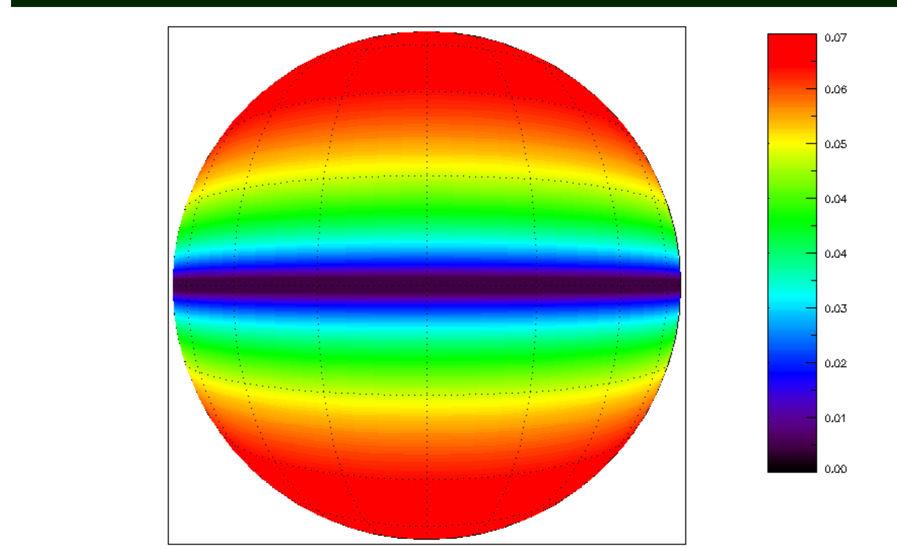
1311.7004

0902.0720

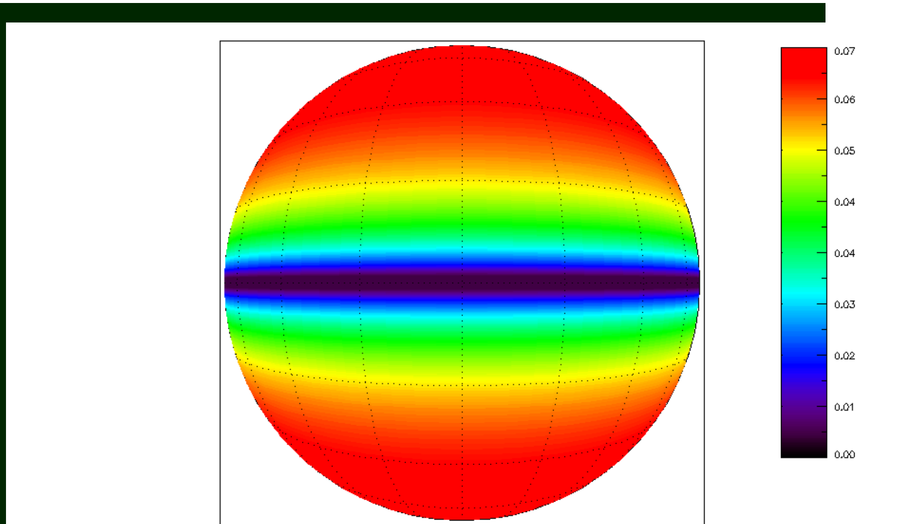
Temperature maps

Pure dipole



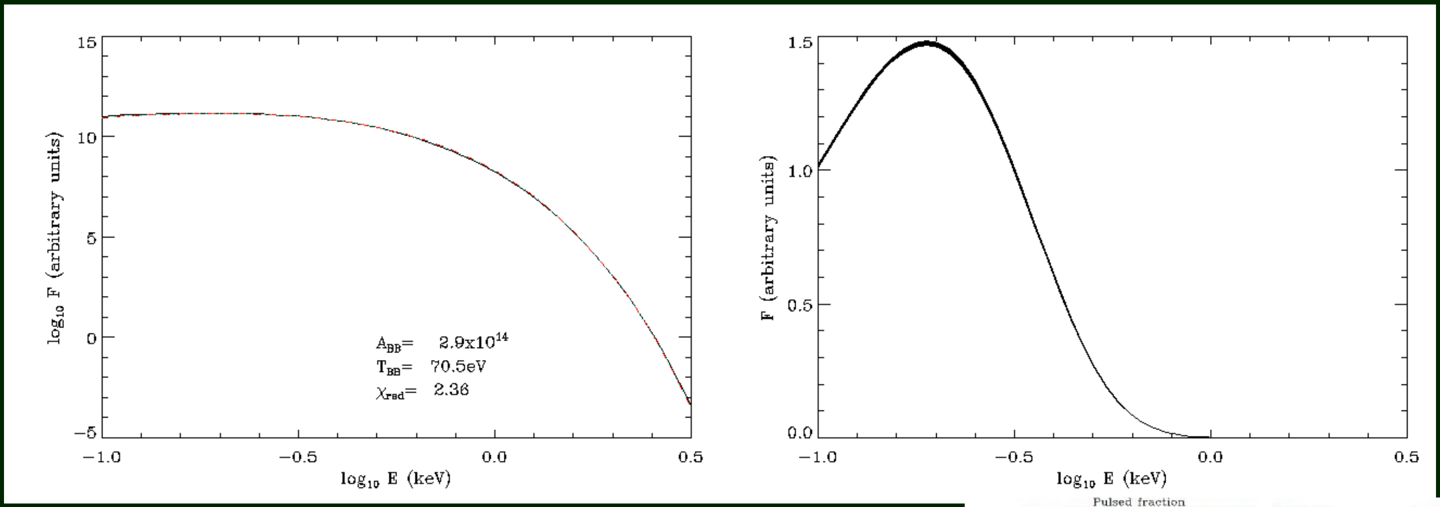


Dipole + octupole

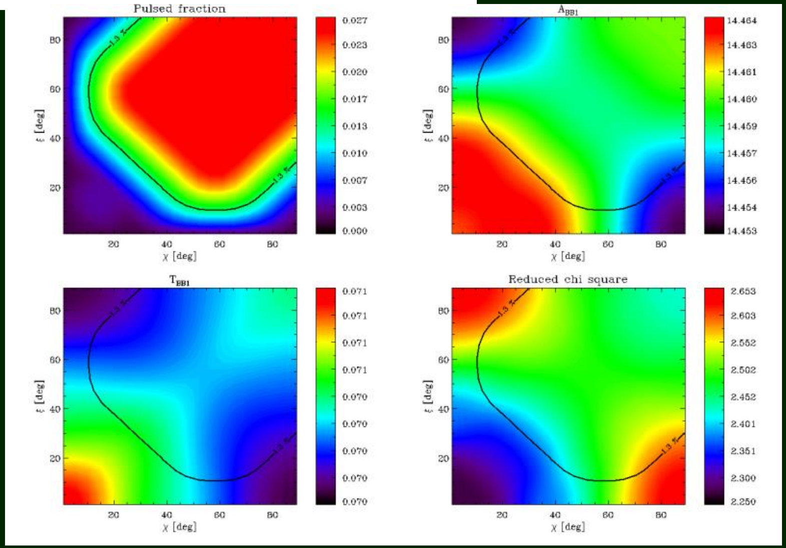


Dipole+octupole+l5

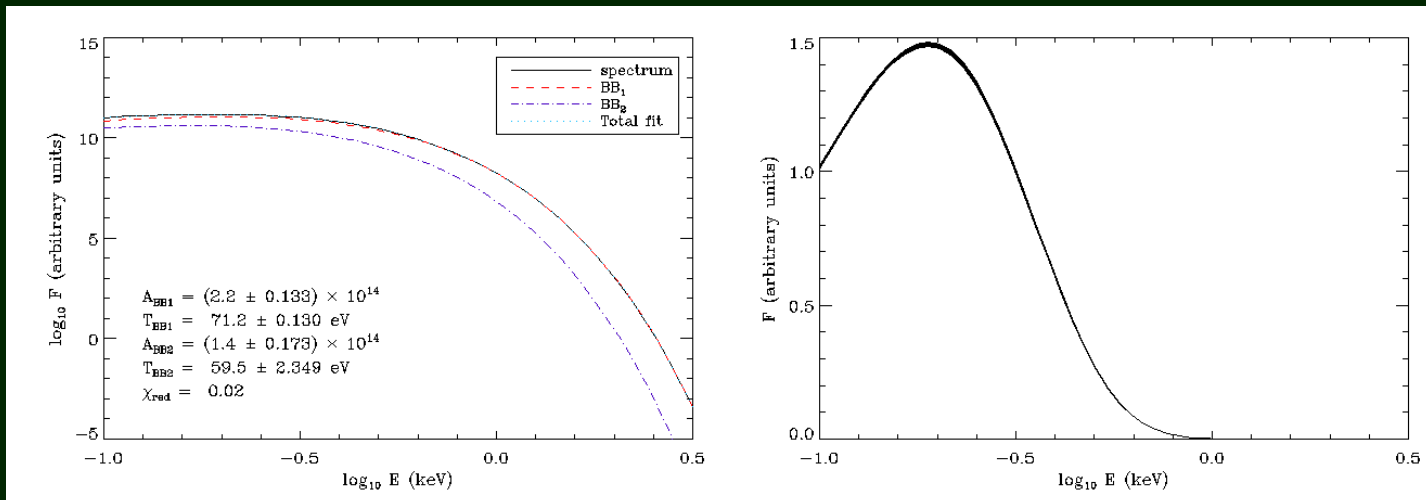
Spectral fits: single blackbody



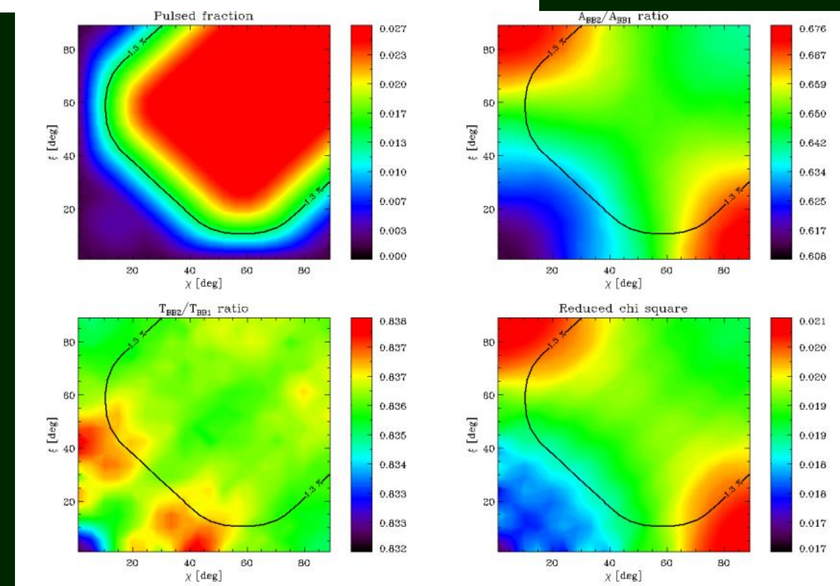
Single black body
does not provide
a good fit, even using,
in addition, a line,
or condensed surface.



Spectral fits: two black bodies



Formally, two black bodies
is the best fit for 1856.
And for dipole+octupole
we can obtain a very good fit.
But ….



Observational data

Parameter	Single BB	Two BB
N_H [10^{19} cm^{-2}]	$4.8^{+0.2}_{-0.2}$	$12.9^{+2.2}_{-2.3}$
kT_h^∞ [eV]	$61.5^{+0.1}_{-0.1}$	$62.4^{+0.6}_{-0.4}$
R_h^∞ [km]	$5.0^{+0.1}_{-0.1}$	$4.7^{+0.2}_{-0.3}$
kT_s^∞ [eV]	-	$38.9^{+4.9}_{-2.9}$
R_s^∞ [km]	-	$11.8^{+5.0}_{-0.4}$
σ_{sys}	1.5%	0.6%
χ^2_ν	1.12	1.11

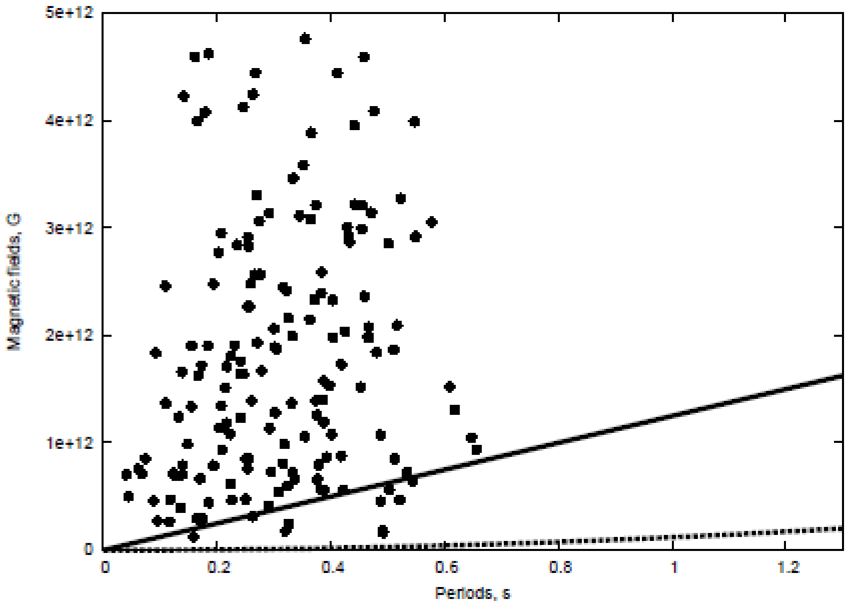
Two black bodies is the best fit.
The colder component corresponds to larger surface area.
This is in contrast with our results for the Hall attractor proposed by GC2013 (dipole + octupole).

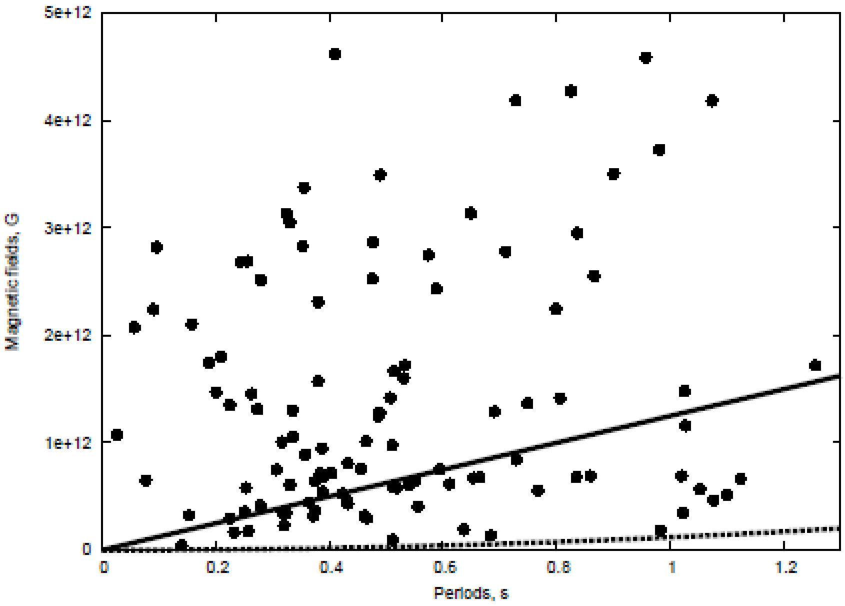
1202.2121

Synthetic populations

Constant field

Exponential decay

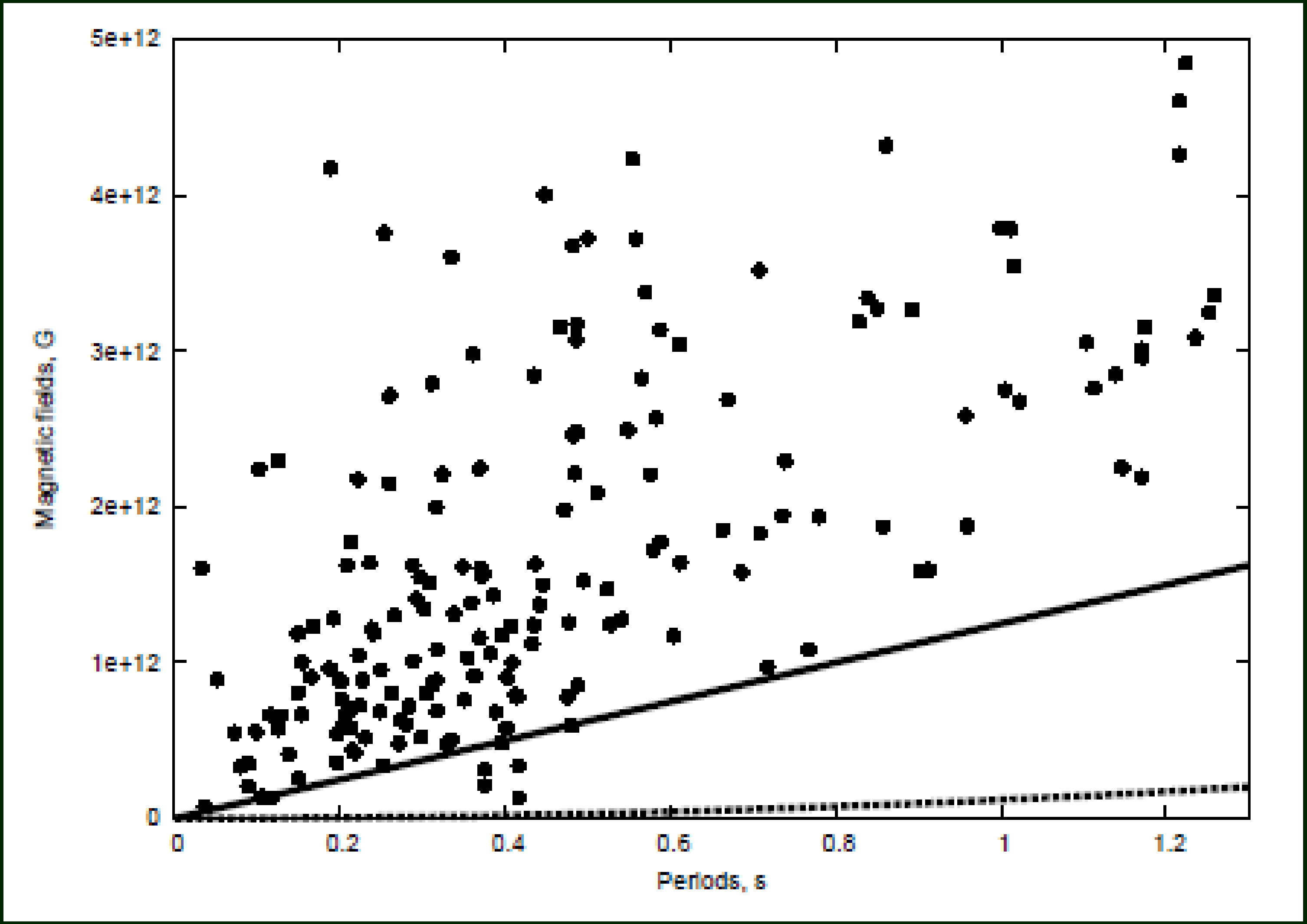




$$P_0 = P\sqrt{1 - \frac{t_{\rm true}}{\tau}}.$$

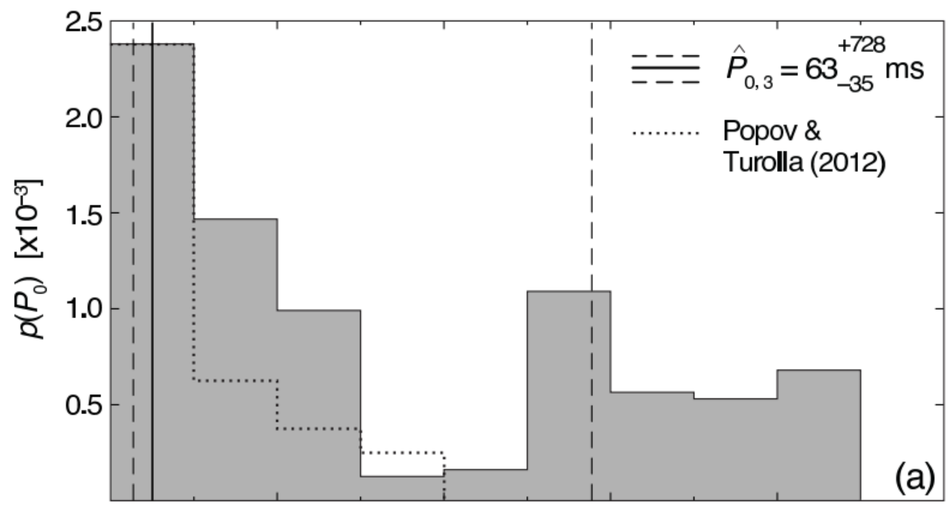
Igoshev, Popov 2013

Fitting the field decay

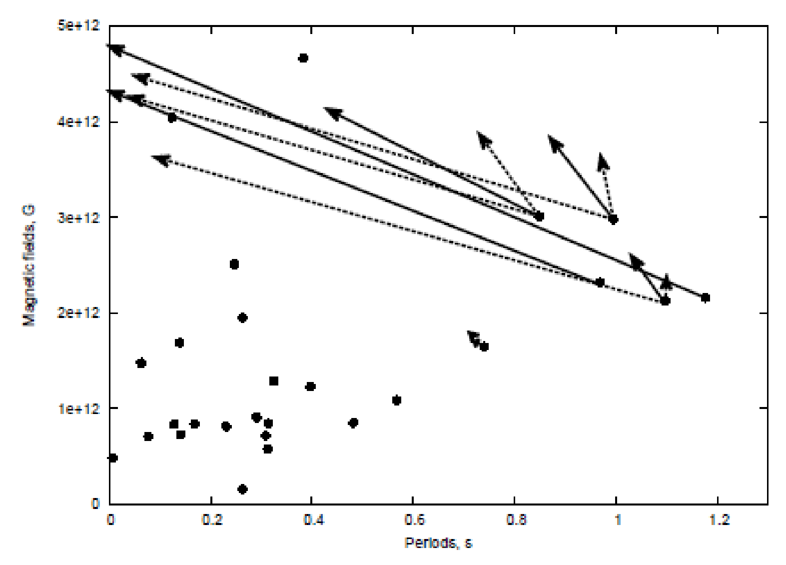


Igoshev, Popov 2013

Another option: emerging field



The problem is just with few (6) most long-period NSs.
Is it possible to hide them when they are young, and make them visible at the age ~few million years?



Yes! Emerging magnetic field!!!

Then we need correlations between different parameters.

Igoshev, Popov 2013